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WITH FOUR YEARS OF DATA FROM THE ICETOP DETECTOR*

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SEARCH FOR SOURCES OF HIGH-ENERGY NEUTRONS WITH FOUR YEARS OF DATA FROM THE ICETOP DETECTOR

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ABSTRACT

IceTop is an air-shower array located on the Antarctic ice sheet at the geographic South Pole. IceTop can detect an astrophysical flux of neutrons from Galactic sources as an excess of cosmic-ray air showers arriving from the source direction. Neutrons are undeflected by the Galactic magnetic field and can typically travel 10 (E/PeV) pc before decay. Two searches are performed using 4 yr of the IceTop data set to look for a statistically significant excess of events with energies above 10 PeV (10^{16} eV) arriving within a small solid angle. The all-sky search method covers from -90° to approximately -50° in declination. No significant excess is found. A targeted search is also performed, looking for significant correlation with candidate sources in different target sets. This search uses a higher-energy cut (100 PeV) since most target objects lie beyond 1 kpc. The target sets include pulsars with confirmed TeV energy photon fluxes and high-mass X-ray binaries. No significant correlation is found for any target set. Flux upper limits are determined for both searches, which can constrain Galactic neutron sources and production scenarios.

Key words: astroparticle physics – cosmic rays – methods: data analysis

1. INTRODUCTION

The Galactic magnetic field (GMF) strongly affects the arrival distribution of charged cosmic rays, thereby obscuring their sources. A compact source of high-energy neutrons would manifest as a point source in cosmic-ray arrival directions since neutrons are not deflected by magnetic fields. Secondary neutral particles are an expected signature of hadronic acceleration in Galactic sources. Neutral particles would be produced as the cosmic-ray protons and nuclei undergo pp and $p\gamma$ collisions and photodisintegration, respectively, on the ambient photons and cosmic rays within the dense environment surrounding their source (see, e.g., Candia et al. 2002; Crocker et al. 2005; Civasinni et al. 2006; Anchordoqui et al. 2007). For example, neutrons result from charge-exchange

interactions,

$$p\gamma \rightarrow n\pi^+,$$

where a π^+ emerges with the proton's positive charge and the neutron retains most of the energy. For interacting proton primaries, photons resulting from π^0 decays take a small fraction of the proton energy. The production of neutrons exceeds the production of photons at the same energy (Crocker et al. 2005).

It is plausible that known Galactic sources could produce high-energy neutron fluxes, based on the measured TeV energy photon flux. For some Galactic sources, the energy flux of TeV photons is greater than $1 \text{ eV cm}^{-2} \text{ s}^{-1}$ (Hinton & Hofmann 2009). Sources producing particle fluxes with an E^{-2}

differential energy spectrum inject equal energy into each energy decade. If sources in the Galaxy produce PeV photons in addition to TeV photons, the PeV photon energy flux would also exceed $1 \text{ eV cm}^{-2} \text{ s}^{-1}$ at Earth. For sources that produce neutrons by hadronic processes as well, the neutron energy flux would be even higher since the neutron production rate exceeds the photon production rate, as noted previously.

Free neutrons undergo beta decay with a $880.0 \pm 0.9 \text{ s}$ half-life (Olive et al. 2014). Due to this decay, sources will only be visible within about $10 (E/\text{PeV}) \text{ pc}$ of Earth. Since plausible accelerators such as young pulsars are no closer than 100 pc , searches at energies above 10 PeV are the most promising.

A diffuse flux of neutrons could be expected from interactions of cosmic-ray primaries with ambient photons and the interstellar medium. However, at PeV energies this flux would appear all over the sky since the effective range is less than the thickness of the Galactic disk. This complicates a search for correlations with the Galactic plane since an excess signal could not be constrained to a particular region of the sky, for example, Galactic latitudes $|b| < 10^\circ$.

At energies above 10^{18} eV (1 EeV), the Pierre Auger Observatory recently performed a search for neutrons in the southern hemisphere, finding no significant signal excesses or correlations with catalogs of Galactic objects, and established flux upper limits (Aab et al. 2012, 2014). The Telescope Array experiment has established flux limits for point sources above 0.5 EeV in the northern hemisphere (Abbasi et al. 2015). KASCADE (Antoni et al. 2004) and CASA-MIA (Chantell et al. 1997; Borione et al. 1998) found no point sources in the northern hemisphere, also setting flux limits (an all-sky limit in the case of KASCADE). AGASA (Hayashida et al. 1999) and a reanalysis (Bellido et al. 2001) of SUGAR data reported slight excesses toward the Galactic center, although these were later not confirmed by Auger (Aab et al. 2015).

This paper reports the results of two searches for point-like signals in the arrival direction distribution of 4 yr of IceTop data. The two searches are an all-sky search for general hot spots on the sky and a search for correlations with nearby known Galactic sources. In the all-sky search, we look for an excess of events from any direction in the sky, evaluating the significance of any excess using the method of Li & Ma (1983). The observable signature of a neutron flux is an excess of proton-like air showers. The targeted search is treated as a stacked analysis using a set of candidate sources from an astrophysical catalog. It is assumed that many or all of the candidates for a given set are emitting neutrons, so the combined signal should be more significant than that of a single target. In both the all-sky and targeted searches, we set flux upper limits using the procedures of Feldman & Cousins (1998).

This paper is organized as follows. In Section 2, the IceTop detector is described. Section 3 summarizes the reconstruction methods and characteristics of the data set. The analysis methods and details of the search methods are described in Section 4. The search results are presented in Section 5. A discussion of the results (Section 6) concludes the paper.

2. ICECUBE/ICETOP

IceTop is the surface air-shower array of the IceCube Neutrino Observatory at the geographical South Pole located 2835 m above sea level (Abbasi et al. 2013). Its final configuration consists of 81 stations covering 1 km^2 with an

Table 1

Detector Configurations and Their Respective Number of Events and Effective Snow Attenuation Lengths for All Years Used in This Analysis

Configuration	Live Time (days)	Number of Events $N_{>10 \text{ PeV}}$ ($N_{>100 \text{ PeV}}$)	Snow Depth (m)
IC79	327.3	291,738 (2986)	2.1
IC86-1	342.0	305,138 (3173)	2.25
IC86-2	332.3	306,868 (3025)	2.25
IC86-3	362.2	329,743 (3374)	2.3
Total	1368.8	1,233,487 (12,558)	...

average station separation of 125 m . Detector construction started in 2005 and finished in 2010. A single station consists of two light-tight tanks separated by 10 m . Each tank is 1.8 m in diameter, 1.3 m in height, and filled with transparent ice to a height of 0.9 m . A tank contains two optical sensors, each consisting of a 10-inch Hamamatsu photomultiplier tube together with electronic boards for detection, digitization, and readout (Abbasi et al. 2009, 2010). The two sensors are operated at different gains for increased dynamic range. The IceTop trigger condition requires at least three stations to have recorded hits within a $5 \mu\text{s}$ time window (Abbasi et al. 2013). IceTop detects showers at a rate of approximately 30 Hz with a minimum primary particle energy threshold of about 400 TeV . Its surface location near the shower maximum makes it sensitive to the full electromagnetic component of the shower in addition to the muonic component.

Cosmic-ray reconstruction relies on the optical detection of Cerenkov radiation within tanks of ice emitted by secondary particles produced by cosmic-ray interactions in the upper atmosphere. Information from individual tanks, including position, deposited charge, and pulse timing, is used to infer the air-shower direction, core location, and shower size estimate S_{125} , which is related to the cosmic-ray primary energy (Aartsen et al. 2013a).

