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CAN IBEX IDENTIFY VARIATIONS IN THE GALACTIC ENVIRONMENT OF THE SUN USING ENERGETIC NEUTRAL ATOMS?

Priscilla C. Frisch1, Jacob Heerikhuisen2, Nikolai V. Pogorelov2, Bob DeMajistre3, Geoffrey B. Crew4, Herbert O. Funsten5, Paul Janzen6, David J. McComas7,12, Eberhard Moebius8, Hans-Reinhard Mueller8, Daniel Brett Reisenfeld9, Nathan A. Schwadron10, Jonathan D. Slavin11, and Gary Paul Zank2

1 Department of Astronomy and Astrophysics, University of Chicago, Chicago, IL 60637, USA; frisch@oddjob.uchicago.edu
2 CSPAR, University of Alabama, Huntsville, AL 35809, USA; Jacobh@ucer.edu, nikolaip@ucer.edu, zank@email.cspar.uah.edu
3 Applied Physics Laboratory, Johns Hopkins University, Laurel, MD 20723, USA; Bob.DeMajistre@jhuapl.edu
4 Los Alamos National Laboratory, Los Alamos, NM 87545, USA; hfunsten@lanl.gov
5 Los Alamos National Laboratory, Los Alamos, NM 87545, USA; hfunsten@lanl.gov
6 Department of Physics and Astronomy, University of Montana, Missoula, MT 59812, USA; paul.janzen@umontana.edu, dan.reisenfeld@umontana.edu
7 Southwest Research Institute, San Antonio, TX 78227, USA; DMComas@swri.edu
8 Space Science Center, University of New Hampshire, Durham, NH 03824, USA; eberhard.moebius@unh.edu
9 Department of Physics and Astronomy, Dartmouth College, Hanover, NH 03755, USA; Hans-Reinhard.Mueller@Dartmouth.edu
10 Boston University, Boston, MA 02215, USA; nathanas@bu.edu
11 Harvard-Smithsonian Center for Astrophysics, Harvard, Cambridge, MA 02138, USA; jslavin@cfa.harvard.edu

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ABSTRACT

The Interstellar Boundary Explorer (IBEX) spacecraft is providing the first all-sky maps of the energetic neutral atoms (ENAs) produced by charge exchange between interstellar neutral H\(^+\) atoms and heliospheric solar wind and pickup ions in the heliosphere boundary regions. The “edge” of the interstellar cloud presently surrounding the heliosphere extends less than 0.1 pc in the upwind direction, terminating at an unknown distance, indicating that the outer boundary conditions of the heliosphere could change during the lifetime of the IBEX satellite. Using reasonable values for future outer heliosphere boundary conditions, ENA fluxes are predicted for one possible source of ENAs coming from outside of the heliopause. The ENA-production simulations use three-dimensional MHD plasma models of the heliosphere that include a kinetic description of neutrals and a Lorentzian distribution for ions. Based on this ENA-production model, it is then shown that the sensitivities of the IBEX 1.1 keV skymaps are sufficient to detect the variations in ENA fluxes that are expected to accompany the solar transition into the next upwind cloud. Approximately 20% of the IBEX 1.1 keV pixels appear capable of detecting the predicted model differences at the 3σ level, with these pixels concentrated in the Ribbon region. Regardless of the detailed ENA production model, the success of the modeled B·R ∼ 0 directions in reproducing the Ribbon locus, together with our results, indicates that the Ribbon phenomenon traces the variations in the heliosphere distortion caused by the relative pressures of the interstellar magnetic and gaseous components.

Key words: ISM: clouds – ISM: kinematics and dynamics – ISM: magnetic fields – ISM: structure – plasmas – Sun: heliosphere

1. INTRODUCTION

The dynamic heliosphere varies with the properties of the surrounding interstellar cloud and the solar wind. A theoretical study of the expected heliosphere response to different types of interstellar clouds shows that both the overall dimensions and hydrogen filtration should vary substantially with variations in the physical properties of circumheliospheric interstellar material (ISM; Müller et al. 2006, 2008). The presence of an interstellar magnetic field (ISMF) causes heliosphere asymmetries that can diagnose the properties of the surrounding ISM, but which are partially offset by the charge-exchange coupling of interstellar H\(^+\) and protons upstream of the heliopause (Pogorelov et al. 2009c; Opher et al. 2009; Ratkiewicz et al. 2008; Izmodenov 2009). The velocity discontinuity observed between interstellar gas inside of the heliosphere (e.g., Witte 2004) and ISM toward the nearest stars in the upwind direction (36 Oph and α Cen; Landsman et al. 1984; Lallement et al. 1995; Wood et al. 2000; Linsky & Wood 1996) is usually interpreted to indicate that the Sun is immersed in a separate interstellar cloud today, but will enter a separate cloud sometime in the next ~4000 years. The Interstellar Boundary Explorer (IBEX) is for the first time mapping heliospheric energetic neutral atoms (ENAs), formed by charge exchange between solar wind ions and pickup ions with neutral interstellar atoms (McComas et al. 2009a, 2009b; Funsten et al. 2009b; Fuselier et al. 2009; Schwadron et al. 2009). In this paper, we show that, depending on the source of ENAs observed, IBEX is capable of detecting the variable heliosphere boundary conditions that might accompany the expected (someday, it could be soon) solar transition into a new interstellar environment in the upwind direction.

