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AN ULTRA-FAST X-RAY DISK WIND IN THE NEUTRON STAR BINARY GX 340+0

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ABSTRACT

We present a spectral analysis of a brief Chandra/HETG observation of the neutron star low-mass X-ray binary GX 340+0. The high-resolution spectrum reveals evidence of ionized absorption in the Fe K band. The strongest feature, an absorption line at approximately 6.9 keV, is required at the 5σ level of confidence via an F-test.

Photoionization modeling with XSTAR grids suggests that the line is the most prominent part of a disk wind with an apparent outflow speed of \( v = 0.04c \). This interpretation is preferred at the 4σ level over a scenario in which the line is H-like Fe XXVI at a modest redshift. The wind may achieve this speed owing to its relatively low ionization, enabling driving by radiation pressure on lines; in this sense, the wind in GX 340+0 may be the stellar-mass equivalent of the flows in broad ionization line quasars. If the gas has a unity volume filling factor, the mass outflow rate in the wind is over \( 10^{-3}M_\odot \text{yr}^{-1} \), and the kinetic power is nearly \( 10^{39} \text{erg s}^{-1} \) (or, 5–6 times the radiative Eddington limit for a neutron star). However, geometrical considerations—including a small volume filling factor and low covering factor—likely greatly reduce these values.

Key words: accretion, accretion disks – X-rays: binaries

1. INTRODUCTION

Accretion onto neutron stars with low magnetic fields is expected to be similar to accretion onto black holes, or at least black holes with low spin parameters. Timescales and line shifts should simply scale with mass, and with the depth of the inner edge of the disk within the potential well. Observations appear to confirm this expectation: quasi-periodic oscillations scale in the expected manner (e.g., Wijnands & van der Klis 1999), and relativistic lines from neutron stars place interesting limits on stellar radii but are not as extreme as lines from spinning black holes (e.g., Cackett et al. 2008, 2010; Di Salvo et al. 2009; Papitto et al. 2009; Miller et al. 2013).

Accretion flows onto neutron stars and black holes might be even more similar, far from the compact object. X-ray disk winds from stellar-mass black holes are emerging as important facets of the accretion flow; in some cases, the mass-loss rate can rival or exceed the accretion rate in the inner disk (e.g., King et al. 2012, 2015; Miller et al. 2015). Such winds appear to arise between \( 10^{-5} \text{–} 10^{-3}GM/c^2 \) from the black hole, depending on particulars. Ionized X-ray absorption is commonly observed in neutron star binaries, but evidence of outflowing gas has lagged the rapid progress being made in stellar-mass black holes. Nevertheless, evidence of disk winds in neutron star X-ray binaries is growing, with flows in GX 13+1 (e.g., Ueda et al. 2004) and IGR J17480–2446 (Miller et al. 2011) marking two prominent examples.

In some moderate-resolution spectra of active galactic nuclei (AGNs) extremely fast outflows with very high mass fluxes have been observed (e.g., Tombesi et al. 2010). These winds may contribute significantly to feedback and the co-evolution of massive black holes and host galaxies. In stellar-mass black holes, outflow speeds of \( 0.01–0.05c \) are detected at gratings resolution (e.g., King et al. 2012, 2015; Miller et al. 2015). Such flows have not previously been reported in the steady emission of neutron star X-ray binaries (but see Pinto et al. 2014 for evidence of fast outflows during bursts).

GX 340+0 is a neutron star X-ray binary, located close to the Galactic Plane. It is known to be a low-mass X-ray binary and a “Z” source based on its behavior in X-ray color–color diagrams. Cackett et al. (2010) analyzed an XMM-Newton observation of GX 340+0 and found evidence of a relativistic iron line from the inner disk, strongly impacted by X-ray absorption consistent with H-like Fe XXVI (also see D’Ai et al. 2009). Lavagetto et al. (2004) also found X-ray absorption in a BeppoSAX spectrum of GX 340+0; this was modeled with a Gaussian at 6.8 keV. This energy lies between Fe XXV (6.70 keV) and Fe XXVI (6.97 keV) and would imply a large velocity shift. In order to better understand the nature of the ionized X-ray absorption in GX 340+0 at high resolution, we have analyzed archival Chandra/HETG spectra of the source.