Snow accumulates on the top of stations with time, attenuating the electromagnetic portion of the shower, lowering S_{125} . This accumulation occurs in a nonuniform way due to wind patterns around nearby structures. Snow depth measurements for each tank are performed twice a year, allowing for depth interpolation at the time of an event. An exponential correction factor is applied during event reconstruction to the signal of each tank such that the corrected tank signal $S_{125} = S_{125}^{\text{snow}} \exp(x/\lambda_{\text{eff}})$. Here, S_{125}^{snow} is the detected signal in the tank, x is the slant depth through the snow above the tank, and λ_{eff} is the effective attenuation length due to the snow. Values for λ_{eff} are selected such that the resulting S_{125} distributions for each year are consistent. The attenuation length changes over time as the snow depth generally increases across the entire array (Rawlins et al. 2015a).

3. RECONSTRUCTION METHODS AND DATA SET

This analysis uses 4 yr of IceTop experimental data collected between 2010 May and 2014 May. For the first year of data (IC79), 73 stations were deployed; for each of the remaining 3 yr (IC86), IceTop operated in its final 81-station configuration.

Event reconstructions are performed using the standard IceTop reconstruction method (Abbasi et al. 2013). The values for the snow attenuation length λ_{eff} differ for each year and are

listed in Table 1. The shower core location on the ground is determined by a signal-weighted likelihood fit to the shower front, with a typical resolution better than 10 m at the highest energies. The primary arrival direction is determined from a fit to the arrival time distributions of signals in the tanks. The angular resolution is the space angle that includes 68% of reconstructed events that would arrive from a fixed direction. This value varies between 0.2 and 0.8 depending on energy and primary mass (Rawlins et al. 2015b). Above 10 PeV, the typical angular resolution, defined as the angle from the true event direction that contains 68% of reconstructed event directions, is better than 0.5, which is taken as the representative value in the analysis.

The shower size estimate S_{125} is determined by fitting the tank signals for the expected signal at 125 m from the shower core location. The relationship between S_{125} and primary cosmic-ray energy is determined by comparison with Monte Carlo simulations for zenith angles less than 37° (Rawlins et al. 2015b). The energy resolution above 2 PeV is better than 0.1 in \log_{10} of the energy (Abbasi et al. 2013).

Events are selected by requiring a good fit to the shower lateral distribution, a reconstructed core location lying within 400 m of the array center (not near the array boundary), and a cut on zenith angle within 37°. Requiring the reconstructed cores within 400 m yields a fiducial area $A = 5.02 \times 10^5 \text{ m}^2$. For the final event selection for the all-sky search, we select energies above 10 PeV, and 100 PeV for the targeted search, resulting in 1,233,487 and 12,558 events, respectively. The total live time is 1363.8 days. Table 1 lists the live time, number of events for each energy threshold, and effective snow attenuation length for each year.

The targeted search uses a higher-energy cut since most astrophysical objects of interest for this search lie at Galactic distances of order 1 kpc or greater. This cut is also motivated by the fact the lower-energy neutrons will not typically survive from 1 kpc and that lower energy contains only background contributions.

4. SEARCH METHODS

For both search methods, top-hat search windows are drawn on the sky. This procedure allows for selecting events using a hard cut on the space angle between the event direction and the window center. The locations of these search windows are described in the following sections with more detailed information about the two searches. The radius of the search window in both searches is based on the actual IceTop point-spread function and is chosen such that it optimizes the sensitivity to a point source. Point-source sensitivity is optimized by choosing a window size χ based on the angular resolution. The point-spread function is taken to be $p(\theta) = (\theta/\sigma^2) \exp(-\theta^2/2\sigma^2)$, where $\sigma = \psi/1.51$. Here, θ is the space angle between the reconstructed and true arrival directions and ψ is the angular resolution. Using top-hat search windows, the sensitivity is optimized with $\chi = 1.59\sigma = 1.05\psi$, or 0.52.

To find a signal excess within a search window, one must first know the expected number of events without signal, i.e., the background expectation value. The background value for each search window is determined by time-scrambling the data set many times. Each time-scrambled set has the same number of events as the data set. For each event, we keep its zenith and azimuth angles in detector coordinates and randomly select

another time in the data set within a 24 hr window centered on the time of the event. The search window content of the background expectation map is taken as the mean content of 10^3 and 10^6 time-scrambled maps for the all-sky and targeted searches, respectively.

4.1. All-sky Search

In the all-sky search, we look for excesses within search windows located in all parts of the sky within the field of view of IceTop. These windows are centered on the pixels of a high-resolution HEALPix (Górski et al. 2004) map. N_{side} is a parameter used to define and generate the map's pixels, with higher values generating higher-resolution maps. We select a map defined by $N_{\text{side}} = 128$, which provides 19,800 points within the IceTop field of view and simply provides central locations from which to draw the search windows. The typical spacing between adjacent window locations in this map is 0.46. Although window overlap will cause correlations between neighboring windows, this ensures that all events are counted. The data are first binned using a HEALPix map ("bin map") with higher resolution ($N_{\text{side}} = 256$) than the search window map. The content of a given search window is the sum of contents of those pixels in the bin map whose centers fall within the search window. The summed content of a search window is labeled n (n_b) for the data set (background).

Statistical significance of signals within search windows is based on the observed number of events n and the background expectation value n_b . The significance value of a given search window is calculated using the Li–Ma method (Li & Ma 1983) shown in Equation (1),

$$S = \frac{n - n_b}{|n - n_b|} \sqrt{2} \left(n \ln \left(\frac{n + \alpha n}{n_b + \alpha n} \right) + \frac{n_b}{\alpha} \ln \left(\frac{n_b + \alpha n_b}{n_b + \alpha n} \right) \right)^{1/2}, \quad (1)$$

where we have replaced the Li–Ma parameters N_{on} and N_{off} with n and n_b/α , respectively. The Li–Ma method is used only for the all-sky search. Typically α is the ratio of time spent observing on-source to the time spent observing an equivalent off-source solid angle. Here, the parameter α is taken to be the ratio n_b/ξ , where ξ is the sum of the contents of all search windows lying within $\pm 90^\circ$ in R.A. and ± 0.52 in decl. of the search window of interest, excluding the content value n_b of the search window itself. This definition of α provides a local estimate of N_{off} for each search window. IceTop observes large-scale anisotropy in cosmic-ray arrival directions for energies above roughly 1 PeV (Aartsen et al. 2013b); for example, a large deficit in the cosmic-ray arrival direction distribution is observed from 30° to 120° in R.A. The estimate of N_{off} should be representative of the expected cosmic-ray flux in the vicinity of the search window, so this definition for α eliminates bias due to averaging over the field of view.

4.2. Targeted Search

The targeted search is performed to look for correlations of event directions with known nearby Galactic objects. We calculate the Poisson probability $p(n, n_b)$ for observing n or more events within the search window expecting n_b for each object. Fisher's method (Fisher 1925) combines a set of independent probabilities to determine a single measure of

significance P_F for the set. For a sequence of p -values p_1, p_2, \dots, p_n , their product is $\pi = \prod_{i=1}^n p_i$. Fisher’s method allows us to calculate the chance probability that a product π of n p -values obtained uniformly randomly would be less than or equal to the product π_{obs} of the n p -values observed: $P_F(\pi \leq \pi_{\text{obs}})$.

