

# Evaluating mantle and crustal processes using isotope geochemistry

By

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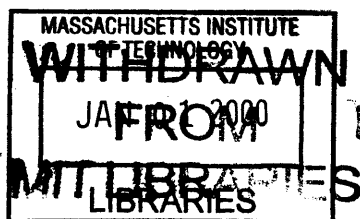
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Lindgren



*To Eliana*

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# **Evaluating mantle and crustal processes using isotope geochemistry**

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## **Abstract**

Geochemical studies are fundamental for understanding how the dynamic Earth works and evolves. These studies place constraints on the composition, formation, age, distribution, evolution and scales of geochemically distinct reservoirs such as the Earth's crust, mantle and core. In this dissertation the strategy has been to work on a broad range of topics to evaluate crustal and mantle processes. This study presents Re-Os systematics to constrain the composition, formation and age of the lower continental crust and the mantle lithosphere, examines melt inclusion from oceanic island basalts to evaluate the scale of the mantle heterogeneities, and uses U-series isotope to constrain geodynamic parameters, such as the upwelling velocities and porosities of mantle plumes.

The lower continental crust plays a pivotal role in understanding the composition and evolution of the continental crust and the petrogenesis of continental basalts. This chapter presents Re/Os isotope measurements which allow us to further our understanding of these problems. Two well-characterized suites of lower crustal xenoliths from Northern Queensland, Australia, which have average major and trace element compositions similar to bulk lower crust, were analyzed for Re/Os isotope systematics. From this data, we infer that the lower crust has 1 to 2 times as much Os, about half of the Re and is less radiogenic in  $^{187}\text{Os}/^{188}\text{Os}$  than the upper continental crust. Our data show that assimilation and fractional crystallization (AFC) are important processes in the formation of the lower crust and lead to dramatic changes in the Os isotopic composition of basalts that pond and



fractionate there. Because of this, the Re-Os system cannot be relied upon to yield accurate mantle extraction ages for continental rocks.

Chapter 2 examines the Re-Os isotopic composition of the Horoman massif, Japan. These data indicate that the Os isotope composition is controlled by the Re content, through radiogenic ingrowth, while the Re content is governed by the extent of depletion in “basaltic component” of the ultramafic rocks. Re-Os systematics suggest that depletion model ages of  $\approx 1.8$  Ga represent the age of the melting event. The colinearity between mafic and ultramafic rocks in the Re-Os isochron diagram defines an apparent age of  $\approx 1$  Ga. The similar “ages” determined by Re-Os and Sm-Nd isotopes and the high Re/Os ratios in the most fertile peridotites plotting to the right of the geochron, indicate that the mafic layers and the ultramafic rocks are genetically related by a refertilization process which took place  $\approx 1$  Ga ago. The Re-Os systematics for other ophiolitic massifs indicate that refertilization of the lithospheric mantle seems to be a more widespread process than previously thought.

Previous studies have suggested that melting processes are responsible for the trace element variability observed in olivine-hosted basaltic melt inclusions. Melt inclusions from four individual lava samples representing three mantle end-members HIMU, EMI and EMII (two from Mangaia, Cook Islands, one from Pitcairn, Gambier chain, and one from Tahaa, Society chain), have heterogeneous Pb isotopic compositions, even though the erupted lavas are isotopically homogeneous. The range of Pb isotopic compositions from individual melt inclusions in a single lava flow spans 50% of the world-wide range observed for ocean island basalts (OIB). The melt inclusion data can be explained by two-component mixing for each island. Our data imply that magmas with different isotopic compositions existed in the volcanic plumbing system prior to or during melt aggregation.

Evaluation of U-series disequilibrium, trace element composition and He, Sr, Nd and Pb isotopes of Galapagos lavas indicates that magma mixing between plume and asthenospheric melts has been the main process responsible for the geochemical variation observed in the archipelago. Correlations between He isotopes and Ti/Ti\*, K/Rb and Nb/La ratios suggest that the mantle plume has positive anomalies of Nb and Ti and negative anomalies of K.  $^{230}\text{Th}$  excesses measured in the lavas indicate that the basalts from Galapagos originated completely or partially in the garnet stability field. Mantle upwelling velocity for the Galapagos plume (Fernandina) ranges from  $\approx 1$  to 3 cm/y with a maximum

porosity of 0.3%, indicating that Galapagos is a mildly buoyant plume. Very slow mantle upwelling rates and very low porosity for Pinta (0.5 to 1 cm/y and 0.1%) and Floreana (0.1 cm/year and <0.1%) islands, support the hypothesis that the movement of the plume across the 91°50' transform fault into a younger and thinner lithosphere produced slow upwelling and small extents of melting.

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# Introduction

The long term goal of studying basalts, crustal rocks and mantle peridotites is to date and deconvolve the superimposed effects of processes responsible for the geochemical variations found in the Earth's crust and mantle. This dissertation attempts to use geochemical tracers such as major elements, trace elements and radiogenic isotopes to add constraints to the geologic processes responsible for the composition, formation and age of various geochemical reservoirs, such as the lower continental crust and the continental mantle lithosphere. Moreover, this study makes use of geochemical tracers in oceanic island basalts to assess the scale of mantle heterogeneities and to define geodynamic parameters during upwelling and melting of mantle plumes.

Chapter 1 applies Re-Os systematics to determine the composition, formation and age of the lower continental crust. The study of the lower continental crust is a key component in estimating the composition and age of the continental crust as a whole (Rudnick and Fountain, 1995). Whereas the upper continental crust is accessible to sampling, the deeper portions of the crust are relatively inaccessible. Lower crustal xenoliths are intact samples from the present-day lower continental crust (Taylor and McLennan, 1995, Rudnick, 1992). Therefore, the study of such xenoliths is very important for determining the bulk composition and origin of the lower continental crust. This study reports Re-Os isotopic data for two geochemically well-characterized suites of lower crustal xenoliths from Northern Queensland, Australia (Rudnick and Goldstein, 1990 and references therein). These data show that assimilation and fractional crystallization processes (AFC) are responsible for the formation and composition of the lower continental crust, and that whole rock Re-Os isochrons for magmas processed in the lower crust may not give reliable age or source composition information.

Chapter 2 examines the Re-Os systematics of a geochemically well-characterized 140 m section across the layered Horoman peridotite, Japan. The Horoman massif has been the subject of numerous structural, petrologic and geochemical studies (Takazawa et al., 1999a, 1999b and references therein). However, the process (or processes) responsible for the formation of the layering, as well as the origin of the fertile plagioclase lherzolite, still remains controversial. The origin of the layered structure has been explained either by melting processes (Takazawa et al., 1999b; Yoshikawa and Nakamura, 1999; Takahashi, 1992) or by a combination of melting and melt accumulation (Obata and

Nagahara, 1987). The Re-Os systematics seem to support the second hypothesis, where the layered structure in the Horoman massif is produced by a combination of melting and melt accumulation, with the plagioclase lherzolite representing the region of melt accumulation.

Chapter 3 applies an old geochemical tool, Pb isotopes, to a new dimension in isotope geochemistry by analyzing Pb isotopes in single melt inclusions from oceanic island basalts. Studies of oceanic basalts have shown that the mantle is isotopically heterogeneous (Zindler and Hart, 1986; Hart et al, 1992 and references therein). However, the nature, distribution, and scale of these heterogeneities remain uncertain. The aggregation of melts on their way to the surface, and mixing in magma chambers prior to eruption, can obscure the chemical and isotopic signatures of pre-aggregated melts. Studies of olivine-hosted melt inclusions have been successful in defining the chemical composition of pre-aggregated melts (Sobolev and Shimizu, 1993). However, the lack of isotopic information on melt inclusions has made it difficult to distinguish whether these melt compositions represent different extents of melting from a single source or whether they originate from different source compositions. The range in Pb isotopic composition observed in melt inclusions from single lava flows suggests that the geochemical variation in melt inclusions has a large component of source control, and is mainly produced either by some mixing of melts or assimilation and fractional crystallization processes deep in the magma chambers.

Chapter 4 presents the first U decay series isotopes study of lavas from the Galapagos archipelago.  $^{238}\text{U}$ - $^{230}\text{Th}$ - $^{232}\text{Th}$  isotopes place unique constraints on magmatic processes and mantle source compositions.  $^{238}\text{U}$ - $^{230}\text{Th}$  disequilibrium provides quantitative evaluation of the present-day depletion of the mantle source, the extent of melting represented by erupted lavas, and helps to constrain estimates of the mantle porosity and upwelling velocities during melting (Elliot, 1997 and references therein). Although the Galapagos islands have been interpreted as the surface manifestation of mantle plumes, geological, geophysical and geochemical studies have found it challenging to fit Galapagos magmatism to a “standard” hotspot model (Kurz and Geist, 1999). The diffuse nature of the Galapagos plume, and the dynamic melting and equilibrium porous flow models using U decay series isotopes constrain mantle upwelling velocities for the Galapagos plume (Fernandina) to  $\leq 3$  cm/y and maximum porosities of 0.3%. These results indicate that

Galapagos is a mildly buoyant plume, significantly less intense thermally than the Hawaiian hotspot.

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# 1. Chapter One

## Re-Os isotope evidence for the composition, formation and age of the lower continental crust

### *1.1 Abstract*

The lower continental crust plays a pivotal role in understanding the composition and evolution of the continental crust and the petrogenesis of continental basalts. We present Re/Os isotope measurements which allow us to further our understanding of these problems. Two well characterized suites of lower crustal xenoliths from Northern Queensland, Australia, which have average major and trace element compositions similar to bulk lower crust, were analyzed for Re/Os isotope systematics. From this data, we infer that the lower crust has 1 to 2 times as much Os, about half of the Re and is less radiogenic in  $^{187}\text{Os}/^{188}\text{Os}$  than the upper continental crust. Our data show that assimilation and fractional crystallization (AFC) are important processes in the formation of the lower crust and lead to dramatic changes in the Os isotopic composition of basalts that pond and fractionate there. Because of this, the Re-Os system cannot be relied upon to yield accurate mantle extraction ages for continental rocks.

### *1.2 Introduction*

Determining the composition and age of the continental crust is fundamental to an understanding of the evolution of the Earth (Taylor and McLennan, 1995, Rudnick, 1995). The study of the lower continental crust is a key component in estimating the composition and age of the continental crust as a whole (Rudnick and Fountain, 1995). Whereas the upper continental crust is accessible to sampling, and sediments provide a good means to obtain its average composition, the deeper portions of the crust are relatively inaccessible. The sampling of lower crust is limited to granulite facies terrains exposed during major tectonic events, and xenoliths carried from lower crustal depths deep to the surface by alkali basalts. Not all granulitic terranes represent the deep crust, and only those terranes that have undergone isobaric cooling for long periods of time (> 100 Ma.) can be representative of the lower crust (Rudnick and Taylor, 1987). In contrast, lower crustal xenoliths are intact samples from the present-day lower continental crust (Taylor and McLennan, 1995,



Rudnick, 1992). Therefore, the study of such xenoliths is very important for determining the bulk composition and origin of the lower continental crust.

Re-Os differs from other isotopic systems (Pb, Sr, Hf and Nd) because Re and Os are siderophile and chalcophile (rather than lithophile) elements, and because Os is a compatible element (Os is retained in the crystals) during partial melting or fractional crystallization (Hart and Ravizza, 1996). This compatible behaviour leads to high Re/Os ratios and radiogenic Os isotopic compositions in the continental crust (Allègre and Luck, 1980) and makes  $^{187}\text{Os}/^{188}\text{Os}$  ratios in fractionating mantle-derived magmas susceptible to change by crustal contamination (Reisberg et al., 1993; Marcantonio et al., 1995; Roy-Barman and Allègre, 1995; Widom and Shirey, 1991; Martin et al., 1991; Walker et al., 1997). The slow evolution of  $^{187}\text{Os}/^{188}\text{Os}$  in the mantle contrasts with the much more rapid growth of radiogenic Os in most crustal rocks (Allègre and Luck, 1980). Thus, Os isotopes should be a clear discriminator between crustal and mantle material (Walker et al., 1991; Walker et al., 1994; Horan et al., 1995). The Re-Os isotope systematics of the lower continental crust are important for evaluating

1) the processes responsible for the formation, composition, and age of the lower crust (Taylor and McLennan, 1995; Rudnick, 1992; Zandt and Ammon, 1995; Halliday et al., 1993; Rudnick and Fountain, 1995; Rudnick, 1995; Rudnick and Presper, 1990;

2) the average Re-Os isotopic composition and concentration of the bulk continental crust (Rudnick, 1992);

3) the sensitivity of Re-Os as an indicator of crustal contamination, and the extent to which basalts erupted through the continental crust have chemically exchanged with the deepest crust (Martin, 1989; Lambert et al., 1989; Walker et al., 1991; Marcantonio et al., 1993; Walker et al., 1994; Horan et al., 1995)

We report Re-Os isotopic data for two geochemically well-characterized suites of lower crustal xenoliths from Northern Queensland, Australia (Rudnick et al., 1986; Rudnick and Taylor, 1987; Rudnick and Williams, 1987; Rudnick and Goldstein, 1990; Rudnick 1990; Rudnick and Taylor, 1991; Kempton and Harmon, 1992). In this paper, we use the Re-Os isotopic systematics of these lower crustal xenoliths to show that assimilation and fractional crystallization processes (AFC) are responsible for the formation

and composition of the lower continental crust. We assess the sensitivity of the Re-Os isotopic system as an indicator of crustal contamination, and test whether there is decoupling between Os and other isotopic systems during assimilation processes. Finally, we show that whole rock Re-Os isochrons for magmas processed in the lower crust may not give reliable age information.

### ***1.3 Geologic Background***

Two suites of granulite-facies xenoliths, Chudleigh and McBride, from North Queensland have been the subject of extensive petrological and geochemical studies (Rudnick et al., 1986; Rudnick and Taylor, 1987; Rudnick and Williams, 1987; Rudnick and Goldstein, 1990; Rudnick 1990; Rudnick and Taylor, 1991; Kempton and Harmon, 1992). The Chudleigh xenoliths were carried in <2 Ma basalts erupted through Paleozoic rocks of the Tasman fold belt. Three types of xenoliths are present: plagioclase-rich, pyroxene-rich and transitional between these two groups. Mineralogical studies indicate that the xenoliths underwent isobaric cooling at moderate to high pressure ( $\approx 0.7$  to 1.5 GPa.). These xenoliths are mafic in composition ( $\text{SiO}_2 < 51.2$  wt%), and show good correlations between isotopic compositions, trace element contents and Mg# (molar  $\text{Mg}/(\text{Mg}+\text{Fe})$  ratios). These correlations were used to argue that the samples are genetically-related cumulates from basaltic liquids, something which is very unusual to find in lower crustal xenolith suites (Rudnick et al., 1986; Rudnick and Goldstein, 1990; Rudnick and Taylor, 1991; Kempton and Harmon, 1992). Rudnick et al. (1986), conclude that the age of the Chudleigh xenoliths is relatively young (<100 Ma), probably related to the Cenozoic igneous activity which occurs throughout eastern Australia (Rudnick et al., 1986).

The McBride xenoliths were carried by <3 Ma basalts erupted through the Proterozoic Georgetown inlier. These xenoliths range in composition from mafic to felsic, although mafic compositions dominate. Therefore, their chemical compositions reflect equally diverse protoliths, ranging from mafic cumulates, to mafic and felsic melts, and restites of the anatexis of metapelitic sediments (Rudnick and Taylor, 1987; Rudnick and Goldstein, 1990; Rudnick 1990; Rudnick and Taylor, 1991; Kempton and Harmon, 1992). As in Chudleigh, the xenoliths underwent isobaric cooling at moderate to high pressure ( $> 0.7$  GPa.). The age of these rocks has been determined by U-Pb zircon techniques (Rudnick and Williams, 1987). Most of the ages show xenoliths formation times  $\approx 300$  Ma ago, during calc-alkaline igneous activity. In contrast, some of the xenoliths apparently

formed during the Proterozoic  $\approx$  1570 Ma. Therefore the protoliths for the McBride xenoliths formed during large-scale mixing between Proterozoic crust and basalts during late-Paleozoic ( $\sim$  300 Ma) convergent margin magmatism (Rudnick and Taylor, 1987; Rudnick and Goldstein, 1990; Rudnick 1990; Rudnick and Taylor, 1991; Kempton and Harmon, 1992).

#### ***1.4 Analytical Techniques***

The samples were washed in ultrasonic bath with Millipore water, and dried on a hotplate at 50°C. Approximately 50 g of material was selected from each sample; the samples were wrapped in plastic and hammered to small chips, and the chips were powdered in an agate shatterbox.

The analytical techniques for Re-Os isotopes are described in Ravizza and Turekian (1989), Hauri and Hart (1993) and Peucker-Ehrenbrink et al., (1995). Re and Os were determined on separate powder splits. For the analysis of Os, 15 g of sample powder was digested by a NiS fire-assay technique. The powder is spiked with a solution of enriched  $^{190}\text{Os}$ , and mixed with high purity boric acid, Ni and S powders. The 1.5:1 mixture of flux to sample was melted at 1050 to 1150°C. This step produced a NiS bead that scavenges Os from the fused rock. The bead is dissolved in 6N HCl and the solution is filtered to recover the PGE particulates. The filter is dissolved in an oxidizing  $\text{CrO}_3 + \text{H}_2\text{SO}_4$  acid solution, from which the  $\text{OsO}_4$  is distilled and trapped in HBr acid. The HBr solution is dried down to a 1-2  $\mu\text{l}$  drop and diluted in  $\text{H}_2\text{O}$  until a solution of 1N HBr is obtained. A single bead of Chelex 50 resin is added to the solution, and Os is partitioned into the resin bead at low normality HBr. The bead is deposited in concentrated HBr to liberate the Os, which is then loaded on a Pt filament with  $\text{BaNO}_3$  as an activator. The Os concentrations and isotopic compositions were measured by negative thermal ionization mass spectrometry on NIMA-B (WHOI), with oxygen enhanced emission and analog multiplier detection using a single collector and a dynamic collection routine.

For Re analysis, 0.5 g of sample powder is spiked with a solution of enriched  $^{185}\text{Re}$ , and digested in a mixture of HF-HNO<sub>3</sub>. Re is extracted by simple ion exchange techniques, using AG1X8 resin. The Re is absorbed into the resin from 0.5N HNO<sub>3</sub>, and then eluted in 4N HNO<sub>3</sub>. An identical second step is used to further purify the Re. The Re

measurements were done on a Finnigan MAT Element high-resolution ICP-MS (at WHOI), by rapid peak-hopping using electrostatic scanning.

In-run precision for all but two of the measurements is  $2\sigma < 0.35\%$  for  $^{187}\text{Os}/^{188}\text{Os}$ ,  $2\sigma < 1\%$  for Os and  $2\sigma < 2\%$  for Re contents. The exceptions are samples 83-107 and 83-140, with lower precision for Os concentration ( $2\sigma < 10\%$ ) and isotopic composition ( $2\sigma < 7\%$ ), due to a large blank correction. The analytical blanks ranged from 0.4 to 0.5 pg/g of flux for Os and 5 to 24 pg total blank for Re; the isotopic composition of the blank ranged from 0.4 to 0.47 for  $^{187}\text{Os}/^{188}\text{Os}$ . We used a ratio 1.5:1 of flux to sample. All samples were corrected for blank contribution. Re and Os replicates on separate powder splits ranged from 2 to 11% and from 1 to 18% respectively, excluding sample 85-106, which showed a 32% difference for the Re replicate. The  $^{187}\text{Os}/^{188}\text{Os}$  replicates for the Chudleigh xenoliths are  $< 0.6\%$ , but for the McBride suite, the replicates range from 2 to 5%. The style of the poorer reproducibility in replicates from the McBride suite is suggestive of a “nugget” effect; that is, domains or minerals having higher Re/Os are heterogeneously distributed within the rock and have been out of diffusive equilibrium with the rest of the rock for reasonably long time spans (in this case  $\sim 300$  Ma, the age of the suite; Rudnick and Williams, 1987). We consider high Re/Os nuggets as a possible explanation for the poor reproducibility in the McBride suite because the variation in  $^{187}\text{Os}/^{188}\text{Os}$  ratios and Re contents is large, at similar Os concentrations (Table 1).