2. OBSERVATIONS AND REDUCTION

Chandra has observed GX 340+0 on four occasions, but only spectra from ObsID 1922 show evidence of a disk wind; it is the sole focus of this paper. Observation 1922 started on 2001-08-09 at 12:35:11 (UT). After our reduction, we find that a net exposure time of 5.8 ks was achieved.

The data were reduced using CIAO version 4.7 and the associated calibration files. We downloaded the full observation from the Chandra archive. The standard sequence of routines, tsigdetect, tg_create_mask, tg_resolve_events, and textract were run to produce first-order spectral files. When executing tg_create_mask, we set the parameter “width_factor_hetg” to have a value of 17, rather than the default of 35. This reduces the width of the HEG extraction regions, with the effect of better selecting true first-order HEG photons and enabling the extraction of spectra that carry out to 10 keV. The tools mkgrmf and fullgrf tasks were run to create response files. Finally, add_grating_spectra was run to combine the first-order spectra and responses, and the FTOOL “grppha” was
The spectra were fit using XSPEC version 12.8.2 (Arnaud 1996). The “Churazov” weighting scheme was applied in minimizing the $\chi^2$ fitting statistic. All errors reported in this work reflect parameter values at the 1σ confidence limits.

Cackett et al. (2010) report an interstellar column density of $N_{\text{H,ISM}} = 0.9 - 1.1 \times 10^{22} \text{ cm}^{-2}$ along the line of sight to GX 340+0. D’Ai et al. (2009) report similarly high values. This greatly diminishes the low energy flux from this source. Owing to the lack of sensitivity at low energy, and also owing to chip gaps creating calibration uncertainties in the 3–4 keV range, our analysis was restricted to the combined HEG spectrum and the 4–10 keV band.

An initial fit to the spectrum with an absorbed disk blackbody model ($\text{tbabs} \times \text{diskbb}$; Mitsuda et al. 1984; Wilms et al. 2000) captures the continuum fairly well, but leaves strong residuals in the Fe K band. The resultant fit statistic is $\chi^2/\nu = 686.37/566 = 1.213$. The implied column density is $N_{\text{H}} = 1.5 \times 10^{23} \text{ cm}^{-2}$, and the disk temperature is $kT = 0.85 \text{ keV}$.

Figure 1 depicts this simple model for the spectrum, as well as the data/model ratio, and deviations from the continuum as measured in units of $\Delta \chi^2$. There is a broad flux excess above the continuum in the 5.5–6.5 keV range; this is likely an Fe K emission blurred by effects in the inner accretion disk. D’Ai et al. (2009) and Cackett et al. (2010) both detected this diskline in independent analyses. The most striking and significant feature, however, is an apparent absorption line at 6.9 keV. The addition of a Gaussian model for this line improves the overall fit to $\chi^2/\nu = 641.32/563 = 1.139$.

We performed 2000 Monte Carlo simulations of an absorbed disk continuum spectrum with no line features. We fitted and conducted error bar searches on the resulting spectra with a model including a Gaussian line (emission or absorption) anywhere in the 4–10 keV range. The detection of a line feature with equal or greater strength is 99.6% unlikely. This represents a conservative probability for a blind search; in contrast, our search focused on the Fe K band. The change in $\Delta \chi^2$ indicates a line significant at the 5.6σ level of confidence, as measured by an $F$-test. Counting the number of resolution elements in the 6.7–7.3 keV band as independent trials, the significance is reduced to 5.1σ.

