

OPEN ACCESS

(In)direct detection of boosted dark matter

To cite this article: Kaustubh Agashe *et al* JCAP10(2014)062

View the [article online](#) for updates and enhancements.

Related content

- [Detecting boosted dark matter from the Sun with large volume neutrino detectors](#)
Joshua Berger, Yanou Cui and Yue Zhao
- [Dark matter, shared asymmetries, and galactic gamma ray signals](#)
Nayara Fonseca, Lina Necib and Jesse Thaler
- [Production regimes for Self-Interacting Dark Matter](#)
Nicolás Bernal, Xiaoyong Chu, Camilo Garcia-Cely *et al*.

Recent citations

- [Foraging for dark matter in large volume liquid scintillator neutrino detectors with multiscatter events](#)
Joseph Bramante *et al*
- [First Direct Search for Inelastic Boosted Dark Matter with COSINE-100](#)
C. Ha *et al*
- [Dark matter “transporting” mechanism explaining positron excesses](#)
Doojin Kim *et al*



IOP Astronomy ebooks

Part of your publishing universe and your first choice for astronomy, astrophysics, solar physics and planetary science ebooks.

iopscience.org/books/aas

(In)direct detection of boosted dark matter

Kaustubh Agashe,^a Yanou Cui,^a Lina Necib^b and Jesse Thaler^b

^aMaryland Center for Fundamental Physics, University of Maryland,
College Park, MD 20742, U.S.A.

^bCenter for Theoretical Physics, Massachusetts Institute of Technology,
Cambridge, MA 02139, U.S.A.

E-mail: kagashe@umd.edu, cuiyo@umd.edu, lnecib@mit.edu, jthaler@mit.edu

Received June 20, 2014

Revised September 9, 2014

Accepted September 21, 2014

Published October 24, 2014

Abstract. We initiate the study of novel thermal dark matter (DM) scenarios where present-day annihilation of DM in the galactic center produces boosted stable particles in the dark sector. These stable particles are typically a subdominant DM component, but because they are produced with a large Lorentz boost in this process, they can be detected in large volume terrestrial experiments via neutral-current-like interactions with electrons or nuclei. This novel DM signal thus combines the production mechanism associated with indirect detection experiments (i.e. galactic DM annihilation) with the detection mechanism associated with direct detection experiments (i.e. DM scattering off terrestrial targets). Such processes are generically present in multi-component DM scenarios or those with non-minimal DM stabilization symmetries. As a proof of concept, we present a model of two-component thermal relic DM, where the dominant heavy DM species has no tree-level interactions with the standard model and thus largely evades direct and indirect DM bounds. Instead, its thermal relic abundance is set by annihilation into a subdominant lighter DM species, and the latter can be detected in the boosted channel via the same annihilation process occurring today. Especially for dark sector masses in the 10 MeV–10 GeV range, the most promising signals are electron scattering events pointing toward the galactic center. These can be detected in experiments designed for neutrino physics or proton decay, in particular Super-K and its upgrade Hyper-K, as well as the PINGU/MICA extensions of IceCube. This boosted DM phenomenon highlights the distinctive signatures possible from non-minimal dark sectors.

Keywords: dark matter theory, neutrino detectors, cosmology of theories beyond the SM

ArXiv ePrint: [1405.7370](https://arxiv.org/abs/1405.7370)



Contents

1	Introduction	1
2	Two component dark matter	4
3	Thermal relic abundances and present-day annihilation	5
4	Detecting boosted dark matter	7
4.1	Flux of boosted dark matter	7
4.2	Detection of boosted dark matter	9
4.3	Backgrounds to boosted dark matter	11
4.4	Impact of earth attenuation	15
5	Detection prospects for present and future experiments	16
5.1	Event selection	17
5.2	Signal rates	18
5.3	Background rates	20
5.4	Estimated experiment reach	21
6	Summary of existing constraints	23
7	Conclusions and other possibilities	25
A	Analytic approximations to relic abundances	27
B	Direct detection of non-boosted DM	29
C	Boosted DM scattering off hadrons	31
C.1	Elastic scattering off hadrons	31
C.2	Deep inelastic scattering off hadrons	32

1 Introduction

A preponderance of gravitational evidence points to the existence of dark matter (DM) [1–3]. Under the compelling assumption that DM is composed of one or more species of massive particles, DM particles in our Milky Way halo today are expected to be non-relativistic, with velocities $v_{\text{DM},0} \simeq \mathcal{O}(10^{-3})$. Because of this small expected velocity, DM indirect detection experiments are designed to look for nearly-at-rest annihilation or decay of DM, and DM direct detection experiments are designed to probe small nuclear recoil energies on the order of $\frac{\mu^2}{m_N} v_{\text{DM},0}^2$ (μ is the reduced mass of the DM-nucleus system, m_N is the nucleus mass). In addition, these conventional detection strategies are based on the popular (and well-motivated) assumption that DM is a weakly-interacting massive particle (WIMP) whose thermal relic abundance is set by its direct couplings to the standard model (SM).

In this paper, we explore a novel possibility that a small population of DM (produced non-thermally by late-time processes) is in fact relativistic, which we call “boosted DM”. As a concrete example, consider two species of DM, ψ_A and ψ_B (which need not be fermions),

with masses $m_A > m_B$. Species ψ_A constitutes the dominant DM component, with no direct couplings to the SM. Instead, its thermal relic abundance is set by the annihilation process¹

$$\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B. \quad (1.1)$$

At the present day, non-relativistic ψ_A particles undergo the same annihilation process in the galactic halo today, producing relativistic final state ψ_B particles, with Lorentz factor $\gamma = m_A/m_B$. These boosted DM particles can then be detected via their interactions with SM matter at large volume terrestrial experiments that are designed for detecting neutrinos and/or proton decay, such as Super-K/Hyper-K [5, 6], IceCube/PINGU/MICA [7–9], KM3NeT [10], and ANTARES [11], as well as recent proposals based on liquid Argon such as LAr TPC and GLACIER [12, 13], and liquid scintillator experiments like JUNO [14, 15]. In such experiments, boosted DM can scatter via the neutral-current-like process

$$\psi_B X \rightarrow \psi_B X^{(\prime)}, \quad (1.2)$$

similar to high energy neutrinos. This boosted DM phenomenon is generic in multi-component DM scenarios and in single-component DM models with non-minimal stabilization symmetries), where boosted DM can be produced in DM conversion $\psi_i \psi_j \rightarrow \psi_k \psi_\ell$ [4, 16, 17], semi-annihilation $\psi_i \psi_j \rightarrow \psi_k \phi$ (where ϕ is a non-DM state) [16, 18–21], $3 \rightarrow 2$ self-annihilation [22–24], or decay transition $\psi_i \rightarrow \psi_j + \phi$.

In order to be detectable, of course, boosted DM must have an appreciable cross section to scatter off SM targets. Based on eq. (1.1) alone and given our assumption that ψ_A is isolated from the SM, one might think that ψ_B could also have negligible SM interactions. In that case, however, the dark sector would generally have a very different temperature from the SM sector, with the temperature difference depending on details related to reheating, couplings to the inflaton, and entropy releases in the early universe [25–28]. So if we want to preserve the most attractive feature of the WIMP paradigm — namely, that the thermal relic abundance of ψ_A is determined by its annihilation cross section, insensitive to other details — then ψ_B must have efficient enough interactions with the SM to keep ψ_A in thermal equilibrium at least until $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$ freezes out. Such ψ_B -SM couplings then offer a hope for detecting the dark sector even if the major DM component ψ_A has no direct SM couplings.

As a simple proof of concept, we present a two-component DM model of the above type, with ψ_A/ψ_B now being specified as fermions. The dominant DM component ψ_A has no (tree-level) interactions with the SM, such that traditional DM searches are largely insensitive to it. In contrast, the subdominant DM component ψ_B has significant interactions with the SM via a dark photon γ' that is kinetically-mixed with the SM photon. The two processes related to the (in)direct detection of the ψ_A/ψ_B dark sector are illustrated in figure 1. In the early universe, the process on the left, due to a contact interaction between ψ_A and ψ_B , sets both the thermal relic abundance of ψ_A as well as the production rate of boosted ψ_B in the galactic halo today. The resulting boosted ψ_B population has large scattering cross sections off nuclei and electrons via dark photon exchange, shown on the right of figure 1.

¹To our knowledge, the first use of $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$ to set the relic abundance of ψ_A appears in the assisted freeze-out scenario [4]. As an interesting side note, we will find that assisted freeze-out of ψ_A can lead to a novel “balanced freeze-out” behavior for ψ_B . In appendix A, we show that the relic abundance can scale like $\Omega_B \propto 1/\sqrt{\sigma_B}$ (unlike $\Omega_B \propto 1/\sigma_B$ for standard freeze-out). In this paper, of course, we are more interested in the boosted ψ_B population, not the thermal relic ψ_B population.

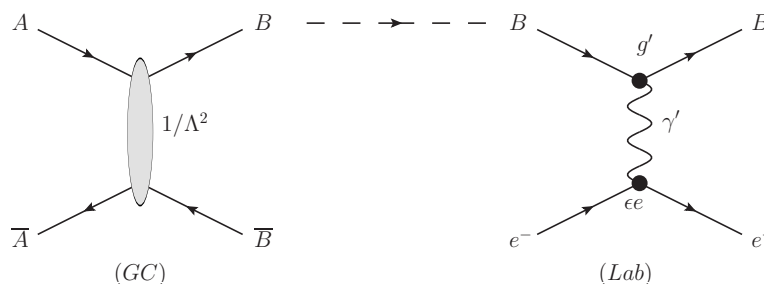


Figure 1. (Left) Production of boosted ψ_B particles through ψ_A annihilation in the galactic center: $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$. This process would be considered “indirect detection” of ψ_A . (Right) Scattering of ψ_B off terrestrial electron targets: $\psi_B e^- \rightarrow \psi_B e^-$. This process would be considered “direct detection” of ψ_B .

Assuming that ψ_B itself has a small thermal relic abundance (which is expected given a large SM scattering cross section), and is light enough to evade standard DM detection bounds, then (direct) detection of boosted ψ_B via (indirect) detection of ψ_A annihilation would offer the best non-gravitational probe of the dark sector.²

Beyond just the intrinsic novelty of the boosted DM signal, there are other reasons to take this kind of DM scenario seriously. First, having the dominant DM component ψ_A annihilate into light stable ψ_B particles (i.e. assisted freeze-out [4]) is a novel way to “seclude” DM from the SM while still maintaining the successes of the thermal freeze-out paradigm of WIMP-type DM.³ Such a feature enables this model to satisfy the increasingly severe constraints from DM detection experiments. A key lesson from secluded DM scenarios [30] is that it is often easier to detect the “friends” of DM (in this case ψ_B) rather than the dominant DM component itself [36]. Second, our study here can be seen as exploring the diversity of phenomenological possibilities present (in general) in multi-component DM scenarios. Non-minimal dark sectors are quite reasonable, especially considering the non-minimality of the SM (with protons and electrons stabilized by separate B - and L -number symmetries). Earlier work along these lines includes, for instance, the possibility of a mirror DM sector [26, 37–39]. Recently, multi-component DM scenarios have drawn rising interest motivated by anomalies in DM detection experiments [40–42] and possible new astrophysical phenomena such as a “dark disk” [43]. Boosted DM provides yet another example of how the expected kinematics, phenomenology, and search strategies for multi-component DM can be very different from single-component DM.

The outline of the rest of this paper is as follows. In section 2, we present the above model in more detail. In section 3, we describe the annihilation processes of both ψ_A and ψ_B , which sets their thermal relic abundances and the rate of boosted DM production today, and we discuss the detection mechanisms for boosted DM in section 4. We assess the discovery prospects at present and future experiments in section 5, where we find that Super-K should

²Because ψ_A has no direct coupling to the SM, the ψ_A solar capture rate is suppressed. By including a finite ψ_A -SM coupling, one could also imagine boosted DM coming from annihilation in the sun. The possibility of detecting fast-moving DM emerging from the sun has been studied previously in the context of induced nucleon decay [29], though not with the large boost factors we envision here which enable detection via Cherenkov radiation. Note, however, that ψ_B particles are likely to become trapped in the sun due to energy loss effects (see section 4.4), limiting solar capture as a viable signal channel.

³For variations such as annihilating to dark radiation or to dark states that decay back to the SM, see for instance refs. [30–35].

already be sensitive to boosted DM by looking for single-ring electron events from the galactic center (GC). We summarize the relevant constraints on this particular model in section 6, and we conclude in section 7 with a discussion of other DM scenarios with similar phenomenology. More details are relegated to the appendices.

2 Two component dark matter

Consider two species of fermion DM ψ_A and ψ_B with Dirac masses $m_A > m_B$, which interact via a contact operator⁴

$$\mathcal{L}_{\text{int}} = \frac{1}{\Lambda^2} \bar{\psi}_A \psi_B \bar{\psi}_B \psi_A. \quad (2.1)$$

This operator choice ensures an s -wave annihilation channel [44], $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$ as in figure 1, which is important for having a sizable production rate of boosted ψ_B today. A UV completion for such operator is shown in figure 11a in appendix B. Other Lorentz structures are equally plausible (as long as they lead to s -wave annihilation).

As an extreme limit, we assume that eq. (2.1) is the sole (tree-level) interaction for ψ_A at low energies and that ψ_A is the dominant DM component in the universe today. We assume that both ψ_A and ψ_B are exactly stable because of separate stabilizing symmetries (e.g. a $\mathbb{Z}_2 \times \mathbb{Z}_2$).

The subdominant species ψ_B is charged under a dark $U(1)'$ gauge group, with charge +1 for definiteness. This group is spontaneously broken, giving rise to a massive dark photon γ' with the assumed mass hierarchy

$$m_A > m_B > m_{\gamma'}. \quad (2.2)$$

We will take the gauge coupling g' of the dark $U(1)'$ to be sufficiently large (yet perturbative) such that the process $\psi_B \bar{\psi}_B \rightarrow \gamma' \gamma'$ efficiently depletes ψ_B and gives rise to a small thermal relic abundance (see eq. (3.5) below).

Via kinetic mixing with the SM photon [45–47] (strictly speaking, the hypercharge gauge boson),

$$\mathcal{L} \supset -\frac{\epsilon}{2} F'_{\mu\nu} F^{\mu\nu}, \quad (2.3)$$

γ' acquires ϵ -suppressed couplings to SM fields. In this way, we can get a potentially large cross section for ψ_B to scatter off terrestrial SM targets, in particular $\psi_B e^- \rightarrow \psi_B e^-$ from γ' exchange (with large g' and suitable ϵ) as in figure 1. In principle, we would need to account for the possibility of a dark Higgs boson H' in the spectrum, but for simplicity, we assume that such a state is irrelevant to the physics we consider here, perhaps due to a Stuckelberg mechanism for the $U(1)'$ [48, 49] or negligible couplings of H' to matter fields.

The parameter space of this model is defined by six parameters

$$\{m_A, m_B, m_{\gamma'}, \Lambda, g', \epsilon\}. \quad (2.4)$$

Throughout this paper, we will adjust Λ to yield the desired DM relic abundance of ψ_A , assuming that any DM asymmetry is negligible. Because the process $\psi_B e^- \rightarrow \psi_B e^-$ has

⁴Via a Fierz rearrangement, we can rewrite this operator as

$$-\frac{1}{4\Lambda^2} \left(\bar{\psi}_A \psi_A \bar{\psi}_B \psi_B + \bar{\psi}_A \gamma^\mu \psi_A \bar{\psi}_B \gamma_\mu \psi_B + \frac{1}{2} \bar{\psi}_A \Sigma^{\mu\nu} \psi_A \bar{\psi}_B \Sigma_{\mu\nu} \psi_B + \bar{\psi}_A \gamma^5 \psi_A \bar{\psi}_B \gamma^5 \psi_B - \bar{\psi}_A \gamma^\mu \gamma^5 \psi_A \bar{\psi}_B \gamma_\mu \gamma^5 \psi_B \right),$$

where $\Sigma^{\mu\nu} = \frac{i}{2} [\gamma^\mu, \gamma^\nu]$.

homogeneous scaling with g' and ϵ , the dominant phenomenology depends on just the three mass parameters: m_A , m_B , and $m_{\gamma'}$. To achieve a sufficiently large flux of boosted ψ_B particles, we need a large number density of ψ_A particles in the galactic halo. For this reason, we will focus on somewhat low mass thermal DM, with typical scales:

$$m_A \simeq \mathcal{O}(10 \text{ GeV}), \quad m_B \simeq \mathcal{O}(100 \text{ MeV}), \quad m_{\gamma'} \simeq \mathcal{O}(10 \text{ MeV}). \quad (2.5)$$

Constraints on this scenario from standard DM detection methods are summarized later in section 6. This includes direct detection and CMB constraints on the thermal relic ψ_B population. In addition, ψ_A can acquire couplings to γ' through a ψ_B -loop, thus yielding constraints from direct detection of ψ_A , and we introduce a simple UV completion for eq. (2.1) in appendix B which allows us to compute this effect without having to worry about UV divergences.

There are a variety of possible extensions and modifications to this simple scenario. One worth mentioning explicitly is that ψ_A and/or ψ_B could have small Majorana masses which lead to mass splittings within each multiplet (for ψ_B this would appear after $U(1)'$ breaking) [50, 51]. As discussed in refs. [52–54], both components in an inelastic DM multiplet can be cosmologically stable, such that the current day annihilation is not suppressed. These splittings, however, would typically soften the bounds on the non-relativistic component of ψ_A/ψ_B from conventional direct detection experiments, since the scattering would be inelastic (either endothermic or exothermic). This is one way to avoid the direct detection of bounds discussed in section 6.

3 Thermal relic abundances and present-day annihilation

To find the relic density of ψ_A/ψ_B , we need to write down their coupled Boltzmann equations. In appendix A, we provide details about this Boltzmann system (see also refs. [4, 55, 56]), as well as analytic estimates for the freeze-out temperature and relic abundance in certain limits. Here, we briefly summarize the essential results.

The annihilation channel $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$ not only determines the thermal freeze-out of the dominant DM component ψ_A but also sets the present-day production rate for boosted ψ_B particles in Milky Way. Considering just the operator from eq. (2.1), the thermally-averaged cross section in the s -wave limit is:

$$\langle \sigma_{A\bar{A} \rightarrow B\bar{B}v} \rangle_{v \rightarrow 0} = \frac{1}{8\pi\Lambda^4} (m_A + m_B)^2 \sqrt{1 - \frac{m_B^2}{m_A^2}}. \quad (3.1)$$

As discussed in appendix A, the Boltzmann equation for ψ_A effectively decouples from ψ_B when $\langle \sigma_{B\bar{B} \rightarrow \gamma'\gamma'v} \rangle \gg \langle \sigma_{A\bar{A} \rightarrow B\bar{B}v} \rangle$. In this limit, the relic density Ω_A takes the standard form expected of WIMP DM (assuming s -wave annihilation):

$$\Omega_A \simeq 0.2 \left(\frac{5 \times 10^{-26} \text{ cm}^3/\text{s}}{\langle \sigma_{A\bar{A} \rightarrow B\bar{B}v} \rangle} \right). \quad (3.2)$$

Notice that in order to get the observed DM relic abundance $\Omega_A \approx 0.2$, the thermal annihilation cross section is around twice the “standard” thermal cross section $3 \times 10^{-26} \text{ cm}^3/\text{s}$ where a Majorana fermion DM with $\simeq 100 \text{ GeV}$ mass is assumed. The slight discrepancy is because our Ω_A is the sum of the abundances of both Dirac particles ψ_A and $\bar{\psi}_A$, and the ψ_A we are interested in has lower mass $\lesssim 20 \text{ GeV}$ (see, e.g., ref. [57]).