## 1.5 Results

The Os isotopic composition and Re and Os concentrations of sixteen samples from two well-characterized lower crustal xenolith suites from North Queensland are reported in Table 1. The two suites of xenoliths from Queensland show differences in Re and Os concentration and Re/Os ratios, but have a similar and large range in  $^{187}\text{Os}/^{188}\text{Os}$  ratios (from 0.1721 to 1.759 for McBride versus 0.1889 to 1.814 for Chudleigh) (Figure 1 a and b). The McBride xenoliths have a larger range of Re (two orders of magnitude), and a smaller range of Os contents and Re/Os ratios than the Chudleigh xenoliths (a factor of less than two, excluding one outlier). Moreover, the Os isotopic ratios of the McBride xenoliths do not correlate with Re/Os ratios (Figure 1). These characteristics are consistent with the extreme lithologic and genetic diversity of the McBride xenoliths; as noted above, these range from paragneisses to mafic and felsic orthogneisses, formed during large-scale mixing between Proterozoic crust and basalts during late-Paleozoic ( $\sim 300$  Ma) convergent

margin magmatism (Rudnick and Taylor, 1987; Rudnick and Goldstein, 1990; Rudnick 1990; Rudnick and Taylor, 1991; Kempton and Harmon, 1992).

In contrast, the Re-Os results for the Chudleigh xenoliths show a strikingly good linear correlation on a Re-Os isochron diagram, defining an apparent age of approximately 260 Ma (Figure 2). However, relatively good positive correlations for present-day  $^{187}\text{Os}/^{188}\text{Os}$  and  $\delta^{18}\text{O}$ , Pb, and Sr isotopic compositions, and negative correlations between  $^{187}\text{Os}/^{188}\text{Os}$  and Nd isotopic ratios contradict the hypothesis that the Re-Os isochron array represents a real age (Figure 3). In particular, the positive correlation between  $^{187}\text{Os}/^{188}\text{Os}$  and  $\Delta 7/4$  Pb (Hart, 1984) implies the existence of an old crustal component in the formation of the lower crust (Figure 3). These correlations between the different isotopes and  $^{187}\text{Os}/^{188}\text{Os}$  ratios suggest that the variation in Os isotopes is not the result of ingrowth from Re/Os, but is produced by mixing or AFC processes, as proposed by earlier studies. These studies showed that the Chudleigh xenoliths are mafic in composition ( $\text{SiO}_2 < 51.2$  wt%), and showing good correlations between Sr, Nd, Pb and  $\delta^{18}\text{O}$  isotopic composition, trace element contents and  $\text{Mg}^\#$ . These correlations reflect a genetically-related suite of cumulates from mafic magma(s) that underwent simultaneous assimilation and fractionation in the deep crust (Rudnick et al., 1986; Rudnick and Goldstein, 1990; Rudnick and Taylor, 1991; Kempton and Harmon, 1992). It is very unusual to find such simple chemical and isotopic systematics in lower crustal xenolith suites. The Chudleigh suite thus provides an unique natural laboratory with which to investigate the chemical changes occurring in mafic magmas that pond and fractionate within the deep continental crust.

## ***1.6 Discussion***

### **1.6.1 The Re-Os composition of the lower continental crust**

The Chudleigh and McBride xenoliths have average major and trace element compositions similar to estimates of the bulk lower continental crust (Rudnick, 1992; Rudnick and Fountain, 1995) (Figure 4), and can therefore be used to place broad constraints on the Re-Os concentration and Os isotopic composition of the lower crust. The xenoliths have Os and Re concentrations which average 102 and 276 ppt ( $\text{Re}/\text{Os} = 2.71$ ), with median values of 49 and 184 ppt ( $\text{Re}/\text{Os} = 3.75$ ) respectively (Table 1, Figure 1). The average and the median are quite different because of one outlier (sample 85-100) with high Os ( $\approx 1000$  ppt) and Re ( $\approx 800$  ppt) concentrations. The median values are

probably more useful in representing the bulk lower crust, but more data are clearly desirable. The concentration-weighted mean  $^{187}\text{Os}/^{188}\text{Os}$  ratio for the lower crust in north Queensland is 0.4 or 0.8, including and excluding the outlier, respectively. Recent analyses of eclogite and granulite facies xenoliths from central Arizona by Esperança et al., (1997), give similar results (Table 1, Figure 1), although these xenoliths have somewhat lower Os contents, which translates into a slightly higher mean Re/Os ratio. We infer from these data that, relative to the upper crust (Esser and Turekian, 1993), the lower crust has one to two times as much Os, about half of the Re, and is probably less radiogenic (having one half to one third the  $^{187}\text{Os}/^{188}\text{Os}$  of the upper crust). However, the Os isotopic composition of the lower continental crust will vary with the age of the crust. If we assume a mean age of 2.3 Ga for the continental crust (Allègre et al., 1983) and the Re-Os composition of the mantle given by Meisel et al. (1996), and our average Re/Os estimates of  $^{187}\text{Re}/^{188}\text{Os} \approx 20$  and 16, respectively (including and excluding the data by Esperança et al. (1997), for the Arizona xenoliths), we calculate the mean  $^{187}\text{Os}/^{188}\text{Os}$  of the bulk lower crust to be between 0.7 to 0.9 (Figure 1b), a range that encompasses the median Os isotopic composition derived from the Queensland xenoliths (Table 1). Importantly, our data establish the presence of material with relatively high Os contents and radiogenic Os isotopic compositions in the lower crust. The variation in Os isotopic composition and Re-Os concentrations in all three xenolith suites illustrates the heterogeneity of the present lower crust, and shows that isotopic homogenization through fluid flow or melting has not been important.

### **1.6.2 Chudleigh xenoliths**

#### *1.6.2.1 The Assimilation and Fractional Crystallization (AFC) model*

We model the Chudleigh xenoliths as cumulates resulting from simultaneous assimilation and fractional crystallization (DePaolo, 1981) of a basaltic melt in the deep continental crust (Figure 3d, e and f; modeling parameters are given in the figure caption). The sigmoidal shape of the curves in the  $^{187}\text{Os}/^{188}\text{Os}$  versus Sr and Nd isotopes shows that during the first stages of AFC the Os isotopic composition is buffered by the Os content in the basaltic magma, whereas during the later stages it is buffered by the Os content and isotopic composition of the assimilant (old continental crust). The fractionation of olivine and/or sulfide decreases the concentration of Os in the magma (Hart and Ravizza, 1996), making it susceptible to contamination with crustal material having low Os concentration (2 ppt, in our model; Figure 3d). Notice that in the initial and final stages

there is a decoupling between  $^{187}\text{Os}/^{188}\text{Os}$  and the other isotopic systems. The compatible behavior of Os relative to Nd, Sr and Pb, and the differences in Os concentration and isotopic composition between mantle melts and evolved crustal materials, accounts for the buffering of Os in the melt; such effects are generally not observed in the other isotopic systems. Moreover, the compatible character of Os explains the large difference between the trends produced by AFC and bulk mixing (Figures 3), and makes AFC a far more effective process than bulk mixing for changing the isotopic composition of the basalt, especially when the contaminant has a low Os concentration.

#### 1.6.2.2 *The age*

Although the Chudleigh xenoliths cannot be dated directly, the chemical systematics described above provide some constraints on their magmatic crystallization age. Because the xenoliths are cumulates, and not crystallized magmas, they exhibit a large range in Sm/Nd ratio, depending upon whether plagioclase or clinopyroxene is the main cumulate phase. Thus the excellent correlations seen today between Nd isotopic composition ( $\epsilon_{\text{Nd}}$  from +9.6 to -6.1), Mg# (63 to 79), Pb ( $^{206}\text{Pb}/^{204}\text{Pb} = 17.8$  to 18.7), Sr ( $^{87}\text{Sr}/^{86}\text{Sr} = 0.70239$  to 0.71467) and  $\delta^{18}\text{O}$  ( $\delta^{18}\text{O} = +6.2$  to +7.0) degrade as the isotopic compositions are calculated at earlier times (Rudnick et al., 1986; Rudnick and Goldstein, 1990; Kempton and Harmon, 1992) (Figure 2, inset). These features led to the conclusion that the cumulates crystallized relatively recently (<100 Ma), and that they are probably related to Cenozoic igneous activity that occurs throughout eastern Australia (Rudnick et al., 1986).

The Re-Os results for the Chudleigh xenoliths show a strikingly good linear correlation on a Re-Os isochron diagram, defining an apparent age of approximately 260 Ma (Figure 2). There are two possible interpretations for this correlation:

- 1) it accurately reflects the igneous crystallization age of the xenoliths.
- 2) it reflects a recent AFC process involving a mantle-like end member and a radiogenic crustal end member.

We prefer the second interpretation for the reasons outlined below.

Parental magmas in the xenolith suite underwent both crystallization (as evidenced by the cumulate textures and compositions) and crustal assimilation (as seen in the Nd, Sr, Pb,  $\delta^{18}\text{O}$  isotopic systems, trace element and Mg# systematics) (Rudnick et al., 1986;

Rudnick and Goldstein, 1990; Kempton and Harmon, 1992). If interpretation 1 is correct, at  $\approx 260$  Ma ago the samples should all have the same initial  $^{187}\text{Os}/^{188}\text{Os}$ , while the Re/Os and the other isotopic systems varied due to progressive contamination. Therefore, this hypothesis assumes that the Os isotopes were either completely immune to crustal assimilation, giving the isotopic composition of the initial melt, or they were homogenized to the isotopic composition of the contaminant. That means that 260 Ma ago, the Os isotopes reflected in the initial or final stage of the sigmoidal curve produced during the AFC process (Figure 3). Neither case is true. The  $(^{187}\text{Os}/^{188}\text{Os})_{260 \text{ Ma}}$  ratios of these xenoliths exhibit a large range, between 0.12 to 0.26, and therefore are not uniform at 260 Ma as expected by interpretation 1 (Figure 5). Even in the case that we assume the initial  $(^{187}\text{Os}/^{188}\text{Os})_{260 \text{ Ma}}$  for the xenoliths to be uniform, it would be required to be approximately  $\approx 0.19$ . This Os isotopic composition is either too high compared to modern mantle-derived basalts (cf. 0.1276 to 0.1565) or too low to represent old continental crust (cf.  $\approx 1.9$ ). Another possibility is that the AFC event was not recent but occurred 260 Ma ago and  $(^{187}\text{Os}/^{188}\text{Os})_{260 \text{ Ma}}$  ratios were modified by contamination. If this would be the case, we would expect that the  $(^{187}\text{Os}/^{188}\text{Os})_{260 \text{ Ma}}$  ratios would correlate with other indicators of contamination such as  $\delta^{18}\text{O}$  isotopes. Figure 5 shows that the initial Os isotope ratios do not correlate with any of the other isotopic systems at  $\approx 260$  Ma. Thus, it is unlikely that the apparent “age” of 260 Ma is real, or that the mixing occurred 260 Ma ago. In contrast, relatively good positive correlations are seen for present-day  $^{187}\text{Os}/^{188}\text{Os}$  and  $\delta^{18}\text{O}$ , Pb, and Sr isotopic compositions and negative correlations between  $^{187}\text{Os}/^{188}\text{Os}$  and Nd isotopic ratios (Figure 3). The inset to Figure 2 shows that the correlations between different isotope systems, as measured by their correlation coefficient, are best within 50 to 100 Ma of the present-day. The general lack of correlations between Mg# and Nd, Pb isotopes at 260 Ma (Figure 2, inset), coupled with the relatively good present-day correlations (Figure 3), strongly suggests that these cumulates did not crystallize at 260 Ma (Rudnick et al., 1986; Rudnick and Goldstein, 1990). We conclude that the positive trend in the Re-Os isochron diagram was produced by AFC processes at a later date.

Our AFC model not only explains the correlation between  $^{187}\text{Os}/^{188}\text{Os}$  and the other isotopic ratios observed in the Chudleigh xenoliths (Figure 3d, e and f), it also reproduces the near-linear trend in the Re-Os isochron diagram (Figure 2), even when the contaminant does not lay on that trend. Our data and modelling show that AFC processes may produce



whole-rock Re-Os isochrons with no age significance, by depletion of Os through olivine-sulfide fractionation, coupled with assimilation of radiogenic crust.

### **1.6.3 Continental basalts, mantle heterogeneity or lower crustal contamination?**

The Chudleigh xenoliths clearly document that the Os isotopic composition of basaltic magmas can be easily affected by crustal assimilation. Small amounts of AFC during basalt-lower continental crust interaction can produce Os isotopic variations that may be erroneously interpreted as evidence for mantle heterogeneity. For example, consider an AFC process involving basalts with a MORB signature ( $^{187}\text{Os}/^{188}\text{Os} \approx 0.1276$ , Os concentration ranging from 10 to 300 ppt; Roy-Barman and Allègre, 1995), interacting with the lower crust ( $^{187}\text{Os}/^{188}\text{Os} \approx 0.804$ , Os  $\approx 50$  ppt). Figure 6 shows that it would require only 1 to 10% AFC, depending on the Os concentration of the basalts to change the isotopic composition of the melt from 0.1276 to 0.1565, a value typical for a more enriched mantle source [using the most primitive OIB samples analyzed for Os isotopes (Hauri and Hart, 1993), and assuming a bulk  $D_{\text{Os}} \approx 10$  (Hart and Ravizza, 1996), and a ratio of mass assimilated to mass fractionated of one half]. Thus, small extents of AFC processes during basalt-lower crust interaction can produce isotopic variations in the melt that may be erroneously interpreted as evidence for mantle heterogeneity.

## ***1.7 Conclusions***

In summary, the lower crust has, on average, higher Os ( $\geq 2$  times) and lower Re concentrations than current estimates of the average upper crust and may thus be comparatively less radiogenic. Materials with relatively high Os concentrations and radiogenic Os isotopic compositions exist in the lower crust and may act as contaminants to basaltic magmas that differentiate there. The Re-Os system is extremely sensitive to AFC processes at relatively small degrees of differentiation (i.e.,  $\ll 20\%$  AFC) and AFC processes are more effective than bulk mixing in changing Os isotopic compositions of melts, due to Os depletion during fractional crystallization. Like other isotopic systems, the Re-Os system is not a good indicator of the crystallization age of rocks that formed through AFC processes.

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NOTE: This chapter is based on an already published work: Saal A. E., Rudnick R.L., Ravizza G. E. and Hart S. R., Re-Os isotope evidence for the composition, formation and age of the lower continental crust. *Nature* 393, 58-61 (1998).

## 1.9 Table Caption

**Table 1:** Re, Os concentration and Os isotopic composition for the Chudleigh and McBride xenolith suites. \* indicates samples and replicate that were run on the same day.  $T_{OsPUM}$  are Os model ages calculated with respect to “bulk earth” ( $^{187}Re/^{188}Os = 0.428$  and  $^{187}Os/^{188}Os = 0.1290$ ; Meisel et al. (1996) and a decay constant of  $1.64 \times 10^{-11} \text{ y}^{-1}$ . Pl = plagioclase and Px = pyroxene. Average and median values for the Arizona suite (Esperança et al, 1997) and average values for the upper continental crust (Esser and Turekian, 1993) are also reported. In some samples, we report duplicates for Re, but we measured the Os concentration and  $^{187}Os/^{188}Os$  ratios only once. In those cases, we calculate the  $^{187}Re/^{188}Os$  and the Os model ages twice, using both Re values and the measured Os concentration and isotopic composition.

## 1.10 Figure Captions

**Figure 1:** a) Os versus Re concentrations, and b) Re/Os versus  $^{187}Os/^{188}Os$  ratios for lower crustal xenoliths from Queensland, Australia. The two mean lower crustal isotopic compositions shown in panel b are weighted according to Os concentrations. The lower square includes an anomalously high Os sample, the upper square excludes this sample. Fields for upper continental crust (without organic and carbonate-shelf sediments) and oceanic basalts (MORB-OIB) were compiled by Bernhard Peucker-Ehrenbrink. The mean and median Re and Os concentration, Re/Os ratios and isotopic composition weighted by Os concentration for Arizona xenoliths (Esperança et al, 1997) are plotted for comparison. The McBride xenoliths show chemical characteristics indicative of a lithologically and genetically diverse suite, and plot in the field defined by upper crustal material. The Chudleigh suite shows characteristics of genetically-related cumulates from basaltic magma undergoing crustal contamination (see text).

**Figure 2:** Re-Os isochron diagram for the Chudleigh suite showing bulk mixing and AFC trends. The data (solid circles) define a linear trend, with an apparent age of  $\approx 260 \text{ Ma}$ . Open circles are duplicate analyses, not included in the isochron regression. The inset shows correlation coefficients (for linear correlations) between Mg#, Nd ( $^{143}Nd/^{144}Nd$ ) and Pb ( $^{206}Pb/^{204}Pb$ ). In contrast to the Re-Os age, the best correlations exist at  $<100 \text{ Ma}$ , suggesting that the crystallization age of these xenoliths was  $<100 \text{ Ma}$ . The AFC model (using the parameters listed in Figure 3 caption) generates a linear trend in the Re-Os

diagram (shown by gray circles), which has no age significance. X's show the bulk mixing trend between the initial melt and the assimilant used in the AFC model. Each gray circle and X corresponds to 5% AFC or bulk mixing, respectively. The York regression for the AFC model was calculated using the same number of points and same distribution in the isochron as the Chudleigh suite, to make the results comparable (the points used are labeled, i.e., 3%). Note that the Chudleigh xenoliths are cumulates, producing low Re/Os ratios during the AFC process. Figures 2, 3 and 6 indicate that careful evaluation of AFC is necessary before age significance is attributed to Re-Os isochrons for lower crustal xenoliths or for basalts (or their cumulates) that may have differentiated in the lower continental crust.