The best-fit line energy is $6.94 \pm 0.02 \text{ keV}$. The rest-frame energy of the He-like Fe xxv resonance line is 6.700 keV, and the rest-frame energy of the H-like Fe xxvi resonance line is 6.970 keV (Verner et al. 1996). Thus, the line could represent absorption in a modest inflow, or in a strong outflow. There are weaker features in the 6.5–6.8 keV range and in the 7.0–7.5 keV range that may aid a self-consistent determination of whether the gas is an inflow or outflow.

To address this question, physical self-consistency is required; line-by-line fitting is not sufficient. We therefore constructed a grid of XSTAR photoionized absorption models (e.g., Kallman & Bautista 2001). A turbulent velocity of 300 km s$^{-1}$ and solar abundances were assumed. A nominal gas density of $n = 10^{14} \text{ cm}^{-3}$ was also assumed. A covering factor must also be specified; we selected $\Omega/4\pi = 0.5$ based on examples of similar absorption in stellar-mass black holes (e.g.,...
Miller et al. (2015). Based on the disk blackbody fit, an input spectrum with \(kT = 0.9\) keV and a 2–30 keV luminosity of \(L = 4.6 \times 10^{37}\) erg s\(^{-1}\) was assumed (based on the unabsorbed flux and an assumed distance of 8.5 kpc; see Penninx et al. 1993). A grid of models spanning a broad range in column density and ionization was generated using the “xstar2xspece” and was included in our subsequent XSPEC fits as a multiplicative table model.

When the grid is allowed to act on the blackbody continuum, it is a 5\(\sigma\) improvement over a model without the absorption, giving \(\chi^2/\nu = 648.7/563 = 1.152\). In this fit, \(N_{\text{H,wind}} = 3.3\) \((7) \times 10^{22}\) cm\(^{-2}\) and \(\log(\xi) = 3.1(1)\). Importantly, a blueshift of \(v = -0.0395(5)c\) is measured. A model wherein the flow is required to have either zero shift or a redshift is significantly worse: \(\chi^2/\nu = 671.86/563 = 1.193\). This is true despite the fact that the ionization parameter moves up to \(\log(\xi) = 5\), as expected if FeXXVI must be very strong relative to FeXXV. A model consisting of \(\nu_{\text{H,wind}} = 2.6 \times 10^9\) cm\(^{-3}\), \(\nu_{\text{H,ISM}} = 4 \times 10^{22}\) cm\(^{-2}\), and an assumed distance of 8.5 kpc; see Penninx 2015 gives \(\chi^2/\nu = 1.193\). This is true despite the fact that the ionization parameter moves up to \(\log(\xi) = 5\), as expected if FeXXVI must be very strong relative to FeXXV. An F-test prefers the blueshifted model over the zero-shift model at the 4.5\(\sigma\) level of confidence. This result does not depend on the interstellar column density; the same blueshift results when a value of \(N_{\text{H,ISM}} = 4 \times 10^{22}\) cm\(^{-2}\) is enforced (though the overall fit is not as good).

This indicates that the absorption is best associated with an rapid outflow. However, it is possible that the overall spectral model is still too simple. Prior work has found evidence of a relativistic diskline in GX 340+0, and the broad flux excess in Figure 1 appears to confirm this feature in the combined Chandra/HETG spectrum. We there added a simple relativistic line, “diskline” (Fabian et al. 1989), in the next model. “Diskline” is characterized in terms of a line energy (restricted to the 6.40–9.7 keV range for Fe-xxvi in our fits), a disk emissivity index \((\beta \propto r^\beta\), bounded to lie in the \(-3 \leq \beta \leq -1\) range), an inner disk radius (measured in units of \(GM/c^2\)), an outer disk radius (fixed at 1000 \(GM/c^2\)), the inner disk inclination (bounded between 20\(^o\) \(\leq \theta \leq 45\(^o\)), and a flux normalization.