In the limit that $m_B \ll m_A$, we have

$$\langle \sigma_{A\bar{A} \rightarrow B\bar{B}} v \rangle \approx 5 \times 10^{-26} \text{ cm}^3/\text{s} \left(\frac{m_A}{20 \text{ GeV}} \right)^2 \left(\frac{250 \text{ GeV}}{\Lambda} \right)^4. \quad (3.3)$$

Note that $m_A \ll \Lambda$ for our benchmark mass $m_A = 20 \text{ GeV}$, so it is consistent to treat the annihilation of ψ_A as coming just from the effective operator in eq. (2.1).

The thermal relic abundance of ψ_B is more subtle. In the absence of ψ_A , the relic abundance of ψ_B would be determined just by the annihilation process $\psi_B \bar{\psi}_B \rightarrow \gamma' \gamma'$, whose thermally-averaged cross section in the s -wave limit is

$$\langle \sigma_{B\bar{B} \rightarrow \gamma' \gamma'} v \rangle_{v \rightarrow 0} = \frac{g'^4}{2\pi} \frac{(m_B^2 - m_{\gamma'}^2)}{(m_{\gamma'}^2 - 2m_B^2)^2} \sqrt{1 - \frac{m_{\gamma'}^2}{m_B^2}}. \quad (3.4)$$

However, the process $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$ is still active even after ψ_A freezes out with a nearly constant ψ_A abundance well above its equilibrium value, which can have impact on the relic abundance of ψ_B . Let $x_{f,B} = m_B/T_{f,B}$, $T_{f,B}$ being the temperature at ψ_B freeze-out. As explained in appendix A, when $\frac{\sigma_B}{\sigma_A} \left(\frac{m_B}{m_A} \right)^2 \gg (x_f)^2$ (i.e. large g'), a good approximation to the relic abundance Ω_B is

$$\frac{\Omega_B}{\Omega_A} \simeq \frac{m_B}{m_A} \sqrt{\frac{\langle \sigma_{A\bar{A} \rightarrow B\bar{B}} v \rangle}{\langle \sigma_{B\bar{B} \rightarrow \gamma' \gamma'} v \rangle}}. \quad (3.5)$$

This $\Omega \propto 1/\sqrt{\sigma}$ behavior is very different from the usual DM abundance relation $\Omega \propto 1/\sigma$. It arises because in this limit, there is a balance between depletion from ψ_B annihilation and replenishment from $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$ conversion. To our knowledge, this “balanced freeze-out” behavior has not been discussed before in the DM literature.

In figure 2a, we show numerical results for Ω_B as a function of g' : for small g' , ψ_B freezes out in the standard way with $\Omega_B \propto 1/\sigma_B$, while for large g' , Ω_B exhibits the $1/\sqrt{\sigma_B}$ scaling from balanced freeze-out. Thus, as long as g' is sufficiently large, then ψ_B will be a subdominant DM component as desired. In figure 2b, we show the full solution to the coupled Boltzmann equations for ψ_A and ψ_B (see eq. (A.1)) for the following benchmark scenario:

$$m_A = 20 \text{ GeV}, \quad m_B = 200 \text{ MeV}, \quad m_{\gamma'} = 20 \text{ MeV}, \quad g' = 0.5, \quad \epsilon = 10^{-3}, \quad (3.6)$$

where we have adjusted $\Lambda = 250 \text{ GeV}$ to yield the cross section $\langle \sigma_{A\bar{A} \rightarrow B\bar{B}} v \rangle = 5 \times 10^{-26} \text{ cm}^3/\text{s}$ needed to achieve $\Omega_A \simeq \Omega_{\text{DM}} \approx 0.2$. For this benchmark, ψ_B has a much smaller abundance $\Omega_B \simeq 2.6 \times 10^{-6} \Omega_{\text{DM}}$. We have chosen the reference masses to be safe from existing constraints but visible with a reanalysis of existing Super-K data, and we have chosen the reference value of g' to be comparable to hypercharge in the SM. The values of $m_{\gamma'}$ and ϵ are also interesting for explaining the muon $g - 2$ anomaly [58, 59].

This model, though simple, exhibits a novel ψ_B freeze-out behavior, and the “balancing condition” behind eq. (3.5) may be interesting to study in other contexts. For much of parameter space of our interest in this paper, the $\Omega_B \propto 1/\sqrt{\sigma_B}$ scaling affects the CMB and direct detection constraints on ψ_B . As discussed in section 6, this scaling implies that the constraints from CMB heating on ψ_B annihilation are largely independent of g' . Similarly, unless there is some kind of inelastic splitting within the ψ_B multiplet, there is a firm direct detection bound on m_B that is also largely independent of g' . Note that the benchmark scenario in eq. (3.6) indeed satisfies these bounds (see the star in figure 9).

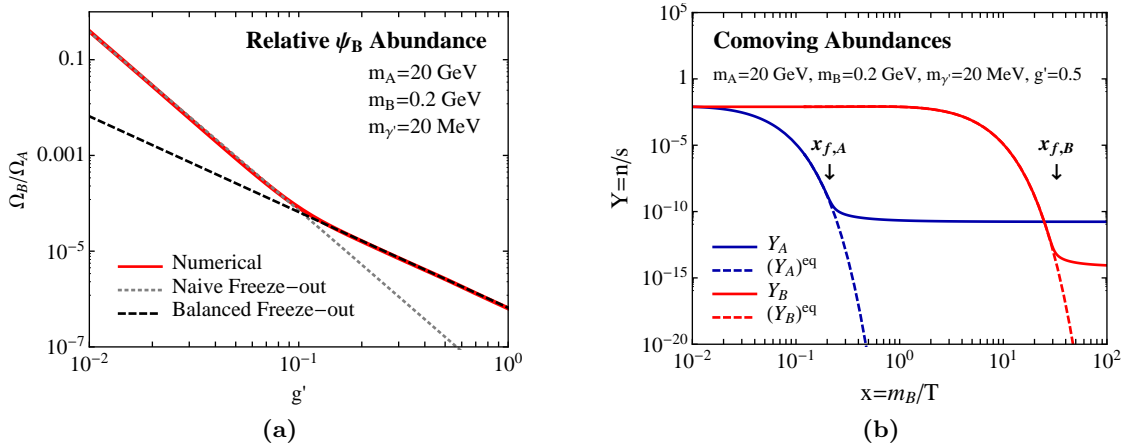


Figure 2. (a) Ratio of the abundances Ω_B/Ω_A as a function of g' , fixing $m_A = 20$ GeV, $m_B = 0.2$ GeV, and $m_{\gamma'} = 20$ MeV. The solid line is the numerical solution of the Boltzmann equation in eq. (A.1), the dotted line is the analytic estimate from assuming independent thermal freeze-out of ψ_A and ψ_B (naive freeze-out), and the dashed line is the analytic estimate from eq. (3.5) (balanced freeze-out). (b) Evolution of the co-moving abundances Y_A and Y_B as a function of $x = m_B/T$ for the benchmark in eq. (3.6). The solid lines show the actual densities per unit entropy, while the dashed lines are the equilibrium curves.

4 Detecting boosted dark matter

With ψ_A being the dominant DM species, the annihilation process $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$ is active in the galactic halo today, producing boosted ψ_B particles. To compute the flux of ψ_B incident on the earth, we can recycle the standard formulas from indirect detection of WIMP DM. Roughly speaking, the (in)direct detection of boosted ψ_B particles from ψ_A annihilation is analogous to the familiar process of indirect detection of neutrinos from WIMP annihilation. For this reason, the natural experiments to detect boosted DM are those designed to detect astrophysical neutrinos. As we will see, ψ_B typically needs to have stronger interactions with the SM than real neutrinos in order to give detectable signals in current/upcoming experiments.

We also want to comment that, due to the small mass and suppressed thermal abundance, the non-relativistic relic ψ_B particles can be difficult to detect through conventional direct and indirect DM searches, even with efficient interaction between ψ_B and SM states. (See section 6 for existing bounds on ψ_B .) Therefore, detecting boosted ψ_B particles may be the only smoking gun from this two-component ψ_A/ψ_B system.

4.1 Flux of boosted dark matter

The flux of ψ_B from the GC is

$$\frac{d\Phi_{GC}}{d\Omega dE_B} = \frac{1}{4} \frac{r_{Sun}}{4\pi} \left(\frac{\rho_{local}}{m_A} \right)^2 J \langle \sigma_{A\bar{A} \rightarrow B\bar{B}} v \rangle_{v \rightarrow 0} \frac{dN_B}{dE_B}, \quad (4.1)$$

where $r_{Sun} = 8.33$ kpc is the distance from the sun to the GC and $\rho_{local} = 0.3$ GeV/cm³ is the local DM density. Since the $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$ annihilation process yields two mono-energetic

boosted ψ_B particles with energy m_A , the differential energy spectrum is simply

$$\frac{dN_B}{dE_B} = 2\delta(E_B - m_A). \quad (4.2)$$

The quantity J is a halo-shape-dependent dimensionless integral over the line of sight,

$$J = \int_{\text{l.o.s}} \frac{ds}{r_{\text{Sun}}} \left(\frac{\rho(r(s, \theta))}{\rho_{\text{local}}} \right)^2, \quad (4.3)$$

where s is the line-of-sight distance to the earth, the coordinate $r(s, \theta) = (r_{\text{Sun}}^2 + s^2 - 2r_{\text{Sun}}s \cos \theta)^{1/2}$ is centered on the GC, and θ is the angle between the line-of-sight direction and the earth/GC axis. Assuming the NFW halo profile [60], we use the interpolation functions $J(\theta)$ provided in ref. [61] and integrate them over angular range of interest. In particular, when trying to mitigate neutrino backgrounds in section 4.3, we will require the $\psi_B e^- \rightarrow \psi_B e^-$ process to give final state electrons within a cone of angle θ_C from the GC.

To illustrate the scaling of the flux, we integrate over a 10° cone around the GC and obtain

$$\Phi_{\text{GC}}^{10^\circ} = 9.9 \times 10^{-8} \text{ cm}^{-2} \text{ s}^{-1} \left(\frac{\langle \sigma_{A\bar{A} \rightarrow B\bar{B}} v \rangle}{5 \times 10^{-26} \text{ cm}^3/\text{s}} \right) \left(\frac{20 \text{ GeV}}{m_A} \right)^2. \quad (4.4)$$

For completeness, the flux over the whole sky is:

$$\Phi_{\text{GC}}^{4\pi} = 4.0 \times 10^{-7} \text{ cm}^{-2} \text{ s}^{-1} \left(\frac{\langle \sigma_{A\bar{A} \rightarrow B\bar{B}} v \rangle}{5 \times 10^{-26} \text{ cm}^3/\text{s}} \right) \left(\frac{20 \text{ GeV}}{m_A} \right)^2. \quad (4.5)$$

These estimates are subject to uncertainties on the DM profile; for example, an Einasto profile would increase the flux by an $\mathcal{O}(1)$ factor [61].

Note that this GC flux estimate is the same as for any mono-energetic DM annihilation products.⁵ Therefore we can estimate the expected bound on the boosted DM-SM cross section by reinterpreting neutrino bounds on DM annihilation. Looking at ref. [62], the anticipated Super-K limit on 1–100 GeV DM annihilating in the Milky Way exclusively to monochromatic neutrinos is $10^{-21} - 10^{-22} \simeq \text{cm}^3/\text{sec}$. This is four to five orders of magnitude weaker than a typical thermal annihilation cross section ($\simeq 10^{-26} \text{ cm}^3/\text{sec}$). Assuming thermal relic ψ_A DM exclusively annihilates to boosted ψ_B particles, we can estimate the bound on the ψ_B -SM cross section by scaling down the charged current neutrino scattering cross section (10^{-38} cm^2 , see eq. (4.19)) by the corresponding factor. This gives an estimated bound of

$$\sigma_{B\text{SM} \rightarrow B\text{SM}} \lesssim 10^{-33} - 10^{-34} \text{ cm}^2, \quad (4.6)$$

which is consistent with the cross section derived later in eq. (4.16) for a benchmark model that is on the edge of detectability.⁶

⁵Up to factors of 2 if the particles considered are Majorana or Dirac, and the number of particles created in the final state.

⁶Our numbers are less consistent with Super-K bounds shown in conference proceedings in ref. [63], which are two orders of magnitude more constraining than expected from ref. [64]. However, the details of the Super-K analysis are not available for direct comparison.

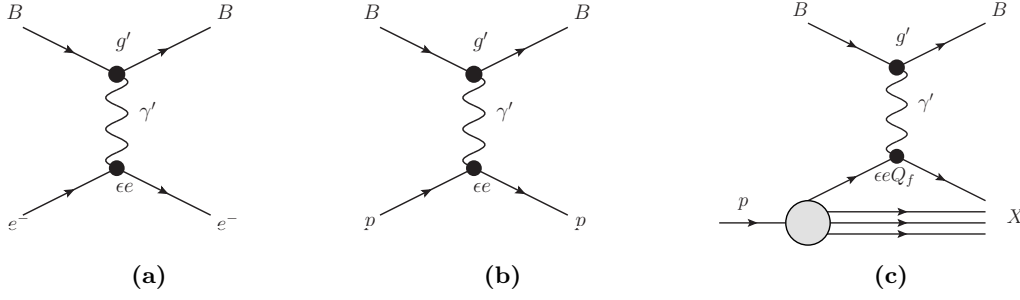


Figure 3. Detection channels for boosted ψ_B in neutrino experiments. (a) Elastic scattering on electrons. (b) Elastic scattering on protons (or nuclei). (c) Deep inelastic scattering on protons (or nuclei). For Cherenkov experiments, we find that the most promising channel is electron scattering.

4.2 Detection of boosted dark matter

The flux of boosted ψ_B particles estimated from eq. (4.4) is rather small.⁷ Therefore, in order to detect boosted ψ_B , one needs a large volume, small background detector sensitive to the (quasi-)elastic scattering process

$$\psi_B X \rightarrow \psi_B X', \quad (4.7)$$

where X and X' are SM states (possibly the same). Because the γ' is kinetically-mixed with the photon, ψ_B can scatter off any SM state X with electromagnetic couplings via t -channel exchange of γ' .⁸ A large scattering cross section favors light $m_{\gamma'}$, large ϵ , and large g' ; the values of $m_{\gamma'} \gtrsim 10$ MeV and $\epsilon \sim 10^{-3}$ in the benchmark in eq. (3.6) are (marginally) consistent with current limits on dark photons [65].

Existing neutrino detectors such as Super-K, IceCube, and their upgrades can be employed to detect boosted DM via eq. (4.7). The strategy is to detect Cherenkov light from the final state charged particles, so the energy of outgoing X' must be above the Cherenkov threshold. In terms of a Lorentz factor, the threshold is

$$\text{Water: } \gamma_{\text{Cherenkov}} = 1.51, \quad \text{Ice: } \gamma_{\text{Cherenkov}} = 1.55, \quad (4.8)$$

where there is typically a stricter analysis threshold E^{thresh} on X' as well, depending on experimental specifics. Furthermore, one needs to distinguish ψ_B scattering from the large background of neutrino scattering events, which we discuss more in section 4.3.

As shown in figure 3, there are three detection channels for boosted ψ_B at a neutrino detector: elastic scattering off electrons, elastic scattering off protons (or nuclei), and deep inelastic scattering (DIS) off protons (or nuclei). As discussed in more detail in appendix C, although the total ψ_B scattering cross section off protons and nuclei can be sizable, the de-

⁷For comparison, the flux of non-relativistic relic ψ_B particles incident on earth is approximately

$$\Phi_{\text{local}} \simeq \frac{\rho_{\text{local}} v_0}{m_B} \frac{\Omega_B}{\Omega_{\text{DM}}} = 2.25 \times 10^3 \text{ cm}^{-2} \text{ s}^{-1} \left(\frac{200 \text{ MeV}}{m_B} \right) \left(\frac{\Omega_B}{10^{-5}} \right).$$

where $v_0 \simeq 220$ km/sec.

⁸There are also subdominant scatterings from weak charges as well.

tectable signal strengths in these channels are suppressed relative to scattering off electrons.⁹ Thus, we focus on the elastic scattering off electrons

$$\psi_B e^- \rightarrow \psi_B e^- \quad (4.9)$$

as the most promising detection channel, though we present signal studies for the other channels in appendix C. At detectors like Super-K, the signal would appear as single-ring electron events coming from the direction of the GC.

We start by discussing the kinematics of scattering off electrons (the same logic would hold for protons). In the rest frame of an electron target with mass m_e , the momenta of incoming and outgoing particles are:

$$\begin{aligned} \text{Incident } \psi_B: p_1 &= (E_B, \vec{p}), & \text{Scattered } \psi_B: p_3 &= (E'_B, \vec{p}'), \\ \text{Initial } e: p_2 &= (m_e, 0), & \text{Scattered } e: p_4 &= (E_e, \vec{q}). \end{aligned} \quad (4.10)$$

For ψ_B coming from nearly-at-rest ψ_A annihilation,

$$E_B = m_A. \quad (4.11)$$

The maximum scattered electron energy occurs when \vec{p} and \vec{p}' are parallel:

$$E_e^{\max} = m_e \frac{(E_B + m_e)^2 + E_B^2 - m_B^2}{(E_B + m_e)^2 - E_B^2 + m_B^2}. \quad (4.12)$$

The minimum detectable energy is set by the analysis threshold (assumed to be above the Cherenkov threshold),

$$E_e^{\min} = E_e^{\text{thresh}} > \gamma_{\text{Cherenkov}} m_e. \quad (4.13)$$

Of course, to have any viable phase space, $E_e^{\max} \geq E_e^{\min}$. From eqs. (4.12) and (4.13), we can also express the viable kinematic region in terms of boost factors γ_e and γ_B (taking $m_A \gg m_B \gg m_e$):

$$\gamma_e^{\min} = \frac{E_e^{\text{thresh}}}{m_e}, \quad \gamma_e^{\max} = 2\gamma_B^2 - 1, \quad \gamma_B = \frac{E_B}{m_B} = \frac{m_A}{m_B}. \quad (4.14)$$

The differential cross section for ψ_B elastic scattering off electrons is:

$$\frac{d\sigma_{Be^- \rightarrow Be^-}}{dt} = \frac{1}{8\pi} \frac{(\epsilon e g')^2}{(t - m_{\gamma'}^2)^2} \frac{8E_B^2 m_e^2 + t(t + 2s)}{\lambda(s, m_e^2, m_B^2)}, \quad (4.15)$$

where $\lambda(x, y, z) = x^2 + y^2 + z^2 - 2xy - 2xz - 2yz$, $s = m_B^2 + m_e^2 + 2E_B m_e$, $t = q^2 = 2m_e(m_e - E_e)$, and one should make the replacement $E_B = m_A$ for our scenario. To give a numerical sense of

⁹The reason is that ψ_B scattering proceeds via t -channel exchange of the light mediator γ' , so the differential cross section peaks at small momentum transfers, while achieving Cherenkov radiation (or DIS scattering) requires large momentum transfers. For elastic scattering, this logic favors electrons over protons in two different ways: an $\mathcal{O}(1 \text{ GeV})$ ψ_B can more effectively transfer momentum to electrons compared to protons because of the heavier proton mass, and protons require a larger absolute momentum transfer to get above the Cherenkov threshold. Compounding these issues, protons have an additional form-factor suppression, identifying proton tracks is more challenging than identifying electron tracks [66–68], and the angular resolution for protons is worse than for electrons at these low energies [68]. We note that liquid Argon detectors are able to reconstruct hadronic final states using ionization instead of Cherenkov light, so they may be able to explore the (quasi-)elastic proton channels down to lower energies, even with smaller detector volumes [12, 13].

the Cherenkov electron signal cross section, integrating eq. (4.15) over the allowed kinematic region for the benchmark in eq. (3.6) yields

$$\sigma_{Be^- \rightarrow Be^-} = 1.2 \times 10^{-33} \text{ cm}^2 \left(\frac{\epsilon}{10^{-3}} \right)^2 \left(\frac{g'}{0.5} \right)^2 \left(\frac{20 \text{ MeV}}{m_{\gamma'}} \right)^2, \quad (4.16)$$

for an experimental threshold of $E_e^{\text{thresh}} = 100 \text{ MeV}$. The approximate scaling is derived in the limit $m_e E_e^{\text{thresh}} \ll m_{\gamma'}^2 \ll m_e E_e^{\text{max}}$, where the dependance on E_B , m_B , and E_e^{thresh} is weaker than polynomial, which holds in the vicinity of the benchmark point but not in general. For completeness, the full cross section for ψ_B -electron scattering without an energy threshold cut is

$$\sigma_{Be^- \rightarrow Be^-}^{\text{tot}} = 1.47 \times 10^{-33} \text{ cm}^2 \left(\frac{\epsilon}{10^{-3}} \right)^2 \left(\frac{g'}{0.5} \right)^2 \left(\frac{20 \text{ MeV}}{m_{\gamma'}} \right)^2. \quad (4.17)$$

Since this cross section is rather high, we have to account for the possibility that ψ_B particles might be stopped as they pass through the earth. In section 4.4, we find that the attenuation of ψ_B particles is mild, so we will treat the earth as transparent to ψ_B particles in our analysis.