**Figure 3:** Isotopic correlations and AFC model results for the Chudleigh xenoliths (black circles).  $^{187}\text{Os}/^{188}\text{Os}$  versus a)  $^{206}\text{Pb}/^{204}\text{Pb}$ , b)  $\Delta 7/4 \text{ Pb}$ , c)  $\delta^{18}\text{O}$ , d) Os concentration, e)  $^{143}\text{Nd}/^{144}\text{Nd}$  and f)  $^{87}\text{Sr}/^{86}\text{Sr}$ . Relatively good correlations exist between  $^{187}\text{Os}/^{188}\text{Os}$ , Os concentration and Pb, Sr, Nd, Pb,  $\delta^{18}\text{O}$  isotopic systems, with the exception of one sample. The aberrant sample (83-114, open circle) is probably showing a “nugget” effect; this explanation is consistent with this sample having the lowest Os isotopic ratios and the highest Os concentration, and plotting outside the correlation defined by the other samples. The AFC curves (gray circles) are marked in 5% increments (except the labeled points), where the ratio of mass assimilated to mass fractionated is  $\approx 0.67$ . The contaminant has intermediate Sr and Nd isotopic compositions ( $^{87}\text{Sr}/^{86}\text{Sr} \approx 0.75$  and  $^{143}\text{Nd}/^{144}\text{Nd} \approx 0.5114$ ) and elemental concentrations (Sr  $\approx 200$  ppm, Nd  $\approx 27$  ppm), based on published data for crust from the Tasman fold belt (Rudnick et al, 1986). This crustal contaminant is assumed to have very low Os ( $\approx 2$  ppt), to test the sensitivity of Os to crustal contamination, and Re concentrations ( $\approx 24$  ppt), and  $^{187}\text{Os}/^{188}\text{Os} \approx 2.6477$  ( $T_{\text{OsCH}} \approx 2$  Ga.), similar to Proterozoic upper continental crust. All of the isotope systems (e.g., Sr, Nd, Pb,  $\delta^{18}\text{O}$  and Os) require a contaminant that is higher in  $\delta^{18}\text{O}$ ,  $\Delta 7/4\text{Pb}$ , Sr, Pb and Os and lower in Nd isotopic composition than any sample seen in either the Chudleigh or McBride xenolith suites; this suggests the presence in the lower crust of (unsampled) old material with high  $\delta^{18}\text{O}$  and high time-integrated Rb/Sr, U/Pb and Re/Os and low Sm/Nd. For the parental basaltic melt, we used the most primitive isotopic composition seen in the xenoliths ( $^{87}\text{Sr}/^{86}\text{Sr} \approx 0.702$ ,  $^{143}\text{Nd}/^{144}\text{Nd} \approx 0.5131$ ) and inferred  $^{187}\text{Os}/^{188}\text{Os} \approx 0.1276$ . We assumed this melt has Sr  $\approx 350$  ppm, Nd  $\approx 14$  ppm, Re  $\approx 250$  ppt, and 7.5 times the Os concentration of the contaminant ( $\approx 15$  ppt). Bulk partition coefficients for Sr and Nd were

estimated assuming cumulates of plagioclase + clinopyroxene + olivine ( $D_{Sr} \approx 0.5$ ,  $D_{Nd} \approx 0.1$ ). The bulk partition coefficient for Os ( $D_{Os} \approx 7$ ), was estimated from the data of Hart and Ravizza, (1996); Re was assumed to be moderately incompatible ( $D_{Re} \approx 0.6$ ). The AFC model presented here is not unique, but serves to illustrate that the trends observed in the data can be explained by an AFC process. In contrast, the same is not true for bulk mixing, which is not successful in reproducing the observed arrays.

**Figure 4:** Comparison between the mean major and trace elements composition of global lower crustal xenoliths with those from the Queensland xenoliths. The global dataset (available on <http://www-ep.es.lnl.gov/germ/>) includes 450 analyses of major elements and 150 analyses of trace elements, excluding the Queensland data. Both means for a variety of elements agree within the standard errors ( $2\sigma$ ) of the compilations, suggests that the global average is a reasonable estimate of lower crust composition. The exception are Cs, Th, U, Pb, Cu and Zn, difference that we attribute to their mobility during secondary processes and to the difficulty in their accurate measurement. The fact that many lower crustal xenoliths are demonstrably older than the magmatic episode that carried them to the surface indicates that the xenoliths are not genetically related to that magmatic event (Rudnick, 1992).

**Figure 5:**  $(^{187}\text{Os}/^{188}\text{Os})_{260\text{Ma}}$  versus  $(^{206}\text{Pb}/^{204}\text{Pb})_{260\text{Ma}}$ ,  $(^{143}\text{Nd}/^{144}\text{Nd})_{260\text{Ma}}$ , and  $(^{87}\text{Sr}/^{86}\text{Sr})_{260\text{Ma}}$ ,  $^{187}\text{Re}/^{188}\text{Os}$  and  $\delta^{18}\text{O}$ . Note that the present-day correlation between the isotopes is lost when they are corrected back to 260 Ma, indicating that the AFC processes did not occur 260 Ma ago. Moreover, the xenoliths exhibit a large range in  $(^{187}\text{Os}/^{188}\text{Os})_{260\text{Ma}}$  from 0.12 to 0.26; this variability would not be expected if the linear trend produced by the Chudleigh xenoliths in the Re-Os isochron diagram define a real age of 260 Ma (see text for discussion).

**Figure 6:**  $^{187}\text{Os}/^{188}\text{Os}$  versus extent of differentiation during an AFC process. This is a hypothetical case where we evaluate how much differentiation it would take for a MOR basalt interacting with the lower crust to give Os isotopic ratios similar to those measured in OIBs. We assumed for MORB  $^{187}\text{Os}/^{188}\text{Os} \approx 0.1276$ , and an Os concentration ranging from 10 to 300 ppt. We used our average Os values for the lower crust ( $^{187}\text{Os}/^{188}\text{Os} \approx 0.804$ ,  $\text{Os} \approx 50$  ppt). We also assumed that the ratio between mass assimilated and mass fractionated is equal to 0.5 and the bulk partition coefficient for Os is  $D_{Os} \approx 10$ . Symbols are given at every 1% of AFC. Thus, it would require only 1 to 10% AFC, depending on

the initial basalt Os concentration, to change the MORB Os isotopic composition of 0.1276 to that of OIBs  $\approx$  0.1565.



**Table 1: Re and Os concentrations and Os isotopic compositions of granulite-facies lower crustal xenolith suites from Northern Queensland, Australia.**

Sample	Lithology	Re (ppt)	Os (ppt)	$^{187}\text{Re}/^{188}\text{Os}$	$^{187}\text{Os}/^{188}\text{Os}$	$T_{\text{OsPUM}}$ (Ma)
<b>Chudleigh suite</b>						
83-107	Pl-rich cumulate	153.2	3.20	265	1.249	258
		146.0		252		271
83-114*	Pl-rich cumulate	109.6	49.7	10.66	0.1875	349
		89.8	53.8	8.08	0.1889	478
83-127	Pl-rich cumulate	132.2	16.55	39.5	0.3056	275
		146.6		43.9		248
83-131	Pl-rich cumulate	104.7	9.39	55.9	0.4251	324
83-140	Pl-rich cumulate	259	3.85	391	1.814	263
83-110	Px-rich cumulate	70.8	31.3	11.27	0.3104	1,013
83-115	Px-rich cumulate	292	15.29	97.9	0.613	302
BC	Transitional cum.	297	51.01	28.7	0.2703	305
<b>McBride suite</b>						
83-160	Felsic granulite	219	61.9	19.33	1.180	3,299
83-162	Felsic granulite	80.1	56.2	7.23	0.5013	3,250
85-100*	Mafic granulite	734	1,043	3.41	0.1729	897
		808	931	4.21	0.1652	586
85-108*	Mafic granulite	1,141	106.5	55.0	0.6230	550
		1,210	104.5	59.3	0.6079	494
85-106	Mafic cumulate	50.8	47.4	5.26	0.1993	884
		34.5		3.58		1,349
83-159	Mafic residue	422	55.9	41.9	1.2959	1,692
85-114*	Mafic residue	234	84.0	16.28	1.7597	5,970
		230	86.2	15.46	1.6590	5,910
83-157	Metasediment	64.1	44.5	8.00	1.2832	8,649
<b>Crustal Averages</b>						
	North Queensland Mean	276	102	13.54	0.418	1,330
	<b>North Queensland Median</b>	<b>184</b>	<b>49</b>	<b>19.70</b>	<b>0.804</b>	<b>2,100</b>
	Central Arizona Mean	286	45	32.33	0.548	796
	Central Arizona Median	161	31	26.42	0.6710	1,259
	Average Upper Crust	400	50	47.62	1.9256	2,278

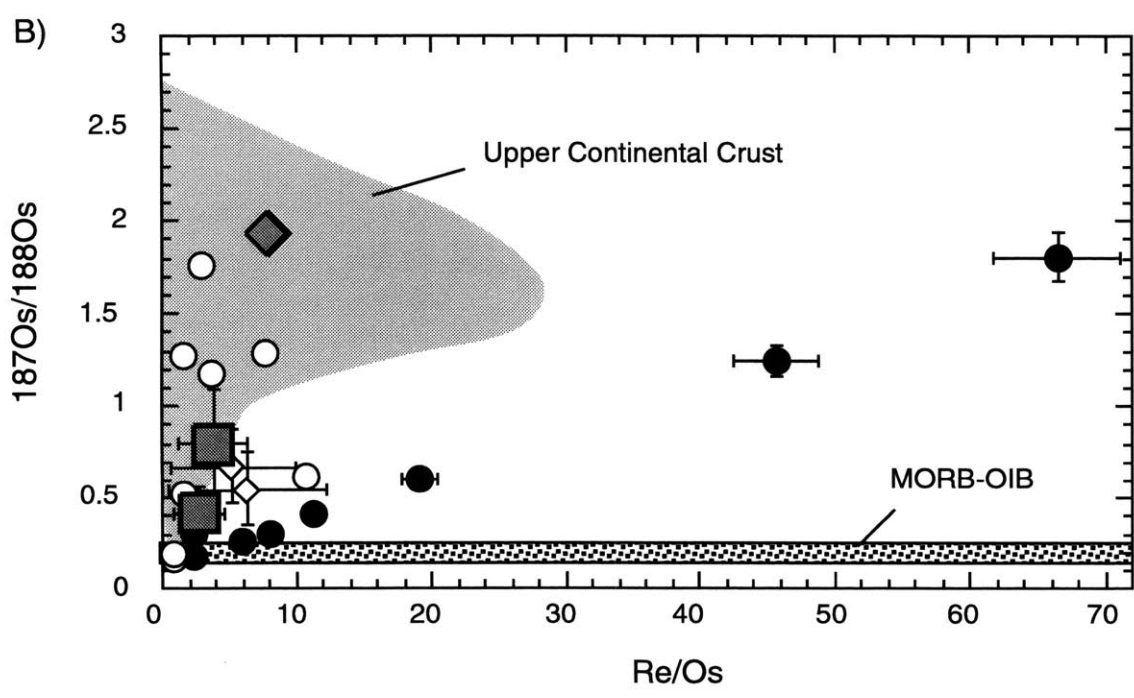
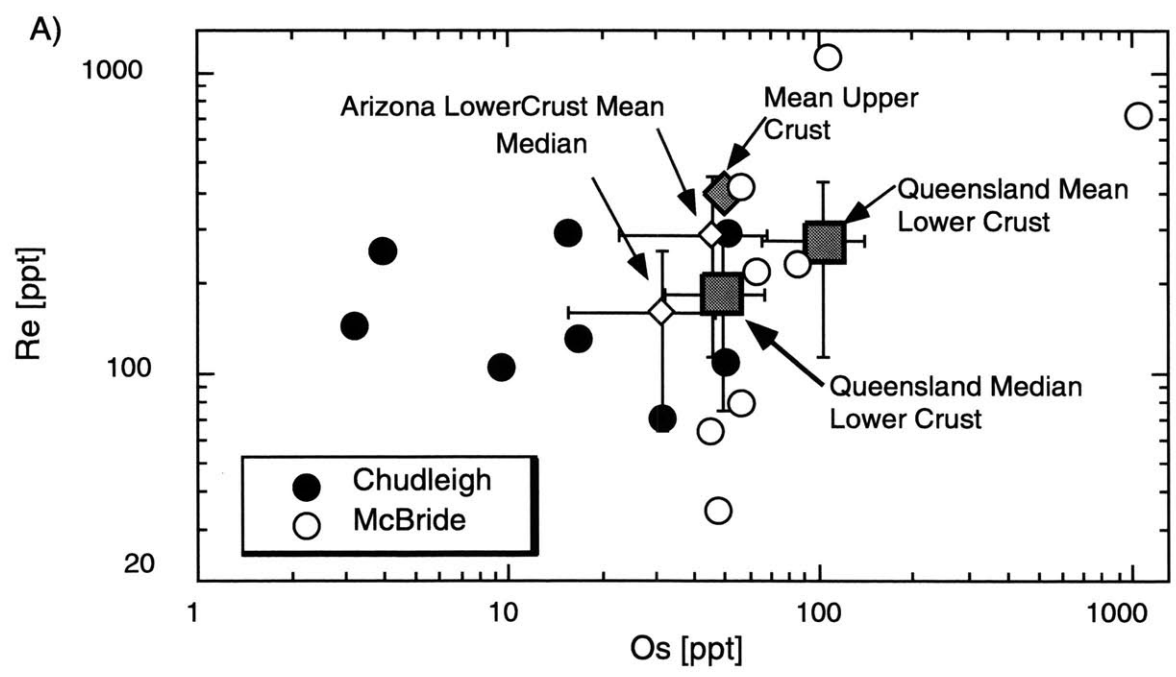