A model consisting of \(\text{tabsh} \times \text{xstar_abs} \times \text{diskline + diskbb}\) improves the fit at the 5.6\(\sigma\) level of confidence to \(\chi^2/\nu = 599.1/558 = 1.073\). The absorption is not substantially changed, a blueshift of \(v = -0.0395(6)c\) is again measured \((N_{\text{H,wind}} = 2.8 \pm 0.1 \times 10^{22}\) cm\(^{-2}\), \(\log(\xi) = 3.1 \pm 0.1\)). The measured diskline parameters include \(E = 6.4^{+0.1}_{-0.2}\) keV, \(q = -3^{+0.1}_{-0.1}\), \(R_{\text{in}} = 6.0^{+0.3}_{-0.2}\), \(\theta = 20^{+1}_{-1}\), and a normalization of \(K = 0.101(1)\) (translating into an equivalent width of \(W = 310 \pm 30\) eV). This model is preferred over one including a diskline but disallowing blueshifts at the 4\(\sigma\) level of confidence \(\chi^2/\nu = 618.0/558\).

Weak residuals remain in the Fe K band, suggesting that the wind is more complex than our single-zone model. A model with three zones achieves only modest improvements in the fit statistic, giving \(\chi^2/\nu = 590.1/552 = 1.069\). This fit accounts for weaker residuals in the Fe K band, but also fits some apparent lines in the 4–5 keV band that can be associated with He-like Ca XIX and H-like Ca XX (see Figures 2 and 3). We regard this model as our best-fit model; it is fully described in Table 1. The best model wherein all three components cannot be blueshifted gives \(\chi^2/\nu = 637.9/552\) (a 6.6\(\sigma\) difference for a change of one free parameter); the best three-component model wherein only one component is disallowed a blueshift gives \(\chi^2/\nu = 621.4/552\) (a 5.3\(\sigma\) difference for a change of one free parameter).

The data do not require re-emission from the wind; dynamical broadening of such emission can potentially give radius constraints (e.g., Miller et al. 2015, 2016). Winds do not have to be launched with the local escape speed; rather, they can be accelerated continually, or once certain conditions obtain. However, within this framework, an outflow velocity of \(v = 0.04c\) corresponds to a launching radius of \(r \approx 1250 \text{GM/c}^2 \approx 2.6 \times 10^9\) cm (assuming a neutron star of 1.4 \(M_\odot\)).

The mass outflow rate in each component can be calculated by starting with the standard formula, \(M = 4\pi r^2 \rho v\). Adjusting for a non-spherical flow, and writing in terms of number density, the equation becomes \(M = \Omega m_p n r^2\) (where \(\Omega\) is the covering factor, \(\mu = 1.23\) is the mean atomic weight, \(m_p\) is the mass of the proton, and \(n\) is the number density). Using the ionization parameter (recall, \(\xi = L/n\), it is possible to write the mass outflow rate without assuming a density: \(M = \Omega m_p (L/\xi)\). The kinetic power in the outflow is then \(J_{\text{kin}} = 0.5 M c^2\).

For the fast \(v = 0.04c\) component detected in GX 340+0, the mass outflow rate is \(M \approx 1.1 \times 10^{21}\) g s\(^{-1}\), or \(M = 1.8 \times 10^{-5} \text{M}_\odot\) yr\(^{-1}\). This is a very high mass flux; the implied infall rate at the inner disk is just \(5 \times 10^{17}\) g s\(^{-1}\), assuming an efficiency of \(\eta \approx 0.1\). The implied power in the fast component is even more extreme: \(J_{\text{kin}} \approx 8 \times 10^{38}\) erg s\(^{-1}\). This implies that the mechanical power exceeds the radiative eddington limit by a factor of 5–6.