In figure 4a, we show the normalized, logarithmic electron spectrum for different benchmarks, including the one from eq. (3.6). The electron energy E_e peaks at relatively low values due to the t -channel γ' , as discussed further in footnote 9. We note that the position of the peak depends both on m_B and $m_{\gamma'}$, though the dominant effect of $m_{\gamma'}$ is to change the overall signal cross section (not visible in this normalized plot). The angular distribution of the recoil electron is shown in figure 4b. The signal is very forward peaked, as expected from $m_B \gg m_e$. This is advantageous when looking for boosted DM from the GC, since the recoil electrons' direction is tightly correlated to that of the ψ_B 's.

In figure 5, we compare the energy profile of the signal to the observed background electron events at SK-I [69]. Using the benchmark model in eq. (3.6), we plot the (logarithmic) energy spectrum of the yearly signal event yield within a cone of 10° around the GC:

$$\frac{dN}{d \log E_e} = E_e \frac{dN}{dE_e} = \Delta T N_{\text{target}} \Phi_{\text{GC}}^{10^\circ} E_e \frac{d\sigma_{Be^- \rightarrow Be^-}}{dE_e}, \quad (4.18)$$

where $\Phi_{\text{GC}}^{10^\circ}$ is defined in eq. (4.4), ΔT is a year, and N_{target} is the number of targets (electrons) at Super-K. Anticipating the analysis of section 5.2, we also plot a more realistic spectrum obtained by convolving the signal scattering cross section $\psi_B e^- \rightarrow \psi_B e^-$ with the shape of the DM halo. This convolved spectrum matches nicely to the naive spectrum from eq. (4.18), as expected given the peaked nature of the angular spectrum in figure 4b, with signal losses at low energies arising because less energetic electrons can be more easily deflected outside the search cone. Once the background from ref. [69] is scaled by the appropriate factor of $(\pi(10^\circ)^2)/(4\pi) \approx 8 \times 10^{-3}$, the signal for this benchmark is visible above the background, though the peak location is (accidentally) at a similar location.

4.3 Backgrounds to boosted dark matter

The major background to the boosted DM signal comes from atmospheric neutrinos, which are produced through interactions of cosmic rays with protons and nuclei in the earth's atmosphere. Atmospheric neutrino energy spectrum peaks around 1 GeV and follows a power law

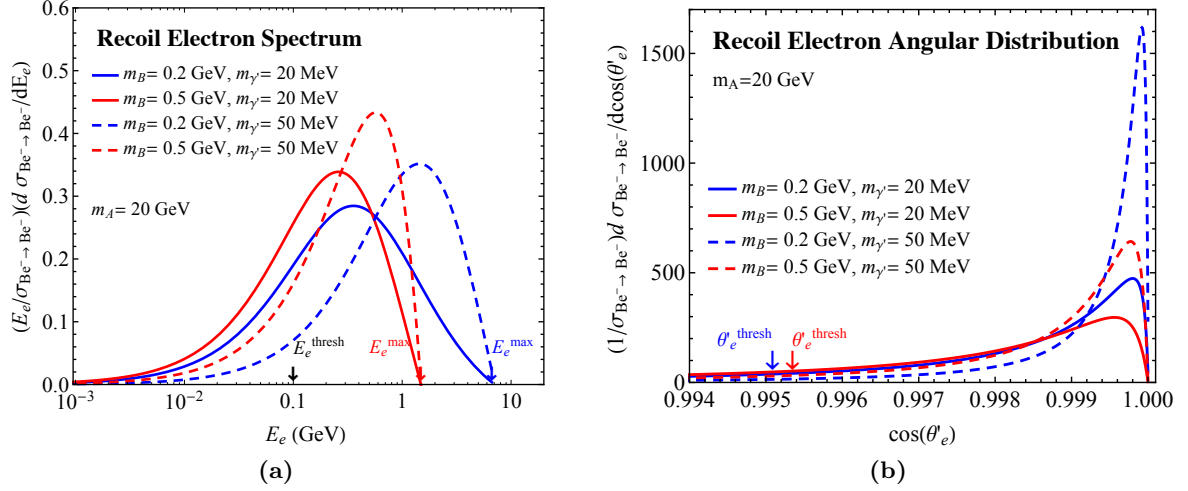


Figure 4. (a) Normalized recoil electron spectrum for different benchmark scenarios. Also indicated is the maximum scattered electron energy, given by eq. (4.12) as well as the experimental threshold of Super-K in the Sub-GeV category (See eq. (5.2)). (b) Recoil electron angular distribution for the same signal benchmarks, assuming a ψ_B particle coming directly from the GC. The cutoff angle θ_e^{thresh} is obtained by substituting the 100 MeV energy threshold into eq. (4.22).

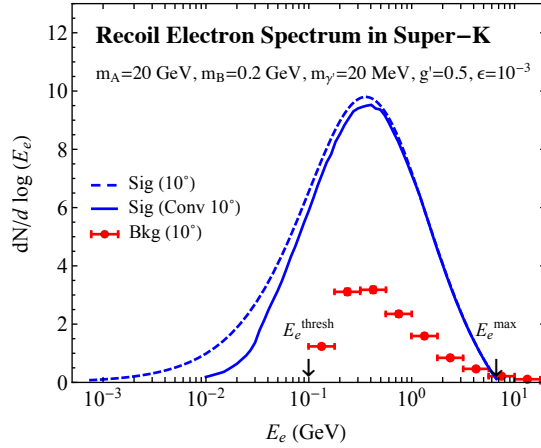


Figure 5. Energy spectrum of signal and background events, normalized to the expected event yield over one year. The blue dashed line corresponds to the naive formula in eq. (4.18) for the number of signal events in a 10° search cone. The solid blue line is spectrum obtained from the convolution in eq. (5.5). The background spectrum of CC ν_e and $\bar{\nu}_e$ events comes from Super-K [69], scaled by a factor $\pi(10^\circ)^2/(4\pi)$ to account for the nominal 10° search cone. Note that data is available only for $E_e > 100$ MeV, which is the same experimental threshold given in eq. (5.2). Also indicated is the maximum scattered electron energy, given by eq. (4.12).

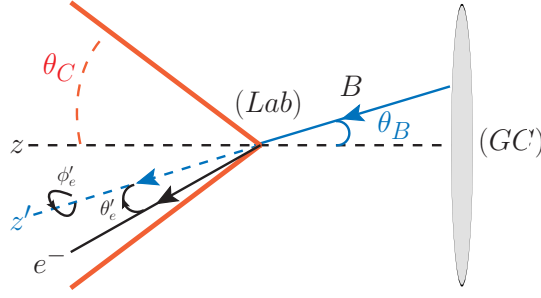


Figure 6. Angles involved in boosted DM detection. When a ψ_B particle arrives at an angle θ_B from the GC, it scatters to produce an electron at angle θ_e with respect to z (θ'_e and ϕ'_e with respect to z'). To better isolate the signal from the uniform atmospheric neutrino background, we impose a search cone of half-angle θ_C .

$E^{-2.7}$ at higher energies [70]. The scattering process $\psi_B e^- \rightarrow \psi_B e^-$ with an energetic outgoing electron faces a large background from charged-current (CC) electron-neutrino scattering $\nu_e n \rightarrow e^- p$ when the outgoing proton is not detected, as well as $\bar{\nu}_e p \rightarrow e^+ n$ since Cherenkov-based experiments cannot easily distinguish electrons from positrons. For $\mathcal{O}(1 \text{ GeV})$ neutrinos, the CC cross sections are [71]

$$\sigma_{\text{CC}}^{\nu_e} \approx 0.8 \times 10^{-38} \text{ cm}^2 \left(\frac{E_\nu}{\text{GeV}} \right), \quad (4.19)$$

$$\sigma_{\text{CC}}^{\bar{\nu}_e} \approx 0.3 \times 10^{-38} \text{ cm}^2 \left(\frac{E_\nu}{\text{GeV}} \right). \quad (4.20)$$

While smaller than the expected signal cross section in eq. (4.16), the atmospheric neutrino flux is much higher than the boosted ψ_B flux. The neutral current process $\nu_e e^- \rightarrow \nu_e e^-$ can also mimic the signal but it is subdominant to the CC interaction due to m_e/m_p suppression [71].

There are a number of discriminants one could use to (statistically) separate our signal from the neutrino background.

- *Angular restriction:* boosted ψ_B particles have a definite direction because they come from the GC. In galactic coordinates, the atmospheric neutrino background has no preferred direction. Therefore, one can impose that the detected electron falls within a cone of half-opening angle θ_C with respect to the GC. As shown in figure 6, there are two relevant axes to consider: the z -axis connecting the earth to the GC and the z' -axis in the direction that the ψ_B travels along. Through $\psi_B e^- \rightarrow \psi_B e^-$ scattering, a ψ_B particle coming from an angle θ_B ($\theta'_B = 0$) will yield a final state electron with scattering angle (θ'_e, ϕ'_e) , with

$$\cos \theta_e = \cos \theta_B \cos \theta'_e - \sin \theta_B \sin \phi'_e \sin \theta'_e, \quad (4.21)$$

$$\cos \theta'_e = \frac{(m_A + m_e)}{\sqrt{m_A^2 - m_B^2}} \frac{\sqrt{E_e - m_e}}{\sqrt{E_e + m_e}}, \quad (4.22)$$

and ϕ'_e uniformly distributes between 0 and 2π . To the extent that the electron energy is large and $m_A \gg m_B \gg m_e$, we have $\cos \theta_e \approx \cos \theta_B$. As we will see in section 5.4, the optimum angle θ_C to maximize the signal acceptance while minimizing the neutrino background is around 10° , assuming perfect angular resolution.

- *Energy restriction:* boosted ψ_B particles have a mono-energetic spectrum ($E_B = m_A$), compared to the continuous atmospheric neutrino energy spectrum. This implies a correlation between the measured E_e and $\cos \theta_e$. That said, we suspect that the typical angular resolution of neutrino experiments is not fine enough to make use of this feature. In fact, more important than energy resolution is to have a low energy threshold, since as shown in figure 4, the signal cross section peaks at small E_e .
- *Absence of muon excess:* the process $\psi_B e^- \rightarrow \psi_B e^-$ does not have a corresponding muon signature, whereas the neutrino CC process $\nu_e n \rightarrow e^- p$ is always accompanied by $\nu_\mu n \rightarrow \mu^- p$. So an electron excess from boosted DM should *not* have a correlated excess in muon events. One can also require fully-contained events to reduce the cosmic ray muon background.
- *Anti-neutrino discrimination:* the anti-neutrino background $\bar{\nu}_e p \rightarrow e^+ n$ is in principle reducible since it involves a final state positron instead of an electron. Super-K cannot perfectly distinguish ν_e from $\bar{\nu}_e$ events, but as of the SK-IV analyses, they have used likelihood methods to separate these two categories by studying the number of decay electrons in each process. The purity of the ν_e sample is 62.8% and that of $\bar{\nu}_e$ is 36.7% [72, 73]. We have not used this feature in our current analysis. Adding gadolinium to Super-K would help tagging neutrons from the $\bar{\nu}_e$ CC process, and thus might improve the purity of these samples [74].
- *Multi-ring veto:* the process $\psi_B e^- \rightarrow \psi_B e^-$ leads to electron-like single-ring events only, without correlated multi-ring events. In contrast, neutrino CC process $\nu_e n \rightarrow e^- p$ can lead to multi-ring events when the outgoing proton energy is above Cherenkov threshold [68], or when the scattering is inelastic so that other charged hadronic states such as π^\pm are produced. Argon-based detectors could improve the background discrimination since they can detect the hadronic final states from neutrino scattering better than water-based experiments. We note that for some extreme parameters (increasing g' or ϵ), it is possible for ψ_B to interact twice (or more) within the detector, also creating a potential multi-ring signal (or a lightly-ionizing track in a scintillator detector). That said, for such high cross sections, the signal would be heavily attenuated while traversing the earth (see section 4.4). Another potential disadvantage of a multi-ring veto is that we might miss out on interesting signals such as ψ_B scattering accompanied by γ' bremsstrahlung ($\psi_B e^- \rightarrow \psi_B e^- \gamma' \rightarrow \psi_B e^- e^+ e^-$).
- *Solar neutrino/muon veto:* solar neutrinos dominate the background under around 20 MeV [75], though one can of course preform an analysis in solar coordinates and exclude events from the sun. In addition, there is a background from muons that do not Cherenkov radiate but decay to neutrinos in the detector volume; these are relevant in the range of 30–50 MeV and can be mitigated through fiducial volume cuts [76]. To avoid both of these complications, we will use a cut of $E_e > 100$ MeV in our analysis below. Of course the threshold of 100 MeV in Super-K could be brought down as low as 50 MeV (where solar neutrino backgrounds start to dominate). The potential advantage of looking in the 50–100 MeV range is that the backgrounds from atmospheric neutrinos are lower. The main disadvantage is the degradation of the angular resolution of the detector [77].

The first two points favor detectors with excellent angular resolution and low energy thresholds on the outgoing electron. The next three points mean that one could distinguish the

boosted DM signal from neutrinos coming from WIMP DM annihilation in the GC; boosted DM only gives a single-ring electron signal whereas neutrinos from WIMPs would give equal contributions to an electron and muon signal, both single- and multi-ring events, and equal contributions to a neutrino and anti-neutrino signal. The last point suggests the interesting possibility of looking for boosted DM from the sun due to DM solar capture, though in the particular model we study in this paper, the solar capture rate is too small to be visible, and any boosted DM particles from the sun would face considerable solar attenuation (see section 4.4 below). The above criteria can be thought of as a general algorithm for background rejection, while specifics can be tailored to a particular experiment. For instance, “multi-ring veto” does not apply to PINGU where Cherenkov rings cannot be reconstructed and all non- μ -like events are classified as “cascade events”.

4.4 Impact of earth attenuation

As seen in eq. (4.17), the signal cross section $\sigma_{Be^- \rightarrow Be^-}$ is relatively high, so as they cross the earth, the ψ_B particles might get deflected and lose energy. This is a particularly important effect for Northern Hemisphere experiments like Super-K, where a typical ψ_B would have to traverse through $\sim 10^5$ km (75% of the earth’s diameter). The dominant cause for energy loss is (minimum) ionization of atoms. While not relevant for detection, the main source of angular deflection is scattering off nuclei. In the following, we base our discussion on the standard analysis of particle propagation through matter as developed in the PDG [78].

First we estimate the ψ_B ’s energy loss. Just as for a heavy charged particle traversing the earth (see, e.g., ref. [79]), the main energy loss mechanism is through ionization. For $\beta\gamma$ factors of 10–100, a muon loses ≈ 1 GeV of energy per meter of rock [78]. A muon scatters off nuclei via a t -channel γ exchange, while a ψ_B scatters off nuclei via the exchange of a γ' . We can approximate the length required for a ψ_B to lose 1 GeV by scaling the couplings and the propagator of the ψ_B - e^- scattering process to those of the μ - e^- scatterings:

$$L_B \approx L_\mu \epsilon^2 \frac{g'^2}{e^2} \left(\frac{t}{t - m_{\gamma'}^2} \right)^2, \quad (4.23)$$

where $t = 2m_e(m_e - E_e) \approx -10^{-2}$ GeV² for our key benchmark in eq. (3.6). In this case, ψ_B loses ≈ 1 GeV per 9×10^8 cm, giving rise to a total expected loss of 1 GeV per trip through the earth ($R_\oplus = 6.4 \times 10^8$ cm). Since 1 GeV of energy loss is never more than $\simeq 10\%$ of the ψ_B ’s initial energy in the parameter space of interest, we will assume the earth is transparent to ψ_B ’s for the rest of the analysis. Accounting for the energy loss is approximately equivalent to shifting the plots in figures 7 and 9 by the energy loss on the m_A axis. The parameter space of small m_A is the most affected, but that region is already constrained by CMB bounds as shown in section 6.

Turning to the angular distribution, the dominant source of deflection is from elastic scattering off of nuclei. Note that ψ_B - e^- scattering processes lead to very small angles of deflection because of the mass hierarchy $m_B \gg m_e$; indeed for the key benchmark in eq. (3.6), the maximum possible deflection per scatter is 0.14° . In contrast, Coulomb-like scattering of ψ_B particles off of nuclei can give rise to a more substantial deflection (including full reversal). The mean-square change in angle per collision process is

$$\begin{aligned} \langle \theta_B^2 \rangle &\simeq 2 - 2 \langle \cos(\theta_B) \rangle, \\ \langle \cos(\theta_B) \rangle &= \frac{1}{\sigma_{BN \rightarrow BN}} \int_0^{E_N^{\max}} \cos(\theta_B(E_N)) \frac{d\sigma_{BN \rightarrow BN}}{dE_N} dE_N \simeq \cos(0.2^\circ), \end{aligned} \quad (4.24)$$

where $\sigma_{BN \rightarrow BN}$ is the scattering cross section of ψ_B 's off a nucleus N (see eq. (C.6)), and E_N^{\max} is defined similarly to eq. (4.12). In the last step of eq. (4.24), we have inserted the benchmark value from eq. (3.6). Treating the deflection of ψ_B particles as a random walk through the earth, the total deflection is

$$\langle \theta_{\text{total}}^2 \rangle^{1/2} = \langle \theta_B^2 \rangle^{1/2} \sqrt{\frac{\ell_{BN \rightarrow BN}}{R_{\oplus}}}, \quad (4.25)$$

where the quantity under the square root is the number of steps (interactions). The mean free path to interact with a nucleus of charge number Z and atomic number A is

$$\ell_{BN \rightarrow BN} = \frac{1}{n \sigma_{BN \rightarrow BN} \left(\frac{Z}{26}\right)^2 \left(\frac{55.84}{A}\right)} = 1.5 \times 10^7 \text{ cm}, \quad (4.26)$$

where n is the number density. Under the conservative assumption that the earth is entirely made of iron (the benchmark A and Z values above), and taking the mass density of earth to be $\rho = 5.5 \text{ g/cm}^3$, the number of scatters is ≈ 64 for the benchmark in eq. (3.6), giving a total deflection of:

$$\langle \theta_{\text{total}}^2 \rangle^{1/2} = 1.6^\circ. \quad (4.27)$$

We checked that for different values of m_A , m_B , and $m_{\gamma'}$, the total deflection does not vary much compared to eq. (4.27). Since this deflection is small compared with the search cone of 10° that is used in section 5.2, we neglect the angular deflection of ψ_B 's in our analysis.

Interestingly, if a signal of boosted DM is found, we could potentially use the earth attenuation to our advantage by correlating candidate signal events with the position of the GC with respect to the experiment. Indeed, with high enough statistics, the effect of earth shadowing would give rise to time-dependent rates, energies, and angles for ψ_B scattering. As mentioned in footnote 2, solar attenuation would have an adverse effect on possible boosted DM signals from the sun. Since the radius of the sun is 100 times larger than that of the earth, the ψ_B particles would lose a factor of 100 more energy, so we would need $m_A \gtrsim \mathcal{O}(100 \text{ GeV}-1 \text{ TeV})$ for ψ_B particle to escape the sun. Alternatively, for a smaller scattering cross section of ψ_B particles with the SM, the sun might then be a viable source of signal [80].