Figure 1

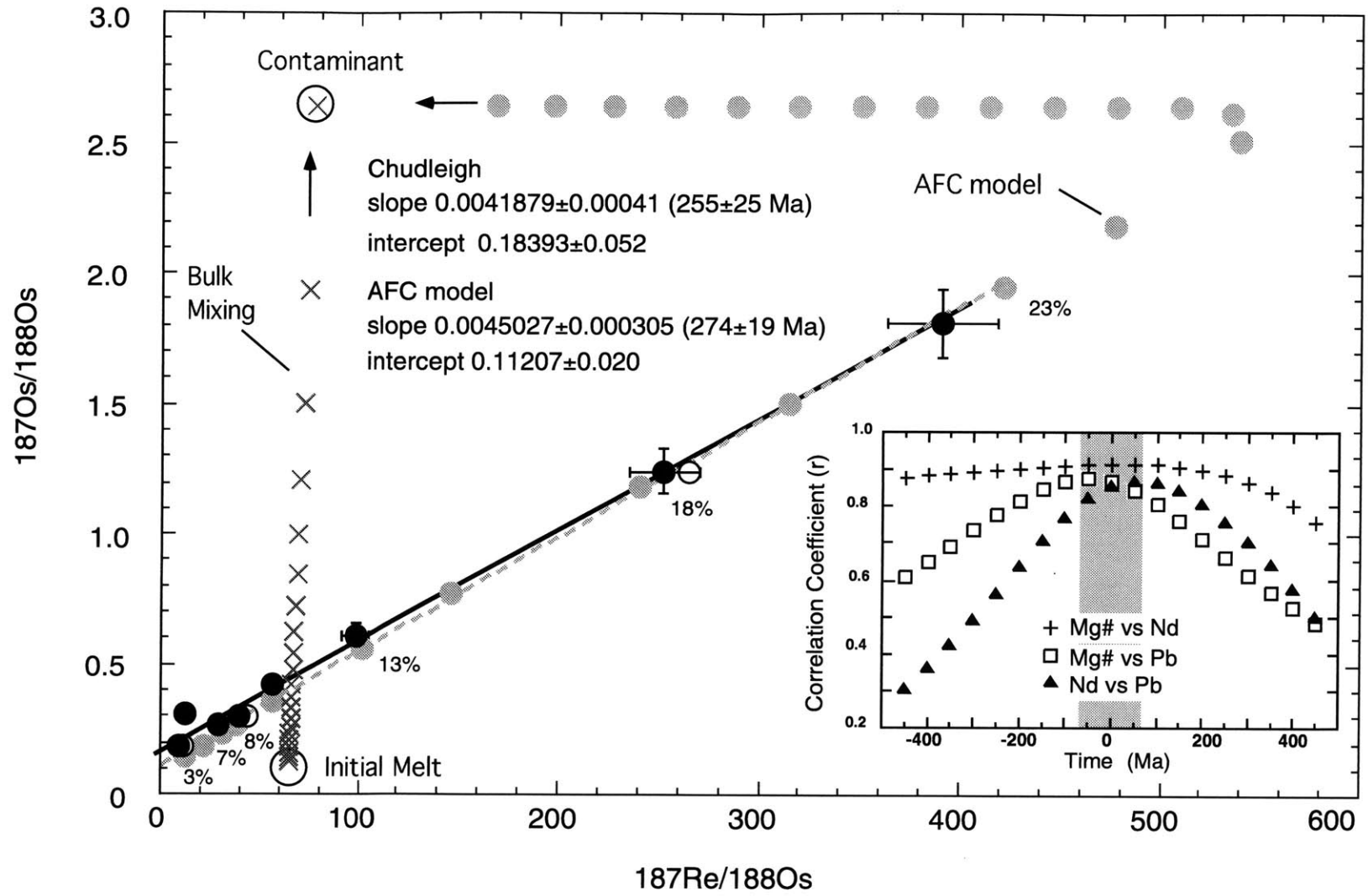


Figure 2

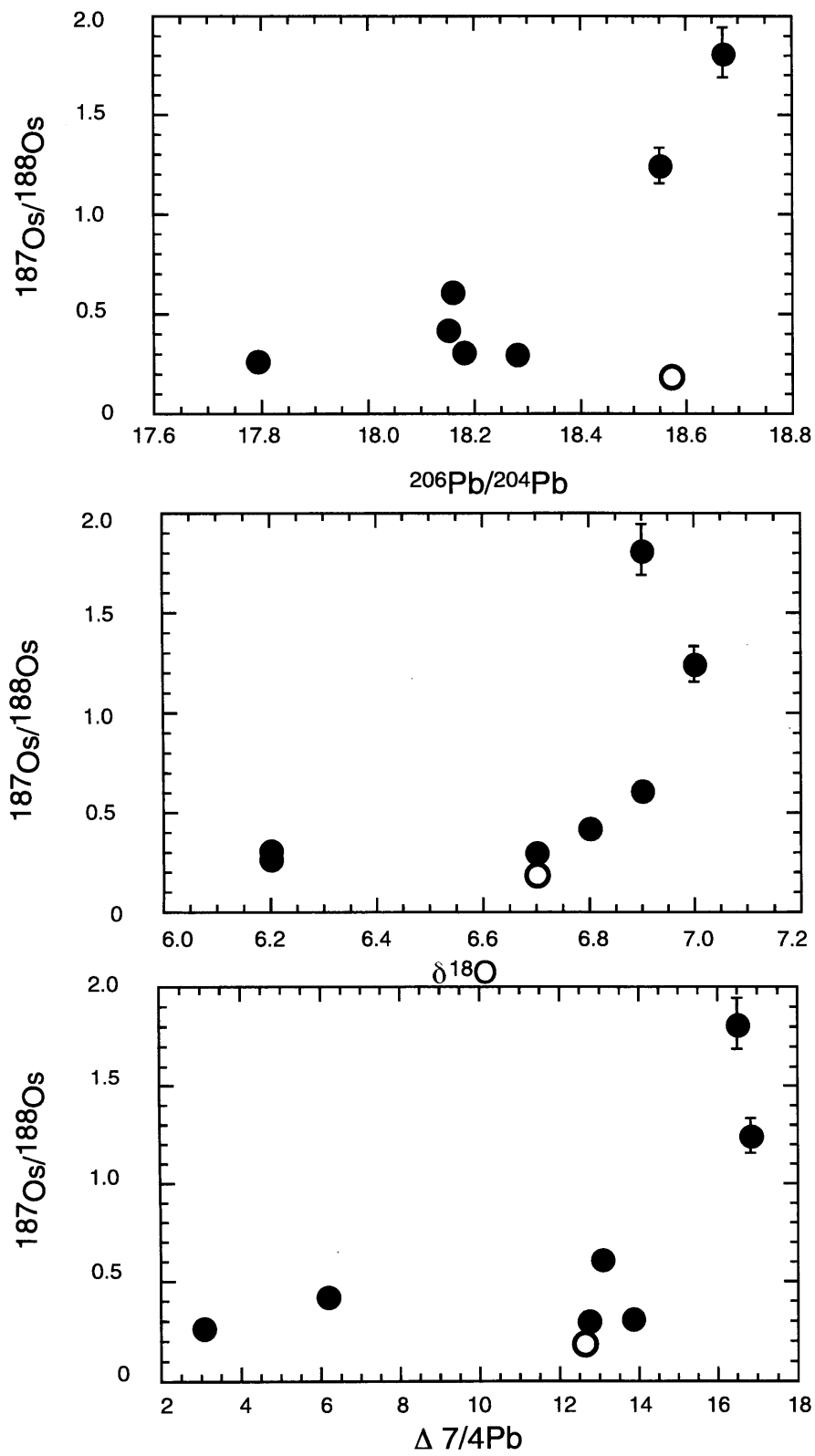


Figure 3

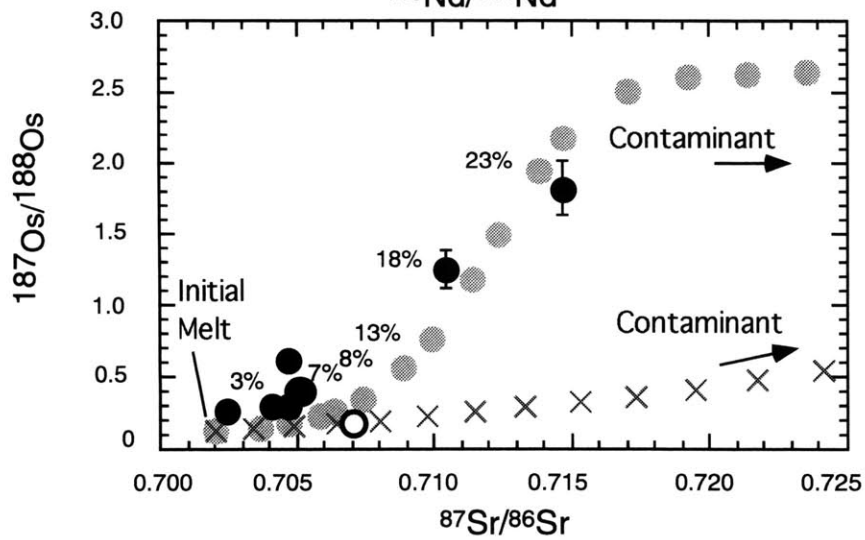
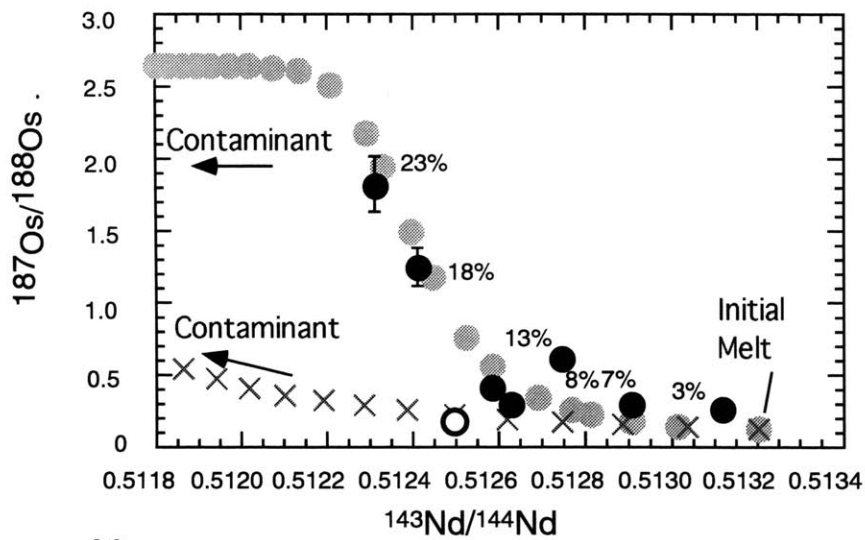
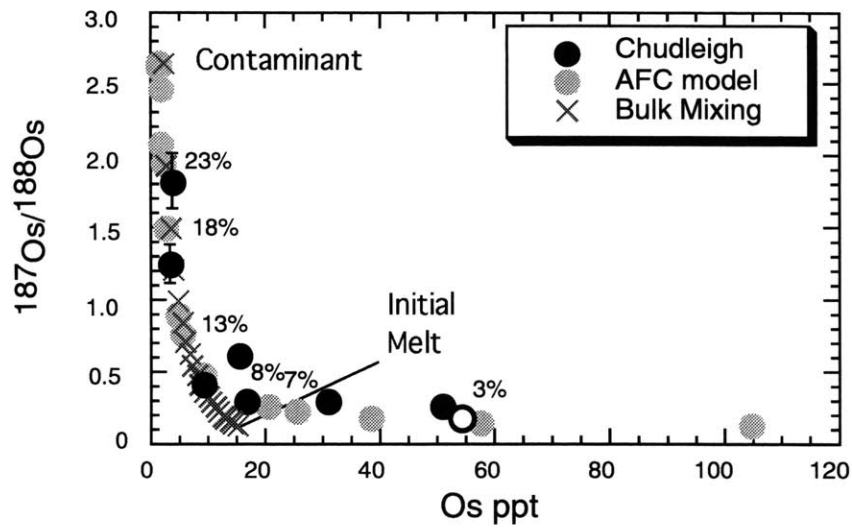


Figure 3 continuation

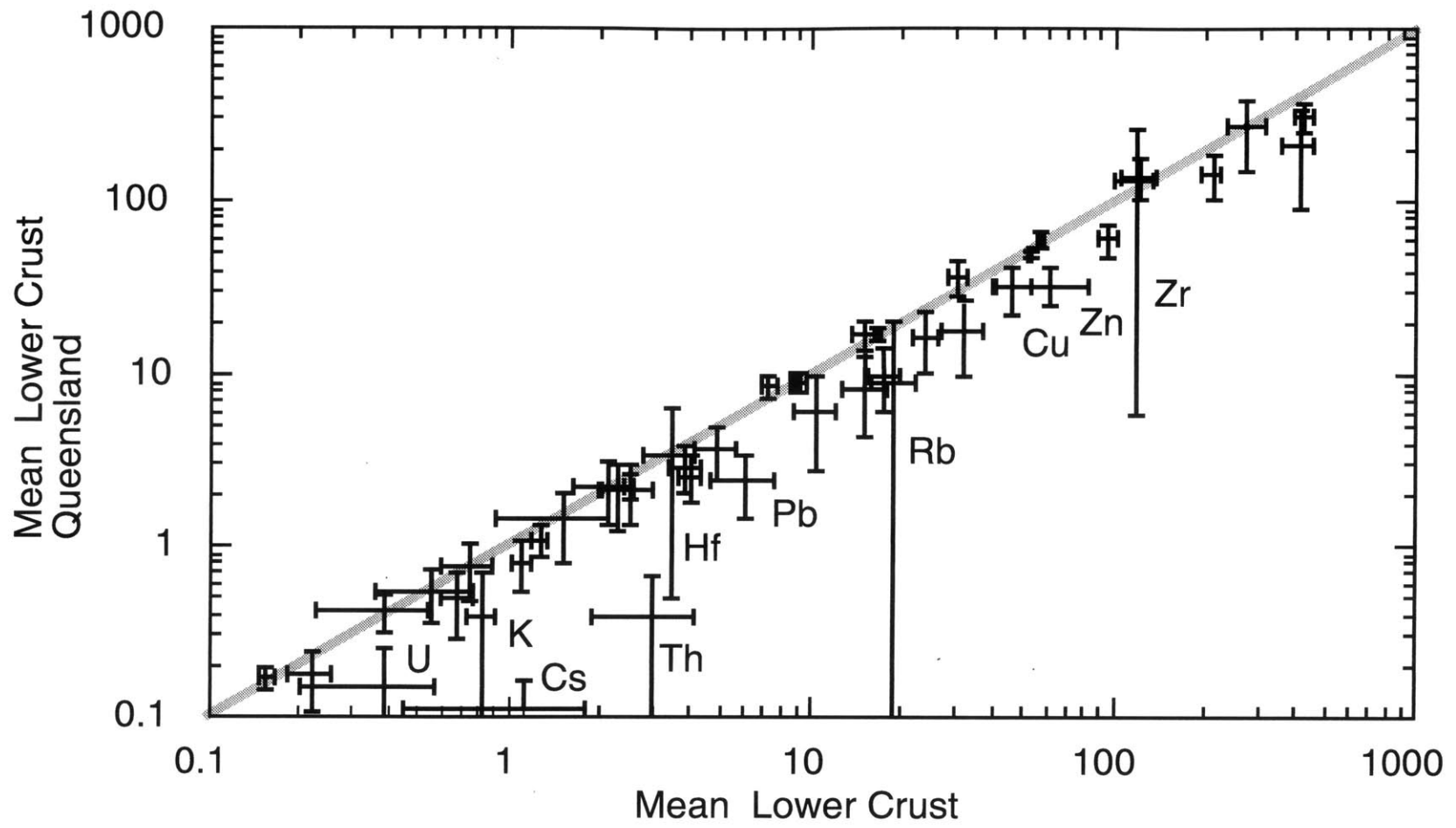


Figure 4

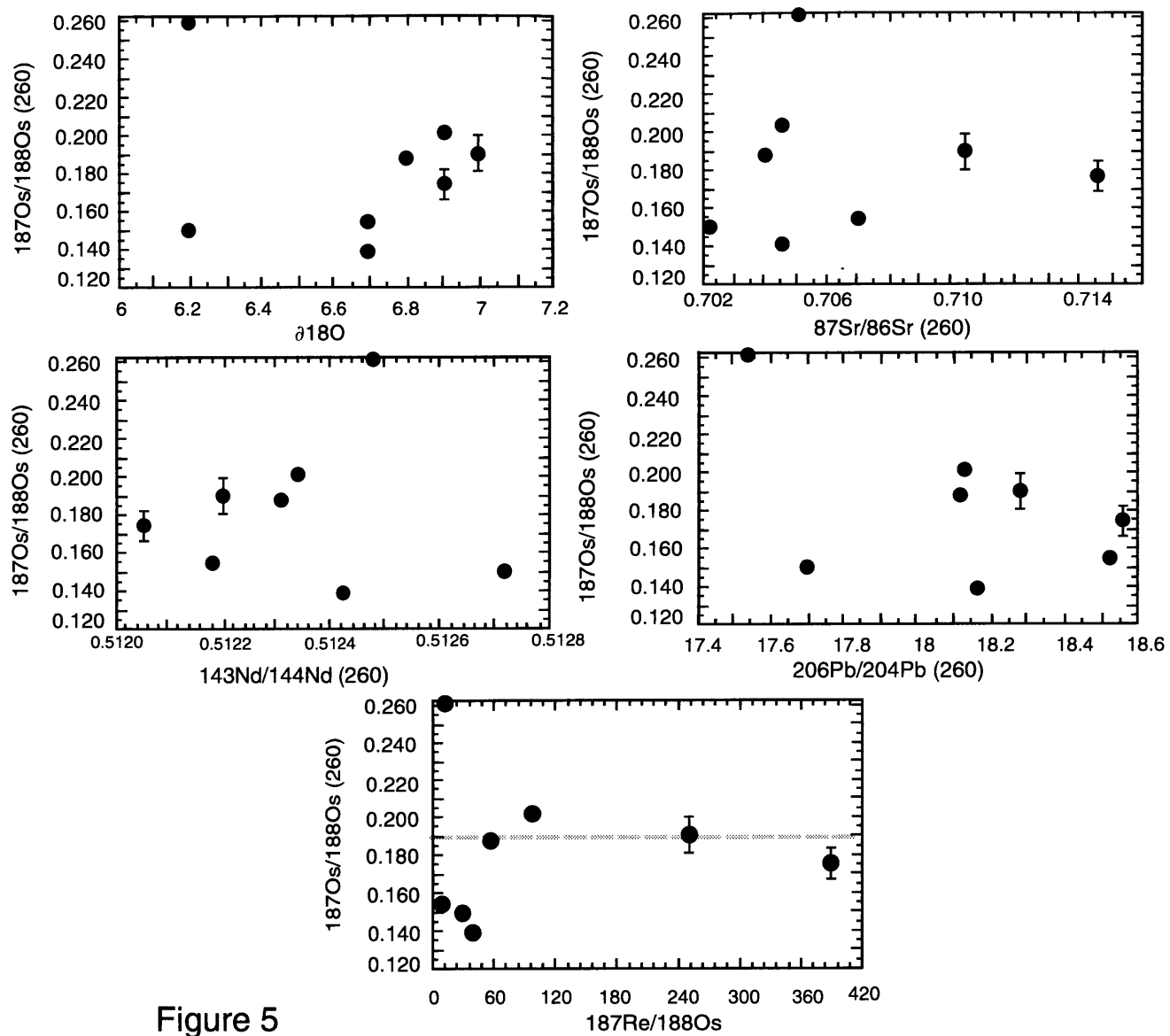


Figure 5

# AFC: Basalt-Lower Crust interaction

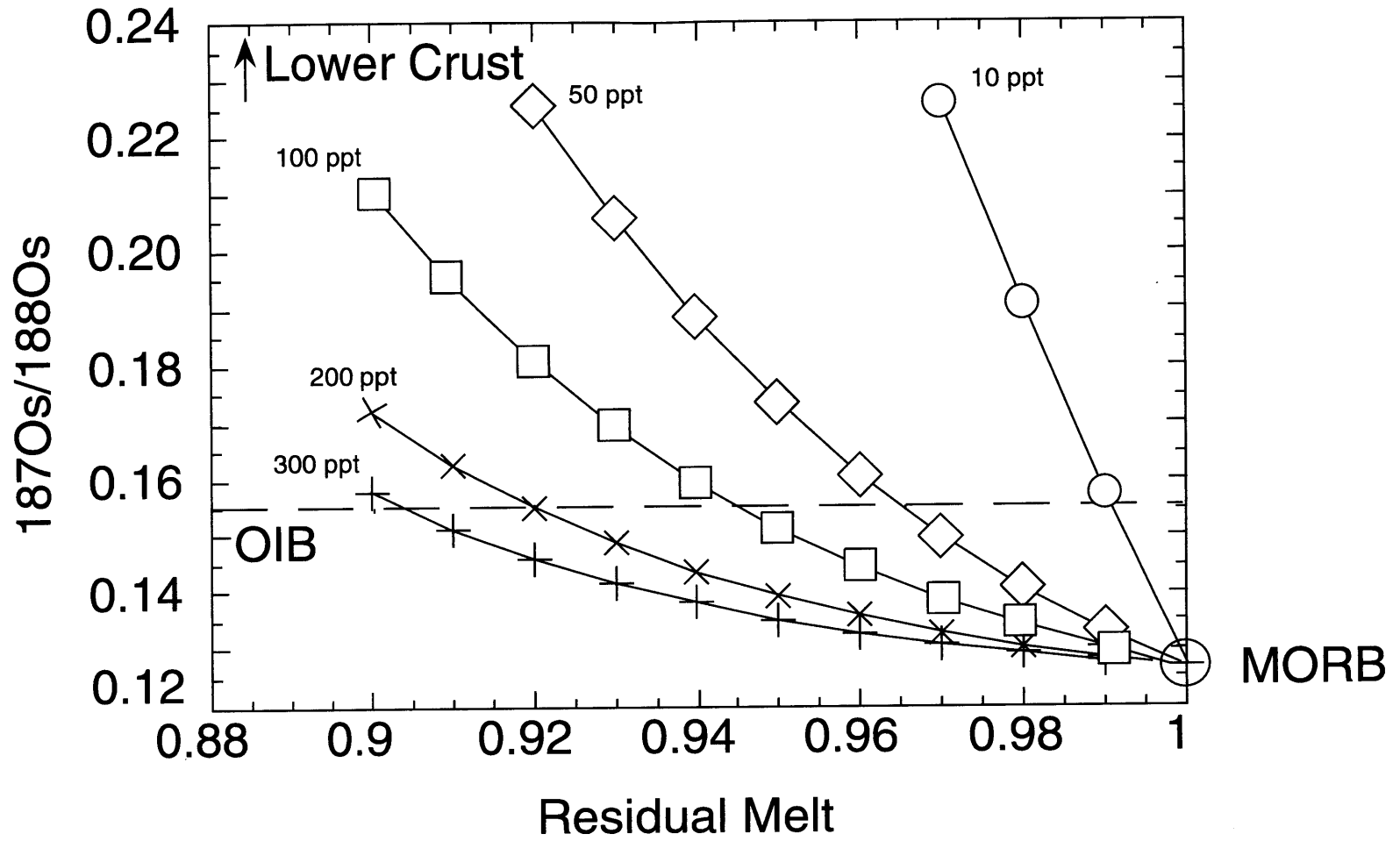


Figure 6



## 2. Chapter Two

### **Re-Os isotopic composition of the Horoman peridotite: Implications for the genesis of the Plagioclase Lherzolite and the layered structure**

#### *2.1 Abstract*

We report the Re-Os isotopic composition of the Horoman peridotite. Twenty samples were analyzed from a geochemically well-characterized 140-m section across the layered sequence ranging from plagioclase lherzolite to lherzolite to harzburgite (the Bozu section). The range in  $^{187}\text{Os}/^{188}\text{Os}$  ratios observed in the 140-m Bozu section is similar to that reported for other peridotitic massifs, suggesting that the process(es) responsible for the Re-Os isotope variation at the meter or the whole massif scales are the same. The Re-Os isotope systematics in the Horoman peridotite indicate that the Os isotope composition is controlled by the Re content, through radiogenic ingrowth, while the Re content is governed by the extent of depletion in “basaltic component” of the ultramafic rocks. Re depletion model ages of  $\approx 1.8$  Ga suggest that the age of melting is much older than  $\approx 1$  Ga, previously estimated based on Nd isotopes. However, the colinearity of data for mafic layers and the ultramafic samples in a Re-Os isochron diagram defines an apparent age of  $\approx 1$  Ga, similar to the ages defined by Nd isotopes. The colinearity between mafic and ultramafic rocks, the similar “ages” determined by Re-Os and Sm-Nd isotopes and the high Re/Os ratios in the most fertile peridotites plotting to the right of the geochron, indicate that the mafic layers and the ultramafic rocks are genetically related by a refertilization process which took place  $\approx 1$  Ga ago. The Re-Os systematics for other peridotite massifs indicate that refertilization of the lithospheric mantle seems to be a more widespread process than previously thought. Confirmation of the refertilization hypothesis could have important implications for our understanding of how the continental lithosphere forms and grows.

#### *2.2 Introduction*

The Horoman peridotite is a fault-bounded mantle slice, 8 X 10 X 3 km, emplaced in the southern end of the high-temperature low-pressure Hidaka Metamorphic Belt in

Hokkaido, Japan (Niida, 1974; 1984). The Horoman massif has been the subject of numerous structural, petrologic and geochemical studies (Nagasaki, 1966; Komatsu and Nochi, 1966; Niida, 1974; 1984; Arai and Takahashi, 1989; Obata and Nagahara, 1987; Takahashi et al., 1989; Takahashi, 1991; 1992; Ozawa and Takahashi, 1995; Frey et al., 1991, Takahashi and Arai, 1989; Takazawa et al., 1992; 1994; 1996, 1999a, 1999b; Takazawa 1996; Yoshikawa et al., 1993; Yoshikawa and Nakamura, 1998; Morishita and Kodaera, 1998). However, the process (or processes) responsible for the formation of the layering, as well as the origin of the fertile plagioclase lherzolite, still remains controversial.

Several hypotheses have been proposed to explain the formation of the layering:

1) polybaric melting of an upwelling fertile mantle (plagioclase lherzolite), followed by subsolidus thinning and folding of large-scale simple compositional bands (Takazawa, 1996; Yoshikawa and Nakamura, 1999)

2) formation of cumulates through fractional crystallization (Niida, 1984).