A supplemental measure of significance P_G is provided by Good’s method (Good et al. 1955), which allows for weights to be assigned to each probability. In a similar way to Fisher’s method, for a sequence of p -values p_i with weight w_i , the weighted product $\pi_w = \prod_{i=1}^n p_i^{w_i}$. Good’s method allows us to calculate the chance probability that a product π_w of n p -values obtained uniformly randomly with weights w_i would be less than or equal to the product $\pi_{w,\text{obs}}$ of the n p -values observed: $P_G(\pi_w \leq \pi_{w,\text{obs}})$. Here, these weights are proportional to the object’s recorded electromagnetic flux listed in the catalog, its relative exposure to IceTop, and an expected flux attenuation factor. This factor is equal to the survival probability for a neutron with energy equal to the median energy of an E^{-2} energy spectrum between 100 PeV and 1 EeV to arrive from the distance of a candidate source object. The weights are normalized such that their sum is 1 for each target set.

Treating the unweighted and weighted probabilities (P_F, P_G) as individual test statistics, we calculate the fraction of time-scrambled data sets with corresponding values less than that observed with the data. This post-trials fraction is an unbiased indicator of the correlation probability between the data set and each source set. Both the weighted and unweighted probabilities and corresponding post-trials fractions are reported. The unweighted probability is independent of the assumption that neutron emission is proportional to the electromagnetic emission and of how the flux, relative exposure, and decay probability are used to construct the object weight.

4.3. Target Catalogs

We consider three distinct classes: millisecond pulsars (Manchester et al. 2005) (msec), γ -ray pulsars (Abdo et al. 2013) (γ -ray), and high-mass X-ray binaries (HMXB) (Liu et al. 2007). The msec catalog⁵⁵ provides a comprehensive list of rotation-powered pulsars. The γ -ray catalog is the second Fermi-LAT pulsar catalog. The HMXB catalog⁵⁶ represents a comprehensive selection of X-ray sources, composed of a compact object orbiting a massive OB class star. These classes are considered candidate sources due to their independent evidence for high-energy particle production and high flux measured at Earth. The Galactic center lies outside the zenith angle cut, lies well beyond the effective neutron range even at energies of a few hundred PeV, and is not considered a candidate in this search.

Only objects with known distances are included in the final catalog selection. Distances for each candidate are cross-checked with the TeVCat catalog.⁵⁷ Most objects are eliminated from each catalog by the zenith angle cut and by requiring that the distance is known. Sources that further appear in multiple sets are retained only in the smaller set, resulting in 17 objects in the γ -ray set, 16 objects in the msec set, and 20 objects in the HMXB set as shown in Tables 2–4, respectively. The columns in each table are the object

designation, R.A., decl., distance, electromagnetic flux as recorded in the catalog, relative exposure value to IceTop, survival probability for a neutron with energy equal to the median energy of an E^{-2} energy spectrum between 100 PeV and 1 EeV, and normalized weight value. Figure 1 shows the locations of each object in equatorial coordinates. The Galactic plane is depicted by a green band to illustrate the preferential association of the γ -ray pulsar and HMXB sets with that part of the sky.

4.4. Flux Upper Limit Calculation

Flux upper limits are calculated for both the all-sky and targeted searches using

$$F_{\text{UL}} = 1.39 s_{\text{UL}}/\zeta, \quad (2)$$

where s_{UL} is the upper limit on the number of signal events in the search window and $\zeta = TA\cos(\theta)\epsilon$ is the exposure of IceTop, where T is the live time, $A\cos(\theta)$ is the projected detector area exposed to the search window, which depends on the zenith angle θ , and ϵ is the reconstruction efficiency (taken as 95% according to Monte Carlo studies). The signal upper limit s_{UL} is calculated using a 90% Feldman–Cousins confidence level (Feldman & Cousins 1998) based on n and n_b for the search window. The factor 1.39 is a compensation factor to include signal events that fall outside the search window. The search window includes only 71.8% of signal events based on the top-hat window and the assumed IceTop point-spread function; therefore, s_{UL} is scaled by $1/0.718 = 1.39$.

The flux upper limit can be rewritten as

$$F_{\text{UL}} = 0.776 (s_{\text{UL}}/\cos(\theta)) (\text{km}^{-2} \text{yr}^{-1}), \quad (3)$$

by substituting $TA\epsilon = 1.79 \text{ km}^2 \text{ years}$. For an assumed E^{-2} energy spectrum over the 100 PeV–1 EeV energy decade, the median energy is 181.8 PeV. The median energy flux upper limit in ($\text{eV cm}^{-2} \text{s}^{-1}$) over this energy range can be written as

$$F_{\text{UL}}^E = 0.447 (s_{\text{UL}}/\cos(\theta)). \quad (4)$$

Over the 10 PeV–1 EeV energy decades, the median energy is 19.80 PeV, so the conversion factor between the particle flux and median energy flux upper limits is

$$1 \text{ part. km}^{-2} \text{yr}^{-1} = 0.0628 \text{ eV cm}^{-2} \text{s}^{-1}. \quad (5)$$

An important point to note is that Equation (4) assumes an E^{-2} energy spectrum *as measured at Earth*, which is related to the source energy spectrum only after accounting for neutron decay factors that depend on the source distance. Figure 2 shows the attenuation factor of the energy spectrum injected at the source due to decay during propagation for representative distances. For a sufficiently distant source, the source spectrum would be harder than that observed at Earth. The lower-energy portions of the spectrum are increasingly suppressed with distance as these neutrons are removed.

The attenuation curves in Figure 2 have a strong effect on the sensitivity of the searches. The all-sky search uses a 10 PeV energy threshold; thus, it is sensitive only to sources at extremely close distances due to the large number of events lying near threshold. The targeted search is sensitive mostly to higher neutron energies closer to EeV energies, which are capable of crossing larger Galactic distances. For example, a

⁵⁵ <http://www.atnf.csiro.au/research/pulsar/prscat>.

⁵⁶ <http://heasarc.gsfc.nasa.gov/w3browse/all/hmxbcat.html>.

⁵⁷ <http://tevcat.uchicago.edu>

Table 2
Characteristics of the Fermi γ -ray Catalog

Object Name	R.A. (deg)	Decl. (deg)	Distance (kpc)	Energy Flux in the Range 0.1–100 GeV ($\text{erg cm}^{-2} \text{s}^{-1}$)	Relative Exposure	Survival Probability ^a	Normed Weight
J0101-6422	15.30	−64.38	0.55	1.047e-11	0.902	0.72	0.026
J1016-5857	154.09	−58.95	2.9	5.444e-11	0.857	0.18	0.032
J1028-5819	157.12	−58.32	2.33	2.426e-10	0.851	0.248	0.199
J1048-5832	162.05	−58.53	2.74	1.958e-10	0.853	0.194	0.126
J1105-6107	166.36	−61.13	4.98	4.89e-11	0.876	0.0509	0.008
J1112-6103	168.06	−61.06	12.2	2.034e-11	0.875	<0.001	<0.001
J1119-6127	169.81	−61.46	8.4	7.148e-11	0.879	0.0066	0.002
J1124-5916	171.16	−59.27	4.8	6.168e-11	0.860	0.057	0.012
J1125-5825	171.43	−58.42	2.62	8.9e-12	0.852	0.209	0.006
J1357-6429	209.26	−64.49	2.5	3.388e-11	0.903	0.22	0.027
J1410-6132	212.59	−61.53	15.6	2.63e-11	0.879	<0.001	<0.001
J1418-6058	214.68	−60.97	1.6	3.017e-10	0.874	0.38	0.39
J1420-6048	215.03	−60.80	5.61	1.698e-10	0.873	0.035	0.020
J1509-5850	227.36	−58.85	2.62	1.273e-10	0.856	0.209	0.088
J1513-5908	228.48	−59.14	4.21	3.243e-11	0.858	0.0807	0.009
J1531-5610	232.87	−56.18	2.09	1.94e-12	0.831	0.287	0.002
J1658-5324	254.66	−53.40	0.93	2.893e-11	0.803	0.57	0.052

Note.