The discovery of an unexpected “Ribbon” of ENA emission, in directions where the ISMF draping over the heliosphere is thought to be perpendicular to the sightline, showed that IBEX may be detecting plasma-neutral interactions beyond the heliopause. The similar spectra of ENAs in the Ribbon and adjacent sightlines suggest that the Ribbon represents a selection effect rather than an ENA population with a re-energized origin. McComas et al. (2009a) and Schwadron et al. (2009) noted that the Ribbon is organized by the most likely direction of the external ISMF and proposed several different potential sources of the Ribbon, including the possibility that the Ribbon might be created from a population of anisotropic suprathermal ions gyrating around the ISMF just outside the heliopause. These ions could be indigenous to the outer heliosheath (beyond the heliopause) or more likely would arise from ENAs that propagated out from

12 University of Texas, San Antonio, TX, USA.
the supersonic solar wind and/or inner heliosheath (between the termination shock and the heliopause); these authors noted that the problem with this idea is that the relatively confined pitch angle distributions would need to be maintained long enough to create “secondary ENAs,” which likely takes several years on average. Subsequently, Heerikhuisen et al. (2010) incorporated this idea, where outward propagating ENAs create an anisotropic population of pickup ions in the outer heliosheath that seed “secondary ENA” production several hundred AU upwind of the heliopause, into a quantitative model. While it is still uncertain how long the ion ring beam takes to scatter into a shell distribution (Florinski et al. 2010), with marginally stable ring distributions predicted by some models for the distribution of the pitch angles of pickup ions in the outer heliosheath (Gamayunov et al. 2010), the Heerikhuisen et al. simulation assumes that the re-neutralization time is essentially instantaneous compared to the scattering time so that the new secondary ENA will have a preferred direction that is perpendicular to the local ISMF direction. In this model, IBEX then sees these secondary ENAs where the gyration plane of the ion is aligned with the sightline to IBEX, i.e., where the sightline is perpendicular to the ISMF direction.

IBEX data are uniquely qualified to simultaneously test both the direction of the ISMF at the heliosphere and the density of interstellar neutrals. The ISMF drapes around the heliosphere, rotating by ∼30° between “infinity” and the heliopause near the nose. The Ribbon is ∼20° wide and at least 270° long, possibly forming a complete circle in the sky. IBEX only sees ENAs with momentum vectors directed back toward the inner heliosphere. The long mean free paths of ENAs, e.g., ∼200 AU for 1.1 keV ENAs in n(p+) ∼0.1 cm−3 plasma, allow detection in the inner heliosphere of secondary ENAs formed in regions beyond the heliopause with elevated interstellar densities and relatively isotropic ENA velocities (compared to the outward radial flows for primary ENAs produced in the supersonic solar wind, although not compared to the inner heliosheath ion populations). The solar wind contributing to ENA production includes both core ions from the expanding solar corona and pickup ions formed inside of the heliosphere by charge exchange between interstellar neutrals and the core solar wind.

The predictive capabilities of global heliosphere models have improved significantly to accommodate observational constraints placed by the 10 AU difference in the termination shock distances found by the Voyager 1 and Voyager 2 spacecraft (Stone 2008; Opher et al. 2009), the ∼5° offset between the upwind directions of interstellar H² and He⁰ flowing into the heliosphere determined from Solar and Heliospheric Observatory/ SWAN and Ulysses/GAS data (Lallement et al. 2005; Witte 2004; Frisch 2010, where the He⁰ upwind direction must first be converted to J2000 coordinates for this comparison), the properties of the ISM surrounding the heliosphere (Slavin & Frisch 2008), and now the IBEX data on ENA fluxes and the Ribbon. Although the IBEX Ribbon itself was not predicted by models of ENA formation in the heliosphere, the global heliosphere models provided the ISMF orientation that matches well with the configuration of the Ribbon in the sky (Schwadron et al. 2009; Pogorelov et al. 2009c).