3) melting and melt segregation of an homogeneous mantle, where the depleted rocks represent residues of melting and the plagioclase lherzolite represents the region of melt accumulation (Obata and Nagahara, 1987)

4) melt segregation processes governed by suction associated with local melting along a fracture, during continuous ascent of a mantle peridotite (Takahashi, 1992).

In summary, the origin of the layered structure has been explained either by melt depletion processes or by a combination of melting and melt accumulation. Thus, depending on the model considered, the plagioclase lherzolite has been considered as an unmelted section of a fertile peridotite or as a region of melt accumulation.

The Re-Os isotope system provides a unique view of mantle processes. Parent and daughter elements for other isotopic systems, (Rb-Sr, Sm-Nd and U-Pb) behave incompatibly during mantle melting, being mainly controlled by silicate phases. In contrast, Re behaves as a moderately incompatible element, while Os is highly compatible during melting, and both Re and Os have chalcophile and siderophile affinities (Walker et al., 1989; Reisberg et al., 1991; Luck and Allègre, 1991; Martin, 1991; Carlson and Irving, 1994; Hart and Ravizza, 1997; Shirey and Walker, 1997, Handler et al., 1997, Burton et al., 1998). Consequently, mantle peridotites have high Os concentrations, and this is

presumed to make the Os isotope system more resistant to young metasomatic disturbances than the other isotopic systems. Thus, Re-Os isotope systematics can be a more reliable indicator of processes such as melting in peridotitic massifs than Sr, Nd and Pb isotopes (Reisberg et al., 1991; Luck and Allègre, 1991; Reisberg and Lorand, 1995; Roy Barman et al., 1996; Burnham et al., 1998).

The long term goal of studying the continental mantle lithosphere is to date and deconvolve the superimposed effects of processes responsible for the geochemical variations found in the peridotite massif. To deconvolve the complex history of peridotitic massif is one of the few avenues available to investigate the percolation and extraction of melts from the mantle. In this paper we report Re-Os isotope data for the Horoman peridotite. We selected a geochemically well-characterized 140 m section across the layered peridotite with rock types ranging from plagioclase lherzolite to lherzolite to harzburgite (the Bozu section); two mafic layers were also included. Takazawa and co-workers have studied this section and the mafic layers from the Horoman massif in detail, reporting major, trace element and isotopic compositions of whole rocks and mineral separates (Takazawa, 1996; Takazawa et al., 1992a; 1992b; 1993; 1994; 1995; 1996, 1999a, 1999b). The main objective of this work is to determine the origin of the plagioclase lherzolites (unmelted fertile peridotite or a region of melt impregnation and accumulation?) and the role played by the mafic layers.

### **2.2.1 Geological Background**

The Horoman peridotite is a massif tectonically emplaced in the south-western tip of the high-temperature low-pressure Hidaka metamorphic belt, Hokkaido, Japan (Figure 1). To the west, the Horoman ultramafic complex is in fault contact with an peridotite section composed of greenschist to amphibolite-facies metamorphic rocks (Miyashita, 1983; Komatsu et al., 1994). To the east and south, the massif is in tectonic contact with a metamorphic sequence, ranging from biotite-schists to granulites, and intruded by granitoids, and representing a continental or island arc crustal section (Komatsu and Nochi, 1966; Niida, 1974; Osanai et al., 1991; Shimura et al., 1992; Komatsu et al, 1994, and references therein). A Rb-Sr age of  $23 \pm 1.2$  Ma obtained for a phlogopite-bearing spinel lherzolite is essentially identical to the oldest ages reported for the Hidaka metamorphic belt. This age may indicate the time of uplift for both the Horoman mantle sliver and the Hidaka metamorphic belt (Yoshikawa et al., 1993).

The Horoman complex consists of a cyclically layered sequence of plagioclase lherzolite, lherzolite and harzburgite with subordinate amounts of dunite, gabbro and pyroxenite (Figure 2). The massif has been divided into two main zones (Niida, 1974; 1984):

1) The Lower Zone, approximately 2 km thick, consists of cyclic layers of plagioclase lherzolite to lherzolite to harzburgite, with some dunites; the cycle repeats every 100 to 500 m. The abundance of mafic layers is sparse, with the exception of thin bands in the plagioclase lherzolite.

2) The Upper Zone, approximately 1 km thick, is characterized by a fine-scale layering (centimeter scale) of plagioclase lherzolite and harzburgite as the dominant rock types, with subordinate dunite.

In the Upper Zone, plagioclase lherzolite and mafic layers are more abundant than in the Lower Zone (Takazawa, 1996; Takazawa et al., 1999a). Moreover, the equilibration temperature, calculated using the composition of orthopyroxene porphyroclast cores, increases from the Lower Zone to the Upper Zone by approximately 100°C (Ozawa and Takahashi, 1995; Takazawa et al., 1996). The association of high equilibration temperature, and the high proportion of mafic layers and plagioclase lherzolite in the Upper Zone, suggests a genetic link between these rock types (Ozawa and Takahashi, 1995).

The compositional layering and the presence of fine grain size mineral aggregates (seams) define the foliation planes in the Horoman peridotite. The megascopic structures are characterized by a monoclinical layered structure in the northern part, a synclinal structure in the central part and a domal structure in the southern part of the massif (Niida, 1984). The border between the Upper and Lower Zones is continuous, without evidence for a thrust or shear zone (Niida, 1974; Frey et al., 1991; Takazawa, 1996; Takazawa et al., 1999a). In the same way, the boundary between the compositional layers is sharp on mesoscopic scale, but transitional over a few centimeters on a microscopic scale.

Petrographic studies show that the plagioclase in the plagioclase-lherzolite usually occurs in seams associated with spinel, pyroxenes, olivine and Ti-pargasite. The seams share a similar mineralogical association with the mafic layers (Niida, 1984; Takahashi and Arai, 1989; Takazawa et al., 1999a; 1999b). Moreover, in the plagioclase lherzolite, the seams and orthopyroxene porphyroclasts contain two-pyroxene, Cr-spinel symplectites.

Takahashi and Arai (1989) and Ozawa and Takahashi (1995) interpreted the symplectites as the breakdown product of garnet formed by subsolidus reaction during decompression. This is also consistent with the garnet pyroxenite composition of the Type I mafic layers (Takazawa et al., 1999a). Thus, the mafic layers and the seams from the plagioclase lherzolite seem to have a common mineralogy suggesting a genetic link between them.

Based on petrography and mineralogy, Takahashi (1991) divided the Horoman complex into three genetically different peridotitic suites:

1) The Main Harzburgite-Lherzolite suite, consisting of a harzburgite, spinel lherzolite and plagioclase lherzolite layered complex, with transitional boundaries, representing residues after various degree of magma extraction.

2) The Banded Dunite-Harzburgite suite, consisting of dunites and harzburgites with olivine orthopyroxenites, with sharp layer boundaries, representing cumulates from a Mg-rich magma.

3) The Spinel-rich Dunite-Wehrlite suite, consisting of dunites with large oblate chromian spinels representing cumulates from magma segregated from the surrounding Main Harzburgite-Lherzolite suite.

The earlier division of the complex given by Niida (1974; 1984) considers only the Main Harzburgite-Lherzolite suite of Takahashi (1991), and this is the main focus of our work.

In the Horoman peridotite, there are five different petrographic types of mafic layers (Niida, 1984; Shiotani and Niida, 1997; Takazawa et al., 1999a), of which two are dominant (Takazawa et al., 1999a):

Type I : plagioclase + Ti augite + olivine + orthopyroxene + Ti pargasite (or kaersutite) + green spinel + titaniferous magnetite + ilmenite + sulfide

Type II : plagioclase + Cr diopside + olivine + orthopyroxene + pargasite + spinel + magnetite + sulfide

The mafic layers vary from a few centimeters to several meters in thickness, are usually oriented parallel to the foliation plane of the peridotite and typically are in sharp contact with the ultramafic host (Niida, 1984, Takazawa et al., 1999a). Sometimes in the plagioclase-lherzolite, it is possible to observe that the large seams of mafic compositions

grade into Type I mafic layers (Niida, 1984), suggesting a genetic relationship between the seams and the Type I layers.

The primary mineral assemblages in the mafic layers indicate multiple episodes of melt injection over a wide range of pressure, from the garnet stability field (Type I) to the plagioclase stability field (Type II) (Takazawa et al., 1999a). The sequence of the mafic injection is controversial. Takazawa, (1996) initially considered a simple sequence for intrusion of the mafic layers, starting first in the garnet and ending in the plagioclase stability field. Field observations support this sequence, with Type I layers being highly discontinuous, indicating that they have experienced more deformation than the Type II layers, which are traceable for more than a kilometer. Lately, Takazawa et al. (1999a), proposed a second possible sequence of intrusion based on Nd model ages. They suggest that Type II layers intrude at low pressure, early in the evolution of the Horoman complex and Type I layers intrude later at high pressure, during the subduction of the Horoman complex.

### **2.2.2 Sampling the Bozu section**

We chose the Bozu section for our study, as it is a representative 140 m section from the Lower Zone. This section is a continuously exposed lithologic sequence ranging from harzburgite to lherzolite to plagioclase lherzolite, and it has been geochemically well-characterized, with integrated geochemical observations made on a scale of  $\mu\text{ms}$  to tens of meters (Takazawa et al., 1992; 1996; Takazawa, 1996; Takazawa et al., 1999b) (Figure 3a). Based on this study of the Bozu section, Takazawa and co-workers concluded that the layered structure was produced by melting of an upwelling fertile mantle (plagioclase lherzolite) approximately 1 Ga ago. In the lithosphere, the Horoman massif suffered metasomatism by a light-REE-rich melt/fluid, producing enrichment of incompatible elements in harzburgite, lherzolite and in a special type of plagioclase-lherzolite that Takazawa et al, 1999b defined as E-type (Figure 3b). The E-Type plagioclase lherzolite is enriched in MgO, Mg<sup>#</sup> and in the very incompatible trace elements (LREE), but is depleted in TiO<sub>2</sub>, Al<sub>2</sub>O<sub>3</sub>, CaO, Na<sub>2</sub>O and HREE contents similar to that of the lherzolites. During uplift of the massif, the peridotite was transformed from the garnet to spinel to plagioclase-facies mineralogy by subsolidus reaction.

We selected two harzburgite, two lherzolites, four E-type plagioclase lherzolites and ten plagioclase lherzolites for Re-Os isotopic systematics. In addition, we analyzed two

Type I mafic layers for Re-Os isotopes; one from the Lower Zone and one from the Upper Zone. With the exception of the plagioclase-lherzolites, where we analyzed all the samples available, the peridotite samples are all from the center of their respective layers, to avoid transitional areas.

### ***2.3 Analytical Techniques***

The 200 to 500 gram samples were sliced into several slabs using a diamond saw, and they were later polished using silicon carbide sand paper to remove contamination from the saw. The samples were washed with Millipore water in an ultrasonic bath and dried on a hotplate at 50°C. Finally, the samples were wrapped in plastic and hammered to small chips, and the chips were powdered in an agate shatterbox (Takazawa et al., 1999b).

The analytical techniques for Re-Os isotopes were described in Ravizza and Turekian (1989) and Hauri and Hart (1993). Re and Os were determined on separate powder splits. For the analysis of Os, 1 g of sample powder was fused by the NiS fire-assay technique. The sample powder was spiked with  $^{190}\text{Os}$  and mixed with high purity boric acid, Ni and S powders. A 4:1 or 2:1 mixture of flux to sample, for peridotite and mafic layer respectively, were fused at 1050°C to 1150°C. This step produced a NiS bead that scavenges Os from the melted rock. The bead is dissolved in 6N HCl and the solution is filtered. The filter is dissolved in an oxidizing  $\text{CrO}_3 + \text{H}_2\text{SO}_4$  acid solution, from which the  $\text{OsO}_4$  is distilled and trapped in HBr. The HBr solution is dried down to a 1-2  $\mu\text{l}$  drop and diluted with  $\text{H}_2\text{O}$  until a of 1N HBr slution is obtained. A single bead of Chelex 50 resin is added to the HBr solution to extract the Os. The Chelex bead is deposited in concentrated HBr, liberating the Os, and this is then loaded onto a Pt filament with  $\text{BaNO}_3$  as an activator. The Os concentrations and isotopic compositions were measured by negative thermal ionization mass spectrometry on NIMA-B (WHOI), with oxygen-enhanced emission and single collector analog detection using an electron multiplier and a dynamic collection routine.

For Re analysis, 1 g of sample powder is spiked with  $^{185}\text{Re}$  and digested in a mixture of HF- $\text{HNO}_3$ . Re is extracted by a simple ion exchange technique using AG1X8 anion resin. The Re is absorbed onto the resin from 0.5N  $\text{HNO}_3$  and eluted in 4N  $\text{HNO}_3$ . An identical second step is used to further purify the Re. The Re measurements were done

on a Finnigan MAT Element ICP-MS (at WHOI) by rapid peak hopping using electrostatic scanning.

In-run precision is  $2\sigma < 0.35\%$  for  $^{187}\text{Os}/^{188}\text{Os}$ ,  $2\sigma < 1\%$  for Os concentration and  $2\sigma < 2\%$  for Re concentration. Analytical blanks ranged from 0.2 to 0.26 for  $^{187}\text{Os}/^{188}\text{Os}$ , 1.2 to 1.4  $\text{pg g}^{-1}$  of flux for Os concentration, and 2 to 3  $\text{pg}$  total procedural blank for Re concentration. All samples have been corrected for blank contribution; The correction is insignificant for Re and Os contents or Os isotopic composition, with the exception of samples BZ-116 and BZ-125 which have an estimated  $2\sigma$  error of  $\approx 25\%$  for Re concentration due to the blank correction. Re and Os replicates on separated powder split differ by  $<10\%$  for Re,  $<2\%$  for Os concentration, and  $<0.5\%$  for  $^{187}\text{Os}/^{188}\text{Os}$ .

## 2.4 Results

Results for 20 ultramafic and mafic samples analyzed from the Horoman peridotite are listed in Table 1; whole rock major and trace elements compositions are reported by Takazawa et al. (1999a, 1999b). In the Bozu section, the  $^{187}\text{Os}/^{188}\text{Os}$  ratios of the ultramafic rocks ranges from 0.1158 to 0.1283, while the Re and Os concentrations vary from 7 to 350 ppt and from 2.9 to 5.2 ppb respectively. Thus, while the Re content varies by a factor of 50, the Os concentration varies by about a factor 2.

There are two important observations that we can make directly from this data set:

1) The  $^{187}\text{Os}/^{188}\text{Os}$  ratios obtained over a 140 m section of peridotite is as large as the range in Os isotopic composition reported for a number of peridotite massifs such as Beni Boussera, Baldisero, Lanzo, Lherz and Ronda (Roy-Barman et al., 1996; Reisberg et al., 1991; Reisberg and Lorand 1995), (Figure 4a).

2) The lowest and highest Os isotopic ratios found for the Horoman suite correspond to two samples separated by a distance of only 5 m. The samples are an E-type and a N-type (normal) plagioclase lherzolite respectively (Figure 4b).

The  $^{187}\text{Os}/^{188}\text{Os}$  ratios correlate negatively with the extent of melt depletion calculated for each rock type (Figure 5, also compare 3a and 4b). There is a gradual decrease in the Os isotopic composition from plagioclase lherzolites to lherzolites to harzburgites, suggesting a progressive time-integrated depletion in Re/Os.



$^{187}\text{Os}/^{188}\text{Os}$  ratios and Re contents define a negative correlation with MgO concentrations; however, the Os content does not correlate with either  $^{187}\text{Os}/^{188}\text{Os}$  ratios, or Re or MgO concentrations (Figure 6). The negative correlation between Re and MgO contents suggests that Re behaves as a moderately incompatible element during melting. This is supported by the positive correlation between Re and  $\text{Al}_2\text{O}_3$ , Sc and Yb contents, and between Re and the modal proportion of clinopyroxene (Figure 7). Moreover, the well-defined correlation (one to one) with the zero intercept between Re and Yb suggest that both elements have similar bulk partition coefficients (in agreement with the conclusion of Hauri and Hart, 1997). On the other hand, the small variation in Os concentration (factor of  $\approx 2$ ) indicates that Os behaves as a highly compatible element during mantle melting (Morgan, 1986; Hart and Ravizza, 1996).

The  $^{187}\text{Os}/^{188}\text{Os}$  ratios for Horoman peridotites define a positive correlation with Re content (Figure 7) and  $^{187}\text{Re}/^{188}\text{Os}$  ratios (Figure 8). The  $^{187}\text{Os}/^{188}\text{Os}$  ratios for the ultramafic samples are lower than the Os isotopic composition estimated for the Primitive Upper Mantle (PUM =  $0.1290 \pm 0.0009$ ; Meisel et al., 1996). Moreover, most of the samples have  $^{187}\text{Re}/^{188}\text{Os}$  ratios below PUM (PUM = 0.4 - 0.428; McDonough and Sun, 1995; Meisel et al., 1996; Morgan, 1986, Luck and Allègre, 1983). However, five N-Type plagioclase lherzolite plot to the right of the geochron and four of the five samples have higher  $^{187}\text{Re}/^{188}\text{Os}$  ratios than the estimates for PUM.

Two mafic layers were also analyzed for Re-Os isotopes (Table 1). In both cases they correspond to the Type I mafic layers (garnet pyroxenites) defined by Takazawa et al., (1999a). The Os isotopic ratios of these samples, ranging from 0.788 to 2.001, are much higher than the nearly chondritic ratios of the ultramafic rocks. The  $^{187}\text{Re}/^{188}\text{Os}$  ratios range from 34.8 to 86.36, a variation that is closely related to the Os isotopic range. In the mafic layers, the Re contents (3.5 - 0.23 ppb) are higher than the Os concentrations (0.245 - 0.035), and the range in Re content is a factor of 2 higher than that for Os.

Table 1 show model ages for the Horoman samples calculated in two different ways. One method ( $T_{MA}$ ), as defined by Luck and Allègre (1984), extrapolates the  $^{187}\text{Re}/^{188}\text{Os}$  back in time until the Os isotopic composition of the sample equals that for a model of the isotopic evolution of the mantle. This model is sensitive to problems of Re modification by metasomatism or melt addition. These processes may change the Re but are unlikely to change the Os of the peridotites. The high concentration of Os in peridotite

compared to basalts or fluids preclude any modification of Os contents, or even the isotopic composition, if the metasomatic process is relatively young. Thus, to avoid potential problems with open system Re behaviour, Walker et al. (1989), calculated “Re depletion model ages” ( $T_{RD}$ ). The calculations for this model are done in the same way as above, but assumes that the Re/Os ratios of the sample is equal to zero (Table 1 and Figure 9). Thus, this method gives minimum model ages for the Re depletion event. Two main observations can be made from the model ages of the Horoman samples:

1) Assuming a mantle source equal to PUM,  $T_{MA}$  gave implausible ages for some of the ultramafic rocks indicating that the Re-Os systems has not remained closed since melting. This is especially the case for the N-Type plagioclase lherzolites, which have  $^{187}\text{Re}/^{188}\text{Os}$  higher than PUM. In these cases, the model ages are in the future, indicating mainly the addition of Re to the peridotite. In contrast, the two mafic layers give reasonable model ages of 1.3 to 1.16 Ga, similar to the ages of  $\approx 1$  Ga reported for the depletion event in the Horoman peridotites by Takazawa (1996) and Yoshikawa and Nakamura (1999) using Nd isotopes. However, Takazawa et al. (1999a), based on Nd isotopes, considered the age for the Type I mafic layers to be young (80 Ma), while the Type II layers are considered to be old ( $\approx 0.8$ ) Ga. These results are contrary to the field observations which show that the Type I layers are very discontinuous and have experienced more deformation than the Type II layers, some of which are traceable for more than a kilometer (Takazawa et al., 1999a). The field relationship and the Re-Os model ages are more consistent with the hypothesis that the Type I mafic layers are ancient, as previously proposed by Takazawa (1996). Moreover, if the Type I mafic layers are 80 Ma old cumulates from a MORB source as proposed by Takazawa et al. (1999a), it would require an unrealistically high  $^{187}\text{Re}/^{188}\text{Os}$  ratio of  $\approx 1400$  to produce the measured  $^{187}\text{Os}/^{188}\text{Os} \approx 2$ . This would also require subsequent major modification of the Re/Os ratios, lowering it to the observed value of 86.