^a Calculated using the median energy of an E^{-2} spectrum between 100 PeV and 1 EeV.

Table 3
Characteristics of the Msec Catalog

Object Name	R.A. (deg)	Decl. (deg)	Distance (kpc)	Energy Flux at Sun ($\text{erg kpc}^{-2} \text{s}^{-1}$)	Relative Exposure	Survival Probability ^a	Normed Weight
J1017-7156	154.46	−71.94	0.26	1e+35	0.951	0.86	0.5
B0021-72F	6.02	−72.08	4	8.8e+33	0.951	0.09	0.005
J1125-6014	171.48	−60.24	1.94	2.3e+33	0.868	0.314	0.004
J1910-5959A	287.93	−59.97	4.5	1.6e+32	0.866	0.068	<0.001
J1103-5403	165.89	−54.06	3.16	3.7e+32	0.809	0.151	<0.001
J1216-6410	184.03	−64.17	1.71	4.9e+32	0.900	0.360	0.001
J1933-6211	293.39	−62.20	0.63	8.3e+33	0.885	0.69	0.032
J1740-5340A	265.19	−53.68	3.4	1.2e+34	0.806	0.13	0.008
J2129-5721	322.34	−57.35	0.4	9.9e+34	0.842	0.8	0.4
J1431-5740	217.76	−57.67	4.07	2.2e+32	0.845	0.0877	<0.001
J0711-6830	107.98	−68.51	1.04	3.3e+33	0.931	0.537	0.010
J1629-6902	247.29	−69.05	1.36	9.9e+32	0.934	0.443	0.003
J2236-5527	339.22	−55.46	2.03	2.8e+32	0.824	0.297	<0.001
J1757-5322	269.31	−53.37	1.36	8e+32	0.803	0.443	0.002
J1435-6100	218.83	−61.02	3.25	1.1e+32	0.875	0.143	<0.001
J1337-6423	204.38	−64.38	6.3	2.9e+31	0.902	0.023	<0.001

Note.

^a Calculated using the median energy of an E^{-2} spectrum between 100 PeV and 1 EeV.

suppression factor S can be defined as the ratio between the number of neutrons with an injected E^{-2} spectrum observed after including attenuation and the number observed not including attenuation for the same E^{-2} spectrum. For an E^{-2} spectrum between 10 PeV and 1 EeV, removal of half the neutrons from the observed spectrum, or $S = 0.5$, corresponds to a propagation distance of about 0.15 kpc. Between 100 PeV and 1 EeV, $S = 0.5$ corresponds to a distance of about 1.25 kpc. Generally speaking, the sensitivity of any neutron search will be shifted toward the higher-energy portion of the injected energy spectrum at the source due to decay, unless sources are sufficiently close that decay does not significantly modify the energy spectrum. This can be seen in Figure 3, which shows an example E^{-2} energy spectrum modified by the distance-dependent decay attenuation.

These flux limits are time-averaged values based on the IceTop exposure ζ . Particularly for the objects in the targeted source sets, it is possible that transient fluxes may temporarily exceed these limits. The energy flux limits derived from Equation (4) are strongly dependent on the assumption that an injected E^{-2} energy spectrum at the source is not strongly modified in the energy range the limit applies to by neutron decay en route.

5. RESULTS

5.1. All-sky Search

Figures 4 and 5 show the differential and cumulative distributions of the 19,800 Li–Ma values compared to the isotropic expectation. In both figures, the blue and green lines

Table 4
Characteristics of the HMXB Catalog

Object Name	R.A. (deg)	Decl. (deg)	Distance (kpc)	Energy Flux in the Range 2–10 keV (μJy)	Relative Exposure	Survival Probability ^a	Normed Weight
1H 0739-529	116.85	-53.33	0.52	0.7	0.802	0.73	0.007
1H 0749-600	117.57	-61.10	0.4	0.7	0.875	0.8	0.008
GRO J1008-57	152.44	-58.29	5	1200	0.851	0.05	0.9
RX J1037.5-5647	159.40	-56.80	5	3.3	0.837	0.05	0.002
1A 1118-615	170.24	-61.92	5	0.1	0.882	0.05	<0.001
4U 1119-603	170.31	-60.62	9	10	0.871	0.005	<0.001
IGR J11215-5952	170.44	-59.86	8	42	0.865	0.008	0.005
2S 1145-619	177.00	-62.21	2.3	4	0.885	0.25	0.02
1E 1145.1-6141	176.87	-61.95	8	4	0.883	0.008	<0.001
4U 1223-624	186.66	-62.77	3	9	0.889	0.2	0.02
1H 1249-637	190.71	-63.06	0.3	2.2	0.891	0.8	0.03
1H 1253-761	189.81	-75.37	0.24	0.6	0.968	0.87	0.008
1H 1255-567	193.65	-57.17	0.11	0.8	0.840	0.94	0.01
4U 1258-61	195.32	-61.60	2.4	0.3	0.880	0.24	0.001
2RXP J130159.6-635806	195.50	-63.97	5.5	6.3	0.890	0.037	0.003
SAX J1324.4-6200	201.11	-62.01	3.4	0.4	0.883	0.1	<0.001
2S 1417-624	215.30	-62.70	6	2	0.889	0.03	<0.001
SAX J1452.8-5949	223.21	-59.82	9	0.045	0.864	0.005	<0.001
XTE J1543-568	226.00	-56.77	10	8	0.836	0.003	<0.001
1H 1555-552	238.59	-55.33	0.96	1.7	0.822	0.56	0.013

Note.

^a Calculated using the median energy of an E^{-2} spectrum between 100 PeV and 1 EeV.