In the discussions below, we rely on the ENA-production models quantified by Heerikhuisen et al. (2010) and Pogorelov et al. (2009b) to predict the ENA fluxes for a heliosphere immersed in the next interstellar cloud versus the present-day surrounding interstellar cloud. The ENA-production simulations use three-dimensional MHD plasma models of the heliosphere that include a kinetic description of neutrals and a Lorentzian distribution for solar wind protons to approximate the suprathermal population of pickup ions in the heliosheath region. The interstellar neutrals, which are coupled self-consistently to the plasma component in the MHD heliosphere models, act to symmetrize the heliosphere. Any asymmetry in the quiescent plasma distribution (e.g., created by ISMF) results in variations in the number of charge-exchange events, creating new ions that are decelerated by the heliopause and that therefore mitigate the asymmetry (Pogorelov et al. 2008a, 2009b, 2009c). The local interstellar material (LISM) flow, for the assumed ISMF strength, is subfast magnetosonic (Table 1). This results in the absence of a bow shock in front of the heliopause and increases the width of the region where the ISMF deviates from its unperturbed orientation. Since the local ISMF direction varies as interstellar protons approach the heliopause, the mean free path for the charge-exchange interaction must also be included self-consistently in any Ribbon production models.

The Ribbon ENAs are formed upwind of the heliopause in the Heerikhuisen et al. (2010) model, so the outer boundary conditions set by the ISM have a direct impact on predicted ENA fluxes and provide a means of estimating ENA variations from a cloud transition, regardless of whether the model is correct in detail. Any model that reproduces the observed B · R ∼ 0 of the Ribbon, which is seen where the ISMF (B) is perpendicular to the sightline (R), should provide useful insights into the deformation of the heliosphere due to altered boundary conditions from the variable interstellar wind. Our conclusions here rely explicitly on the assumption that the Heerikhuisen et al. (2010) model provides a viable description of ENA production, both for the cloud we are in today and for the nearby cloud observed toward α Cen and 36 Oph in the upwind direction.

We focus on the 1.1 keV data, because the contrast between Ribbon ENA fluxes and diffuse ENA fluxes is stronger at this energy, and partly because of the enhanced outward flow of ENAs from core solar wind ions that have typical energies near 1 keV. The ENA spectra are explicitly predicted by the heliosphere models, however, this spectral information is not used here. The observed fluxes of ∼4.5 keV ENAs are an order of magnitude lower than the 1.1 keV fluxes and the mean free paths are ∼50% larger. Since the outflowing ion fluxes will decrease as ∼R−2 with distance from the Sun R, both parent ion densities and the resulting 1 AU ENA densities are lower when production regions are further from the Sun.

2. PROPERTIES OF THE UPWINDISM

The boundary conditions of the heliosphere are set by the ISM and vary over geologically short timescales. Interstellar clouds within ∼50 pc flow past the Sun with heliocentric velocities that cluster around ∼28 km s⁻¹ (after projection effects are removed), or ∼1 pc per 35,000 yr (Frisch & Slavin 2006). If the nearby ISM is in pressure equilibrium, then models of the interstellar cloud around the heliosphere (Slavin & Frisch 2008) combined with ISM data (Redfield & Linsky 2004) yield an estimate for the typical cloud length of ∼1 pc. Cloud column densities for 23 cloud components toward stars within 10 pc yield a range of cloud lengths 0.06–3 pc, giving typical cloud crossing times for the Sun of ∼1450–2.8 × 10⁸ yr.³¹

³¹ These estimates assume an equilibrium thermal pressure of ∼3 × 10⁻¹³ dynes cm⁻² for the present-day cloud, uniform magnetic, and cosmic ray pressures, D/H ∼1.6 × 10⁻⁵, and a uniform proton density of 0.08 cm⁻³.

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Properties of Circumheliospheric ISM used for Models

<table>
<thead>
<tr>
<th>Quantity</th>
<th>LIC</th>
<th>Model 1 Original</th>
<th>Model 2 Today</th>
<th>Model 3 Next (Same ISMF)</th>
<th>Model 4 Next (New ISMF)</th>
</tr>
</thead>
<tbody>
<tr>
<td>(n(H^\circ) \text{ (cm}^{-3}\text{)})</td>
<td>0.19</td>
<td>0.15</td>
<td>0.15</td>
<td>0.23</td>
<td>0.23</td>
</tr>
<tr>
<td>(n(p^\circ) \text{ (cm}^{-3}\text{)})</td>
<td>0.06</td>
<td>0.06</td>
<td>0.06</td>
<td>0.08</td>
<td>0.08</td>
</tr>
<tr>
<td>(V \text{ (km s}^{-1}\text{)})</td>
<td>26.3</td>
<td>26.4</td>
<td>26.4</td>
<td>28.8</td>
<td>28.8</td>
</tr>
<tr>
<td>(T \text{ (K)})</td>
<td>6300</td>
<td>6530</td>
<td>6530</td>
<td>5400</td>
<td>5400</td>
</tr>
<tr>
<td>(</td>
<td>B</td>
<td>\text{ (}\mu\text{G)})</td>
<td>[2.7]</td>
<td>3</td>
<td>3</td>
</tr>
<tr>
<td>(B) direction, (\lambda, \beta)</td>
<td>237°, 30°</td>
<td>224°, 41°</td>
<td>224°, 41°</td>
<td>252°, 42°</td>
<td></td>
</tr>
<tr>
<td>(B) direction, (\ell, b)</td>
<td>22°, 41°</td>
<td>36°, 53°</td>
<td>36°, 53°</td>
<td>40°, 32°</td>
<td></td>
</tr>
<tr>
<td>No. of Magneto sonic Mach number</td>
<td>1.0</td>
<td>1.1</td>
<td>1.1</td>
<td>0.8</td>
<td>0.8</td>
</tr>
</tbody>
</table>

Notes.