2) The Re depletion model ages ( $T_{RD}$ ) for the sample with the lowest  $^{187}\text{Os}/^{188}\text{Os} = 0.1158$  yields a minimum age for the melt depletion event of 1.86 Ga (Table 1 and Figure 9). This age constrains the age of melting to be much older than that previously estimated by Takazawa (1996) and Yoshikawa and Nakamura (1999).

The positive correlation defined by the Horoman peridotites in the isochron diagram ( $^{187}\text{Os}/^{188}\text{Os}$  versus  $^{187}\text{Re}/^{188}\text{Os}$ ) yields an apparent “age” of  $1.04 \pm 0.4$  Ga (Figure 8). If the

two mafic layers are included in the regression, the apparent “age” is  $1.16 \pm 0.28$  Ga (Figure 10). These apparent “ages” again are very similar to ages reported for the ultramafic rocks based on Nd isotopes, by Takazawa (1996) and Yoshikawa and Nakamura (1999). Moreover, it suggests that Type I mafic layers are  $\approx 1$  Ga, much older than 80 Ma.

## ***2.5 Discussion***

### **2.5.1 What controls the Re/Os systematics of the Horoman peridotite?**

The correlation between  $^{187}\text{Os}/^{188}\text{Os}$  ratios and Re content (Figure 7) and the lack of correlation of Os isotopes with Os concentrations (Figure 6) suggest that the isotopic composition of the samples is mainly controlled by the variation in Re content. The large variation of Re (a factor of 50) compared to that of Os (factor of less than 2) indicates that the Re/Os ratios in Horoman are mainly controlled by the variations in Re. Thus, the Os isotopic composition is mainly related to the variation of Re/Os ratios (Re content!) through ingrowth.

Given that the variation in Re content controls the Os isotopic composition, the next step is to determine what governs the variation in Re content. The negative correlation between Re and MgO and the positive correlation between Re and  $\text{Al}_2\text{O}_3$ , Sc and Yb contents, and the modal proportion of clinopyroxene, indicates that Re is mainly controlled by silicate phases. The variation in  $\text{Al}_2\text{O}_3$ , Sc and Yb contents, and the modal proportion of clinopyroxene in the ultramafic rocks indicates a change in the proportion of a basaltic component in the Horoman peridotites. Thus, the observed Re variation in Horoman was not produced by an exotic sulfide component as proposed by Burnham et al. (1998) in the Pyrenees, but by the variation in the proportion of a basaltic component between the different rock types. In contrast to the suggestion of Burnham et al. (1998) that Re is controlled by sulfides, in the Horoman peridotite the correlation of Re with Yb contents defines a linear trend that tends toward the origin. This simple observation indicates that both elements had similar bulk partition coefficient (D) during melting (Figure 7). Similar D's (bulk partition coefficient) for Yb and Re, assuming that the Re content is controlled by sulfides, will require an extraordinary set of coincidences. For example, the variation in the proportions of sulfides and silicate phases in the mantle residue have to be compensated by the change of  $D_{\text{min/melt}}$  (partition coefficient between a mineral phase and the melt) for Re and Yb between melt-crystal and melt-sulfide to keep the similarity in D's during melting.

Moreover, Burton et al. (1998), by analyzing mineral separates from spinel and garnet peridotites, confirmed that the Re budget is almost exclusively controlled by silicate phases, especially garnet when present. The work of Burton et al. (1998) and the correlation found between Yb and Re for the Horoman peridotites, is confirmed by the experimental data of Righter and Hauri (1998) which proves that Re is compatible in garnet during melting. Although garnet is not present in Horoman peridotite, in the plagioclase lherzolite, the seams and orthopyroxene porphyroclasts contain two-pyroxene, Cr-spinel symplectites. Takahashi and Arai (1989) and Ozawa and Takahashi (1995) interpreted the symplectites as the breakdown product of garnet formed by subsolidus reaction during decompression.

In summary, the Os isotope ratios in the Horoman massif are controlled mainly by the variation in Re content, and the Re concentration is governed by the variation of a silicate phase (mainly garnet), which from now on we will call the “basaltic component”.

### **2.5.2 What is the process (or processes) responsible for the variation in the basaltic component in the Horoman peridotite?.**

There are two simple processes that can produce the observed variation in “basaltic component”: melt removal (melting) or melt addition (refertilization) to the peridotites. We can make use of the Re-Os systematics to differentiate between these two processes, and to evaluate the origin of the layered structure and the plagioclase lherzolite. There are two main hypotheses:

a) the layered structure is produced by melting of a fertile mantle represented by the N-Type plagioclase lherzolite (Takazawa et al., 1999b; Yoshikawa and Nakamura, 1999; Takahashi, 1992)

b) the layered structure results from a combination of melting and melt accumulation, where the N-Type plagioclase lherzolite is the zone of melt accumulation (Obata and Nagahara, 1987).

Clearly melting has played a fundamental role in defining the geochemical variation of the Horoman peridotites (Takazawa et al, 1999b; Yoshikawa and Nakamura, 1999 and references therein). Thus, the decrease in Re and Os isotopic composition on going from the N-Type plagioclase lherzolite to the harzburgite can, in principle, be explained by progressive extraction of a basaltic component. This hypothesis can explain the correlation between Re, Os isotopic composition, MgO, modal clinopyroxene and  $Al_2O_3$ , Sc and Yb

contents. However, the melt removal hypothesis has difficulty explaining the following observations:

1) The high Re/Os ratios observed in five N-Type plagioclase lherzolites plotting to the right of the geochron (Figure 8). Re/Os ratios higher than PUM and Chondrites Re/Os are difficult to explain by melt depletion processes. Although the value for the PUM Re/Os can be questioned, all the work done up to now has concluded that the Re/Os ratios and Os isotopic composition for PUM is chondritic (McDonough and Sun, 1995, Meisel et al., 1996 and references therein). Moreover, new platinum group element (PGE) data for the Horoman peridotite, especially the supra-chondritic Pd/Ir ratios, do not support the melting hypothesis (Rehkämper et al, 1997).

2) the small scale layering of the peridotitic massif. The samples with the highest and the lowest Os isotopic composition are only 5 m apart (Figure 4b). This range in Os isotopic composition observed in the Horoman peridotites overlaps with the reported range in Os isotopes for a compilation of peridotite massifs (Figure 4a). Although it can be argued that until now there is not a good understanding of the scale of melt extraction, to explain a variation in depletion of basaltic component from 8% to 19% over a 5 m distance seems excessive (Figure 3a and 4b).

These two obstacles for the simple melt removal can be easily explained by the refertilization (melt addition) model. The addition of a basaltic component to a peridotite can explain the high Re/Os ratios in the five plagioclase lherzolites, plotting to the right of the geochron, and can also explain the variation in Os isotopic composition over very short distances. Moreover, such refertilization processes would account for the correlation between Re, Os isotopic composition, MgO, modal clinopyroxene and  $Al_2O_3$ , Sc and Yb contents.

In the Horoman peridotite the linear variation of  $Na_2O$  as a function of MgO and the high content of HREE in the N-Type plagioclase lherzolites (compared to melting models) suggest that melting did not take place in the spinel stability field (Takazawa et al., 1999b; Yoshikawa and Nakamura, 1999). To model the  $Na_2O$  and HREE contents, Takazawa et al., (1999b) considered melting in the garnet stability field. Because the variation of HREE in the Horoman rocks requires a bulk  $D < 1$ , consideration of melting in the garnet field required a very small amount of garnet in the initial fertile peridotite (Takazawa et al., 1999b). However, no model has been presented to explain the variation of all the major

elements. In the Horoman massif, the linear variation in major elements as a function of MgO and the high HREE content in the most fertile plagioclase lherzolites can also be explained by refertilization process, as we show below in point 2.5.3.

The main reason that previous studies have rejected the hypothesis of refertilization in peridotitic massifs is that peridotites are not enriched in incompatible trace elements relative to estimates for primitive mantle. There is no physical reason why the addition of the “basaltic component” should stop at a level just necessary to compensate for the initial depletion (Burnham et al., 1998, Takazawa et al., 1999b). This is an important point, and the answer can give important clues regarding melting processes, and on the type of “basaltic component” that produces the refertilization. To answer this question, we make two simple observations: First, previous works such as Johnson et al. (1990), have shown that melting of ocean ridge mantle approaches a fractional melting process, where the melt readily separates from the source, and accumulates prior to eruption. If this is the case, after 20 to 25% of melting (common estimates for harzburgites-lherzolites) the level of depletion would have stripped most of the incompatible trace elements from the residual mantle, especially the most incompatible elements. Second, there is no reason why the refertilization has to be identified with a basaltic melt phase. The “basaltic component” could be in the form of cumulates derived from basalt that passes through the lithosphere, cools and undergoes fractional crystallization. Thus, our “basaltic component” can be a mafic cumulate. The association between peridotites and mafic cumulates is common in peridotite massifs. Peridotitic massifs such as Horoman contain mafic layers that have been classified as mafic cumulates (Takazawa et al., 1999a). Clearly, to enrich a  $\approx 20\%$  depleted peridotite in incompatible trace elements beyond the levels in primitive mantle will require such a large addition of a mafic cumulate ( $\gg 20\%$ ) that the mixture will no longer be a peridotite. In reality, certain elements that are moderately incompatible during melting of the peridotite and highly compatible in the crystallizing minerals forming the cumulates from the percolating basalts, can be enriched beyond the primitive mantle values during the refertilization. For elements that are highly incompatible during melting, their concentration in the depleted peridotite will be so low that it will be impossible to refertilize to values above primitive mantle. In contrast, elements that are moderately incompatible during melting can reach concentrations above the values for primitive mantle by the addition of reasonable amount of mafic cumulates. However, only some elements are moderately incompatible during melting and compatible in the crystallizing phases forming the

cumulates. For example, let us assume that the peridotite melts in the spinel stability field, and that the percolating basalt forms cumulate of plagioclase or garnet; therefore Sr, HREE and Re will be moderately incompatible during melting and strongly compatible (Sr in plagioclase and HREE and Re in garnet) in the cumulate respectively. Thus, these elements can provide some clues regarding the refertilization process.

Several observations indicate that mixing between depleted peridotites, such as the E-Type plagioclase lherzolite, and mafic layers (Type I) provides a mechanism for the formation of the N-Type plagioclase lherzolites in the Bozu section.

1) The mineralogy of the seams in the plagioclase lherzolites are similar to that of the Type I mafic layers. The present mineralogy of the seams and mafic layers do not represent primary igneous phases, but is a mineral association produced by subsolidus reaction at lower pressure (Takazawa et al., 1999a). The primary mineralogy of the mafic layers can be calculated based on their major and trace element compositions. Thus, Takazawa et al., (1999a) determined that Type I mafic layers were garnet-pyroxenites. The presence of garnet in the mafic layers is consistent with the presence of orthopyroxene-clinopyroxene-spinel symplectites in the seams, produced by decompression of garnet during subsolidus reaction (Takahashi and Arai, 1989). Moreover, garnet is the main phase that controls the Re content of the rock (Richter and Hauri, 1998; Burton et al, 1998), suggesting that the addition of garnet-pyroxenite to a depleted peridotite is responsible for the high Re/Os ratios in the plagioclase lherzolite.

2) The high Re/Os ratios of the Type 1 mafic layers support this hypothesis, explaining the N-Type plagioclase lherzolite plotting to the right of the geochron by addition of garnet pyroxenite to a depleted peridotite.

3) The colinearity of the Type I mafic layers and the ultramafic rocks in a Re-Os isochron diagram suggest that these rocks are genetically related. The Re-Os isotopes of these rocks define an age of  $\approx 1.16 \pm 0.28$  Ga., similar to the reported age based on Nd isotopes (Takazawa 1996; Yoshikawa and Nakamura, 1999); this also supports the hypothesis that the mafic layers and the peridotites are genetically linked (Figure 10)

4) The field observation that Type I mafic layers are dismembered in the peridotite, passing transitionally from mafic layers to form seams in the plagioclase lherzolite (Niida, 1984), supports the hypothesis of refertilization.

All of these observations endorse the Type I mafic layers as a suitable mafic end member for the mixing hypothesis. The field relationships between the mafic layers and the plagioclase lherzolite are especially supportive of this hypothesis.

### **2.5.3 Can the mixing hypothesis explain the major and trace element contents of the fertile plagioclase lherzolite?**

The question now is can the hypothesis of mixing between Type I mafic layers and E-Type plagioclase lherzolite explain the major and trace element contents of the N-type plagioclase lherzolite? To answer this question, we will model the N-Type plagioclase lherzolite as a mixture between E-Type plagioclase lherzolite and Type I mafic layers. It is a very difficult task to define the composition of the mixing components. Our simple approach is to use the mean composition of both the E-Type plagioclase lherzolite and the Type I mafic layers analyzed by Takazawa et al, (1999a and b). Moreover, by averaging the measured compositions we avoid the problem of dealing with any internal chemical differentiation which has affected the Type I layers (Takazawa et al., 1999a). Because the most depleted samples have undergone strong metasomatism (Takazawa et al., 1992), it is difficult to define a pristine depleted peridotite end-member (Figure 11). Therefore, we will avoid incompatible elements affected by metasomatism and model only the major elements, and the trace elements from Gd to Lu, including Cr and Sc. Note that the major elements K and P have been affected by metasomatism.

First, the data are normalized following the scheme of Allègre et al, (1995) so that each element is given a comparable weight. Each parameter used is normalized to its standard deviation, multiplied by the concentration range and divided by the mean concentration and the analytical error. Allègre et al. (1995), argued that the elements that have larger ranges in concentration will carry more weight in the regression. Clearly, the elements that have a larger concentration range will be better indicators of the processes responsible for the geochemical variation. To model the mixing process, we consider the simple model presented by Cantagrel et al. (1984); they express the mixing equation

$$Ca*(1-Xb)+ Cb*Xb=Cm, \text{ as}$$

$$(Cb-Ca)*Xb= (Cm-Ca)$$

where Ca, Cb and Cm are the concentrations of an element in the depleted peridotite, the mafic layer, and the mixture (fertile peridotite) respectively; and Xb is the proportion of



mafic layer in the mixture. We have defined the composition of the depleted lherzolite, the mafic layers and the mixture (fertile peridotite), so we can plot (Cm-Ca) as a function of (Cb-Ca), where each point represents one element. If the fertile peridotite is produced by mixing between a depleted peridotite and a mafic layer, the points in the figure will define a line passing through the origin with a slope (Xb). The slope represents the proportion of the mafic component in the mixture (Figure 12). Thus, we apply a simple linear regression and obtained the proportion of the mafic component in the mixture, for each N-Type plagioclase lherzolite. For the uncertainty in the slope (proportion of mafic layers), we propagated the analytical errors reported by Takazawa et al, (1999a and b) and the error in the linear regression.

Figure 13 show the calculated values for all the elements normalized to their measured values for each of the 10 N-Type plagioclase lherzolites. Surprisingly, the mixing model reproduces the measured major and trace element composition for the fertile plagioclase lherzolite quite well simply by using the same two end member component and only changing the proportion of the mafic component in the mixture. As we previously mentioned, K and P are the two elements that are not well reproduced by this model. We believe that the main reason for this is the K and P contents chosen for the depleted peridotite end member are too high, as result of metasomatism (Figure 11).

#### **2.5.4 The age of the melting and mixing events**

The ages of the melting and mixing events are not well defined. Up to now the age of the depletion event was considered to be  $\approx 1$  Ga, based on Nd isotopes (Takazawa 1996; Yoshikawa and Nakamura, 1999). In contrast, Re-Os depletion model ages indicates that the minimum age for the depletion event has to be older than 1.8 Ga (Figure 9, Table 1). The ultramafic and mafic rocks from the Horoman massif define an apparent age of  $1.16 \pm 0.28$  Ga in the Re-Os isochron diagram (Figure 10), similar to the age of  $\approx 1$  Ga based on Nd isotopes reported by Takazawa (1996) and Yoshikawa and Nakamura, (1999). Moreover, the Rb-Sr isotope systematics on whole rocks reported by Yoshikawa and Nakamura, (1999) also define an apparent age of  $\approx 1$  Ga. Thus, the consistency of the  $\approx 1$  Ga age defined by the three isotopic systems suggests that this age has a geologic meaning. The colinearity between the mafic layers and the peridotites in the Re-Os isochron diagram and the  $\approx 1.1$ -1.3 Ga Re-Os model ages for the mafic layers suggest that the mafic layers are genetically related to the peridotite through a refertilization process  $\approx 1$

Ga ago. Another interpretation is that the melt depletion and refertilization are two parts of the same process as envisioned by Obata and Nagahara (1987). These authors considered melting and melt segregation of an homogeneous mantle, where the depleted rocks represent residues of melting and the plagioclase lherzolite represents the region of melt accumulation. If this is the case, 1Ga is the time of depletion and refertilization, and the Re depletion model age older than 1.8 Ga indicates that the Horoman peridotite was a depleted peridotite previous to the depletion and refertilization event at 1Ga.

In synthesis, the Re-Os isotope systematics indicate that the age of melting event in Horoman peridotites is  $\geq 1.8$  Ga (Re depletion model age), much older than the ages  $\approx 1$  Ga previously estimated based on Nd isotopes (Takazawa et al., 1999b; Yoshikawa and Nakamura, 1999). Re-Os isotope systematics, major and trace elements contents, and field and petrographic information are consistent with the hypothesis that the N-Type plagioclase lherzolite are produced by refertilization of the depleted peridotite by the addition of mafic cumulates (Type I) at  $\approx 1$  Ga. Refertilization of a depleted peridotite seems a more suitable process than melting to explain the variation in major elements and Os isotopes at the scale of a few meters. We propose that the refertilization process has played an important role in the formation of the layering, especially in the Horoman massif where the plagioclase lherzolite forms 60% of the total outcrop (Takazawa et al., 1999a).

#### **2.5.5 Is Horoman an unusual place or does the refertilization process play a role in other peridotite massifs?**

Several peridotite massifs have been studied with Re-Os systematics: Ronda, the Pyrenees (Lherz), Beni Bousera, Lanzo (Reisberg et al, 1991; Reisberg and Lorenz, 1995, Roy-Barman et al, 1996 Burnham et al. 1998, among others). However, with the exception of Ronda, Re-Os isotopic data combining both peridotite and the associated mafic rocks are sparse. Although the data are meager, the Re-Os systematics seem to indicate that a refertilization processes could have played an important role in these orogenic lherzolite massifs. Figure 14 shows whole rock Re-Os isochron diagrams for several peridotite massifs. In all of these cases, the most fertile samples have high Re/Os ratios, plotting to the right of the geochron. Moreover, the mafic layers are co-linear with the ultramafic rocks defining ages that are similar to the ages determined by Nd isotopes. The fact that both the Re-Os and Sm-Nd isotope systematics give similar ages suggests that the linear array on the isochron diagram is not only an indication of mixing but may also represent the age of refertilization. In Ronda, Reisberg et al, (1991), analyzed both garnet-rich and

spinel-rich mineral bands; when plotted on the isochron diagram, only the garnet-rich bands are colinear with the ultramafic rocks. These garnet-rich seams in Ronda have been interpreted to be the product of tectonic disaggregation of mafic layers in an isotopically depleted peridotite (Reisberg et al, 1989). These observations from Ronda mimic the observations made in Horoman. Two different processes have been invoked to explain the colinearity of the mafic layers and the ultramafic rocks in the Re-Os isochron diagram for the Ronda massif (Reisberg et al, 1991, Roy-Barman et al., 1996):

1) The marble cake model of Allègre and Turcotte (1986), where the mafic layers are remnants of subducted oceanic crust and the colinearity has no age significance.