5 Detection prospects for present and future experiments

We now assess the detection prospects for boosted DM at present and future detectors for neutrinos and/or proton decay. In table 1, we summarize the (approximate) capacities/sensitivities of some of the representative relevant experiments, given in terms of the detector volume V_{exp} , electron energy threshold E_e^{thresh} , and angular resolution θ_e^{res} . From this table, we can already anticipate which experiments are going to be best suited for boosted DM detection.

Due to the relatively small flux of boosted DM, a larger volume detector, such as IceCube, KM3NeT, or ANTARES would be favored in order to catch more signal events. However, the energy threshold for the original IceCube are much too high for our purposes (and similarly for KM3NeT and ANTARES), since the energy transferred to the outgoing electron is suppressed due to the t -channel γ' (see figure 4). Even the $\simeq 1 \text{ GeV}$ threshold of PINGU is not ideal, though it will have some sensitivity.

Experiment	Volume (MTon)	E_e^{thresh} (GeV)	θ_e^{res} (degree)	Refs.
Super-K	2.24×10^{-2}	0.01	3°	[69]
Hyper-K	0.56	0.01	3°	[81]
IceCube	10^3	100	30°	[82, 83]
PINGU	0.5	1	23° (at GeV scale)	[8]
MICA	5	0.01	30° (at 10 MeV scale)	[9, 84]

Table 1. List of experiments studied in this paper, their angular resolutions θ_e^{res} on the Cherenkov-emitted electron direction, and the typical minimum energy threshold E_e^{thresh} of the detected electron. We note here that IceCube has too high of an energy threshold for our analysis, but we are interested in its future low-energy extensions such as PINGU and MICA. For PINGU, we have scaled the nominal volume (1 MTon) down by a factor of 2 to estimate particle identification efficiency. The MICA values are speculative at present, since there is not yet a technical design report.

So although Super-K/Hyper-K have smaller detector volumes, their low energy threshold is better matched to the boosted DM signal. In addition, Super-K/Hyper-K have excellent angular resolution,¹⁰ which makes it possible to optimize the θ_C search cone criteria. Ultimately, MICA would offer better coverage in the energy range of interest. It is also worth mentioning that the proposed experiments for proton decay based on large scale liquid Argon detectors [12, 13] can also be sensitive to boosted DM due to their low thresholds and large volume. As mentioned in footnote 9, liquid Argon detectors may also have sensitivity to the proton scattering channel as well.

In the following subsections, we discuss event selection, signal/background rates, and expected signal significance in the above experiments. For signal-only studies of the subdominant channels involving ψ_B scattering off protons/nuclei, see appendix C.

5.1 Event selection

As discussed in sections 4.2 and 4.3, the leading boosted DM signal comes from elastic scattering off electrons ($\psi_B e^- \rightarrow \psi_B e^-$) and the leading background is from atmospheric neutrinos (mostly $\nu_e n \rightarrow e^- p$). In principle, one could use the full multivariate information about the kinematics of the outgoing electron to separate signal and background. In order to keep the analysis simple, we will do a cut-and-count study to estimate the sensitivity.

To isolate events coming from the GC, we will use the search cone θ_C described in figure 6. The dominant background from CC ν_e scattering of atmospheric neutrinos is assumed to be uniform across the sky, so the background in a search cone of half-angle θ_C scales proportional to θ_C^2 . Of course, one cannot take θ_C to be too small, otherwise the signal acceptance degrades. To optimize for the signal significance in section 5.4, we will convolve the angular dependence of halo J -factor and the angular dependence of the $\psi_B e^- \rightarrow \psi_B e^-$ cross section to figure out the optimum θ_C . Anticipating that result, we will find

$$\theta_C = \max\{10^\circ, \theta_e^{\text{res}}\}, \quad (5.1)$$

where 10° applies to the high resolution experiments (Super-K/Hyper-K), and the other experiments are limited by their angular resolutions.

¹⁰More accurately, the angular resolution of fully contained 1 ring Multi-GeV electrons is 1.5° while that of fully contained 1 ring Sub-GeV electrons is less than 3.3° as shown in ref. [85].

From eqs. (4.12) and (4.13), we have minimum and maximum electron energies E_e^{\min} and E_e^{\max} for the signal. Ideally, one would adjust the energy selection for a given value of m_A and m_B , and try to push the analysis threshold E_e^{\min} to be as low as possible. To be conservative, we will take the standard Super-K event categories for fully-contained single-ring electron events (see e.g. ref. [73])

$$\text{Sub-GeV: } \{100 \text{ MeV}, 1.33 \text{ GeV}\}, \quad (5.2)$$

$$\text{Multi-GeV: } \{1.33 \text{ GeV}, 100 \text{ GeV}\}, \quad (5.3)$$

without attempting to do finer energy binning. For Super-K, Hyper-K, and MICA, we will use both categories as separate event selections; for the Sub-GeV category we will choose only zero-decay events. PINGU has a higher energy threshold and cannot reconstruct Cherenkov rings nor efficiently separate μ -like and e -like events near threshold, so we will only use the Multi-GeV category, while also adding in backgrounds from multi-ring events and μ -like events; we will also scale the PINGU effective volume down by a factor of 2 to account for an estimated reconstruction efficiency of $\sim 50\%$ [8]. Note that the 100 MeV lower bound of the Sub-GeV category is above the nominal 10 MeV threshold of Super-K, so there is room for improved signal acceptance. Similarly, when the 1.33 GeV upper bound of the Sub-GeV category is above E_e^{\max} , then we are overestimating the background.

5.2 Signal rates

Imposing the θ_C and energy range requirements, the number of signal electron events is:

$$\begin{aligned} N_{\text{signal}}^{\theta_C} &= \Delta T N_{\text{target}} (\Phi_{\text{GC}} \otimes \sigma_{Be^- \rightarrow Be^-})|_{\theta_C} \\ &= \frac{1}{2} \Delta T \frac{10 \rho_{\text{Water/Ice}} V_{\text{exp}} r_{\text{Sun}}}{m_{\text{H}_2\text{O}}} \frac{1}{4\pi} \left(\frac{\rho_{\text{local}}}{m_A} \right)^2 \langle \sigma_{A\bar{A} \rightarrow B\bar{B}v} \rangle_{v \rightarrow 0} \\ &\quad \times \int_0^{2\pi} \frac{d\phi'_e}{2\pi} \int_{\theta'_{\min}}^{\theta'_{\max}} d\theta'_e \sin \theta'_e \frac{d\sigma_{Be^- \rightarrow Be^-}}{d \cos \theta'_e} \int_0^{\pi/2} d\theta_B \sin \theta_B 2\pi J(\theta_B) \Theta(\theta_C - \theta_e), \end{aligned} \quad (5.4)$$

where ΔT is the time duration of the observation, N_{target} is the number of target electrons, Φ_{GC} is the DM flux from the GC, and $\sigma_{Be^- \rightarrow Be^-}$ is the ψ_B -electron scattering cross section (which depends on the energy integration range in eq. (4.15)). The factor of 10 in the second line is the number of electrons per molecule of water. The DM flux and scattering cross section have to be convolved in order to isolate events that pass the θ_C requirement, and the angles in the last line are the same as in figure 6 with θ_e given in eq. (4.21). The integration limits $\theta'_e \in \{\theta'_{\min}, \theta'_{\max}\}$ are given by eq. (4.22) by requiring $E_e \in \{E_e^{\max}, E_e^{\min}\}$ (note the reversal of the limits, and that $\theta'_{\min} = 0$ if eq. (4.12) is more restrictive than the energy categories above).

To get a sense of the expected signal rate, we consider the number of signal events for $\theta_C = 10^\circ$ in the combined categories:

$$\frac{N_{\text{signal}}^{10^\circ}}{\Delta T} = 25.1 \text{ year}^{-1} \left(\frac{\langle \sigma_{A\bar{A} \rightarrow B\bar{B}v} \rangle}{5 \times 10^{-26} \text{ cm}^3/\text{s}} \right) \left(\frac{20 \text{ GeV}}{m_A} \right)^2 \left(\frac{\sigma_{Be^- \rightarrow Be^-}}{1.2 \times 10^{-33} \text{ cm}^2} \right) \left(\frac{V_{\text{exp}}}{22.4 \times 10^3 \text{ m}^3} \right), \quad (5.5)$$

broken down by 21.1/year for Sub-GeV and 4.0/year for Multi-GeV, and the reference cross sections are based on the benchmark in eq. (3.6). In our analysis below, we will always assume that $\langle \sigma_{A\bar{A} \rightarrow B\bar{B}v} \rangle$ takes on the thermal relic reference value.

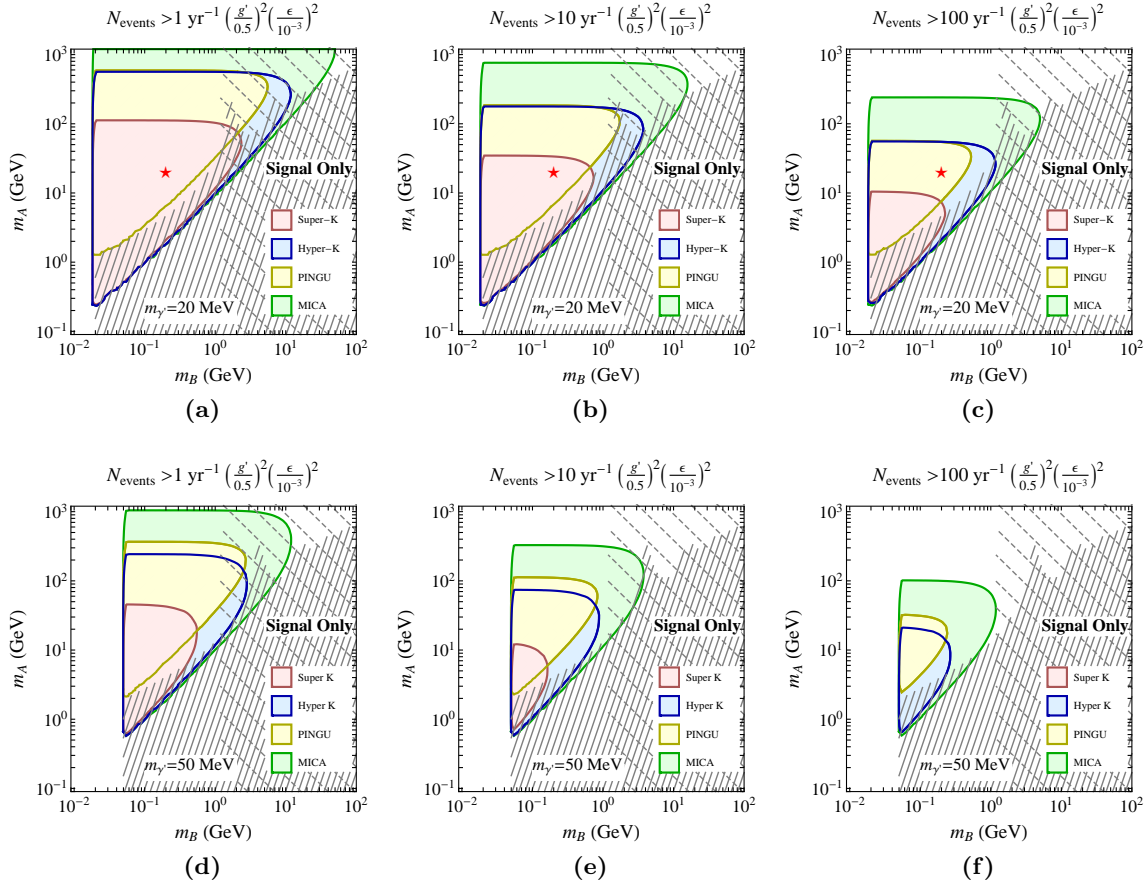


Figure 7. Number of $\psi_B e^- \rightarrow \psi_B e^-$ signal events in Super-K, Hyper-K, PINGU, and MICA in the m_A/m_B plane, for $m_{\gamma'} = 20$ MeV (top) and $m_{\gamma'} = 50$ MeV (bottom). The indicated regions are for 1 (left), 10 (center), 100 (right) detected events in a one year period, normalized to the couplings $\epsilon = 10^{-3}$ and $g' = 0.5$. We have imposed the angular criteria of $\theta_C = 10^\circ$ and the electron energy range of $\{100 \text{ MeV}, 100 \text{ GeV}\}$ ($\{1.33 \text{ GeV}, 100 \text{ GeV}\}$ for PINGU). Also shown are model-dependent constraints on the relic ψ_B population from section 6: the solid gray lines are from CMB heating (shown only for $g' = 0.5$), and the dashed gray lines are from DAMIC direct detection (which are independent of g' , but can be eliminated by adding an inelastic splitting). The red star indicates the benchmark in eq. (3.6).

Because $\sigma_{Be^- \rightarrow Be^-}$ scales homogeneously with g' and ϵ , the number of signal events does as well, so the only non-trivial dependence is on the mass parameters m_A , m_B , and $m_{\gamma'}$. In figure 7, we set two benchmark values $m_{\gamma'} = 20$ MeV and $m_{\gamma'} = 50$ MeV, and show what part of the $m_A - m_B$ parameter space yields

$$\frac{N_{\text{signal}}^{10^\circ}}{\text{year}} = x \left(\frac{g'}{0.5} \right)^2 \left(\frac{\epsilon}{10^{-3}} \right)^2, \quad (5.6)$$

for $x = 1, 10, 100$. These reference values for $m_{\gamma'}$ have been chosen such that the t -channel scattering processes are not overly suppressed by the dark photon mass, and the reference ϵ is close to the maximum allowed by dark photon constraints. In the triangular regions in figure 7, the top edge is set by m_A which controls the DM number density (and therefore the

annihilation rate), the left edge is set by the requirement that $m_B > m_{\gamma'}$, and the diagonal edge is set by the electron energy threshold.

In these figures, we have included model-dependent constraints from CMB heating and direct detection, discussed in the later section 6.¹¹ It is worth emphasizing that both of these constraints are due to the thermal relic ψ_B population, and are independent of the boosted DM phenomenon. Indeed, as discussed at the end of section 2, we could give ψ_B a small Majorana mass splitting, which would eliminate the bound from (elastic) direct direction experiments while not affecting very much the kinematics of boosted $\psi_B e^- \rightarrow \psi_B e^-$ detection. The CMB constraints are more robust since they mainly depend on ψ_B being in thermal contact with the SM via $\psi_B \bar{\psi}_B \rightarrow \gamma' \gamma'$, though the CMB constraints could potentially softened if γ' somehow decays to neutrinos (or to non-SM states).

5.3 Background rates

The atmospheric neutrino backgrounds have been measured by Super-K over a 10.7 year period, during runs SK-I (1489 days), SK-II (798 days), SK-III (518 days) and SK-IV (1096 days), and the final results are summarized in ref. [73]. In the Sub-GeV category, a total of 7755 fully-contained single-ring zero-decay electron events were seen the 100 MeV to 1.33 GeV energy range, giving a yearly background rate of

$$\text{Sub-GeV: } \frac{N_{\text{bkgd}}^{\text{all sky}}}{\Delta T} = 726 \text{ year}^{-1} \left(\frac{V_{\text{exp}}}{22.4 \times 10^3 \text{ m}^3} \right). \quad (5.7)$$

In the Multi-GeV category, 2105 fully-contained single-ring electron events were seen in the 1.33 GeV to 100 GeV energy range [73, 87], yielding

$$\text{Multi-GeV: } \frac{N_{\text{bkgd}}^{\text{all sky}}}{\Delta T} = 197 \text{ year}^{-1} \left(\frac{V_{\text{exp}}}{22.4 \times 10^3 \text{ m}^3} \right). \quad (5.8)$$

To estimate the background for PINGU (which lacks the ability to reconstruct Cherenkov rings), we add in multi-ring and μ -like events in the Multi-GeV category, changing 197 year^{-1} to 634 year^{-1} , which then has to be scaled by the effective PINGU detector volume.

For the boosted DM search, the background is reduced by considering only events where the electron lies in the search cone θ_C . We assume a uniform background distribution from the entire sky, so the background within a patch in the sky of angle θ_C is:

$$N_{\text{bkgd}}^{\theta_C} = \frac{1 - \cos \theta_C}{2} N_{\text{bkgd}}^{\text{all sky}}, \quad (5.9)$$

For $\theta_C = 10^\circ$ relevant for Super-K, we have

$$\text{Sub-GeV: } \frac{N_{\text{bkgd}}^{10^\circ}}{\Delta T} = 5.5 \text{ year}^{-1}. \quad (5.10)$$

$$\text{Multi-GeV: } \frac{N_{\text{bkgd}}^{10^\circ}}{\Delta T} = 0.35 \text{ year}^{-1}. \quad (5.11)$$

Ideally, we would use the full energy dependence of the background in order to optimize the signal/background separation, but given the rather low background rate, we will make the conservative choice to consider the whole Sub-GeV energy range.

¹¹The bump around $m_B = 10 \text{ GeV}$ in the CMB heating bound is due to a Sommerfeld resonance.

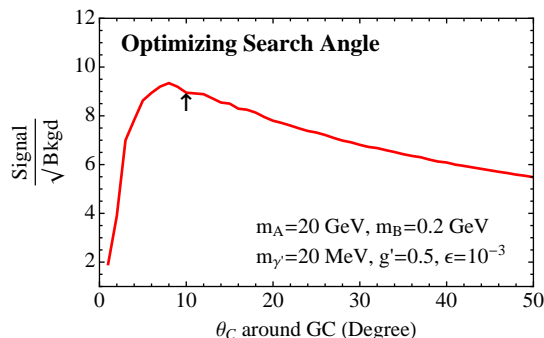


Figure 8. Yearly signal significance in the Sub-GeV category for our benchmark in eq. (3.6) as a function of the search cone angle θ_C . The peak around 10° is seen for other parameter choices as well.

Since one can estimate the background by looking at a side-band away from the GC, the background uncertainties in a θ_C cone should be dominated by Poisson fluctuations. For the all sky background, we note that Super-K saw a $\simeq 10\%$ mismatch between the measured atmospheric background and the Monte Carlo estimate [73, 87], so there is in fact a bit of room beyond Poisson fluctuations to accommodate a boosted DM signal in the current Super-K data.¹²

In order to have a fair comparison of the sensitivities at different experiments, for Hyper-K and MICA, we use the same event selection assuming the same exposure time, based on the available Super-K ~ 10 year data set, and simply scale up the background rate proportional to the detector volume V_{exp} (and adjust θ_C for MICA). As already mentioned, since PINGU has a higher energy threshold and an inability to reconstruct Cherenkov rings, we rescale the full Multi-GeV category (single-ring supplemented by the multi-ring and μ -like events).

5.4 Estimated experiment reach

Given the signal and background rates above, we can find the optimal search cone θ_C to maximize the significance

$$\text{Sig}^{\theta_C} \equiv \frac{N_{\text{signal}}^{\theta_C}}{\sqrt{N_{\text{bkgd}}^{\theta_C}}}. \quad (5.12)$$

In figure 8, we plot the significance as a function of search angle for our benchmark model in eq. (3.6); we checked that other parameter choices show similar behavior. We see that the significance peaks at around 10° , and falls off somewhat slowly after that. For Super-K/Hyper-K with 3° resolution, we can effectively ignore experimental resolution effects and take θ_C at the optimal value. For PINGU and MICA, we approximate the effect of the experimental resolution by taking $\theta_C = \theta_e^{\text{res}}$; a more sophisticated treatment would be to apply Gaussian smearing to the electrons. This is the logic behind eq. (5.1) above.