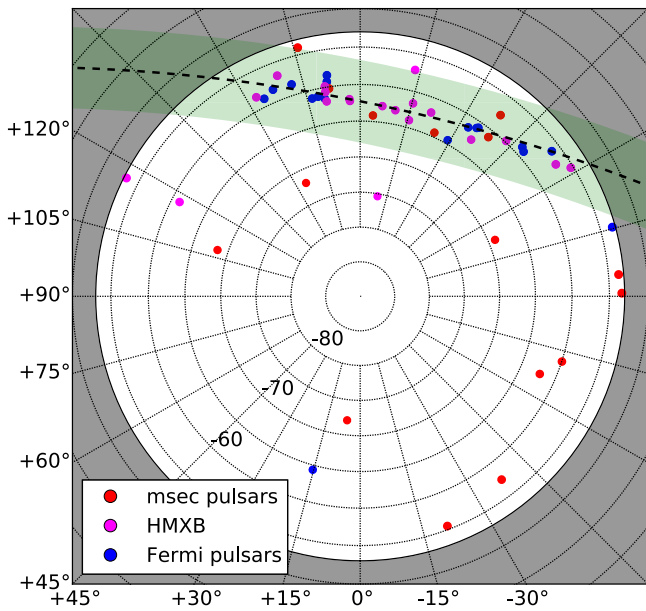


Figure 1. Equatorial polar sky map of each catalog set. The dashed black line indicates the Galactic plane, and the green band shows $b = \pm 5^\circ$. Each circle is 0.5° in radius.

show the Li–Ma significance distribution for the data and isotropy, respectively. There are no Li–Ma values larger than 4. The dashed line shows the Gaussian form expected for the distribution to follow if deviations from isotropy are due only to statistical fluctuations. In Figure 5, the gray shaded region in the cumulative plot shows the 95% containment band for isotropy; the presence of search windows with statistically significant signal excess would extend above and to the right of this band. The absence of such a feature indicates that no statistically significant signal excess is observed and that the

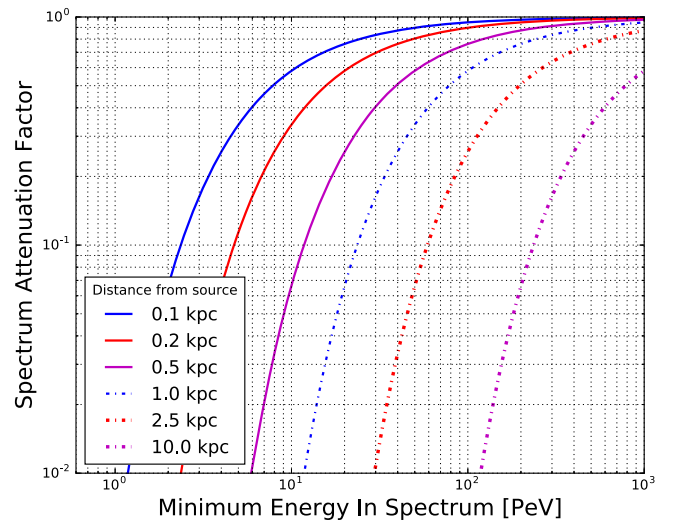


Figure 2. Spectrum attenuation factor due to neutron decay as a function of minimum energy for an E^{-2} spectrum and distance from the source. The attenuation factor is a function of the median energy, which itself depends on the minimum energy.

observed excesses are consistent with fluctuations about the expectation.

Figures 6 and 7 show sky maps of the Li–Ma and flux upper limit values for each search window. No statistically significant clustering on the sky is observed, including the Galactic plane depicted by the black dashed ($b = 0^\circ$) and solid ($b = \pm 5^\circ$) lines. As noted previously, the energies of most events used in this search lie close to the 10 PeV energy cut, which corresponds to a neutron range of order 100 pc. The sphere from which signal could arrive is contained within the Galactic disk, so any excesses arising from cosmic-ray interactions in the disk would be distributed over the entire field of view, not concentrated within a narrow band across the sky.

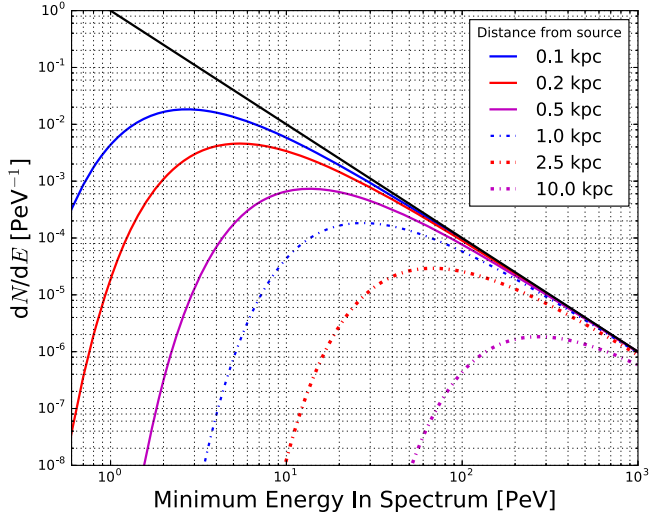


Figure 3. Effect of decay attenuation on an E^{-2} energy spectrum at source as a function of minimum energy and distance from the source.

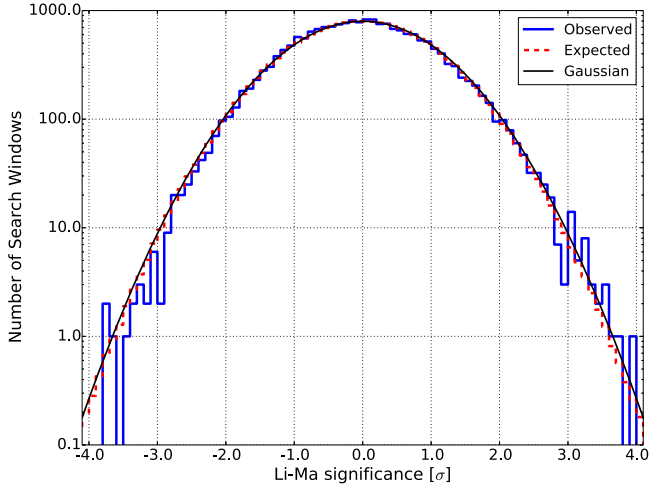


Figure 4. Differential histograms of Li–Ma values (blue) and the isotropic expectation (green). The dashed line shows the Gaussian approximation for the expected Li–Ma distribution in the case in which deviations result only from statistical fluctuations.

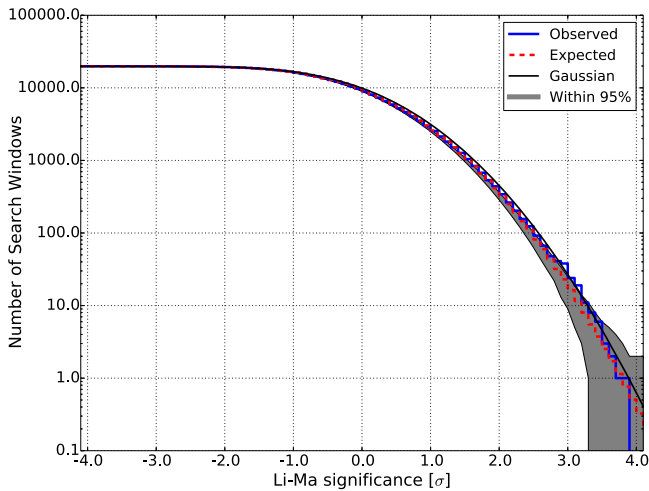


Figure 5. Cumulative histograms of Li–Ma values (blue) and the isotropic expectation (green).

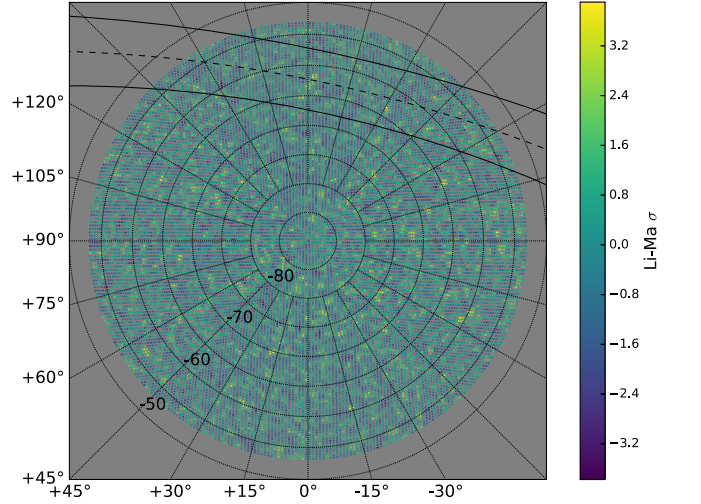


Figure 6. Equatorial polar sky map of Li–Ma values for each search window. The dashed black line indicates the Galactic plane, and the solid black lines depict $b = \pm 5^\circ$.