The estimates can be interpreted as the mass outflow rate and kinetic power divided by the filling factor. If the volume filling factor is, e.g., \(f \approx 10^{-2}\) (commensurate with some AGNs; see, e.g., Blustin et al. 2005), the actual outflow rate is sub-Eddington, and the mass outflow rate is also reduced in direct proportion. If the launching radius derived by associating the observed wind speed with a local escape velocity is used, a very high density value results \((n \approx 1.1 \times 10^{18}\) cm\(^{-3}\)), and \(f = N_{\text{H,wind}}/n r\) implies \(f \approx 1.6 \times 10^{-4}\). This would reduce the mass outflow rate and kinetic power in GX 340+0 by four orders of magnitude. The mass outflow rate would then agree with the inferred mass accretion rate at the inner disk, to within a factor of a few. However, as noted above, there is no requirement that the observed wind speed is a local escape velocity. If the filling factor is not low, it may be the case that GX 340+0 was observed in a super-Eddington phase.

4. DISCUSSION AND CONCLUSIONS

We have analyzed an archival Chandra/HETG spectrum of the neutron star low-mass X-ray binary GX 340+0. The spectra reveal strong evidence of a single strong line close to 6.9 keV and several weaker absorption lines. Self-consistent photo-ionization modeling establishes that the feature is most likely produced in a complex disk wind, with a component that is blueshifted by \(v = 0.04c\). Even if the wind has a low volume filling factor, its mass outflow rate and kinetic power would still rank among the highest—or as the highest—yet detected from a neutron star or stellar-mass black hole. In this section, we discuss the physical processes by which the wind may be driven and compare the outflow to extreme winds detected in other sources.

Disk winds detected in Chandra observations of V404 Cyg and IGR J17091–3634 may offer the best points of comparison for GX 340+0. The 2015 outburst of V404 Cyg was particularly extreme, and it is possible that the mass
accretion rate was highly super-Eddington in some phases. In spectra of V404 Cyg, the observed velocity shifts exceed 0.01c, and the implied mass outflow rate is approximately \( \dot{M}_{\text{wind}} \approx 10^{-5} \, M_\odot \, \text{yr}^{-1} \) (assuming a unity filling factor; King et al. 2015). Spectra of IGR J17091−3624 reveal a wind with two components, with speeds of 0.03c and 0.05c. Here again,
Figure 4. Light curves of GX 340+0 from the RXTE/ASM (1.5–12.0 keV). The left panel spans 100 days, centered on Chandra ObsID 1922 (dashed red line), in which the relativistic wind is detected. The right panel spans 100 days including ObsIDs 6631 and 6632 (dashed blue lines), in which a wind is not apparent. It is clear that ObsID 1922 was obtained at a lower source flux, and it is possible that such intervals should be targeted in the future to better understand the wind in GX 340+0: Dashed, gray horizontal lines indicate the flux level at which ObsID 1922 was obtained.

Unlike the disk winds that are typically observed in stellar-mass black holes, the wind in GX 340+0 has components with an ionization parameter below $\xi = 10^3$ erg cm s$^{-1}$. Simulations have identified this as a threshold below which radiation pressure on lines can drive disk winds (e.g., Proga 2003). Directly tapping into the radiative luminosity may help to explain why the outflow in GX 340+0 has a high speed. In this sense, the observed wind may be similar to the extreme outflows in broad absorption line quasars (BALQSOs; e.g., Arav et al. 2001).

In BALQSOs, however, geometric shielding is required to keep the gas from becoming overionized by X-rays. It is unclear how such shielding might be achieved in an X-ray binary, particularly when the relativistic emission line in this system indicates a clear view of the inner disk (e.g., D’Ai et al. 2009; Cackett et al. 2010). It may be the case that a combination of mechanisms, plausibly including radiative pressure, thermal driving (e.g., Begelman et al. 1983), and magnetic processes (e.g., Blandford & Payne 1982; Proga 2003) act to drive the wind in GX 340+0. It is notable that the inclination angle of $\theta = 35^\circ \pm 1^\circ$ indicated by the relativistic line, is commensurate with the optimal angle for driving magnetocentrifugal winds (Blandford & Payne 1982).