In figures 9 and 10, we show the 2σ sensitively possible with the 10.7 years of Super-K data, using the optimal $\theta_C = 10^\circ$ selection criteria, as well as the estimated reach for Hyper K, PINGU, and MICA for the same period of time. We treat the Sub-GeV and

¹²Associated with the published search for DM from the GC via upward going muons [88], there is also unpublished electron data from SK-I, -II, and -III [89, 90]. For $\cos \theta > 0.8$ ($\theta_C \simeq 37^\circ$), around 600 Sub-GeV fully-contained single-ring zero-decay electron events were observed in a 7.7 year period. This number has subsequently been updated to around 850 events in the full 10.7 year data set [91]. In principle, these could be used to set a stronger bound than we show in this paper, since no statistically significant excess is seen.

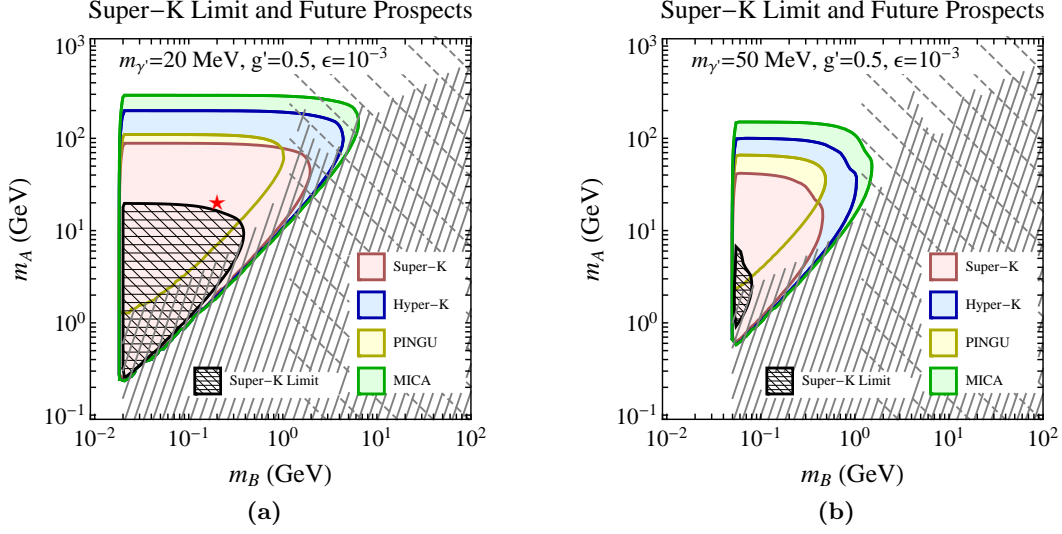


Figure 9. Signal significance at Super-K, Hyper K, PINGU and MICA on the m_A/m_B plane, for $m_{\gamma'} = 20$ MeV (left) and $m_{\gamma'} = 50$ MeV (right), fixing $\epsilon = 10^{-3}$ and $g' = 0.5$. Shown are the 2σ reaches with 10 years of data, taking $\theta_C = 10^\circ$ and adding the significances of the $E_e \in \{100 \text{ MeV}, 1.33 \text{ GeV}\}$ and $E_e \in \{1.33 \text{ GeV}, 100 \text{ GeV}\}$ categories in quadrature (only the latter for PINGU). Also shown is the current 2σ exclusion using all-sky data from Super-K, where we assume a 10% uncertainty on the background. The grey model-dependent limits are the same as in figure 7: the solid gray lines are constraints on ψ_B from CMB heating and the dashed gray lines are from DAMIC. The red star is the benchmark from eq. (3.6).

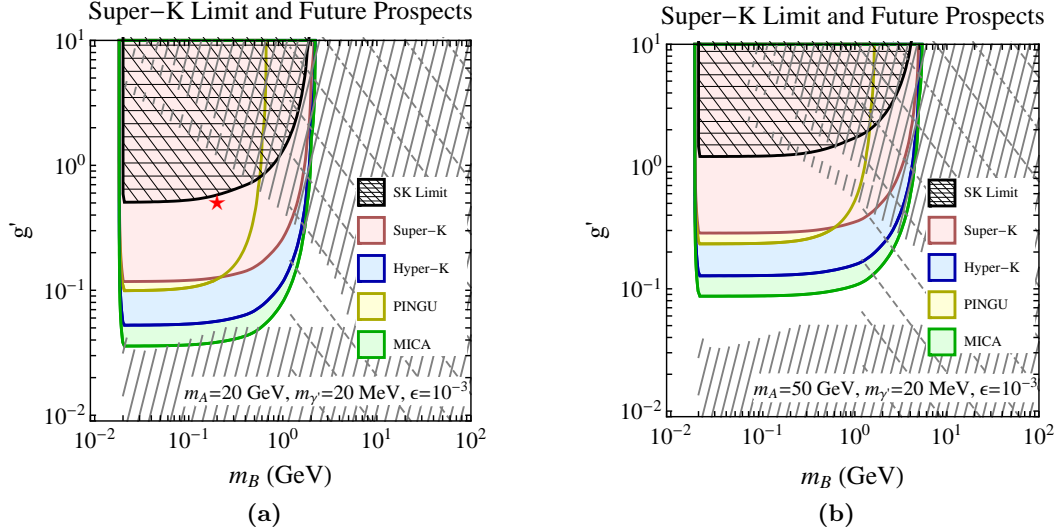


Figure 10. Same as figure 9, but on the g'/m_B plane, for $m_A = 20$ GeV (left) and $m_A = 50$ GeV (right), fixing $\epsilon = 10^{-3}$ and $m_{\gamma'} = 20$ MeV. The spikes in the CMB heating bounds (solid gray lines) are from Sommerfeld resonances.

Multi-GeV categories separately and report the overall significance as the quadrature sum of the significances obtained from the two categories. We also show the current bounds from Super-K that one can place without the θ_C selection (i.e. using the all-sky background), taking $\delta N_{\text{bkgd}}/N_{\text{bkgd}} = 10\%$ to account for systematic uncertainties in the all-sky background. Here, we are only allowing for the two energy categories, and further improvements are possible if one adjusts the energy range as a function of m_A and m_B .

Taken together, these experiments have substantial reach for boosted DM. The prospects for Super-K to find single-ring electron events from the GC are particularly promising, given that the data (with angular information) is already available [92] and one simply needs to change from lab-coordinates to galactic coordinates (as in refs. [89, 90]).

6 Summary of existing constraints

Apart from the measured neutrino fluxes discussed in section 5.3, we know of no model-independent constraints on the boosted DM phenomenon. There are, however, constraints on the particular model described here, and we summarize those constraints in this section. The most relevant bounds are due mainly to the relic ψ_B population left over from thermal freeze-out, which leads to bounds from “Direct detection of non-relativistic ψ_B ” and “CMB constraints on ψ_B annihilation” described below and seen in figures 9 and 10.

- *Limits on the dark photon γ' .* As discussed earlier, dark photon searches have set limits of $m_{\gamma'} \gtrsim \mathcal{O}(10 \text{ MeV})$ and $\epsilon \lesssim 10^{-3}$, assuming the dominant decay mode is $\gamma' \rightarrow e^+e^-$ [65]. For $m_{\gamma'} < \mathcal{O}(100 \text{ MeV})$, beam dump experiments place a bound of roughly $\epsilon \gtrsim 10^{-5}$ [93]. We have used $m_{\gamma'} = 20 \text{ MeV}$ and $\epsilon = 10^{-3}$ as a benchmark in this paper, which yields a detectable boosted DM signal while satisfying the current dark photon bounds. Our benchmark is also within the region of interest for explaining the muon $g - 2$ anomaly [58, 59].
- *Direct detection of non-relativistic ψ_A .* Thermal relic ψ_A particles are subject to constraints from conventional DM direct detection experiments (e.g. XENON, LUX, and CDMS) via their scattering off nuclei. As discussed in more detail in appendix B, the constraints on ψ_A are rather weak since ψ_A has no tree-level interactions with the SM. That said, ψ_A can scatter off nuclear targets via a ψ_B -loop. Since we have approximated the $\psi_A \bar{\psi}_A \psi_B \bar{\psi}_B$ interaction as a contact operator, this loop process is model-dependent. In figure 11a and eq. (B.1), we give an example UV completion involving an extra scalar ϕ that allows us to estimate the ψ_A -nucleon scattering cross section. Due to the loop factor and the mass suppression from $m_\phi \gg m_A$, the limits on ψ_A are safe for most values of the parameter space, as shown in figure 11b. As already mentioned, one could introduce inelastic splitting within the ψ_A multiplet to further soften direct detection constraints [52–54].
- *Direct detection of non-relativistic ψ_B .* Despite the small relic abundance of ψ_B , it has a large ψ_B -nucleon scattering cross section, as calculated in appendix C.1.

$$\sigma_{Bp \rightarrow Bp} = 4.9 \times 10^{-31} \text{ cm}^2 \left(\frac{\epsilon}{10^{-3}} \right)^2 \left(\frac{g'}{0.5} \right)^2 \left(\frac{20 \text{ MeV}}{m_{\gamma'}} \right)^4 \left(\frac{m_B}{200 \text{ MeV}} \right)^2, \quad (6.1)$$

where the scaling assumes $m_B \ll m_p$. Thus, direct detection experiments essentially rule out any elastic ψ_B -nucleon scattering above the detector threshold. Of course, in

the parameter space of our interest, the ψ_B mass is $\leq \mathcal{O}(1 \text{ GeV})$, which is close to or below the threshold of LUX [94] and the low CDMS threshold analysis [95], and the most constraining limits come from CDMSlite [96] and DAMIC [97]. Because of this, light ψ_B particles can evade existing direct detection bounds.

In figure 10, we demonstrate the constraints on the (g', m_B) plane from the DAMIC experiment (which has a lower threshold than CDMSlite), using the effective nuclear cross section

$$\sigma_{Bp \rightarrow Bp}^{\text{eff}} = \frac{\Omega_B}{\Omega_{\text{DM}}} \sigma_{Bp \rightarrow Bp}. \quad (6.2)$$

Essentially, the allowed parameter space is independent of g' and $(m_{\gamma'})^{-4}$, since the expected $\psi_{Bp} \rightarrow \psi_{Bp}$ cross section is so large that any events above the energy threshold of the experiment would be seen. There is also the fact that when g' is $\mathcal{O}(10^{-2})$ and higher, the abundance scales as g'^{-2} (see eq. (3.5)), which cancels with the g'^2 scaling of $\sigma_{Bp \rightarrow Bp}$, yielding a g' -independent bound.¹³ Of course, as with ψ_A , the direct detection bound on ψ_B could be alleviated by introducing inelastic splittings.

It has been recently pointed out that sub-GeV DM might be better constrained by scattering off electrons rather than off nuclei [98], as in recent XENON10 bounds [99]. In our case, these bounds are subsumed by CMB heating bounds discussed below. Note that for $\psi_B e^- \rightarrow \psi_B e^-$, the conventional direct detection process and the boosted DM detection process have very different kinematics, so one should not be surprised that the XENON10 bounds do not influence the boosted DM signal regions.

- *Indirect detection of non-relativistic ψ_B .* The annihilation process $\psi_B \bar{\psi}_B \rightarrow \gamma' \gamma'$ and the subsequent γ' decay to two e^+ / e^- pairs gives rise to a potential indirect detection signal in the positron and diffuse γ -ray channels. The recent constraint on DM annihilation in positron channel from AMS-02 is demonstrated in refs. [100, 101], where the bound is strongest for 2-body final state and weaker when there are more particles in the final state like in our case. The suppressed relic abundance of Ω_B relative to Ω_{DM} helps relieve the constraints on our model. In addition, at the sub-GeV mass which we are interested in, the background uncertainty of the above indirect detection limit is large due to solar modulation. The CMB considerations below give stronger constraints for the parameter range of our interest. The diffuse γ -ray signal from e.g. inverse Compton scattering of e^\pm produced from ψ_B annihilation has a smaller cross section and also faces large background uncertainty in the sub-GeV region. In fact, the γ -ray search for DM at Fermi, for instance, has a lower energy cutoff at $\sim 4 \text{ GeV}$ [102]. Indirect detection signals from ψ_A annihilation have to go through higher order or loop processes, and are much suppressed.
- *CMB constraints on ψ_B annihilation.* With a light mass of $m_B \lesssim \mathcal{O}(1 \text{ GeV})$, thermal ψ_B annihilation in the early universe may be subject to bounds from CMB heating [103]. The CMB constrains the total power injected by DM into ionization, heating, and excitations. For the dominant DM component with relic density $\Omega_{\text{DM}} \approx 0.2$, the bound is directly imposed on the quantity:

$$p_{\text{ann,DM}} = f_{\text{eff}} \frac{\langle \sigma v \rangle}{M_\chi}, \quad (6.3)$$

¹³The $\psi_{Bp} \rightarrow \psi_{Bp}$ cross section scales as $(m_{\gamma'})^{-4}$, but we have checked that these bounds do not soften until $m_{\gamma'}$ is higher than $\mathcal{O}(1 \text{ GeV})$, which is not the regime we are studying in this paper.

where f_{eff} is the fraction of the annihilation power that goes into ionization, which depends on the annihilation channel and its energy scale. Though ψ_B is a small fraction of total DM, it does annihilate into γ' which subsequently decays via $\gamma' \rightarrow e^+e^-$. Therefore, the CMB spectrum constrains

$$p_{\text{ann},\psi_B} = f_{\text{eff}} \frac{\langle \sigma_{B\bar{B} \rightarrow \gamma' \gamma' v} \rangle}{m_B} \left(\frac{\Omega_B}{\Omega_{\text{DM}}} \right)^2 \simeq f_{\text{eff}} \langle \sigma_{A\bar{A} \rightarrow B\bar{B}} \rangle \frac{m_B}{m_A^2}, \quad (6.4)$$

where the last relation is obtained using eq. (3.5) for Ω_B/Ω_A , which is valid for large values of g' (typically for $g' \gtrsim 0.1$) as explained in the appendix A. These limits are illustrated in figure 10 for $f_{\text{eff}} = 1$, which is a conservative assumption. Due to the presence of a light γ' , there can be an extra Sommerfeld enhancement factor to the $\langle \sigma_{B\bar{B} \rightarrow \gamma' \gamma' v} \rangle$ in eq. (6.4). For the parameter space we consider, we expect that this enhancement saturates at CMB time, which leads to an extra factor of [104]

$$S = \frac{\pi}{\epsilon_v} \frac{\sinh \frac{12\epsilon_v}{\pi\epsilon_\phi}}{\cosh \frac{12\epsilon_v}{\pi\epsilon_\phi} - \cos \left(2\pi \sqrt{\frac{6}{\pi^2\epsilon_\phi} - \left(\frac{6}{\pi^2}\right)^2 \frac{\epsilon_v^2}{\epsilon_\phi^2}} \right)}, \quad \epsilon_v = \frac{4\pi v}{g'^2}, \quad \epsilon_\phi = \frac{4\pi m_{\gamma'}}{g'^2 m_B}. \quad (6.5)$$

This enhancement contributes at low velocities, so we do not expect it to change the picture at freeze out, but it would be relevant in the CMB era where $v \approx 10^{-3}$. For our current parameter space, $S \approx 1$ until high values of $g' = 1$ where it becomes $\mathcal{O}(10)$. We incorporate the enhancement in the calculation of our CMB limits, as can be seen from the resonance peaks in figure 10.

- *BBN constraints on ψ_B annihilation.* The energy injection from ψ_B annihilation in the early universe can also alter standard BBN predictions [105, 106]. The constraints from hadronic final states are the most stringent, comparable to or even somewhat stronger at $\mathcal{O}(1 \text{ GeV})$ than those from the CMB heating as discussed above [105]. However, as we focus on $m_{\gamma'}$ of $\mathcal{O}(10 \text{ MeV})$, the production of hadronic final states (n, p, π) from the leading annihilation channel $\psi_B \bar{\psi}_B \rightarrow \gamma' \gamma'$ followed by γ' decay are not kinematically possible. The subleading channel $\psi_B \bar{\psi}_B \rightarrow q \bar{q}$ is ϵ^2 suppressed. Thus, the major energy injection to BBN is mostly electromagnetic from $\gamma' \rightarrow e^+e^-$, and the related constraint in this case are much weaker than the CMB bound we have considered above [105].
- *Dark matter searches at colliders.* By crossing the Feynman diagrams in figure 3, we see that ψ_B can be produced at colliders such as LEP, Tevatron, and the LHC. If ψ_B were to interact with SM electrons or quarks via a heavy mediator, then collider searches would provide a stronger bound than direct detection at these low DM masses. However, this complementarity is lost when the interaction is due to a light mediator [107–109], which applies to our case where ψ_B interacts with SM states via an $\mathcal{O}(10 \text{ MeV})$ dark photon. In addition, compared to the irreducible main background from electroweak processes, e.g. $e^+e^- \rightarrow Z^{(*)} \rightarrow \nu \bar{\nu}$, the production cross section of ψ_B is suppressed by $\epsilon^2 \lesssim 10^{-6}$, so the collider constraints on our model are rather weak.

7 Conclusions and other possibilities

In this paper, we presented a novel DM scenario which incorporates the successful paradigm of WIMP thermal freeze-out, yet evades stringent constraints from direct and indirect detection experiments, and predicts a novel signal involving boosted DM. The example model

features two DM components, ψ_A and ψ_B . The heavier particle ψ_A (which is the dominant DM component) experiences assisted freeze-out [4] by annihilating into the lighter particle ψ_B (which is the subdominant DM component). The whole dark sector is kept in thermal contact with the SM in the early universe via kinetic-mixing of a dark photon with the SM photon. Only ψ_B couples directly to the dark photon (and hence to the SM), so the dominant DM component ψ_A can largely evade current DM detection bounds. If such a scenario were realized in nature, then the leading non-gravitational signal of DM would come from annihilating ψ_A particles in the galactic halo producing boosted ψ_B particles that could be detected on earth via neutral-current-like scattering via the dark photon. In large volume neutrino or proton-decay detectors, the smoking gun for this scenario would be an electron signal pointing toward the GC, with no corresponding excess in the muon channel. Liquid argon detectors could potentially detect boosted DM through (quasi-)elastic proton scattering, as well as improve the rejection of the dominant neutrino CC backgrounds by vetoing on hadronic activity. Future experiments that use LArTPC technology for tracing the particle paths [110, 111] will provide both directionality and better background discrimination.

This phenomenon of boosted DM is generic in scenarios with multiple DM components. In fact, models with a single component DM could also potentially give rise to the same signature. If the stabilization symmetry is \mathbb{Z}_3 , then the semi-annihilation process $\psi\psi \rightarrow \bar{\psi}\phi$ (where ϕ is a non-DM state) is allowed [112–114]. For $m_\phi = 0$, the outgoing $\bar{\psi}$ would have energy $E_\psi = (5/4)m_\psi$. In the limit $m_\psi \gg m_e$, $\gamma_\psi = 1.25$ implies a maximum $\gamma_e^{\max} = 2\gamma_\psi^2 - 1 = 2.125$, which is above the Cherenkov threshold in water (and ice). Of course, the \mathbb{Z}_3 symmetry is not consistent with ψ being charged under a $U(1)'$, so additional model building would be necessary to get a sufficiently large scattering with the SM. But this example shows why non-minimal dark sectors tend to have some production cross section for boosted DM.