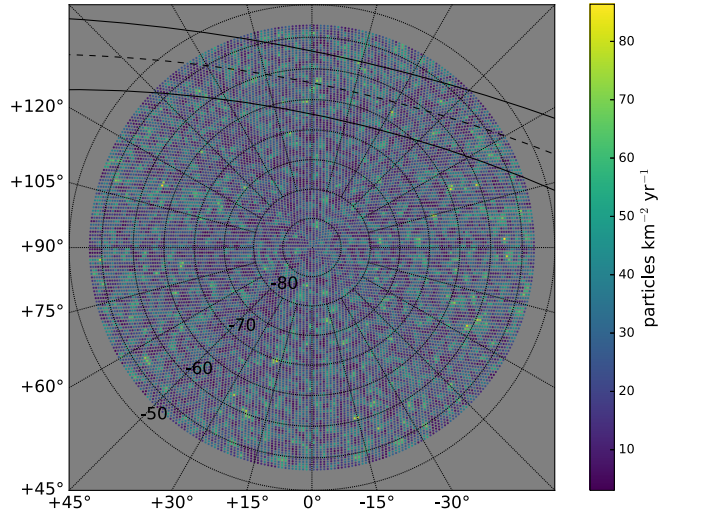


Figure 7. Equatorial polar sky map of flux upper limit values for each search window. The dashed black line indicates the Galactic plane, and the solid black lines depict $b = \pm 5^\circ$.

Figure 8 shows the mean flux upper limit as a function of declination for the all-sky search. The limits are strongest near the South Pole due to the maximal exposure, but there is greater uncertainty on the mean since there are fewer search windows in declination bands closest to the pole.

5.2. Targeted Search

Table 5 lists the correlation probabilities for each catalog, with the corresponding post-trials probability in parentheses. No significant correlation is observed with any catalog. Tables 6–8 give details of each object. The columns in each table are the object designation, observed number of events within the search window, background estimate in the window, particle flux above 100 PeV according to Equation (3), energy flux above 100 PeV according to Equation (4), and Poisson probability $p(n, n_b)$ for observing n events with an expectation number n_b . These flux limits assume an E^{-2} energy spectrum as measured at Earth. The most significant object in each catalog is highlighted in bold.

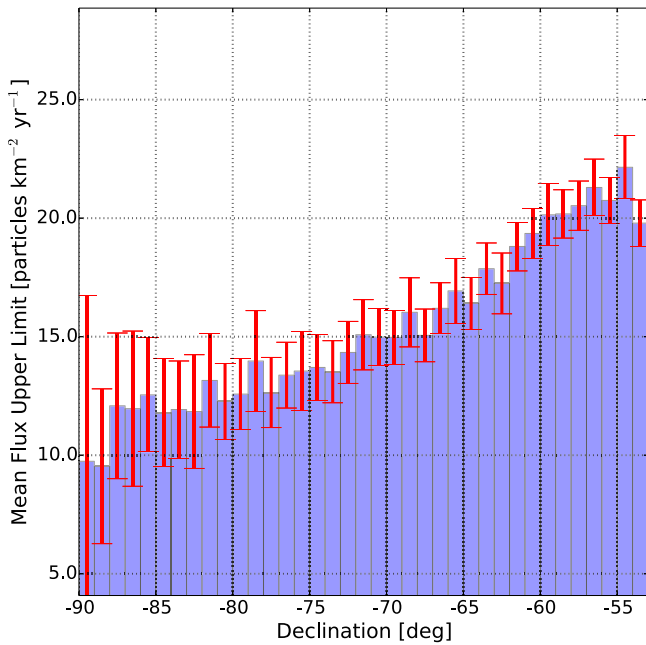


Figure 8. Mean flux upper limit (90% C.L.) for 1° declination bins for the all-sky search. The error bars indicate the statistical uncertainty on the mean value since there are many search windows within each declination band.

Table 5
Targeted Search Results with Each Catalog

Catalog	Unweighted P_F	Weighted P_G
γ -ray	0.999 (0.976)	0.910 (0.776)
msec	0.809 (0.408)	0.888 (0.778)
HMXB	0.999 (0.988)	0.946 (0.971)

Note. Values in parentheses give the post-trials probability.

The catalog probabilities do not appear to be distributed uniformly between 0 and 1, since at least one probability value would be expected to lie below 0.809 in over 99.5% of sets of six uniformly distributed random samples. There exists an underfluctuation in the data along $b = 0^\circ$ compared to the

background expectation, as illustrated in Figure 9. The preferential clustering of the γ -ray pulsar and HMXB catalogs along the Galactic plane, combined with this underfluctuation, acts to drive the catalog probabilities to higher values. Since typically $n < n_b$, the individual Poisson p -values are close to 1. This is checked by rotating these catalogs by a prescribed amount in R.A. and expecting lower catalog probabilities due to higher n and similar n_b in the windows. These values are shown in Table 9 for different rotation values.

We also note that there are four pairs of objects that lie within 1° of each other. In all four cases, the objects in each pair are distinct from each other, lie at different distances, and are from different catalogs. We find consistent results with Table 5 when we mask the object with the farther distance.

6. SUMMARY AND DISCUSSION

IceTop does not observe a statistically significant point source of cosmic-ray arrival directions. Using Equation (5), the all-sky mean flux upper limits for individual declination bands correspond to energy fluxes between about 0.6 and $1.2 \text{ eV cm}^{-2} \text{ s}^{-1}$ between 100 PeV and 1 EeV assuming an E^{-2} neutron energy spectrum as measured at Earth, which are comparable to TeV photon fluxes for Galactic objects (Hinton & Hofmann 2009). These flux limits are the first neutron flux upper limits in the southern hemisphere for energies in the 10 PeV to 1 EeV energy decades. Again, it is important to note that neutron decay en route will modify the energy spectrum as illustrated in Figure 2, so the source spectrum would be generally softer than that constrained. The limits in both searches are strongly dependent on the assumption that an injected E^{-2} spectrum is not significantly modified by decay, as noted in Section 4.4. For the all-sky search, this restricts the applicability of the limits within a small volume around Earth. For the targeted search, there are a number of objects that lie within 1 kpc, so their limits are most compatible with the base assumption.