Figure 4 shows the 1.5–12.0 keV RXTE/ASM light curve of GX 340+0, spanning intervals near the observation in which we have detected a strong wind, and two subsequent Chandra observations that appear to lack strong absorption. The wind is detected in the observation with the lowest flux, though the level is only 10%–20% below the observations lacking wind absorption. Hardness ratios might reveal more information, but the ASM ratios are insensitive owing to the very high column density along the line of sight to GX 340+0. Triggered observations based on MAXI light curves and hardness ratios may be able to reveal links between the state of the disk and wind production in GX 340+0.

### REFERENCES


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### Table 1

Spectral Fitting Results

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<td>$N_{\text{H, wind}}$ (10$^{22}$ cm$^{-2}$)</td>
<td>$4.3^{+1.1}_{-1.0}$</td>
<td>$8.2^{+1.9}_{-1.4}$</td>
<td>$4.7^{+10.8}_{-1.4}$</td>
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<tr>
<td>log($\xi$)</td>
<td>$4.4^{+0.2}_{-0.6}$</td>
<td>$2.3^{+0.8}_{-0.2}$</td>
<td>$2.8^{+0.2}_{-0.1}$</td>
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</tr>
<tr>
<td>$v/c$ (10$^{-3}$)</td>
<td>$-3.6^{+0.7}_{-1.1}$</td>
<td>$-13.3^{+0.7}_{-1.0}$</td>
<td>$-39.7^{+0.8}_{-1.0}$</td>
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</tr>
<tr>
<td>$N_{\text{H,ISM}}$ (10$^{22}$ cm$^{-2}$)</td>
<td>---</td>
<td>---</td>
<td>---</td>
<td>10.1$^{+1.3}_{-0.3}$</td>
</tr>
<tr>
<td>$kT$ (keV)</td>
<td>---</td>
<td>---</td>
<td>---</td>
<td>0.86(1)</td>
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<tr>
<td>diskbb norm.</td>
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<td>---</td>
<td>---</td>
<td>3000$^{+700}_{-300}$</td>
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<tr>
<td>$E_{\text{diskline}}$ (keV)</td>
<td>---</td>
<td>---</td>
<td>---</td>
<td>6.40$^{+0.03}_{-0.03}$</td>
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<tr>
<td>$q$</td>
<td>---</td>
<td>---</td>
<td>---</td>
<td>$3.0^{+0.3}_{-0.5}$</td>
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<tr>
<td>$R_{\text{in}}GM/c^2$</td>
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<td>---</td>
<td>---</td>
<td>7.3$^{+0.1}_{-0.3}$</td>
</tr>
<tr>
<td>$\theta$</td>
<td>---</td>
<td>---</td>
<td>---</td>
<td>35(1)</td>
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<tr>
<td>Norm. (10$^{-2}$)</td>
<td>---</td>
<td>---</td>
<td>---</td>
<td>1.2$^{+0.1}_{-0.3}$</td>
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Note. The best-fit model to the HEG spectrum of GX 340+0 is detailed above. The overall fit is good, though not formally acceptable: $\chi^2/\nu = 590.07/552 = 1.069$. Spectral continuum and wind properties are separated for clarity. The continuum was described using a disk blackbody model. A “diskline” component was included to describe a relativistic emission line. The wind was modeled using three XSTAR components, characterized in terms of the wind column density, ionization parameter, and velocity shift ($N_{\text{H,wind}}, \xi$, and $v/c$).

If the filling factor of the wind in GX 340+0 is not small, the outflow power may actually exceed the observed radiative luminosity. This would likely signal super-Eddington accretion in GX 340+0. The wind would then be driven by electron scattering pressure. At least one ultra-luminous X-ray (ULX) source is powered by a neutron star (Bachetti et al. 2014), and population considerations suggest that a number of ULXs may harbor neutron stars (King & Lasota 2016). However, a super-Eddington flow is difficult to reconcile with the detection of a relativistic diskline: the central engine should be blocked by a super-Eddington photosphere. Moreover, the observed column density in the outflow is well below $N_H \approx 10^{24}$ cm$^{-2}$ (see Table 1).