- a The 1 AU solar wind parameters used for all models are \(n(p^\circ) = 7.4 \text{ cm}^{-3}\), \(T = 51,100 \text{ K}, V = 450 \text{ km s}^{-1}\), and \(|B\)\text{radial} = 37.5 \mu\text{G}.
- b These values for the ISM forming the heliosphere boundary conditions are based on Model 26 in Slavin & Frisch (2008) and Witte (2004).
- c Schwadron et al. (2009) used this model (Pogorelov et al. 2008a) in the initial analysis of IBEX data.
- d This model reproduces the IBEX Ribbon (Heerikhuisen et al. 2010).
- e The next-cloud model, assuming the same ISMF direction as the today-model, Model 2.
- f The next-cloud model, assuming an ISMF direction that differs from Model 2.
- g Determined by assuming that thermal and magnetic pressures are equal.

The velocity of the ISM at the heliosphere, which we term here the heliospheric ISM but which is also known as either the local interstellar cloud (LIC) or the circumheliospheric ISM, is best set by the velocity of interstellar He\(^{\circ}\) observed inside of the heliosphere by the Ulysses satellite (Witte 2004). Interstellar He\(^{\circ}\) experiences minimal filtration in the heliosheath regions (<2%; Müller & Zank 2004) and is not subject to radiation pressure, so that variation in the interstellar He\(^{\circ}\) velocity as it traverses the heliosphere is minimal. The first comparisons between the velocities of interstellar gas in the heliosphere and outside of the heliosphere, toward stars in the upwind direction, showed differences of over 3 km s\(^{-1}\) (Adams & Frisch 1977). When the 26.3 km s\(^{-1}\) He\(^{\circ}\) velocity is projected toward the nearest star in the upwind direction \(\alpha\) CMa, 1.3 pc away and 50° from the heliosphere nose, the projected heliospheric ISM velocity is 170 km s\(^{-1}\), in contrast to the observed cloud velocity from the unsaturated Fe\(^{+}\) and Mg\(^{+}\) lines of \(-18.0 \pm 0.2 \text{ km s}^{-1}\) (Linsky & Wood 1996). The heliospheric nose direction is given by the upwind direction of He\(^{\circ}\) flowing through the heliosphere, or \(\lambda, \beta = 255.4 \pm 0.5, +5.1 \pm 0.2\) (epoch J2000; M. Witte 2004, private communication). The star 36 Oph is 6 pc away and 10° from the heliospheric nose. The projected heliospheric ISM velocity in this direction is \(-25.9 \text{ km s}^{-1}\) versus the observed cloud velocity from the Fe\(^{+}\) and Mg\(^{+}\) lines of \(-28.1 \pm 0.2 \text{ km s}^{-1}\) (Wood et al. 2000). The limit on a cloud component at the heliospheric ISM velocity toward 36 Oph is \(\sim 6 \times 10^{16} \text{ cm}^{-2}\), giving an upper limit to the heliospheric ISM edge\(^{14}\) in this direction of 0.1 pc which will be traversed in less than 4000 years.

The interstellar cloud observed toward 36 Oph and \(\alpha\) Oph is known as the G-cloud (Lallement & Bertin 1992; Frisch et al. 2002; Redfield & Linsky 2008), since it is close to the Sun in the galactic hemisphere. Although other possible clouds have been suggested as the next cloud to be encountered by the heliosphere (Frisch 2003), the G-cloud remains the most likely future heliospheric environment.