As noted previously, hadronic production of photons by protons with an E^{-2} spectrum will inject equal power into each energy decade, and the neutron production at least equals the photon production. At present, these flux upper limits do not

Table 6
Targeted Search Results for the Fermi γ -Ray Catalog

Object Name	Observed n	Background Estimate n_b	F_{UL} ($\text{km}^{-2} \text{ yr}^{-1}$)	F_{UL}^E ($\text{eV cm}^{-2} \text{ s}^{-1}$)	Poisson Probability p
J1016-5857	3	2.62	4.35	2.51	0.487
J1028-5819	1	1.80	2.44	1.41	0.835
J1048-5832	5	2.77	6.57	3.79	0.147
J1105-6107	2	3.79	2.19	1.26	0.892
J1112-6103	3	3.79	3.29	1.90	0.729
J1125-5825	2	2.65	3.02	1.74	0.742
J1124-5916	2	1.73	3.78	2.18	0.517
J1119-6127	3	2.71	4.16	2.40	0.508
J0101-6422	3	2.81	3.98	2.29	0.534
J1357-6429	2	2.34	3.09	1.78	0.679
J1410-6132	1	2.75	1.79	1.03	0.936
J1418-6058	2	2.81	2.83	1.63	0.770
J1420-6048	2	2.62	2.98	1.72	0.737
J1509-5850	0	2.37	0.83	0.48	1.000
J1513-5908	0	1.81	1.07	0.62	1.000
J1531-5610	0	2.78	0.68	0.39	1.000
J1658-5324	1	2.59	2.06	1.19	0.925

Table 7
Targeted Search Results for the Msec Pulsar Catalog

Object Name	Observed n	Background Estimate n_b	F_{UL} ($\text{km}^{-2} \text{yr}^{-1}$)	F_{UL}^E ($\text{eV cm}^{-2} \text{s}^{-1}$)	Poisson Probability p
J0711-6830	2	2.54	2.85	1.64	0.720
J1103-5403	2	1.99	3.77	2.17	0.591
J1017-7156	2	2.23	3.01	1.74	0.654
J1125-6014	2	1.80	3.68	2.12	0.537
J1216-6410	3	2.67	4.10	2.36	0.499
B0021-72F	4	1.95	5.42	3.12	0.133
J1337-6423	5	3.02	6.00	3.46	0.188
J1435-6100	1	2.49	1.94	1.12	0.917
J1431-5740	3	1.84	5.13	2.96	0.281
J1629-6902	0	2.89	0.57	0.33	1.000
J2236-5527	4	2.72	5.54	3.19	0.289
J1933-6211	6	3.20	7.26	4.18	0.106
J1910-5959A	2	3.21	2.58	1.49	0.830
J2129-5721	1	2.48	2.04	1.18	0.916
J1740-5340A	1	2.52	2.10	1.21	0.919
J1757-5322	3	2.18	5.08	2.93	0.372

Table 8
Targeted Search Results for the HMXB Catalog

Object Name	Observed n	Background Estimate n_b	F_{UL} ($\text{km}^{-2} \text{yr}^{-1}$)	F_{UL}^E ($\text{eV cm}^{-2} \text{s}^{-1}$)	Poisson Probability p
1H 0739-529	0	2.51	0.82	0.47	1.000
1H 0749-600	1	2.44	1.98	1.14	0.913
GRO J1008-57	1	2.82	1.82	1.05	0.941
RX J1037.5-5647	2	3.22	2.66	1.53	0.832
IGR J11215-5952	2	1.95	3.56	2.05	0.579
4U 1119-603	2	1.83	3.63	2.09	0.546
1A 1118-615	1	2.08	2.18	1.26	0.876
1E 1145.1-6141	3	2.64	4.21	2.43	0.492
2S 1145-619	3	2.31	4.49	2.59	0.408
4U 1223-624	1	3.39	1.45	0.84	0.966
1H 1249-637	1	2.09	2.14	1.23	0.877
1H 1253-761	2	3.41	2.19	1.26	0.855
1H 1255-567	2	2.01	3.61	2.08	0.598
2RXP J130159.6-635806	2	2.87	2.71	1.56	0.781
4U 1258-61	1	2.93	2.28	1.31	0.947
SAX J1324.4-6200	1	2.49	1.93	1.11	0.917
2S 1417-624	4	2.86	5.01	2.89	0.322
SAX J1452.8-5949	3	1.69	5.15	2.97	0.239
XTE J1543-568	3	2.81	4.29	2.47	0.532
1H 1555-552	0	2.42	8.35	4.81	1.000

strongly constrain the TeV photon production mechanism, or the shape of the parent energy spectrum. No significant correlation is found with known nearby Galactic objects characterized by GeV–TeV energy photon emission and plausibly capable of producing PeV neutrons.

The nonobservation of PeV neutrons may simply indicate that these objects are not producing neutrons at these energies, or that typical Galactic neutron sources are not near Earth. Local PeV neutron production in the Galaxy could simply be episodic or transient, for example, occurring during supernova explosions or other extremely high energy particle production events. Alternatively, the sources may emit particle jets continuously, but their number may be few and the jets are not oriented toward Earth. Individual sources could emit weakly but be densely distributed.

Additionally, the environment around any sources may not be sufficiently dense to facilitate neutron production by cosmic-

ray interaction such that the primaries escape the acceleration region into interstellar space before interacting and producing neutrons. In this case, neutrons decay in interstellar space relatively near the primary source producing secondary protons (Bednarek & Protheroe 1997). These secondary protons then propagate diffusively in the GMF, so sources that are sufficiently far away will not manifest a point-source signal of cosmic-ray neutrons, but could contribute to a proton signal that is smeared on the sky and not necessarily pointing back to the original source; this argument was presented by Bossa et al. (2003) when they considered EeV neutrons from the Galactic center. At PeV energies, neutrons would penetrate much less into the surrounding medium, so any potential signal from the resulting protons would be strongly suppressed by the scattering effects of the GMF and masked by the background cosmic-ray flux.

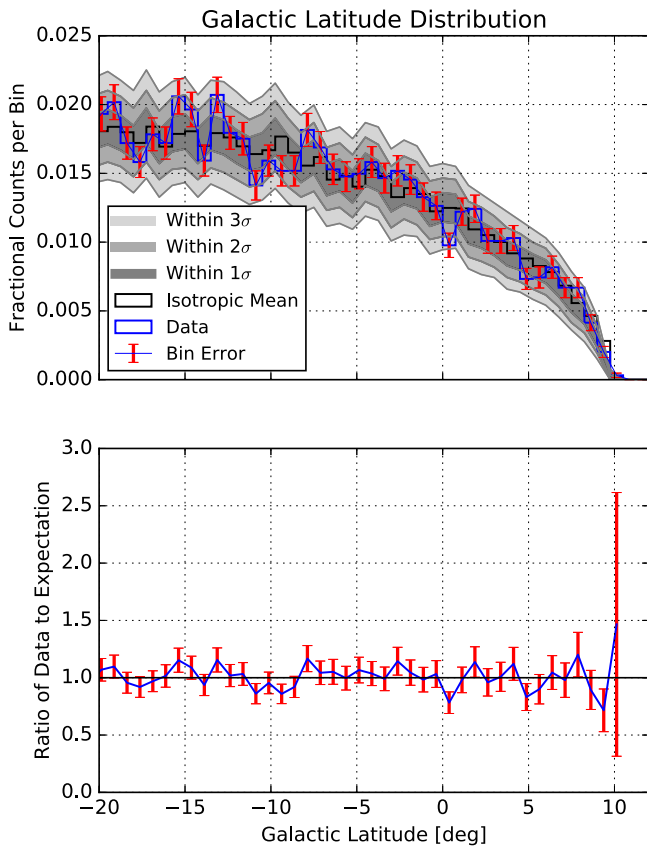


Figure 9. Distribution of number of events above 100 PeV as a function of Galactic latitude. In the top frame, the blue histogram shows the data; the black line shows the isotropic expectation from 10,000 time-scrambled data sets. The red error bars show the Poisson uncertainty in the data histogram. The gray shaded bands depict the 68%, 95%, and 99% containment bands for isotropy in each latitude bin. The bottom frame shows the ratio between the data and isotropic expectation.