2) The mafic layers represent both, cumulates formed during an ancient melting event and residues of earlier mafic layers that reequilibrate during that ancient melting event (Reisberg et al., 1991).

We discard the first model because a systematic study of the mafic layers in Ronda indicates that they are cumulates (Garrido and Bodinier, 1999). Moreover, the agreement between the Re-Os and Sm-Nd ages supports the hypothesis that the ages are significant. The second model is similar to our model for refertilization with the difference that Reisberg et al (1991) considered the isochron to record a melt depletion event, even though they proposed that the most fertile garnet peridotite was produced by refertilization through tectonic disaggregation of mafic layers. These authors also discarded mixing as an explanation because some of the mafic layers were not co-linear with the ultramafic rocks. Not all mafic layers are formed at the same time and therefore not all of them need be responsible for the refertilization process. In Ronda, the best examples of refertilization are the garnet-rich seams produced by disaggregation of garnet-rich mafic layers, which are colinear with the ultramafic rocks (Figure 14). Although Beni Boussera and Lanzo are less convincing examples, due to the few samples analyzed, they show the same tendency as that observed in Ronda, with the mafic layers being colinear with the ultramafic rocks. In the Pyrenees, only ultramafic rocks have been analyzed for both Re-Os isotopic systematics. There are Os isotopic data for mafic layers, there are no reported Re contents. We predict that the Pyrenees will show the same tendencies as Horoman, Ronda and the other peridotite massifs.

If the hypothesis of refertilization is confirmed, it will have important implications for understanding how the continental crust forms and grows. Any hypothesis of how the

crust grows will be different whether we consider that the ages represent melting events of asthenospheric material, or refertilization events of an old depleted lithosphere. Both processes will have consistency between the ages observed in the crustal material and those obtained from the peridotitic massifs. Thus, the ages could indicate either stabilization of both young crust and lithospheric mantle or the refertilization of a depleted lithospheric mantle already stable. For example, Griffin et al. (1998), concluded that subcontinental mantle lithosphere (SCML) becomes progressively less depleted from Archean through Proterozoic to Phanerozoic times. The broad correlation between the SCML composition and the crustal age implies quasi-contemporaneous formation of the crust and its lithospheric mantle. Moreover, this broad correlation would suggest that the extent of melting progressively decreased from the Archean to the Phanerozoic. In contrast to Griffin et al., (1998) hypothesis, we speculate that the correlation between depletion of the SCML and the mantle ages does not indicate formation and stabilization of the continental crust and its lithosphere, but indicates a process of refertilization of the SCML and a resetting of an already stable continental area. Thus, the correlation between mantle lithosphere and crustal ages and extent of depletion of the lithosphere may not have any implicit indication of a correlation between the extent of melting and time.

## ***2.6 Conclusions***

Re-Os isotope systematics in Horoman peridotites indicate that the Os isotopes are mainly controlled by variations in Re content, while Re contents are controlled by the variation of a “basaltic component” in the ultramafic rocks.

Re-Os isotopes, major and trace elements contents, and field and petrographic information are consistent with the hypothesis that the Horoman N-Type plagioclase lherzolites are produced by refertilization of a depleted peridotite, by the addition of a garnet-rich mafic cumulate.

Refertilization of depleted peridotite seems a more suitable process than melting to explain large variations in major elements and Os isotopes over very short distances. For example, the total range in Os isotopic composition occurs between two samples only 5 m apart. Therefore, we propose that refertilization processes have played an important role in the formation of the layering, especially in the Horoman massif, where the fertile plagioclase lherzolite forms 60% of the total outcrop.

Re depletion model ages of  $\approx 1.8$  Ga suggests that the age of melting is much older than  $\approx 1$  Ga previously estimated based on Nd isotopes (Takazawa et al., 1999b; Yoshikawa and Nakamura, 1999). On a Re-Os isochrone diagram the Type I mafic layers and the peridotitic samples are colinear and define an apparent age of  $\approx 1$  Ga, similar to the Nd ages; this suggests that the linear trend in the isochron diagram is not only an indication of mixing but has an age significance, and may represent the age of refertilization.

If the refertilization hypothesis is confirmed, it will have important implications for the understanding of how the continental lithosphere forms and grows.

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## 2.8 Table caption

**Table 1:** Re and Os concentration and Os isotopic composition of 20 samples from the Horoman Peridotite.  $T_{MA}$  are model ages calculated with respect to “primitive upper mantle”  $^{187}\text{Re}/^{188}\text{Os} = 0.428$  and  $^{187}\text{Os}/^{188}\text{Os} = 0.1290$  (Meisel et al., 1996; McDonough and Sun, 1995) and a decay constant of  $1.64 \times 10^{-11} \text{ y}^{-1}$ .  $T_{RD}$  are Re depletion model ages assuming  $^{187}\text{Re}/^{188}\text{Os} = 0$

## 2.9 Figure Caption

**Figure 1:** Simplified geologic map of the Hidaka metamorphic belt. After Komatsu et al., 1986; Takazawa et al., 1999a.

**Figure 2:** Geologic map of the Horoman peridotite. Note the position of the Bozu section in the Lower Zone. The Boundary between the Lower and the Upper Zone is marked by a dashed line. Inset shows the location of the studied area. After Takazawa et al., 1999a.

**Figure 3:** a) Degree of melting along the Bozu section, Horoman peridotite. Dashed vertical lines indicate the contacts between lithologies. The shaded regions in the middle of both lherzolite and plagioclase lherzolite layers stand for phlogopite-bearing lherzolite and E-Type plagioclase lherzolite respectively. The extent of melting is calculated by mass balancing the FeO and MgO contents between initial source and residual peridotite for each analyzed sample. Figure after Takazawa et al. (1999b); Takazawa (1996). b) Chondrite normalized REE contents for representative harzburgite (BZ-125L), lherzolite (BZ-143) and E and N-type plagioclase lherzolite (BZ-260 and BZ-253 respectively) from the Bozu section. Symbols as in Figure 3a. C1 chondrite values from Anders and Grevesse (1989). Data from Takazawa et al., (1999b). Open circle, N-type plagioclase lherzolite; close circle, E-type plagioclase lherzolite; open square, lherzolite; crossed square, harzburgite.

**Figure 4:** a)  $^{187}\text{Os}/^{188}\text{Os}$  in Horoman peridotites and in peridotites from different orogenic lherzolite massifs, compiled by Roy-Barman et al. (1996). Note that the range in Os isotopic composition obtained in a 140 section from the Horoman massif is similar to that observed in the compilation of other peridotite massifs (compilation does not include mafic layers). b)  $^{187}\text{Os}/^{188}\text{Os}$  versus distance along the Bozu section. Note that the extreme variation in Os isotopic composition occurs between an E- and N-Type plagioclase lherzolites separated by only 5m. Symbols as in Figure 3.

**Figure 5:**  $^{187}\text{Os}/^{188}\text{Os}$  versus degree of melting for samples from the Bozu section. The inverse correlation between Os isotopes and the extent of melt depletion indicate that both parameters are related by a single process: melting or refertilization of a depleted peridotite. The extent of depletion is calculated by mass balancing the FeO and MgO contents between initial source and residual peridotite (Takazawa et al., 1999b).

**Figure 6:**  $^{187}\text{Os}/^{188}\text{Os}$  ratios, Re and Os contents versus MgO content for samples from the Bozu section, Horoman peridotite. The samples define a negative correlation between Re, Os isotopes and MgO, but show no correlation between MgO and Os contents. This figure suggests that Re behaves as an incompatible element, while Os is highly compatible. Moreover, the negative correlation between  $^{187}\text{Os}/^{188}\text{Os}$  and MgO indicates that the Os isotopes are controlled by Re content through ingrowth.

**Figure 7:**  $^{187}\text{Os}/^{188}\text{Os}$  ratios,  $\text{Al}_2\text{O}_3$ , Yb, Sc contents and modal clinopyroxene versus Re content. The observed correlations indicate that the variation in  $^{187}\text{Os}/^{188}\text{Os}$  ratios is controlled by Re content through ingrowth, and the Re content is controlled by the variation in the proportion of “basaltic component” of the peridotite. Modal proportions of clinopyroxene (in %) have been calculated by mass balancing the mineral compositions with that of the whole rock (Takazawa et al., 1999b). Dashed line represents Re and Os content for PUM. Dotted square represent the composition of the PUM. Data from McDonough and Sun (1995).

**Figure 8:**  $^{187}\text{Os}/^{188}\text{Os}$  versus  $^{187}\text{Re}/^{188}\text{Os}$ . The samples define a rough positive correlation with an apparent “age” of  $1.04 \pm 0.4$  Ga. ( $2\sigma$ , York regression). Five N-Type plagioclase lherzolites have high  $^{187}\text{Re}/^{188}\text{Os}$  ratios plotting to the right of the geochron. Values for PUM and Geochron from Meisel et al. (1996). Values for Chondrite and Geochron from Smoliar et al. (1996). Dashed line as drawn infers a second stage deviation from a depleted mantle; the isochron cross the geochron at a lower values than that for present PUM, indicating that if the age of 1Ga represents a melting event, the mantle source is a depleted mantle.

**Figure 9:** Frequency diagram for Re depletion model ages for Horoman peridotite. The model ages were calculated using PUM values  $^{187}\text{Re}/^{188}\text{Os} = 0.428$  and  $^{187}\text{Os}/^{188}\text{Os} = 0.1290$  (Meisel et al., 1996). The Re depletion model ages reach values of up to  $\approx 1.8$  Ga., much older than the ages previously reported based on Nd isotopes. The Re depletion model ages should be considered as minimum ages.

**Figure 10:**  $^{187}\text{Os}/^{188}\text{Os}$  -  $^{188}\text{Re}/^{188}\text{Os}$  isochron diagram for Horoman samples including the two Type I mafic layers. All of the samples define an apparent “age” of  $1.16 \pm 0.28$  Ga (York regression), similar to the  $\approx 1$  Ga ages determined using Nd isotopes (Yoshikawa and Nakamura, 1999; Takazawa 1996). The colinearity between the mafic layers and the ultramafic rocks, especially for those plagioclase lherzolite plotting to the right of the geochron (see inset) suggest that the mafic and ultramafic rocks are genetically linked.

**Figure 11:** Primitive mantle normalized trace and major element contents for representative samples of the N-Type plagioclase lherzolite, E-Type plagioclase lherzolite and Type I mafic layers. We have also plotted the average for both the depleted E-Type plagioclase lherzolite and the mafic layers used as end-member in the mixing models (thick lines in gray). Note that in the depleted end-member (E-Type plagioclase lherzolite), the elements more incompatible than Eu have been modified (enriched) by metasomatic processes. Data from Takazawa et al., (1999a; 1999b). Plagioclase lherzolite E-Type (BZ-260); N-Type (BZ-250, BZ-254, BZ-256 and 62210); mafic layer (G2). Primitive Mantle values from McDonough and Sun, (1995).

**Figure 12:** Cm-Ca versus Cb-Ca. diagram to determine the proportion of the end member component in a mixture (Cantagrel et al., 1984). Each point represents one element used in the mixing model. Cm, Ca, Cb are the concentrations of each element in the mixture (N-Type plagioclase lherzolite), the depleted peridotite (E-Type plagioclase lherzolite) and in the mafic layer (Type I) respectively. “Xb” represents the slope of the line, which is the proportion of the mafic layer in the mixture. In this case we used one variable regression.

**Figure 13:** Calculated trace and major element contents from the mixing model, normalized by their measured value for each of the N-Type Plagioclase Lherzolites. Note that with the exception of P and K, the model reproduces the values for major and trace elements more compatible than Eu. We used the same end-members for all the samples, changing only the proportion of mafic layers in the mixture.

**Figure 14:** Re-Os isochron diagrams for a) Ronda, b) Beni Bouseras and Lanzo, and c) Pyrenees peridotites and mafic layers. In all of these cases, the most fertile peridotites in each massif have high  $^{188}\text{Re}/^{188}\text{Os}$ , plotting to the right of the geochrone. Moreover, the mafic layers are colinear with the peridotites in each massif, suggesting that the mafic and ultramafic rocks are genetically related probably through refertilization processes. The case of Ronda is interesting, where garnet-rich bands, believed to be disaggregated garnet-

bearing mafic layers, plot as the most fertile end member of the linear trend defined by the peridotites. Therefore, if the age has any significance, it may indicate the age of refertilization as in the case of the Horoman massif.

**Table 1**

Sample #	Distance (m)	Re (ppb)	Os (ppb)	Re/Os	<sup>187</sup> Os/ <sup>188</sup> Os	<sup>187</sup> Re/ <sup>188</sup> Os	T <sub>RD</sub> (Ga)	T <sub>MA</sub> (Ga)
Plagioclase Lherzolite								
N-Type Bz-263	132	0.205	4.57	0.045	0.1240	0.217	0.71	1.44
N-Type Bz-262	127	0.288	2.89	0.100	0.1253	0.480	0.53	-4.54
N-Type Bz-261	120	0.318	3.22	0.099	0.1259	0.476	0.44	-4.12
E-Type Bz-260	117	0.099	5.18	0.0191	0.1200	0.0918	1.27	1.61
E-Type Bz-259	115	0.147	3.11	0.047	0.1210	0.228	1.14	2.41
E-Type Bz-258	109	0.154	2.92	0.053	0.1201	0.254	1.25	3.04
E-Type Bz-257	103	0.151	4.65	0.0324	0.1158	0.156	1.86	2.90
N-Type Bz-256	101	0.353	2.91	0.121	0.1267	0.585	0.32	-0.89
N-Type Bz-255	97.7	0.270	3.48	0.077	0.1283	0.373	0.10	0.79
N-Type Bz-254	94.7	0.253	3.12	0.081	0.1226	0.391	0.90	9.66
N-Type Bz-253	89.3	0.310	4.02	0.077	0.1256	0.371	0.48	3.51
N-Type Bz-252	87.9	0.318	4.50	0.071	0.1260	0.341	0.43	2.07
N-Type Bz-251	86.2	0.237	3.44	0.069	0.1235	0.332	0.78	3.40
N-Type Bz-250	84.3	0.223	2.44	0.092	0.1246	0.441	0.63	-24.93
Lherzolite								
143	72.3	0.128	3.98	0.0322	0.1226	0.155	0.90	1.41
134	39.9	0.143	4.24	0.0339	0.1190	0.163	1.40	2.25
134R		0.161	4.31	0.0373	0.1196	0.180	1.32	2.27
Harzburgite								
116	9.1	0.0087	3.86	0.0022	0.1193	0.0108	1.37	1.41
125	4.1	0.0066	4.61	0.0014	0.1207	0.0069	1.17	1.19
Type 1 Mafic Layers								
G9		0.233	0.035	6.65	0.788	34.8		1.16
G2		3.53	0.245	14.4	2.010	86.4		1.32