The scenario examined here assumes that the Sun transitions directly from the heliospheric ISM to the G-cloud seen toward 36 Oph and \(\alpha\) Cen. We determine the properties of the next cloud by assuming that the thermal and magnetic pressures in the heliospheric ISM are equal and that the heliospheric ISM and G-clouds are in pressure equilibrium. The G-cloud temperature is found from the mass-dependent Doppler broadened widths of interstellar absorption lines, and is 5400 ± 500 K toward \(\alpha\) Cen, and 5900 ± 500 toward 36 Oph (where the cloud column density is also larger by 70%). The cooler G-cloud temperatures are thus compensated by neutral densities that are 20% larger than the heliospheric ISM. The 9% higher velocity of the G-cloud, in comparison to the velocity of the ISM now surrounding the Sun, will increase the interstellar ram pressure even if thermal pressure and ionization levels remain constant. We test the sensitivity of ENA emission to the ISMF direction using two separate ISMF directions for the G-cloud. The first assumption is that the directions of the ISMF in the heliospheric ISM and G-clouds are the same (Models 2 and 3 in Table 1). For Model 4, we make an arbitrary\(^{15}\) assumption for the G-cloud ISMF direction, which is 28° different from today’s field but less than the \(\sim 30°\) rotation of the ISMF between the ISM and the heliopause for the upwind direction. The detailed heliosphere boundary conditions used here for the next cloud are listed in Table 1. We show below that the IBEX Ribbon is highly sensitive to even small variations in the ISMF direction, even when increased H\(^{\circ}\) densities mitigate the influence of the ISMF on heliospheric asymmetries.

3. ENA Fluxes from Encounter with Next Interstellar Cloud

The heliosphere model has been run for the interstellar boundary conditions listed for Models 1–4 in Table 1. The 1.1 keV ENA fluxes predicted by the resulting models are

\(^{14}\) We use \(n(H^\circ) = 0.19 \text{ cm}^{-2}\) for the heliospheric ISM H\(^{\circ}\) density, based on Model 26 in the radiative transfer models of Slavin & Frisch (2008).

\(^{15}\) Caveat: this direction, which is arbitrary from an interstellar viewpoint, was selected to lie on the great circle that divides the hot and cold poles of the cosmic microwave background dipole moment, and which passes through the heliosphere nose region (Frisch 2010).
displayed in Figure 1. Model 1 corresponds to the heliosphere model displayed in Schwadron et al. (2009). Model 2 (from Heerikhuisen et al. 2010) is the updated model we use for the heliosphere today. Models 3 and 4 correspond to the anticipated heliosphere environment in the next interstellar cloud, with Model 3 having a similar ISMF configuration as Model 2, and Model 4 showing a ISMF field with a different direction. The increased H\(^+\) density in Model 3 produces a brighter Ribbon, while the ram and thermal pressures slightly increase and the magnetic pressure slightly decreases. The magnetic field in Model 2 plays a bigger role in deforming the heliosphere than in Model 3. Model 4 has the Ribbon shifted significantly because of the different ISMF direction.

The differences between the predicted ENA fluxes for Model 2 (the “today-cloud”) and Model 3 (the G-cloud assuming the same ISMF direction as the today-cloud) are directly displayed in Figure 2. The ENA flux differences are obtained by subtracting the predicted fluxes of Model 2 (\(F_{\text{Mod2}}\)) from the predicted fluxes of Model 3 (\(F_{\text{Mod3}}\), left). The most significant difference between the two models is the ram pressure of the neutrals, which is a factor of 1.8 larger for Model 3 versus Model 2. Over most of the sky, the higher flux of interstellar H\(^+\) into the heliosphere for Model 3 generates larger ENA fluxes, with the differences approaching the brightest observed fluxes. However, the red pixels in Figure 2 (left) show regions where the today-cloud has higher fluxes than the next cloud and represent the small shift in the Ribbon position due to the increase in the ratio of thermal to magnetic pressures, \(P_{\text{therm}}/P_{\text{mag}}\), in Model 3 compared to Model 2. The effect of increased H\(^+\) densities and thermal pressures are also seen in the increased ENA emissivity of the eastern flank of the heliosphere in the nose direction, where there is a bulge in total pressure (magnetic and thermal; see Figure 1 in Heerikhuisen et al. 2010). Figure 2 (right) shows the flux differences between Model 2 and Model 4, where the ISMF direction in the G-cloud has been allowed to vary by 20\(^\circ\). This modest variation in the ISMF direction, while retaining the same field strength, leads to an obvious shift in the Ribbon location and shows that variations in the direction of the ISMF draping over the heliosphere should be apparent in the ENA data.

The ENA-production model used here predicts that the Ribbon and non-Ribbon regions respond differently to variations in the interstellar density, because the Ribbon also directly traces heliosphere asymmetries created by the ISMF. Figure 3 shows the model differences modified by the ratio of \(n(H^+)\) in Models 2 and 3 (0.65). The left figure shows 0.65 * \(F_{\text{Mod3}}\) – \(F_{\text{Mod2}}\) and the right figure shows 0.65 * \(F_{\text{Mod4}}\) – \(F_{\text{Mod2}}\). The background blue regions, where difference counts are \(\sim 0\), indicate regions where the ENA fluxes are linearly related to the neutral interstellar density. The variations in the Ribbon colors show that the Ribbon ENAs trace the distortion of the heliosphere, which depends on the asymmetries introduced by the relative interstellar, ram, and thermal pressures, and which changes with different ISM conditions.