Table 9

Results of R.A. Rotation Tests for γ -Ray Pulsar and HMXB Catalogs

Probability	+45°	+90°	+180°
HMXB P_F	0.070 (0.002)	0.991 (0.828)	0.978 (0.741)
HMXB P_G	0.262 (0.129)	0.164 (0.116)	0.946 (0.982)
γ -ray P_F	0.803 (0.408)	0.126 (0.016)	0.991 (0.866)
γ -ray P_G	0.074 (0.025)	0.586 (0.311)	0.845 (0.645)

Note. Values in parentheses give the post-trials probability.

At higher energies, for example, between 10 and 100 PeV, this process could further enrich the cosmic-ray proton fraction above that which is directly accelerated at the source. The knee in the cosmic-ray spectrum is observed around 4 PeV, which is interpreted as an indication of a maximum attainable rigidity of typical Galactic cosmic-ray sources and of associated changes in elemental composition (see, e.g., Hörandel 2005; Blasi 2014). It is plausible that the maximum attainable energy for the proton energy spectrum in nearby sources may not extend well above the knee energy, although for heavier compositions this scales with the nuclear charge Z . Above 10 PeV, the cosmic-ray flux becomes progressively heavier with energy and with a decreasing proton fraction, which is roughly 20% at 10 PeV (Apel et al. 2013; Rawlins et al. 2015b). This suggests that such secondary enrichment may be

unlikely since a recovering proton fraction is not observed at energies between roughly 10 and a few hundred PeV.

The nonobservation of a PeV neutron flux does not necessarily preclude the existence of a PeV photon flux. The neutron energy spectrum at lower energies becomes increasingly modified by decay. PeV photons, on the other hand, have an absorption length considerably larger than the neutron decay distance and will maintain an unmodified energy spectrum that more resembles the injected spectrum at the source. PeV photons could still plausibly be produced by nonhadronic processes, such as inverse-Compton scattering from a high-energy electron population in or near Galactic sources (see, e.g., Schlickeiser 1989; Balbo et al. 2011; Nozawa et al. 2011; Kohri et al. 2012), although there are flux upper limits in the northern (Chantell et al. 1997; Borione et al. 1998; Feng et al. 2015b; Kang et al. 2015a, 2015b) and southern (Aartsen et al. 2013c) hemispheres. These photon limits, except for (Kang et al. 2015a), are for energies of order 1 PeV or below, whereas this analysis is most sensitive at energies above 100 PeV.

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REFERENCES

- Abreu, P., Aglietta, M., Ahlers, M., et al. 2012, *ApJ*, 760, 148
Aab, A., Abreu, P., Aglietta, M., et al. 2014, *ApJL*, 789, L34
Aab, A., Abreu, P., Aglietta, M., et al. 2015, *ApJ*, 804, 15
Aartsen, M. G., Abbasi, R., Abdou, Y., et al. 2013a, *PhRvD*, 88, 042004
Aartsen, M. G., Abbasi, R., Abdou, Y., et al. 2013b, *ApJ*, 765, 55
Aartsen, M. G., Abbasi, R., Abdou, Y., et al. 2013c, *PhRvD*, 87, 062002
Abbasi, R., Abdou, Y., Abu-Zayyad, T., et al. 2010, *Nucl. Instr. Meth., A*, 618, 139
Abbasi, R., Abdou, Y., Ackermann, M., et al. 2013, *Nucl. Instr. Meth., A*, 700, 188
Abbasi, R., Abe, M., Abu-Zayyad, T., et al. 2015, *ApJ*, 804, 133
Abbasi, R., Ackermann, M., Adams, J., et al. 2009, *Nucl. Instr. Meth., A*, 601, 294
Abdo, A., Ajello, M., Allafort, A., et al. 2013, *ApJS*, 208, 17
Anchordoqui, L., Beacom, J., Goldberg, H., Palomares-Ruiz, S., & Weiler, T. 2007, *PhRvD*, 75, 063001

- Antoni, T., Apel, W., Badea, A., et al. 2004, *ApJ*, 608, 865
- Apel, W., Arteaga-Velazquez, J., Bekk, K., et al. 2013, *APh*, 47, 54
- Balbo, M., Walter, R., Ferrigno, C., & Bordas, P. 2011, *A&A*, 527, L4
- Bednarek, W., & Protheroe, R. 1997, *PhRvL*, 79, 2616
- Bellido, J., Clay, R., Dawson, B., & Johnston-Hollitt, M. 2001, *Aph*, 15, 167
- Blasi, P. 2014, *CRPhy*, 15, 329
- Borione, A., Catanese, M., Chantell, M., Covault, C., & Cronin, J. 1998, *ApJ*, 493, 175
- Bossa, M., Mollerach, S., & Roulet, E. 2003, *J. Phys. G: Nucl. Part. Phys.*, 29, 1409
- Candia, J., Epele, L., & Roulet, E. 2002, *Aph*, 17, 23
- Cavasinni, V., Grasso, D., & Maccione, L. 2006, *Aph*, 26, 41
- Chantell, M., Covault, C., Cronin, J., et al. 1997, *PhRvL*, 79, 1805
- Crocker, R., Fatuzzo, M., Jokipii, J., Melia, F., & Volkas, R. 2005, *ApJ*, 622, 892
- Feldman, G., & Cousins, R. 1998, *PhRvD*, 57, 3873
- Feng, Z., Apel, W., Arteaga-Velazquez, J., et al. 2015, Proc. ICRC, 823
- Fisher 1925, *Statistical Methods for Research Workers* (Edinburgh: Oliver and Boyd)
- Good, I. J. 1955, *Journ. Royal Stat. Soc. B*, 17, 264
- Górski, K., Hivon, E., Banday, A., et al. 2004, *ApJ*, 622, 759
- Hayashida, N., Nagano, M., Nishikawa, D., Ohoka, H., & Sakaki, N. 1999, *Aph*, 10, 303
- Hinton, J., & Hofmann, W. 2009, *ARA&A*, 47, 523
- Hörandel, J. 2005, *IJMPA*, 20, 6753
- Kang, D., Apel, W., Arteaga-Velazquez, J., et al. 2015a, Proc. ICRC, 810
- Kang, D., Feng, Z., Apel, W., et al. 2015b, Proc. ICRC, 812
- Kohri, K., Ohira, Y., Ioka, K., et al. 2012, *MNRAS*, 424, 2249
- Li, T.-P., & Ma, Y.-Q. 1983, *ApJ*, 272, 317
- Liu, Q., van Paradijs, J., & van den Heuvel, E. 2007, *A&A*, 455, 1156
- Manchester, R., Hobbs, G., Teoh, A., & Hobbs, M. 2005, *ApJ*, 129, 1993
- Nozawa, S., Kohyama, Y., Itoh, N., et al. 2011, (arXiv:1103.4284)
- Olive, K. A., et al. Particle Data Group 2014, *ChPhC*, 38, 090001 and 2015 update
- Rawlins, K., Aartsen, M., Abraham, K., et al. 2015a, Proc. ICRC, 628 (arXiv:1510.5225)
- Rawlins, K., Aartsen, M., Abraham, K., et al. 2015b, Proc. ICRC, 334 (arXiv:1510.5225)
- Schlickeiser, R. 1989, *A&A*, 213, L23