It is intriguing to consider other scenarios where DM mostly annihilates to other stable states in the dark sector. For example, if both ψ_A and ψ_B are charged under the $U(1)'$ and the mass hierarchy is

$$m_A > m_{\gamma'} > m_B, \quad (7.1)$$

then the annihilation $\psi_A\bar{\psi}_A \rightarrow \gamma'\gamma'$ would be followed by the decay $\gamma' \rightarrow \psi_B\bar{\psi}_B$, and the boosted ψ_B particles could again be detected via t -channel γ' exchange with the SM. Of course, now ψ_A itself has tree-level γ' exchange diagrams with the SM, but if ψ_A has a Majorana mass splitting (allowing it to evade direct detection bounds), boosted DM would again be the dominant mode for DM discovery.¹⁴

The above scenario is particularly interesting in light of the gamma ray excess recently seen in the GC [116]. In the context of DM, this signal could be explained via cascade decays $\psi_A\bar{\psi}_A \rightarrow \gamma'\gamma'$ followed by $\gamma' \rightarrow \text{SMSM}$ [117–120]. Boosted DM could be produced in the same cascade process, since the dark photon could easily have comparable branching ratios for $\gamma' \rightarrow \text{SMSM}$ and $\gamma' \rightarrow \psi_B\bar{\psi}_B$ when $m_{\gamma'} > m_B$. More generally, it is interesting to contemplate scenarios where ψ_A partially annihilates to boosted ψ_B and partially to SM states. For example, the bremsstrahlung process of $\psi_A\psi_A \rightarrow \psi_B\psi_B\gamma'$, where the γ' decays to an electron-positron pair, can be a source of positrons that can be detected in experiments like AMS-02 [121] or indirectly in Gamma ray telescopes [122]. Of course, if the ψ_B states are not too depleted, then they could give indirect detection signals of their own.

¹⁴There would also be interesting signals for ψ_B in DM production/detection experiments [115].

Finally, it is worth considering the broader experimental signatures possible in the paradigm of DM annihilating to stable dark sector states [32, 35, 43, 123, 124], with simple extensions/variations based on our current model. If $m_B \ll m_e$, then ψ_B acts effectively like dark radiation, which may leave signatures in CMB observables such as N_{eff} [35]. If ψ_A has a non-negligible solar capture cross section, then boosted DM could emerge from the sun. If ψ_B takes up sizable fraction of the total DM abundance (perhaps via a leading asymmetric component), then the fact that ψ_B has strong self-interactions may have implications for small scale structure of DM halos including the known anomalies such as cusp-core and too-big-to-fail problems [125, 126]. The potentially rich structure of the dark sector motivates a comprehensive approach to DM searches.

Acknowledgments

We thank Francesco D’Eramo for early discussions about boosted DM. We also thank Brian Batell, Kfir Blum, Chris Kachulis, Ed Kearns, Kuver Sinha, Tracy Slatyer, Greg Sullivan, Zach Thomas, and Itay Yavin for helpful discussions. We thank our anonymous referee for detailed comments and insightful new detection options. K.A. and Y.C. are supported in part by NSF Grant No. PHY-1315155 and by the Maryland Center for Fundamental Physics. L.N. and J.T. are supported by the U.S. Department of Energy (DOE) under cooperative research agreement DE-FG02-05ER-41360. J.T. is also supported by the DOE Early Career research program DE-FG02-11ER-41741 and by a Sloan Research Fellowship from the Alfred P. Sloan Foundation.

A Analytic approximations to relic abundances

The coupled Boltzmann equations for the evolution of the ψ_A/ψ_B abundances are

$$\begin{aligned} \frac{dn_A}{dt} + 3Hn_A &= -\frac{1}{2}\langle\sigma_{A\bar{A}\rightarrow B\bar{B}}v\rangle\left(n_A^2 - \frac{(n_A^{\text{eq}})^2}{(n_B^{\text{eq}})^2}n_B^2\right), \\ \frac{dn_B}{dt} + 3Hn_B &= -\frac{1}{2}\langle\sigma_{B\bar{B}\rightarrow\gamma'\gamma'v}\rangle(n_B^2 - (n_B^{\text{eq}})^2) - \frac{1}{2}\langle\sigma_{B\bar{B}\rightarrow A\bar{A}}v\rangle\left(n_B^2 - \frac{(n_B^{\text{eq}})^2}{(n_A^{\text{eq}})^2}n_A^2\right), \end{aligned} \quad (\text{A.1})$$

where the factor of $\frac{1}{2}$ arises because ψ_A and ψ_B are Dirac fermions, and n_A refers to the sum of the abundances for ψ_A and $\bar{\psi}_A$ (and similarly for n_B). In terms of the comoving abundance $Y_i = n_i/s$, where s is the entropy of the universe, and $x \equiv m_B/T$, we can rewrite the Boltzmann equations as

$$\frac{dY_A}{dx} = -\frac{\lambda_A}{x^2}\left(Y_A^2 - \frac{(Y_A^{\text{eq}})^2}{(Y_B^{\text{eq}})^2}Y_B^2\right), \quad (\text{A.2})$$

$$\frac{dY_B}{dx} = -\frac{\lambda_B}{x^2}\left(Y_B^2 - (Y_B^{\text{eq}})^2\right) + \frac{\lambda_A}{x^2}\left(Y_A^2 - \frac{(Y_A^{\text{eq}})^2}{(Y_B^{\text{eq}})^2}Y_B^2\right), \quad (\text{A.3})$$

where we have introduced the shorthand notations:

$$\lambda_A \equiv \frac{sx^3}{2H(m_B)}\langle\sigma_{A\bar{A}\rightarrow B\bar{B}}v\rangle, \quad \lambda_B \equiv \frac{sx^3}{2H(m_B)}\langle\sigma_{B\bar{B}\rightarrow\gamma'\gamma'v}\rangle, \quad (\text{A.4})$$

and used the fact the total DM number is not changed by the $\psi_A\bar{\psi}_A \rightarrow \psi_B\bar{\psi}_B$ reaction, i.e.

$$-\langle\sigma_{B\bar{B}\rightarrow A\bar{A}}v\rangle\left(Y_B^2 - \frac{(Y_B^{\text{eq}})^2}{(Y_A^{\text{eq}})^2}Y_A^2\right) = +\langle\sigma_{A\bar{A}\rightarrow B\bar{B}}v\rangle\left(Y_A^2 - \frac{(Y_A^{\text{eq}})^2}{(Y_B^{\text{eq}})^2}Y_B^2\right). \quad (\text{A.5})$$

Obtaining accurate solutions requires solving the above coupled equations numerically. In much of the parameter space of interest, however, it is possible to obtain good analytic approximations based on two effectively decoupled equations. When $m_B < m_A$ and $\lambda_B \gg \lambda_A$, ψ_B typically freezes out of equilibrium well after ψ_A does. Therefore, the evolution of Y_A in eq. (A.2) becomes the conventional Boltzmann equation for one species of DM by taking $Y_B \approx Y_B^{\text{eq}}$ at least up until the ψ_A freeze-out time.¹⁵ In the case of s -wave annihilation of our interest, the relic abundance of ψ_A can be well approximated by the familiar result [127] (with an extra factor of 2 to account for both ψ_A and $\bar{\psi}_A$)

$$Y_A(\infty) \simeq \frac{x_{f,A}}{\lambda_A} = \frac{7.6}{g_{*s}/g_*^{1/2} M_{pl} T_{f,A} \langle \sigma_{A\bar{A} \rightarrow B\bar{B}} v \rangle}, \quad (\text{A.6})$$

where $T_{f,A} = m_B/x_{f,A}$ is the freeze-out temperature for ψ_A , and in the last step we used $sx^3/2H(m_B) = 0.132(g_{*s}/g_*^{1/2})M_{pl}m_B$.

The solution for Y_B is more subtle, but can also be greatly simplified when the freeze-out times of ψ_A and ψ_B are well separated. If $x_{f,B} \gg x_{f,A}$, then we can drop terms suppressed by $(Y_A^{\text{eq}}/Y_B^{\text{eq}})^2$ in eq. (A.3), and we can treat the effect of ψ_A on ψ_B freeze-out by taking $Y_A(x_{f,B}) \simeq Y_A(x_{f,A}) \simeq Y_A(\infty)$. Defining $\Delta \equiv Y_B - Y_B^{\text{eq}}$, we rewrite eq. (A.3) as:

$$\frac{d\Delta}{dx} = -\frac{dY_B^{\text{eq}}}{dx} - \lambda_B x^{-2} \Delta (2Y_B^{\text{eq}} + \Delta) + \lambda_A x^{-2} Y_A^2(\infty). \quad (\text{A.7})$$

Focussing on the epoch when ψ_B starts to deviate from equilibrium, we can apply the ansatz $\Delta = cY_B^{\text{eq}}$, where c is $\mathcal{O}(1)$. The equilibrium distribution for $x \gg 1$ is

$$Y_B^{\text{eq}}(x) \simeq +0.145 \frac{g}{g_{*s}} x^{3/2} e^{-x}, \quad (\text{A.8})$$

$$\frac{dY_B^{\text{eq}}}{dx} \approx -0.145 \frac{g}{g_{*s}} x^{3/2} e^{-x} = -Y_B^{\text{eq}}, \quad (\text{A.9})$$

where we only keep the leading power term in x in the second line. Combining all these, we can rewrite eq. (A.7) as a quadratic equation for Y_B^{eq} ,

$$\lambda_B c(2+c)(Y_B^{\text{eq}})^2 - x_f^2(c+1)Y_B^{\text{eq}} - \lambda_A Y_A^2(\infty) = 0, \quad (\text{A.10})$$

whose real positive solution is

$$Y_B^{\text{eq}}(x) = \frac{(c+1)x^2 + \sqrt{(c+1)^2 x^4 + 4\lambda_B \lambda_A c(c+2)Y_A^2(\infty)}}{2\lambda_B c(2+c)}. \quad (\text{A.11})$$

We can then equate this equation with $Y_B^{\text{eq}}(x) \simeq x^{3/2}e^{-x}$ to solve numerically for $x_{f,B}$.

We can see that by removing the contribution from ψ_A (i.e. the term $\propto \lambda_A Y_A^2(\infty)$) in eq. (A.11), ψ_B freezes out in the standard way. In particular, we have the approximate relation $x_{f,B} \simeq \log \lambda_B - \frac{1}{2} \log x_{f,B}$ which yields

$$Y_B(\infty) \simeq \frac{x_{f,B}}{\lambda_B}, \quad (\text{A.12})$$

¹⁵After ψ_B freezes out, $Y_B \approx Y_B^{\text{eq}}$ is invalid, so the two equations formally “re-couple”. Since Y_A has approached its asymptotic value by then, though, it is insensitive to late-time details.

in analogy with eq. (A.6). We also see that eq. (A.11) approaches the standard freeze-out solution when λ_B decreases and approaches λ_A , such that $\psi_B \bar{\psi}_B \rightarrow \gamma' \gamma'$ freezes out at temperatures comparable to $\psi_A \bar{\psi}_A \rightarrow \psi_B \bar{\psi}_B$; in that regime, the effect of ψ_A on the ψ_B evolution is subdominant since $Y_A^{\text{eq}} < Y_B^{\text{eq}}$ for $m_A > m_B$. Standard freeze-out of ψ_B continues to hold when $\lambda_B \ll \lambda_A$, though the approximate solution eq. (A.11) would not be valid in that regime, since $\Omega_B > \Omega_A$, in contradiction to our ansatz that ψ_A constitutes the major DM component.

More surprising is the case of large λ_B . The $Y_A^2(\infty)$ term in eq. (A.11) dominates when

$$\frac{\lambda_B}{\lambda_A} \left(\frac{m_B}{m_A} \right)^2 \gg x_{f,B}^2, \quad (\text{A.13})$$

where we have estimated $x_{f,A}/x_{f,B} \simeq m_A/m_B$. Taking $Y_B^{\text{eq}}(x_{f,B}) \simeq Y_B(\infty)$, eq. (A.11) reduces to

$$Y_B(\infty) = \sqrt{\frac{\lambda_A}{\lambda_B}} Y_A(\infty). \quad (\text{A.14})$$

This behavior is very strange from the point of view of standard freeze-out, since the abundance of ψ_B scales like $1/\sqrt{\sigma_B}$ (instead of like the expected $1/\sigma_B$). A naive quick way of understanding this behavior is by setting $dY_B/dx \approx 0$ in eq. (A.3) and dropping all Y_i^{eq} terms at late times, which immediately leads to eq. (A.14). We call this “balanced freeze-out”, since the abundance of ψ_B is set by the balance between a depleting term ($\propto \lambda_B Y_B^2$) and a replenishing term ($\propto \lambda_A Y_A^2$). Unlike in ordinary freeze-out where the expansion of the universe plays a key role in setting the abundance, in balanced freeze-out the main effect of the Hubble expansion is simply to drive Y_i^{eq} to zero at late times.

B Direct detection of non-boosted DM

In this paper, we have largely assumed that ψ_A has no couplings to the SM. Given the contact interaction in eq. (2.1), though, ψ_A can interact with the dark photon via ψ_B loops. In this appendix, we consider the direct detection bounds on ψ_A from these loop processes. Of course, as with ψ_B , one can relax direct detection limits by giving ψ_A an inelastic mass splitting.

The loop-induced couplings of ψ_A to the dark photon depend on the UV completion of eq. (2.1), and we consider exchange of a complex scalar ϕ with $U(1)'$ charge as a concrete example. The Lagrangian for this system is

$$\mathcal{L} \supset |D_\mu \phi|^2 - m_\phi^2 |\phi|^2 + \bar{\psi}_B \not{D} \psi_B + (\lambda \bar{\psi}_B \psi_A \phi + \text{h.c.}), \quad (\text{B.1})$$

where $D_\mu = \partial_\mu - ig' A'_\mu$. Integrating out ϕ yields the contact interaction in eq. (2.1) with

$$\frac{1}{\Lambda^2} = \frac{\lambda^2}{m_\phi^2}. \quad (\text{B.2})$$

Through ψ_B - ϕ loops, ψ_A acquires a coupling to the dark photon. In the limit $m_\phi \gg m_A \gg m_B$, the ψ_B - ϕ loop generates the effective dimension six operator

$$\delta \mathcal{L} = \frac{g' \lambda^2 \log(m_B^2/m_\phi^2)}{48\pi^2 m_\phi^2} (\bar{\psi}_A \gamma^\mu \partial^\nu \psi_A F'_{\mu\nu} + \text{h.c.}), \quad (\text{B.3})$$

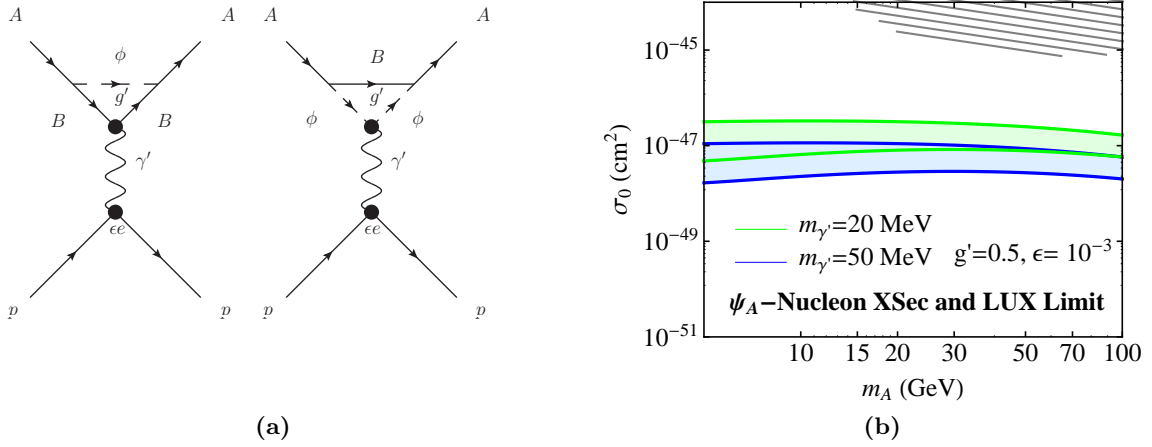


Figure 11. Left: direct detection mechanism for ψ_A via a ψ_B - ϕ loop. Right: scattering cross section of ψ_A on nucleons, sweeping $m_B = 0.1$ GeV–3 GeV and fixing $g' = 0.5$ and $\epsilon = 10^{-3}$. Also shown are the current LUX limit (gray hashes).

which can lead to ψ_A -nucleon scattering as in figure 11a. As discussed in the appendix of ref. [128], the standard dimension five dipole operator $\bar{\psi}_A \Sigma^{\mu\nu} \psi_A F'_{\mu\nu}$ does not appear after integrating out ψ_B and ϕ , because the interactions in eq. (B.1) respect a chiral symmetry acting on ψ_A .

Similar to ref. [128] (but replacing the photon with a dark photon), the dominant effect of eq. (B.3) is to give rise to a charge-charge interaction between DM and a nucleus N . The spin-independent $\psi_A N \rightarrow \psi_A N$ cross section is

$$\frac{d\sigma_{AN \rightarrow AN}}{dE_R} = \frac{m_N (Z\epsilon e)^2}{2\pi v^2} \frac{t^2}{(m_{\gamma'}^2 - t)^2} \left[\frac{g' \lambda^2 \log(m_B^2/m_\phi^2)}{48\pi^2 \frac{m_\phi^2}{m_\phi^2}} \right]^2 F^2(E_R), \quad (\text{B.4})$$

where m_N is the nucleus mass, E_R is the nucleus recoil energy, $t = -2m_N E_R$ is the momentum-transfer-squared, v is the DM velocity, Z is the nucleus charge number, and $F^2(E_R)$ is the nucleus charge form factor. The numerator in the expression above corresponds just to the lowest term in an expansion in t (i.e. small momentum transfer). Spin-independent bounds on DM typically assume equal couplings to neutrons and protons, and can be expressed in terms of an effective nucleon cross section σ_0 , with

$$\frac{d\sigma_{AN \rightarrow AN}}{dE_R} = \sigma_0 \frac{m_N A^2}{2\mu^2 v^2} F^2(E_R), \quad (\text{B.5})$$

where μ is the DM-nucleus reduced mass, and A is the nucleus mass number. Thus, we have

$$\sigma_0 = \frac{\mu^2 (Z\epsilon e)^2}{\pi A^2} \frac{t^2}{(m_{\gamma'}^2 - t)^2} \left[\frac{g' \lambda^2 \log(m_B^2/m_\phi^2)}{48\pi^2 \frac{m_\phi^2}{m_\phi^2}} \right]^2. \quad (\text{B.6})$$

Note that this cross section is momentum dependent, but for simplicity, we will take $E_R \simeq 10$ keV to determine the typical value of t .

For comparing with the bounds from the LUX experiment [94], DM is typically lighter than Xenon in the parameter region of our interest, so $\mu \simeq m_A$. Near the benchmark in eq. (3.6), the DM-nucleon cross section scales roughly as

$$\sigma_0 \approx 2.6 \times 10^{-47} \text{ cm}^2 \left(\frac{\epsilon}{10^{-3}} \right)^2 \left(\frac{g'}{0.5} \right)^2 \left(\frac{250 \text{ GeV}}{\Lambda} \right)^4 \left(\frac{m_A}{20 \text{ GeV}} \right)^2, \quad (\text{B.7})$$

where we have set $\lambda = 1$, ignored the logarithmic dependence on m_ϕ and m_B , and ignored the $m_{\gamma'}$ dependence since $m_{\gamma'}$ is comparable to the typical values of t . Since we adjust Λ (equivalently m_ϕ/λ) to get the right abundance of DM, and since $\Lambda^4 \approx m_A^2 / \langle \sigma_{A\bar{A} \rightarrow B\bar{B}v} \rangle$ from eq. (3.1), we can rewrite this dependence as:

$$\sigma_0 \approx 2.6 \times 10^{-47} \text{ cm}^2 \left(\frac{\epsilon}{10^{-3}} \right)^2 \left(\frac{g'}{0.5} \right)^2 \left(\frac{\langle \sigma_{A\bar{A} \rightarrow B\bar{B}v} \rangle}{5 \times 10^{-26} \text{ cm}^3/\text{sec}} \right). \quad (\text{B.8})$$

In figure 11b, we show the limits of the LUX experiment [94] on the direct detection of ψ_A for different values of $(m_A, m_B, m_{\gamma'})$, and see that these constraints are easily satisfied, though future direct detection experiments would have sensitivity.