The differences in the ENA fluxes predicted by the two next-cloud models are quite obvious in the high-flux regions of the Ribbon, but less obvious (for this color scale) for the directions toward the tail. In order to emphasize the differences in the weaker diffuse ENA emission originating in the low-flux tail region, Figure 4 shows the percentage differences between Models 3 and 2, (\(F_{\text{Mod3}}\) – \(F_{\text{Mod2}}\))/\(F_{\text{Mod2}}\), and Models 4 and 2, (\(F_{\text{Mod4}}\) – \(F_{\text{Mod2}}\))/\(F_{\text{Mod2}}\), with enhanced color scales. The percentage differences in the tail region for Model 3 are larger than for Model 4. The outer heliosheath region around the tail is the most disturbed part of the model results, since the flows are subsonic. Hence, larger relative variations in ENA fluxes are possible.

The capability of IBEX to detect the ENA variations shown in Figures 2–4 rests on the predicted differences in the modeled ENA fluxes compared to the flux uncertainties for the IBEX data. For this comparison, we use the first 1.1 keV ENA flux skymaps in the third energy passband (“ESA 3”) of the six energy passbands in the IBEX-HI neutral atom imager (Funsten et al. 2009a), where fluxes are an order of magnitude larger than at 4.5 keV (ESA 6). Figure 5 (left) shows ESA 3 fluxes, after correction for the Compton–Getting (CG) shift in the energy and spatial distribution of high velocity particles due to the 30 km s\(^{-1}\) orbital motion of the Earth (e.g., Gleeson & Axford 1968). IBEX pixels are \(\sim 7^\circ\). The CG corrections are based on a power-law energy spectra that are derived from adjacent energy passbands, typically \(\sim E^{-1.6}\), which is evaluated over the look-direction and convolved with the energy response of the detector (see the Appendix in McComas et al. 2010, for details on the CG correction to the ENA energies measured by IBEX). The 1\(\sigma\) uncertainties (\(d F_{\mu}\)) on these fluxes have been determined from the Poisson statistics propagated from the measurement uncertainties. The IBEX data are built from data processed by the IBEX Science Operations Center (ISOC) for the first all-sky IBEX map (orbits 11–34). The uncertainties in the ESA 3 fluxes are shown in Figure 5 (right).

\(^{16}\) We have used the IBEX CG corrected data set “fixset_h6d0-id-base-0071-2010-04-09.sav” that is available at the IBEX Science Operations Center (ISOC).
Figure 2. Differences between the ENA fluxes produced by the nominal G-cloud and the nominal LIC. Left: Model 3–Model 2 ($F_{\text{Mod3}} - F_{\text{Mod2}}$). Right: Model 4–Model 2 ($F_{\text{Mod4}} - F_{\text{Mod2}}$). The different distortions of the heliosphere caused by variations in the relative magnetic and thermal ram pressures of the ISM are apparent on the northeast flank of the heliopause in the left figure. The right figure shows that the Ribbon configuration is highly sensitive to variations in the ISMF direction. The differences in the distortion of the heliosphere toward the tail region are marginally visible. Fluxes are given in the units of counts cm$^{-2}$ s$^{-1}$ sr$^{-1}$ keV$^{-1}$.

Figure 3. Differences between the ENA fluxes produced by the nominal G-cloud and the nominal present-day cloud, after normalizing the fluxes in Models 3 and 4 by the ratio of the neutral densities ($0.65$). Left: ($n(H^\circ)_{\text{Mod2}}/n(H^\circ)_{\text{Mod3}})^*F_{\text{Mod3}} - F_{\text{Mod2}}$. Right: ($n(H^\circ)_{\text{Mod2}}/n(H^\circ)_{\text{Mod4}})^*F_{\text{Mod4}} - F_{\text{Mod2}}$. Zero values of the difference (blue) along the Ribbon imply that the Ribbons in the two models overlay each other. The background blue regions, where difference counts are $\sim 0$, indicate the regions where ENA fluxes are linearly related to the neutral interstellar density in these models. Fluxes are given in the units of counts cm$^{-2}$ s$^{-1}$ sr$^{-1}$ keV$^{-1}$.

Figure 4. Percentage differences between the 1.1 keV ENA fluxes produced by the nominal G-cloud and the nominal present-day cloud. Left: $100(F_{\text{Mod3}} - F_{\text{Mod2}})/F_{\text{Mod2}}$. Right: $100(F_{\text{Mod4}} - F_{\text{Mod2}})/F_{\text{Mod2}}$. The right figure shows that the Ribbon configuration is highly sensitive to any variation in the ISMF direction. The measurement uncertainties for 1.1 keV fluxes can be compared to the predicted ENA flux differences for the next cloud versus the present cloud, e.g., $|\Delta F_{\text{model}}| = |F_{\text{Mod3}} - F_{\text{Mod2}}|$, based on the Heerikhuisen et al. (2010) model. For comparison, we preselect data points with $S/N > 3$. In Figure 6, individual pixels in the Ribbon region for both models have values of $|\Delta F_{\text{model}}|/d F_{1\gamma} \gg 5$. The same difference map is plotted in Figure 7, but now color-coded to enhance the differences in the tail region. Lower ENA fluxes toward the tail yield $|\Delta F_{\text{model}}|/d F_{1\gamma} \sim 1–2$ for individual pixels, which is somewhat larger for Model 3 than for Model 4. Groups of 25 pixels would yield a factor of 5 improvement in the signal-to-noise ratio ($S/N$).
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Figure 5. Left: the 1.1 keV IBEX ENA first flux skymap, corrected for the CG effect (see the text). Right: the 1σ statistical uncertainties on these data (see the text).

Figure 6. Differences between ENA fluxes predicted by models constrained by different boundary conditions and compared to the 1σ measurement uncertainties in the ESA 3 energy passband (1.1 keV, see the text). Left: the ratio of the absolute value of Model 3 minus Model 2 fluxes, to 1σ uncertainties (|F_{Model3}−F_{Model2}|/dF_{1σ}). Right: the same as left figure, but showing the absolute value of Model 4 minus Model 2 fluxes (|F_{Model4}−F_{Model2}|/dF_{1σ}). Pixels are left blank where fluxes are insignificant (S/N < 3).

Figure 7. Same plot as Figure 6, except that the color scale is coded to enhance the low-flux areas in the tail region.

of the difference maps, while effectively smoothing the data over ~125 deg², and should still provide a significant test. In order to use ENAs from the tail for identifying the next cloud, either pixels in the tail must be grouped to improve statistics or the comparison should wait for the better statistics of future skymaps.

The predicted ENA flux differences between the today-cloud and the next cloud are testable with IBEX data. Twenty percent of the ESA 3 (1.1 keV) pixels with S/N > 3 test the flux differences between Model 3 and Model 2 at the 3σ level, or |ΔF_{model}|/dF_{1σ} > 3. In addition, 49% of the pixels test these flux differences at the 1σ level, with |ΔF_{model}|/dF_{1σ} > 1 (Figure 8). Similar values are found for comparisons between the predicted flux differences between Model 4 and Model 2, where 18% of the pixels show model differences that are larger than the 3σ ESA 3 flux uncertainties.

We have also evaluated the variations in the 4.5 keV ENA fluxes for the environment of the next cloud (Figure 9). Although the count rates in IBEX-HI ESA 6, at 4.5 keV, are lower by an order of magnitude than at 1.1 keV (Figure 10, left), flux...
detected by IBEX with the measurement uncertainties of the 1.1 keV ENA fluxes today. We predict the flux of ENAs from the next interstellar cloud significantly exceed the measurement uncertainties for 20% of all ESA 3 pixels, while 49% of all pixels test these differences at the 1σ level, for example.

Figure 8. ESA 3 fluxes (ordinate) are compared to the significance of the model differences (abscissa) at each pixel on the sky for Models 3 and 2 (left) and Models 4 and 2 (right). The abscissa shows the differences in the fluxes of the two models, divided by the 1σ measurement uncertainties on the ESA 3 data. Only “good” ESA 3 pixels with S/N > 3 are plotted. The differences between Models 3 and 2 are tested at the 3σ level by 20% of all ESA 3 pixels, while 49% of all pixels test these differences at the 1σ level, for example.

Figure 9. Predicted ENA fluxes at 4.5 keV (ESA 6 energy passband) are shown for Model 2 (left) and Model 3 (right).

Predicted ENA fluxes at 4.5 keV (ESA 6 energy passband) are shown for Model 2 (left) and Model 3 (right).

Variations in the energy dependence and fluxes of ENAs will occur because of the variation of solar wind properties over the solar cycle. These solar cycle contributions fortunately can be modeled in detail using past and present data on the solar wind, and models of the heliosphere response to these disturbances of magnetic pressure and thermal ram pressure that deform the heliosphere are sensitive functions of the boundary conditions imposed by the ISM. If, in addition, the direction of the ISMF shifts by as much as 28°, which is slightly larger than the Ribbon width, then significant differences in the ENA fluxes should be observed in individual pixels near the Ribbon. The heliosphere regions with very low ENA fluxes, such as the tail, provide a test of the cloud properties only if pixels are combined for better statistical significance before comparison with model predictions. The long term capability of IBEX to realistically detect such variations, which will also be superimposed on possible solar cycle variations, requires that the efficiencies in the IBEX sensors (conversion, scattering, sputtering in the conversion subsystem, secondary electron emission at the detector foils, and microchannel plate efficiencies, for instance) either remain stable over years, or alternatively that the instrument performances are tightly monitored. IBEX-HI detector sensitivity is continuously monitored in a number of ways, such as comparison of coincident and non-coincident count rates (Funsten et al. 2005), and periodic gain tests.