C Boosted DM scattering off hadrons

In section 4.2, we focused on the $\psi_B e^- \rightarrow \psi_B e^-$ mode for boosted DM detection. Here, we summarize the signal event rate for boosted DM scattering off protons or nuclei. Since the number of signal events is rather small, we have not pursued a background study, though we remark that the angular pointing for hadronic scattering is rather poor at the low energies we are considering. As discussed in the main text, both event rate and angular resolution for scattering off proton can be improved at liquid Argon detectors.

C.1 Elastic scattering off hadrons

The elastic scattering $\psi_B N \rightarrow \psi_B N$ has similar kinematics to electron scattering (with the obvious replacement $m_e \rightarrow m_N$), except we have to include the electromagnetic form factor. We will express the cross section as a function of the electric and magnetic Sachs form factors G_E and G_M . For protons, we can use the dipole approximation

$$G_E(q^2) = \frac{G_M(q^2)}{2.79} = \frac{1}{(1 + q^2/(0.71 \text{ GeV}^2))^2}. \quad (\text{C.1})$$

To compute the cross section, we use the Rosenbluth formula in the lab frame as cited in [129]

$$\begin{aligned} \frac{d\sigma}{d\Omega} = & \frac{1}{(4\pi)^2} \frac{(\epsilon e)^2 g'^2}{(q^2 - m_{\gamma'}^2)^2} \frac{p'/p}{1 + (E_B - \frac{p E'_B}{p'} \cos \theta)/M} \\ & \times \left(G_E^2 \frac{4E_B E'_B + q^2}{1 - q^2/(4M^2)} + G_M^2 \left((4E_B E'_B + q^2) \left(1 - \frac{1}{1 - q^2/(4M^2)} \right) + \frac{q^4}{2M^2} + \frac{q^2 m_B^2}{M^2} \right) \right). \end{aligned} \quad (\text{C.2})$$

The energies and momenta are defined the same way as eq. (4.10), M is the mass of the proton, and θ is the scattering angle of ψ_B .

The lowest momentum for a proton to Cherenkov radiate is 1.2 GeV, and for our benchmark in eq. (3.6), the proton cross section above this threshold is

$$\sigma_{Bp \rightarrow Bp}^{\text{boost, Cher}} = 1.4 \times 10^{-38} \text{ cm}^2 \left(\frac{\epsilon}{10^{-3}} \right)^2 \left(\frac{g'}{0.5} \right)^2, \quad (\text{C.3})$$

yielding an all-sky event rate of

$$\frac{N_{\text{events}}}{\Delta T} = 1.3 \times 10^{-3} \text{ year}^{-1}. \quad (\text{C.4})$$

Due to the presence of the Cherenkov cutoff and the proton form factor, the elastic scattering rate in eq. (C.3) varies little within the mass range of interest, as given in eq. (2.5). When the transferred energy is above 2.5 GeV, the elastic scattering cross section is rather small, and protons instead typically produce secondary hadronic showers [67]. In that case, one should transition to the DIS calculation below.

As mentioned in footnote 9, large volume liquid Argon detectors are able to detect scattered protons with energies much below the Cherenkov threshold using ionization signals, where the total elastic scattering cross section off protons would be relevant. We find the total elastic $\psi_B p \rightarrow \psi_B p$ scattering cross section for boosted ψ_B with $m_A \gtrsim 1 \text{ GeV}$ to be

$$\sigma_{Bp \rightarrow Bp}^{\text{boost, tot}} = 1.8 \times 10^{-33} \text{ cm}^2 \left(\frac{\epsilon}{10^{-3}} \right)^2 \left(\frac{g'}{0.5} \right)^2 \left(\frac{20 \text{ MeV}}{m_{\gamma'}} \right)^2, \quad (\text{C.5})$$

which is insensitive to m_A, m_B in the boosted ψ_B regime due to the proton form factor. We see that the total elastic scattering rate off proton is much larger than the one with a Cherenkov cutoff, so $\psi_B p \rightarrow \psi_B p$ could potentially be the leading signal detectable at a liquid Argon detector.

Generalizing the previous calculation to a coherent nucleus of charge number Z :

$$\sigma_{BN \rightarrow BN} = 1.2 \times 10^{-30} \text{ cm}^2 \left(\frac{Z}{26} \right)^2 \left(\frac{\epsilon}{10^{-3}} \right)^2 \left(\frac{g'}{0.5} \right)^2 \left(\frac{20 \text{ MeV}}{m_{\gamma'}} \right)^2. \quad (\text{C.6})$$

This same $\psi_B p \rightarrow \psi_B p$ calculation is relevant for direct detection of non-relativistic relic ψ_B . Taking the $q^2 \rightarrow 0$ limit and integrating over all angles, we have the cross section

$$\sigma_{Bp \rightarrow Bp}^{v_B \rightarrow 0} = \frac{(\epsilon e)^2 g'^2}{\pi} \frac{\mu_p^2}{m_{\gamma'}^4}. \quad (\text{C.7})$$

where $\mu_p = m_p m_B / (m_p + m_B)$ is the reduced mass of the dark matter and the proton. This is the basis for eq. (6.1) shown earlier.

C.2 Deep inelastic scattering off hadrons

At sufficiently high energies, ψ_B scattering off hadrons will behave more like deep inelastic scattering (DIS), where the final state is a hadronic shower. The DIS cross section is a convolution of the parton-level cross section with parton distribution functions (PDFs). The parton-level cross section $\hat{\sigma}$ is given by

$$\frac{d\hat{\sigma}}{d\hat{t}} = \frac{1}{8\pi} \frac{(g'\epsilon Q_f)^2}{(\hat{t} - m_{\gamma'}^2)^2} \frac{(\hat{s} - m_B^2)^2 + (\hat{u} - m_B^2)^2 + 2m_B^2 \hat{t}}{(\hat{s} - m_B^2)^2}. \quad (\text{C.8})$$

For the ψ_B -parton system: $\hat{s} + \hat{u} + \hat{t} = 2m_B^2$, $\hat{t} = -Q^2$, and $\hat{s} = (1-x)m_B^2 + xs$. We define x by $p \equiv xP$ where P is the 4-momentum of the initial proton at rest. We define $y \equiv \frac{2P \cdot q}{2P \cdot k} = \frac{-\hat{t}}{\hat{s} - m_B^2}$, which characterizes the fraction of the energy transferred from ψ_B to the parton, since $y = \frac{q^0}{k^0} = 1 - \frac{E'}{E}$ in the rest frame of the initial proton.

From these relations, we get the transferred momentum $Q^2 = xy(s - m_B^2)$, and $dx dQ^2 = \frac{dQ^2}{dy} dx dy = x(s - m_B^2) dx dy$. Including parton distribution functions, and using x/y as variables, we obtain the resulting DIS cross section:

$$\frac{d^2\sigma}{dx dy} = \left(\sum_f x f_f(x, Q) Q_f^2 \right) \frac{(g'\epsilon)^2 s(2x - 2xy + xy^2) + m_B^2(-2x - xy^2 - 2y(1-x))}{8\pi x (xy(s - m_B^2) + m_{\gamma'}^2)^2}, \quad (\text{C.9})$$

where $f_f(x, Q)$ are PDFs with f indicating different flavor of fermion. For numerical evaluation, we use the MSTW2008 NNLO PDFs from ref. [130]. The integration limits of eq. (C.9) are $0 \leq x \leq 1$ and $0 \leq y \leq y_{\max}$, where applying the condition $\cos \theta \leq 1$ we obtain

$$y_{\max} = \frac{4(E_B^2 - m_B^2)(m_B^2 - s)x}{-4E_B^2 m_B^2 + 4E_B^2 m_B^2 x - 4E_B^2 s x - m_B^4 x^2 + 2m_B^2 s x^2 - s^2 x^2}, \quad (\text{C.10})$$

with $s = m_B^2 + M_p^2 + 2M_p E_B$ and $E_B = m_A$. Unlike the familiar case of DIS initiated by nearly massless incoming particles, for the massive ψ_B we consider here, y_{\max} is not trivially 1.

Since the PDFs are only reliable for transferred energies over ~ 1 GeV, we impose $Q^2 \geq (1 \text{ GeV})^2$ as a default cut for numerical integration. Analogous to the discussion for elastic scattering signals, for a particular experiment, there may be harder cut on phase space due to detector threshold energy. For our benchmark in eq. (3.6), the DIS cross section above the 1 GeV threshold is

$$\sigma_{Bp \rightarrow BX} = 1.42 \times 10^{-37} \text{ cm}^2 \left(\frac{\epsilon}{10^{-3}} \right)^2 \left(\frac{g'}{0.5} \right)^2, \quad (\text{C.11})$$

yielding

$$\frac{N_{\text{events}}}{\Delta T} = 3.6 \times 10^{-2} \text{ year}^{-1} \quad (\text{C.12})$$

for the all-sky event rate at Super-K.

References

- [1] F. Zwicky, *Die Rotverschiebung von extragalaktischen Nebeln*, *Helv. Phys. Acta* **6** (1933) 110 [[INSPIRE](#)].
- [2] K.G. Begeman, A.H. Broeils and R.H. Sanders, *Extended rotation curves of spiral galaxies: dark haloes and modified dynamics*, *Mon. Not. Roy. Astron. Soc.* **249** (1991) 523 [[INSPIRE](#)].
- [3] G. Bertone et al., *Particle Dark Matter: Observations, Models and Searches*, Cambridge University Press (2010).
- [4] G. Bélanger and J.-C. Park, *Assisted freeze-out*, *JCAP* **03** (2012) 038 [[arXiv:1112.4491](#)] [[INSPIRE](#)].
- [5] SUPER-KAMIOKANDE collaboration, Y. Fukuda et al., *The Super-Kamiokande detector*, *Nucl. Instrum. Meth. A* **501** (2003) 418 [[INSPIRE](#)].

- [6] K. Abe et al., *Letter of Intent: The Hyper-Kamiokande Experiment — Detector Design and Physics Potential*, [arXiv:1109.3262](#) [INSPIRE].
- [7] ICECUBE collaboration, J. Ahrens et al., *IceCube - the next generation neutrino telescope at the south pole*, *Nucl. Phys. Proc. Suppl.* **118** (2003) 388 [[astro-ph/0209556](#)] [INSPIRE].
- [8] ICECUBE-PINGU collaboration, M.G. Aartsen et al., *Letter of Intent: The Precision IceCube Next Generation Upgrade (PINGU)*, [arXiv:1401.2046](#) [INSPIRE].
- [9] E. Resconi, *The stepping stones to proton decay: IceCube, PINGU, MICA*, talk given at *New Directions in Neutrino Physics*, Aspen, U.S.A., February 2013.
<http://indico.cern.ch/event/224351/contribution/33/material/slides/0.pdf>.
- [10] U.F. Katz, *KM3NeT: Towards a km³ Mediterranean Neutrino Telescope*, *Nucl. Instrum. Meth. A* **567** (2006) 457 [[astro-ph/0606068](#)] [INSPIRE].
- [11] ANTARES collaboration, M. Ageron et al., *ANTARES: the first undersea neutrino telescope*, *Nucl. Instrum. Meth. A* **656** (2011) 11 [[arXiv:1104.1607](#)] [INSPIRE].
- [12] A. Bueno et al., *Nucleon decay searches with large liquid argon TPC detectors at shallow depths: Atmospheric neutrinos and cosmogenic backgrounds*, *JHEP* **04** (2007) 041 [[hep-ph/0701101](#)] [INSPIRE].
- [13] A. Badertscher et al., *Giant Liquid Argon Observatory for Proton Decay, Neutrino Astrophysics and CP-violation in the Lepton Sector (GLACIER)*, [arXiv:1001.0076](#) [INSPIRE].
- [14] Y.-F. Li, J. Cao, Y. Wang and L. Zhan, *Unambiguous Determination of the Neutrino Mass Hierarchy Using Reactor Neutrinos*, *Phys. Rev. D* **88** (2013) 013008 [[arXiv:1303.6733](#)] [INSPIRE].
- [15] Y.-F. Li, *Overview of the Jiangmen Underground Neutrino Observatory (JUNO)*, *Int. J. Mod. Phys. Conf. Ser.* **31** (2014) 1460300 [[arXiv:1402.6143](#)] [INSPIRE].
- [16] F. D’Eramo and J. Thaler, *Semi-annihilation of Dark Matter*, *JHEP* **06** (2010) 109 [[arXiv:1003.5912](#)] [INSPIRE].
- [17] H.S. Cheon, S.K. Kang and C.S. Kim, *Doubly Coexisting Dark Matter Candidates in an Extended Seesaw Model*, *Phys. Lett. B* **675** (2009) 203 [Erratum *ibid.* **B 698** (2011) 324] [[arXiv:0807.0981](#)] [INSPIRE].
- [18] T. Hambye, *Hidden vector dark matter*, *JHEP* **01** (2009) 028 [[arXiv:0811.0172](#)] [INSPIRE].
- [19] T. Hambye and M.H.G. Tytgat, *Confined hidden vector dark matter*, *Phys. Lett. B* **683** (2010) 39 [[arXiv:0907.1007](#)] [INSPIRE].
- [20] C. Arina, T. Hambye, A. Ibarra and C. Weniger, *Intense Gamma-Ray Lines from Hidden Vector Dark Matter Decay*, *JCAP* **03** (2010) 024 [[arXiv:0912.4496](#)] [INSPIRE].
- [21] G. Bélanger, K. Kannike, A. Pukhov and M. Raidal, *Impact of semi-annihilations on dark matter phenomenology - an example of Z_N symmetric scalar dark matter*, *JCAP* **04** (2012) 010 [[arXiv:1202.2962](#)] [INSPIRE].
- [22] E.D. Carlson, M.E. Machacek and L.J. Hall, *Self-interacting dark matter*, *Astrophys. J.* **398** (1992) 43 [INSPIRE].
- [23] A.A. de Laix, R.J. Scherrer and R.K. Schaefer, *Constraints of selfinteracting dark matter*, *Astrophys. J.* **452** (1995) 495 [[astro-ph/9502087](#)] [INSPIRE].
- [24] Y. Hochberg, E. Kuflik, T. Volansky and J.G. Wacker, *The SIMP Miracle*, [arXiv:1402.5143](#) [INSPIRE].
- [25] Z.G. Berezhiani, A.D. Dolgov and R.N. Mohapatra, *Asymmetric inflationary reheating and the nature of mirror universe*, *Phys. Lett. B* **375** (1996) 26 [[hep-ph/9511221](#)] [INSPIRE].

- [26] Z. Berezhiani, D. Comelli and F.L. Villante, *The Early mirror universe: inflation, baryogenesis, nucleosynthesis and dark matter*, *Phys. Lett. B* **503** (2001) 362 [[hep-ph/0008105](#)] [[INSPIRE](#)].
- [27] P. Ciarcelluti, *Cosmology with mirror dark matter*, *Int. J. Mod. Phys. D* **19** (2010) 2151 [[arXiv:1102.5530](#)] [[INSPIRE](#)].
- [28] J.L. Feng, H. Tu and H.-B. Yu, *Thermal Relics in Hidden Sectors*, *JCAP* **10** (2008) 043 [[arXiv:0808.2318](#)] [[INSPIRE](#)].
- [29] J. Huang and Y. Zhao, *Dark Matter Induced Nucleon Decay: Model and Signatures*, *JHEP* **02** (2014) 077 [[arXiv:1312.0011](#)] [[INSPIRE](#)].
- [30] M. Pospelov, A. Ritz and M.B. Voloshin, *Secluded WIMP Dark Matter*, *Phys. Lett. B* **662** (2008) 53 [[arXiv:0711.4866](#)] [[INSPIRE](#)].
- [31] N. Arkani-Hamed, D.P. Finkbeiner, T.R. Slatyer and N. Weiner, *A Theory of Dark Matter*, *Phys. Rev. D* **79** (2009) 015014 [[arXiv:0810.0713](#)] [[INSPIRE](#)].
- [32] L. Ackerman, M.R. Buckley, S.M. Carroll and M. Kamionkowski, *Dark Matter and Dark Radiation*, *Phys. Rev. D* **79** (2009) 023519 [[arXiv:0810.5126](#)] [[INSPIRE](#)].
- [33] Y. Nomura and J. Thaler, *Dark Matter through the Axion Portal*, *Phys. Rev. D* **79** (2009) 075008 [[arXiv:0810.5397](#)] [[INSPIRE](#)].
- [34] J. Mardon, Y. Nomura and J. Thaler, *Cosmic Signals from the Hidden Sector*, *Phys. Rev. D* **80** (2009) 035013 [[arXiv:0905.3749](#)] [[INSPIRE](#)].
- [35] Z. Chacko, Y. Cui, S. Hong and T. Okui, *Cosmological Signals of a Hidden Dark Matter Sector*, work in preparation.
- [36] J.D. Bjorken, R. Essig, P. Schuster and N. Toro, *New Fixed-Target Experiments to Search for Dark Gauge Forces*, *Phys. Rev. D* **80** (2009) 075018 [[arXiv:0906.0580](#)] [[INSPIRE](#)].
- [37] H.M. Hodges, *Mirror baryons as the dark matter*, *Phys. Rev. D* **47** (1993) 456 [[INSPIRE](#)].
- [38] R.N. Mohapatra and V.L. Teplitz, *Mirror dark matter and galaxy core densities of galaxies*, *Phys. Rev. D* **62** (2000) 063506 [[astro-ph/0001362](#)] [[INSPIRE](#)].
- [39] R. Foot, *Seven (and a half) reasons to believe in mirror matter: from neutrino puzzles to the inferred dark matter in the universe*, *Acta Phys. Polon. B* **32** (2001) 2253 [[astro-ph/0102294](#)] [[INSPIRE](#)].
- [40] M. Fairbairn and J. Zupan, *Dark matter with a late decaying dark partner*, *JCAP* **07** (2009) 001 [[arXiv:0810.4147](#)] [[INSPIRE](#)].
- [41] K.M. Zurek, *Multi-Component Dark Matter*, *Phys. Rev. D* **79** (2009) 115002 [[arXiv:0811.4429](#)] [[INSPIRE](#)].
- [42] S. Profumo, K. Sigurdson and L. Ubaldi, *Can we discover multi-component WIMP dark matter?*, *JCAP* **12** (2009) 016 [[arXiv:0907.4374](#)] [[INSPIRE](#)].
- [43] J. Fan, A. Katz, L. Randall and M. Reece, *Double-Disk Dark Matter*, *Phys. Dark Univ.* **2** (2013) 139 [[arXiv:1303.1521](#)] [[INSPIRE](#)].
- [44] Y. Cui, J.D. Mason and L. Randall, *General Analysis of Antideuteron Searches for Dark Matter*, *JHEP* **11** (2010) 017 [[arXiv:1006.0983](#)] [[INSPIRE](#)].
- [45] B. Holdom, *Two $U(1)$'s and Epsilon Charge Shifts*, *Phys. Lett. B* **166** (1986) 196 [[INSPIRE](#)].
- [46] L.B. Okun, *Limits of electrodynamics: paraxphotons?*, *Sov. Phys. JETP* **56** (1982) 502 [[INSPIRE](#)].
- [47] P. Galison and A. Manohar, *Two Z 's or not two Z 's?*, *Phys. Lett. B* **136** (1984) 279 [[INSPIRE](#)].

- [48] E.C.G. Stueckelberg, *Interaction energy in electrodynamics and in the field theory of nuclear forces*, *Helv. Phys. Acta* **11** (1938) 225 [INSPIRE].
- [49] B. K rs and P. Nath, *Aspects of the Stueckelberg extension*, *JHEP* **07** (2005) 069 [[hep-ph/0503208](#)] [INSPIRE].
- [50] D. Tucker-Smith and N. Weiner, *Inelastic dark matter*, *Phys. Rev. D* **64** (2001) 043502 [[hep-ph/0101138](#)] [INSPIRE].
- [51] Y. Cui, D.E. Morrissey, D. Poland and L. Randall, *Candidates for Inelastic Dark Matter*, *JHEP* **05** (2009) 076 [[arXiv:0901.0557](#)] [INSPIRE].
- [52] D.P. Finkbeiner, T.R. Slatyer, N. Weiner and I. Yavin, *PAMELA, DAMA, INTEGRAL and Signatures of Metastable Excited WIMPs*, *JCAP* **09** (2009) 037 [[arXiv:0903.1037](#)] [INSPIRE].
- [53] P.W. Graham, R. Harnik, S. Rajendran and P. Saraswat, *Exothermic Dark Matter*, *Phys. Rev. D* **82** (2010) 063512 [[arXiv:1004.0937](#)] [INSPIRE].
- [54] M. Pospelov, N. Weiner and I. Yavin, *Dark matter detection in two easy steps*, *Phys. Rev. D* **89** (2014) 055008 [[arXiv:1312.1363](#)] [INSPIRE].
- [55] S. Bhattacharya, A. Drozd, B. Grzadkowski and J. Wudka, *Two-Component Dark Matter*, *JHEP* **10** (2013) 158 [[arXiv:1309.2986](#)] [INSPIRE].
- [56] K.P. Modak, D. Majumdar and S. Rakshit, *A Possible Explanation of Low Energy γ -ray Excess from Galactic Centre and Fermi Bubble by a Dark Matter Model with Two Real Scalars*, [arXiv:1312.7488](#) [INSPIRE].
- [57] G. Steigman, B. Dasgupta and J.F. Beacom, *Precise Relic WIMP Abundance and its Impact on Searches for Dark Matter Annihilation*, *Phys. Rev. D* **86** (2012) 023506 [[arXiv:1204.3622](#)] [INSPIRE].
- [58] P. Fayet, *U-boson production in e^+e^- annihilations, psi and Upsilon decays and Light Dark Matter*, *Phys. Rev. D* **75** (2007) 115017 [[hep-ph/0702176](#)] [INSPIRE].
- [59] M. Pospelov, *Secluded U(1) below the weak scale*, *Phys. Rev. D* **80** (2009) 095002 [[arXiv:0811.1030](#)] [INSPIRE].
- [60] J.F. Navarro, C.S. Frenk and S.D.M. White, *The Structure of cold dark matter halos*, *Astrophys. J.* **462** (1996) 563 [[astro-ph/9508025](#)] [INSPIRE].
- [61] M. Cirelli et al., *PPPC 4 DM ID: A Poor Particle Physicist Cookbook for Dark Matter Indirect Detection*, *JCAP* **03** (2011) 051 [Erratum *ibid.* **1210** (2012) E01] [[arXiv:1012.4515](#)] [INSPIRE].
- [62] H. Yuksel, S. Horiuchi, J.F. Beacom and S. Ando, *Neutrino Constraints on the Dark Matter Total Annihilation Cross section*, *Phys. Rev. D* **76** (2007) 123506 [[arXiv:0707.0196](#)] [INSPIRE].
- [63] SPER-KAMIOKANDE collaboration, P. Mijakowski, *Search for neutrinos from diffuse dark matter annihilation in Super-Kamiokande*, *Nucl. Phys. Proc. Suppl.* **229-232** (2012) 546 [INSPIRE].
- [64] J.F. Beacom, N.F. Bell and G.D. Mack, *General Upper Bound on the Dark Matter Total Annihilation Cross section*, *Phys. Rev. Lett.* **99** (2007) 231301 [[astro-ph/0608090](#)] [INSPIRE].
- [65] R. Essig et al., *Dark Sectors and New, Light, Weakly-Coupled Particles*, [arXiv:1311.0029](#) [INSPIRE].
- [66] J.F. Beacom and S. Palomares-Ruiz, *Neutral current atmospheric neutrino flux measurement using neutrino proton elastic scattering in Super-Kamiokande*, *Phys. Rev. D* **67** (2003) 093001 [[hep-ph/0301060](#)] [INSPIRE].
- [67] M. Fechner and C.W. Walter, *The physics impact of proton track identification in future megaton-scale water Cherenkov detectors*, *JHEP* **11** (2009) 040 [[arXiv:0901.1950](#)] [INSPIRE].

- [68] SUPER-KAMIOKANDE collaboration, M. Fechner et al., *Kinematic reconstruction of atmospheric neutrino events in a large water Cherenkov detector with proton identification*, *Phys. Rev. D* **79** (2009) 112010 [[arXiv:0901.1645](#)] [[INSPIRE](#)].
- [69] SUPER-KAMIOKANDE collaboration, Y. Ashie et al., *A Measurement of atmospheric neutrino oscillation parameters by Super-Kamiokande I*, *Phys. Rev. D* **71** (2005) 112005 [[hep-ex/0501064](#)] [[INSPIRE](#)].
- [70] TIBET ASGAMMA collaboration, M. Amenomori et al., *The cosmic-ray energy spectrum around the knee measured by the Tibet-III air-shower array*, *Nucl. Phys. Proc. Suppl.* **175-176** (2008) 318 [[INSPIRE](#)].
- [71] J.A. Formaggio and G.P. Zeller, *From eV to EeV: Neutrino Cross sections Across Energy Scales*, *Rev. Mod. Phys.* **84** (2012) 1307 [[arXiv:1305.7513](#)] [[INSPIRE](#)].
- [72] L.K. Pik, *Study of the neutrino mass hierarchy with the atmospheric neutrino data observed in Super-Kamiokande*, Ph.D. thesis, University of Tokyo, Japan (2012).
- [73] M. Dziomba, *A Study of Neutrino Oscillation Models with Super-Kamiokande Atmospheric Neutrino Data*, Ph.D. thesis, University of Washington, U.S.A. (2012).
- [74] J.F. Beacom and M.R. Vagins, *GADZOOKS! Anti-neutrino spectroscopy with large water Cherenkov detectors*, *Phys. Rev. Lett.* **93** (2004) 171101 [[hep-ph/0309300](#)] [[INSPIRE](#)].
- [75] T.K. Gaisser and M. Honda, *Flux of atmospheric neutrinos*, *Ann. Rev. Nucl. Part. Sci.* **52** (2002) 153 [[hep-ph/0203272](#)] [[INSPIRE](#)].
- [76] SUPER-KAMIOKANDE collaboration, K. Bays et al., *Supernova Relic Neutrino Search at Super-Kamiokande*, *Phys. Rev. D* **85** (2012) 052007 [[arXiv:1111.5031](#)] [[INSPIRE](#)].
- [77] SUPER-KAMIOKANDE collaboration, K. Abe et al., *Solar neutrino results in Super-Kamiokande-III*, *Phys. Rev. D* **83** (2011) 052010 [[arXiv:1010.0118](#)] [[INSPIRE](#)].
- [78] PARTICLE DATA GROUP collaboration, J. Beringer et al., *Review of Particle Physics (RPP)*, *Phys. Rev. D* **86** (2012) 010001 [[INSPIRE](#)].
- [79] I. Albuquerque, G. Burdman and Z. Chacko, *Neutrino telescopes as a direct probe of supersymmetry breaking*, *Phys. Rev. Lett.* **92** (2004) 221802 [[hep-ph/0312197](#)] [[INSPIRE](#)].
- [80] J. Berger, Y. Cui and Y. Zhao, *Detecting Boosted Dark Matter from the Sun with Large Volume Neutrino Detectors*, [arXiv:1410.2246](#) [[INSPIRE](#)].
- [81] HYPER-KAMIOKANDE WORKING GROUP collaboration, E. Kearns et al., *Hyper-Kamiokande Physics Opportunities*, [arXiv:1309.0184](#) [[INSPIRE](#)].
- [82] ICECUBE collaboration, R. Abbasi et al., *Search for Dark Matter from the Galactic Halo with the IceCube Neutrino Observatory*, *Phys. Rev. D* **84** (2011) 022004 [[arXiv:1101.3349](#)] [[INSPIRE](#)].
- [83] ICECUBE collaboration, M.G. Aartsen et al., *Search for neutrino-induced particle showers with IceCube-40*, *Phys. Rev. D* **89** (2014) 102001 [[arXiv:1312.0104](#)] [[INSPIRE](#)].
- [84] D. Cowen, *GeV-scale physics in the south pole icecap*, talk given at *Baryon Number Violation Workshop*, Argonne, U.S.A., April 2013, <https://indico.fnal.gov/getFile.py/access?contribId=118&sessionId=3&resId=0&materialId=slides&confId=6248>.
- [85] F.M. Dufour, *Precise study of the atmospheric neutrino oscillation pattern using Super-Kamiokande I and II*, Ph.D. thesis, Boston University, U.S.A. (2003).
- [86] V.I. Galkin, A.M. Anokhina, E. Konishi and A. Misaki, *On the Capability Of Super-Kamiokande Detector To Define the Primary Parameters Of Muon And Electron Events*, [arXiv:0808.0824](#) [[INSPIRE](#)].

- [87] SUPER-KAMIOKANDÉ collaboration, R. Wendell et al., *Atmospheric neutrino oscillation analysis with sub-leading effects in Super-Kamiokande I, II and III*, *Phys. Rev. D* **81** (2010) 092004 [[arXiv:1002.3471](#)] [[INSPIRE](#)].
- [88] SUPER-KAMIOKANDÉ collaboration, S. Desai et al., *Search for dark matter WIMPs using upward through-going muons in Super-Kamiokande*, *Phys. Rev. D* **70** (2004) 083523 [Erratum *ibid.* **70** (2004) 109901] [[hep-ex/0404025](#)] [[INSPIRE](#)].
- [89] P. Mijakowski, *Direct and Indirect Search for Dark Matter*, Ph.D. thesis, University of Warsaw, Poland (2011).
- [90] P. Mijakowski, *Indirect searches for dark matter particles at Super-Kamiokande*, talk given at 4th Rencontres de Moriond, Cosmology Session, La Thuile, Italy, 11 March 2012, http://moriond.in2p3.fr/J12/transparencies/11_Sunday_pm/mijakowski.pdf.
- [91] R. Wendell, *Atmospheric Results from Super-Kamiokande*, talk given at Neutrino 2014, Boston University, U.S.A., 3 June 2014, <https://indico.fnal.gov/getFile.py/access?contribId=260&sessionId=17&resId=0&materialId=slides&confId=8022>.
- [92] SUPER-KAMIOKANDÉ collaboration, A. Himmel, *Recent results from Super-Kamiokande*, *AIP Conf. Proc.* **1604** (2014) 345 [[arXiv:1310.6677](#)] [[INSPIRE](#)].
- [93] J. Blümlein and J. Brunner, *New Exclusion Limits on Dark Gauge Forces from Proton Bremsstrahlung in Beam-Dump Data*, *Phys. Lett. B* **731** (2014) 320 [[arXiv:1311.3870](#)] [[INSPIRE](#)].
- [94] LUX collaboration, D.S. Akerib et al., *First results from the LUX dark matter experiment at the Sanford Underground Research Facility*, *Phys. Rev. Lett.* **112** (2014) 091303 [[arXiv:1310.8214](#)] [[INSPIRE](#)].
- [95] SUPERCDMS collaboration, R. Agnese et al., *Search for Low-Mass WIMPs with SuperCDMS*, *Phys. Rev. Lett.* **112** (2014) 241302 [[arXiv:1402.7137](#)] [[INSPIRE](#)].
- [96] SUPERCDMS collaboration, R. Agnese et al., *Search for Low-Mass Weakly Interacting Massive Particles Using Voltage-Assisted Calorimetric Ionization Detection in the SuperCDMS Experiment*, *Phys. Rev. Lett.* **112** (2014) 041302 [[arXiv:1309.3259](#)] [[INSPIRE](#)].
- [97] DAMIC collaboration, J. Barreto et al., *Direct Search for Low Mass Dark Matter Particles with CCDs*, *Phys. Lett. B* **711** (2012) 264 [[arXiv:1105.5191](#)] [[INSPIRE](#)].
- [98] R. Essig, J. Mardon and T. Volansky, *Direct Detection of Sub-GeV Dark Matter*, *Phys. Rev. D* **85** (2012) 076007 [[arXiv:1108.5383](#)] [[INSPIRE](#)].
- [99] R. Essig, A. Manalaysay, J. Mardon, P. Sorensen and T. Volansky, *First Direct Detection Limits on sub-GeV Dark Matter from XENON10*, *Phys. Rev. Lett.* **109** (2012) 021301 [[arXiv:1206.2644](#)] [[INSPIRE](#)].
- [100] L. Bergstrom, T. Bringmann, I. Cholis, D. Hooper and C. Weniger, *New limits on dark matter annihilation from AMS cosmic ray positron data*, *Phys. Rev. Lett.* **111** (2013) 171101 [[arXiv:1306.3983](#)] [[INSPIRE](#)].
- [101] A. Ibarra, A.S. Lamperstorfer and J. Silk, *Dark matter annihilations and decays after the AMS-02 positron measurements*, *Phys. Rev. D* **89** (2014) 063539 [[arXiv:1309.2570](#)] [[INSPIRE](#)].
- [102] LAT collaboration, M. Ackermann et al., *Fermi LAT Search for Dark Matter in Gamma-ray Lines and the Inclusive Photon Spectrum*, *Phys. Rev. D* **86** (2012) 022002 [[arXiv:1205.2739](#)] [[INSPIRE](#)].
- [103] M.S. Madhavacheril, N. Sehgal and T.R. Slatyer, *Current Dark Matter Annihilation Constraints from CMB and Low-Redshift Data*, *Phys. Rev. D* **89** (2014) 103508 [[arXiv:1310.3815](#)] [[INSPIRE](#)].

- [104] T.R. Slatyer, *The Sommerfeld enhancement for dark matter with an excited state*, *JCAP* **02** (2010) 028 [[arXiv:0910.5713](#)] [[INSPIRE](#)].
- [105] B. Henning and H. Murayama, *Constraints on Light Dark Matter from Big Bang Nucleosynthesis*, [arXiv:1205.6479](#) [[INSPIRE](#)].
- [106] Z. Berezhiani, A. Dolgov and I. Tkachev, *BBN with light dark matter*, *JCAP* **02** (2013) 010 [[arXiv:1211.4937](#)] [[INSPIRE](#)].
- [107] J. Goodman et al., *Constraints on Dark Matter from Colliders*, *Phys. Rev. D* **82** (2010) 116010 [[arXiv:1008.1783](#)] [[INSPIRE](#)].
- [108] P.J. Fox, R. Harnik, J. Kopp and Y. Tsai, *LEP Shines Light on Dark Matter*, *Phys. Rev. D* **84** (2011) 014028 [[arXiv:1103.0240](#)] [[INSPIRE](#)].
- [109] P.J. Fox, R. Harnik, J. Kopp and Y. Tsai, *Missing Energy Signatures of Dark Matter at the LHC*, *Phys. Rev. D* **85** (2012) 056011 [[arXiv:1109.4398](#)] [[INSPIRE](#)].
- [110] P. Cennini et al., *Detection of scintillation light in coincidence with ionizing tracks in a liquid argon time projection chamber*, *Nucl. Instrum. Meth. A* **432** (1999) 240 [[INSPIRE](#)].
- [111] B. Baller et al., *Liquid Argon Time Projection Chamber Research and Development in the United States*, 2014 *JINST* **9** T05005 [[arXiv:1307.8166](#)] [[INSPIRE](#)].
- [112] G. Bélanger, K. Kannike, A. Pukhov and M. Raidal, *Z_3 Scalar Singlet Dark Matter*, *JCAP* **01** (2013) 022 [[arXiv:1211.1014](#)] [[INSPIRE](#)].
- [113] P. Ko and Y. Tang, *Self-interacting scalar dark matter with local Z_3 symmetry*, *JCAP* **05** (2014) 047 [[arXiv:1402.6449](#)] [[INSPIRE](#)].
- [114] M. Aoki and T. Toma, *Impact of semi-annihilation of Z_3 symmetric dark matter with radiative neutrino masses*, *JCAP* **09** (2014) 016 [[arXiv:1405.5870](#)] [[INSPIRE](#)].
- [115] B. Batell, M. Pospelov and A. Ritz, *Exploring Portals to a Hidden Sector Through Fixed Targets*, *Phys. Rev. D* **80** (2009) 095024 [[arXiv:0906.5614](#)] [[INSPIRE](#)].
- [116] T. Daylan et al., *The Characterization of the Gamma-Ray Signal from the Central Milky Way: A Compelling Case for Annihilating Dark Matter*, [arXiv:1402.6703](#) [[INSPIRE](#)].
- [117] C. Boehm, M.J. Dolan and C. McCabe, *A weighty interpretation of the Galactic Centre excess*, *Phys. Rev. D* **90** (2014) 023531 [[arXiv:1404.4977](#)] [[INSPIRE](#)].
- [118] M. Abdullah et al., *Hidden On-Shell Mediators for the Galactic Center Gamma-Ray Excess*, *Phys. Rev. D* **90** (2014) 035004 [[arXiv:1404.6528](#)] [[INSPIRE](#)].
- [119] A. Martin, J. Shelton and J. Unwin, *Fitting the Galactic Center Gamma-Ray Excess with Cascade Annihilations*, [arXiv:1405.0272](#) [[INSPIRE](#)].
- [120] A. Berlin, P. Gratia, D. Hooper and S.D. McDermott, *Hidden Sector Dark Matter Models for the Galactic Center Gamma-Ray Excess*, *Phys. Rev. D* **90** (2014) 015032 [[arXiv:1405.5204](#)] [[INSPIRE](#)].
- [121] AMS collaboration, M. Aguilar et al., *First Result from the Alpha Magnetic Spectrometer on the International Space Station: Precision Measurement of the Positron Fraction in Primary Cosmic Rays of 0.5–350 GeV*, *Phys. Rev. Lett.* **110** (2013) 141102 [[INSPIRE](#)].
- [122] R. Essig, E. Kuflik, S.D. McDermott, T. Volansky and K.M. Zurek, *Constraining Light Dark Matter with Diffuse X-Ray and Gamma-Ray Observations*, *JHEP* **11** (2013) 193 [[arXiv:1309.4091](#)] [[INSPIRE](#)].
- [123] M. Blennow, E. Fernandez-Martinez, O. Mena, J. Redondo and P. Serra, *Asymmetric Dark Matter and Dark Radiation*, *JCAP* **07** (2012) 022 [[arXiv:1203.5803](#)] [[INSPIRE](#)].
- [124] X. Chu and B. Dasgupta, *A “Pas de Deux” - Dark Radiation Fattens and Puffs-up Dark Matter Halos*, [arXiv:1404.6127](#) [[INSPIRE](#)].

- [125] W.J.G. de Blok, *The Core-Cusp Problem*, *Adv. Astron.* **2010** (2010) 789293 [[arXiv:0910.3538](#)] [[INSPIRE](#)].
- [126] M. Boylan-Kolchin, J.S. Bullock and M. Kaplinghat, *Too big to fail? The puzzling darkness of massive Milky Way subhaloes*, *Mon. Not. Roy. Astron. Soc.* **415** (2011) L40 [[arXiv:1103.0007](#)] [[INSPIRE](#)].
- [127] E.W. Kolb and M.S. Turner, *The Early Universe*, *Front. Phys.* **69** (1990).
- [128] P. Agrawal, S. Blanchet, Z. Chacko and C. Kilic, *Flavored Dark Matter and Its Implications for Direct Detection and Colliders*, *Phys. Rev. D* **86** (2012) 055002 [[arXiv:1109.3516](#)] [[INSPIRE](#)].
- [129] E. Borie, *Muon-proton Scattering*, [arXiv:1207.6651](#) [[INSPIRE](#)].
- [130] A.D. Martin, W.J. Stirling, R.S. Thorne and G. Watt, *Parton distributions for the LHC*, *Eur. Phys. J. C* **63** (2009) 189 [[arXiv:0901.0002](#)] [[INSPIRE](#)].