Variations in the energy dependence and fluxes of ENAs will occur because of the variation of solar wind properties over the solar cycle. These solar cycle contributions fortunately can be modeled in detail using past and present data on the solar wind, and models of the heliosphere response to these variations. Every IBEX skymap is a historical map of the solar cycle because of the energy dependence of ENA travel times and cross sections (McComas et al. 2010), so that unraveling the solar cycle dependence will simultaneously constrain the ENA-production models and improve future predictions of the ENA variations due to the next interstellar cloud.

In this discussion, we considered the scenario where the next cloud is faster, slightly cooler, and more dense than the
heliospheric ISM gas, as expected from pressure equilibrium and observational data. This study is a proof-of-concept, since the properties of the cloud edges are not established. The more extreme possibilities for the next galactic environment to be encountered by the Sun include a hot plasma without neutrals and a cloud interface that is either evaporative or mixed with hot plasma by shear flows. If interstellar clouds within 10 pc are in pressure equilibrium, they will typically fill 20% of the sightline to the stars. The intervening voids evidently will be filled with the low density hot gas that creates the Local Bubble soft X-ray emission, although the emissivity of this plasma is somewhat uncertain because of solar wind contamination (Koutroumpa et al. 2009). Should the heliosphere enter the diffuse plasma attributed to the Local Bubble interior, both interstellar neutrals and exo-heliospheric ENAs will vanish. IBEX and other spacecraft will readily detect this condition. Another possibility for the next solar environment would be an evaporative interface that would form upwind between the heliospheric ISM and hot plasma. Such an interface will show steep increases in the cloud velocity and pressure, and decreases in density, over spatial scales that are determined by the angle between the ISMF and cloud surface (Slavin & Frisch 2008).

The present study considers the ENA fluxes for two separate models of the circumheliiospheric cloud, but ignores possible variations due to changes in the heliosphere configuration during the transitions between the two clouds. The predicted thickness of the conductive boundary on the cloud around the heliosphere, defined as where the temperature falls to 50% of the asymptotic temperature, is 0.32–0.34 pc for an ISMF direction that makes an angle of 30° with the cloud surface (Slavin & Frisch 2008, Models 26 and 27, also see Figure 2, where the cloud edge starts at 3 pc). In the upwind direction, the outflow speeds in the conductive boundary are 20–30 km s$^{-1}$, and opposite to the cloud motion, so that the Sun could traverse the conductive boundary in approximately $\sim$12,000 years for these models. Based on the above models, we expect the change in heliosphere properties for such an environment to be clearly observable in the resulting ENA flux detected by IBEX. A turbulent mixing layer will also produce strong gradients in the temperature and ionization of the surrounding ISM (Slavin et al. 1993). The ENA emissions for a conductive boundary on the surrounding cloud are discussed in detail in Grzedzielski et al. (2010), where the Sun is estimated to emerge from the interface within $\sim$500 years. An alternative possibility is that the G-cloud may be denser than has been assumed in this study. If the interstellar N(Ca +) absorption formed at the G-cloud velocity is entirely within a few parsecs of the Sun, then comparisons between the clouds in the $\alpha$ Cen and $\alpha$ Oph sightlines suggest a tiny cold cloud in addition to the warmer gas (Frisch 2003).

The comparisons in this paper are made without consideration of the solar cycle, although the outer heliosheath regions respond to the variations in the solar wind dynamic pressure and magnetic field that characterize the solar activity cycle (Washimi & Tanaka 1999; Scherer & Fahr 2003; Zank & Müller 2003; Pogorelov et al. 2009a). Although the solar cycle will cause the heliosphere to expand and contract as the solar wind dynamic pressure changes, these pulses travel only a relatively short distance upstream of the heliopause ($\sim$100 AU). The influence of the solar cycle on ENA production and the Ribbon phenomenon is not yet understood. The Ribbon intensity may vary over latitudes due to the ion energy differences and travel times. The extremely low levels of solar activity during the first year of IBEX observations suggest that the solar activity cycle variations must first be understood before reaching a conclusion that we have entered a new interstellar cloud (Pogorelov et al. 2008b; Sternal et al. 2008). As the theoretical models of ENA production become increasingly robust, we expect that studies such as this will yield definitive information on both the heliosphere boundary conditions and the physical properties of the interstellar cloud around the Sun. Finally, while this study has examined only one of the possible sources of the Ribbon current under discussion, the other ideas for producing the Ribbon (McComas et al. 2009a) also generally invoke and seek to match up with the orientation of the external ISMF, so even if another explanation eventually becomes accepted, it may still be possible to directly detect the interstellar transition with IBEX.

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REFERENCES