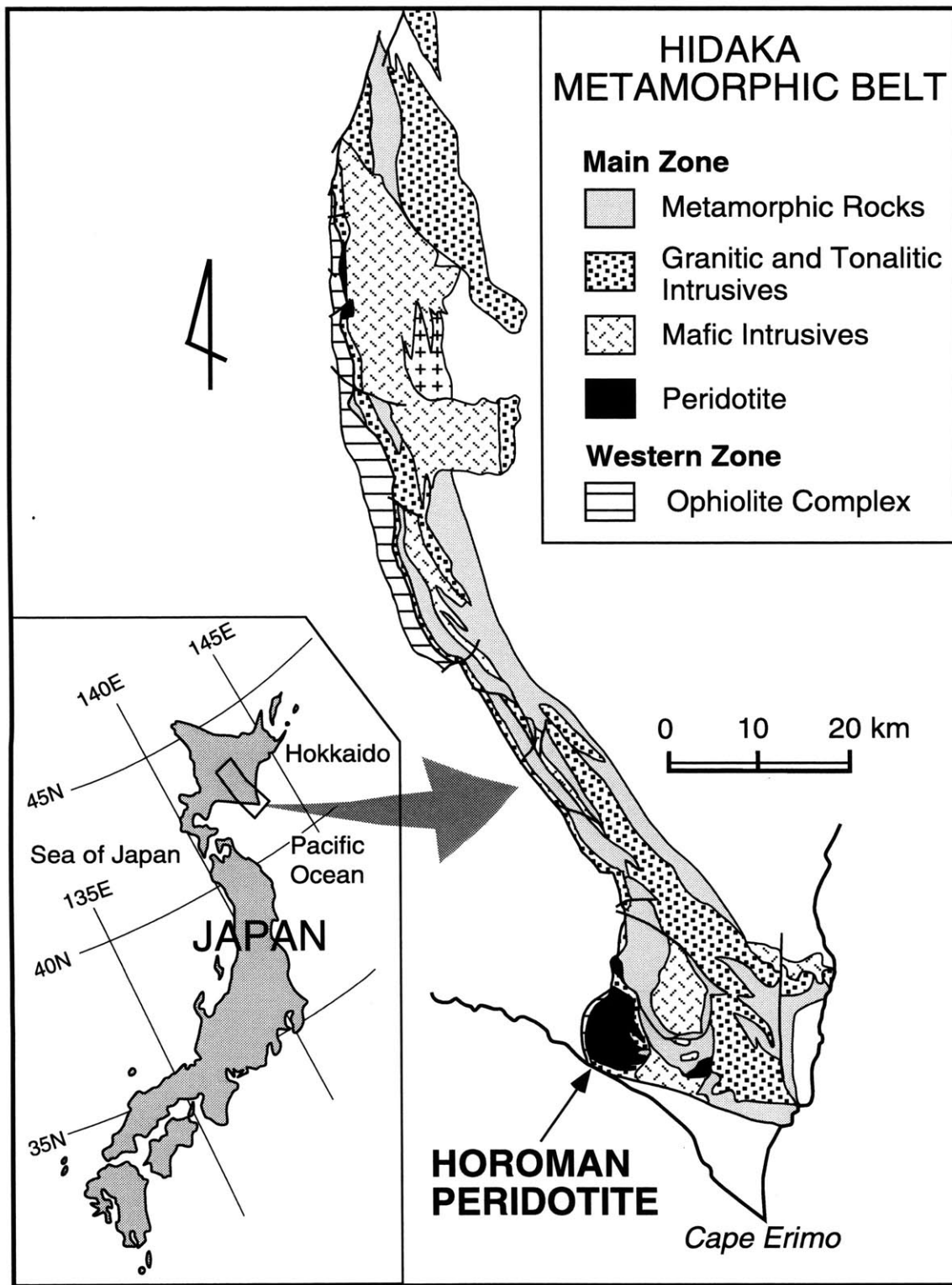


Figure 1



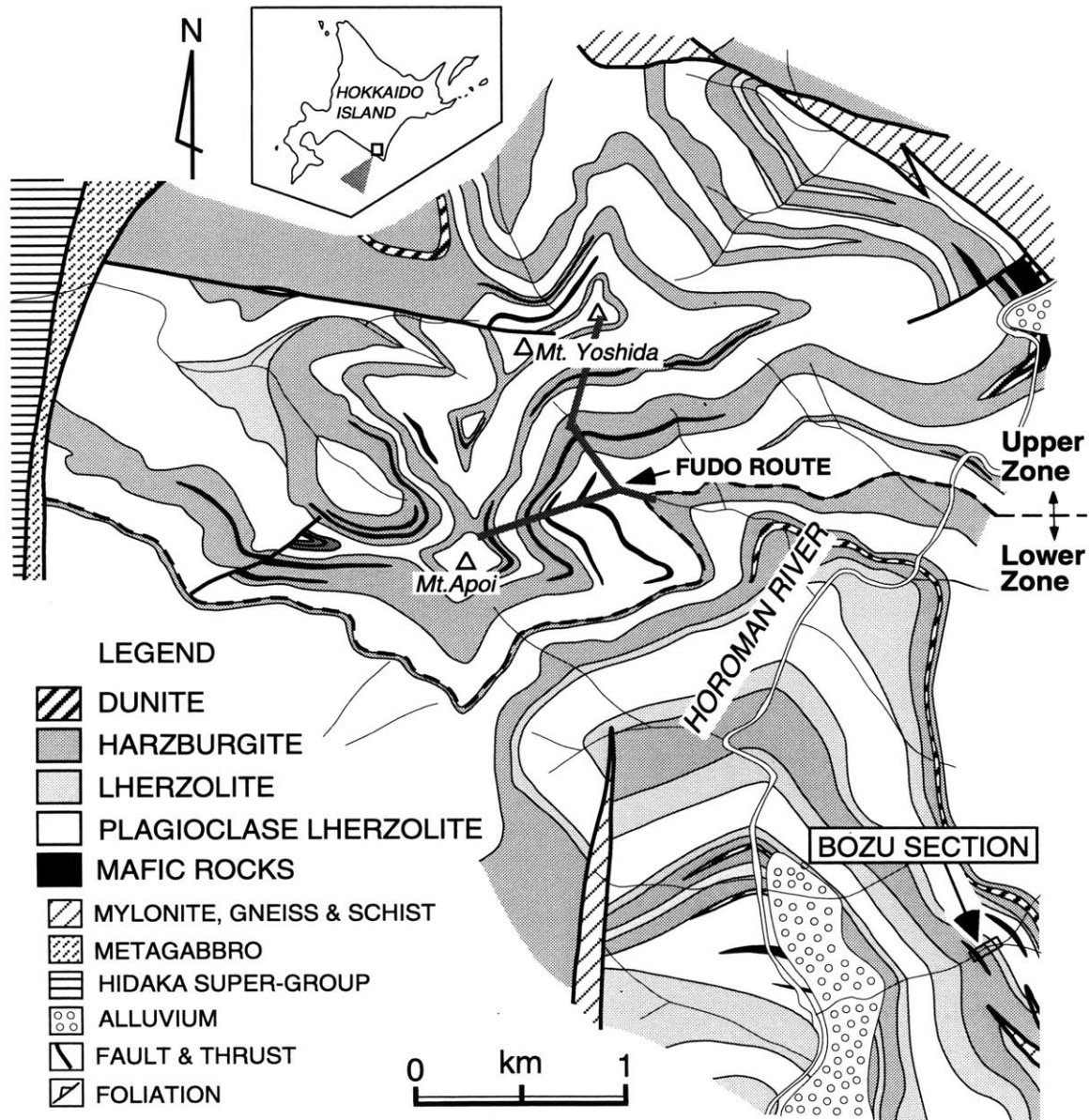


Figure 2

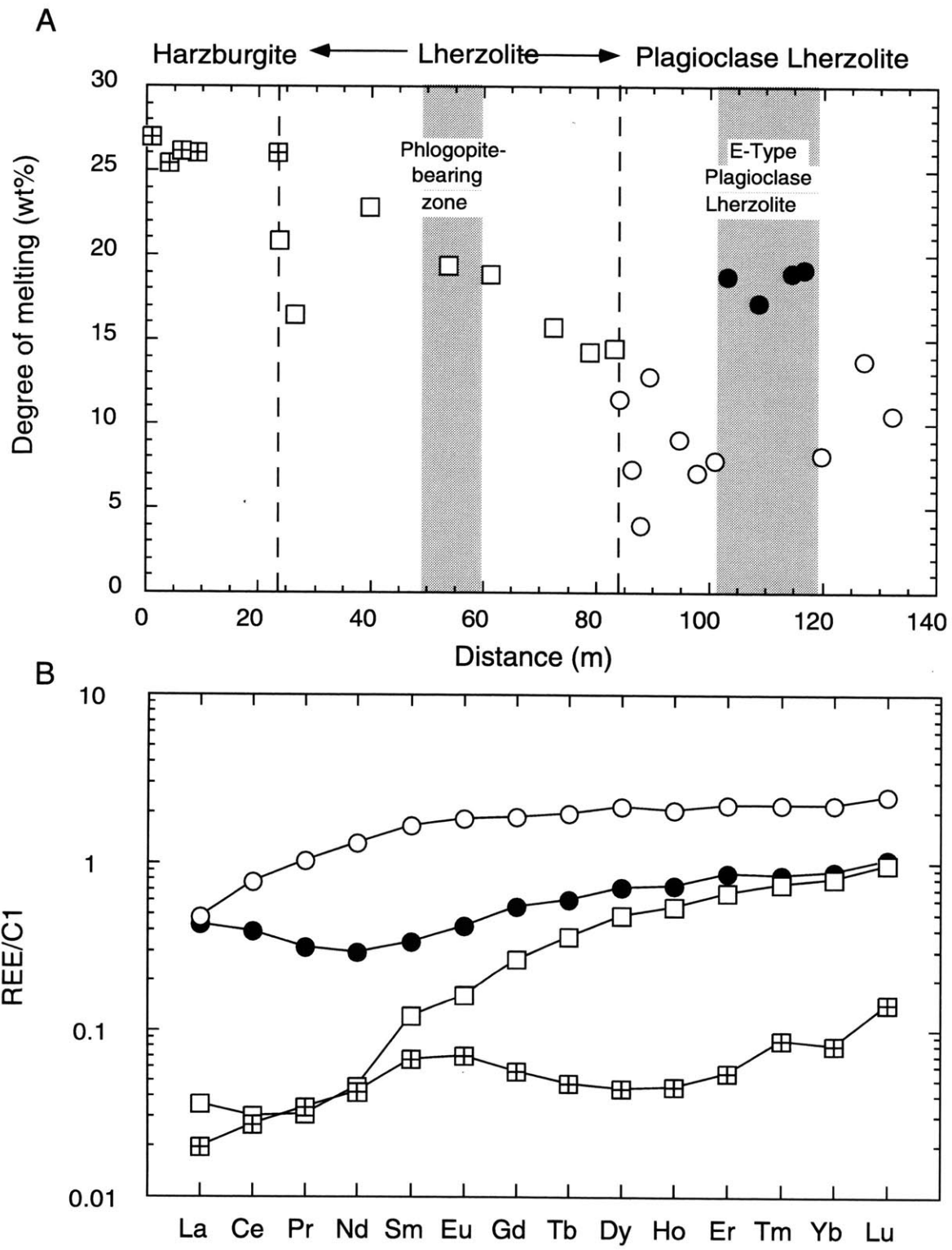


Figure 3

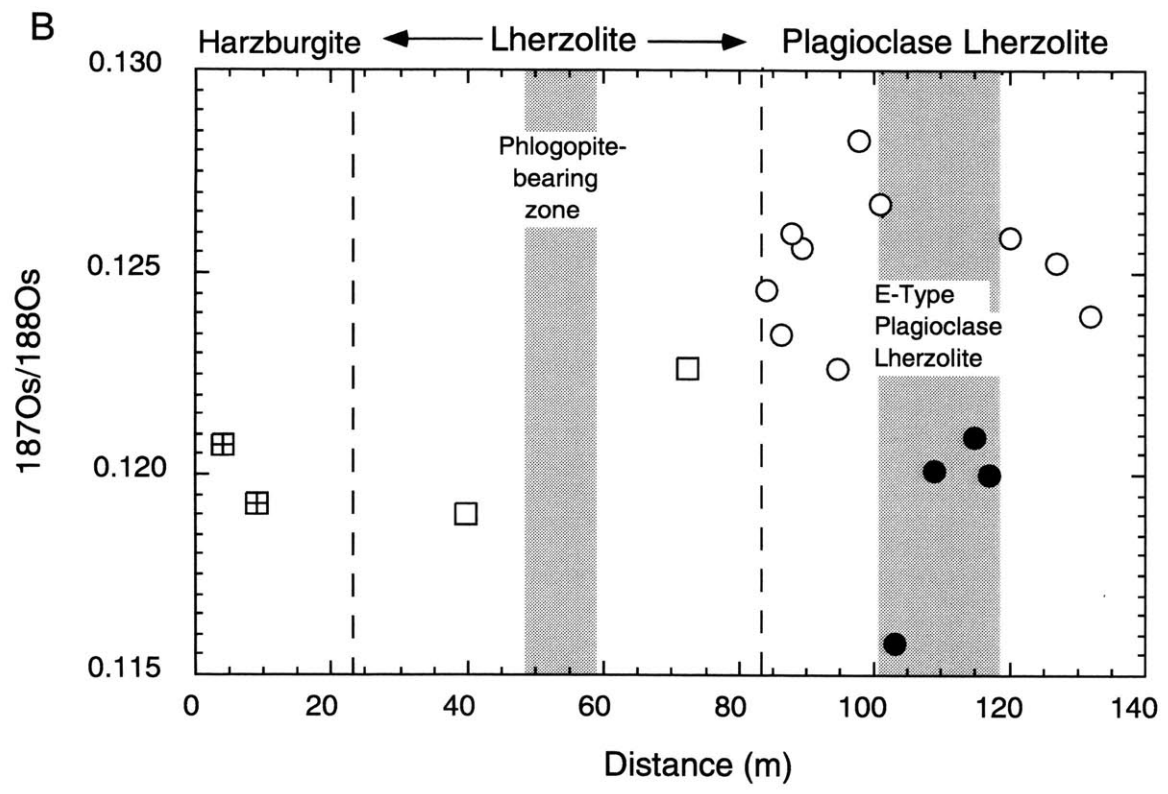
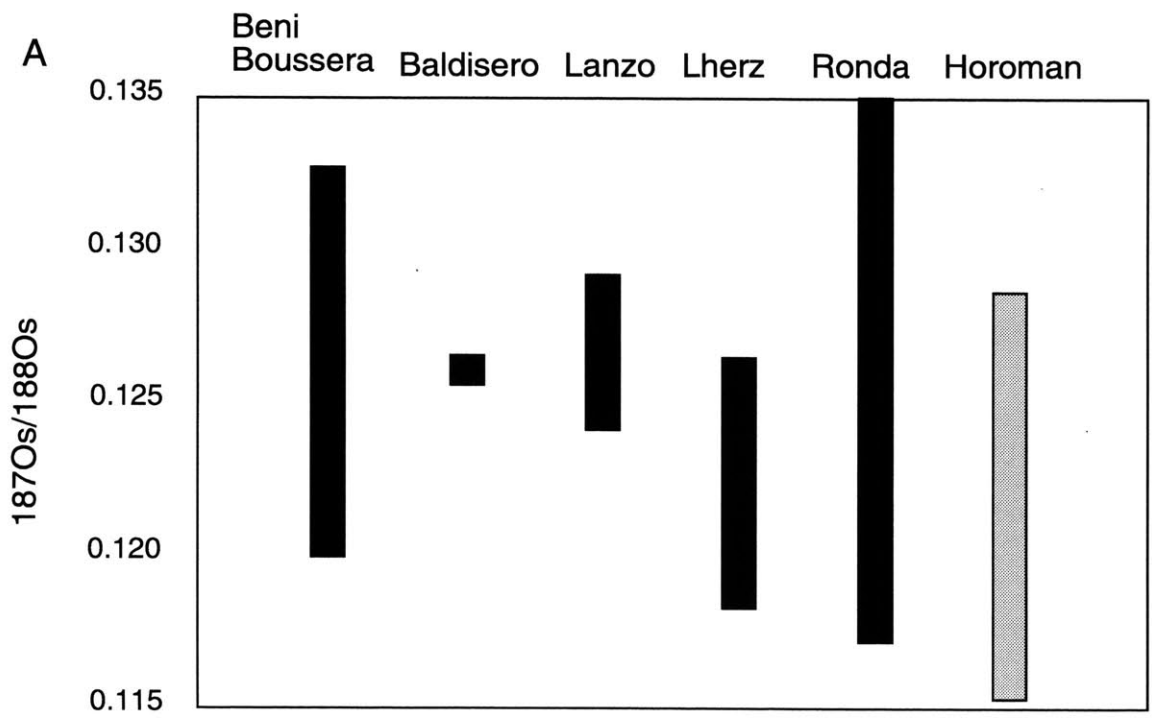


Figure 4

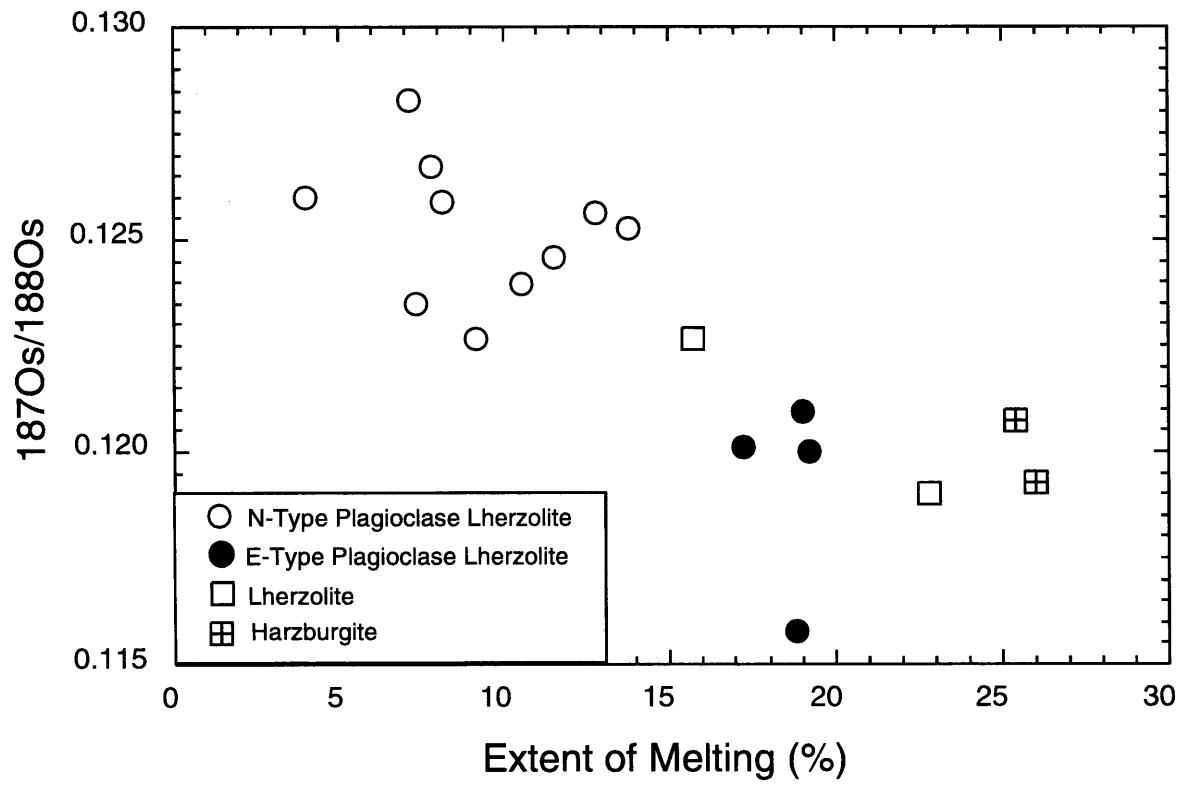


Figure 5

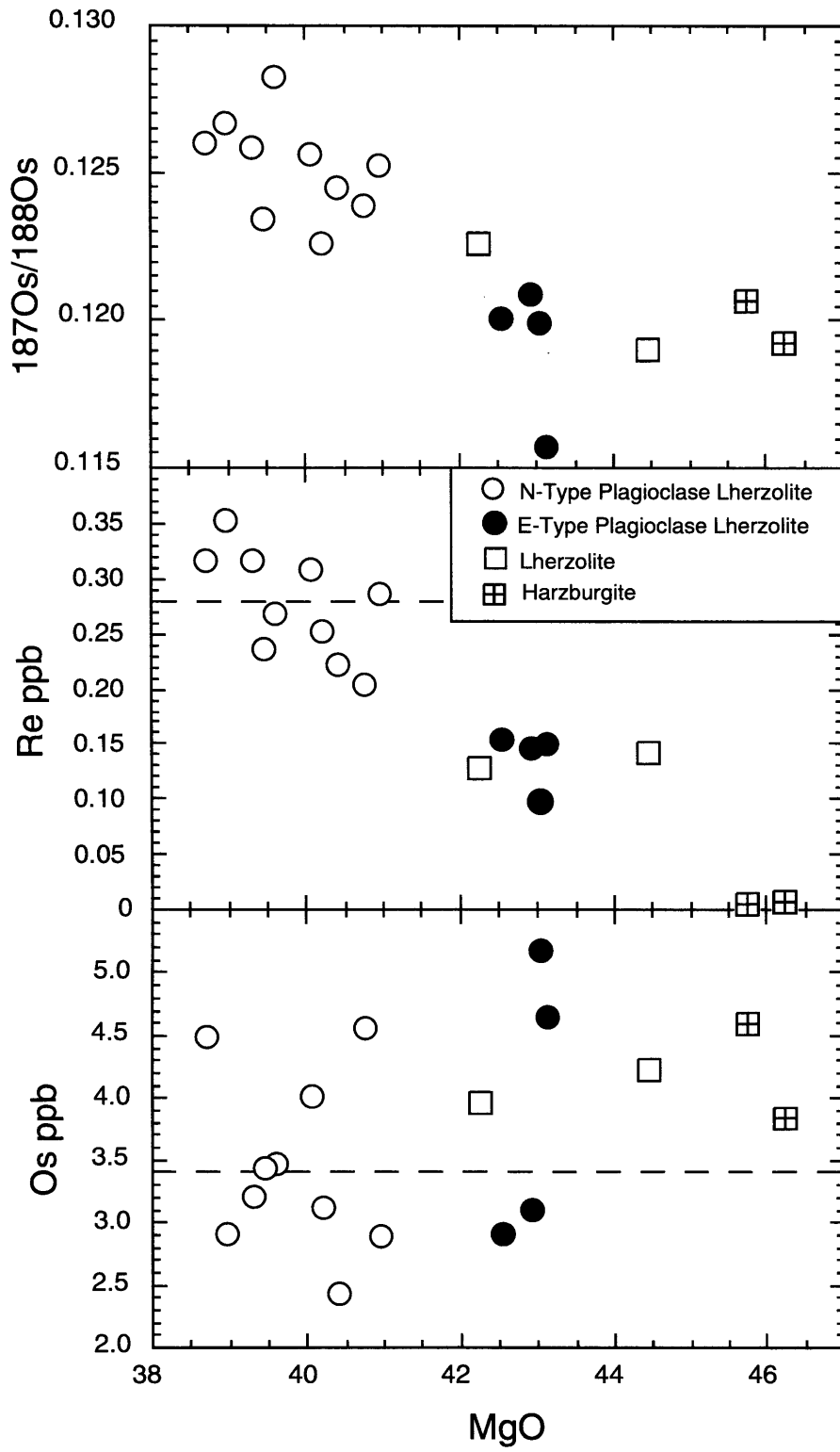


Figure 6

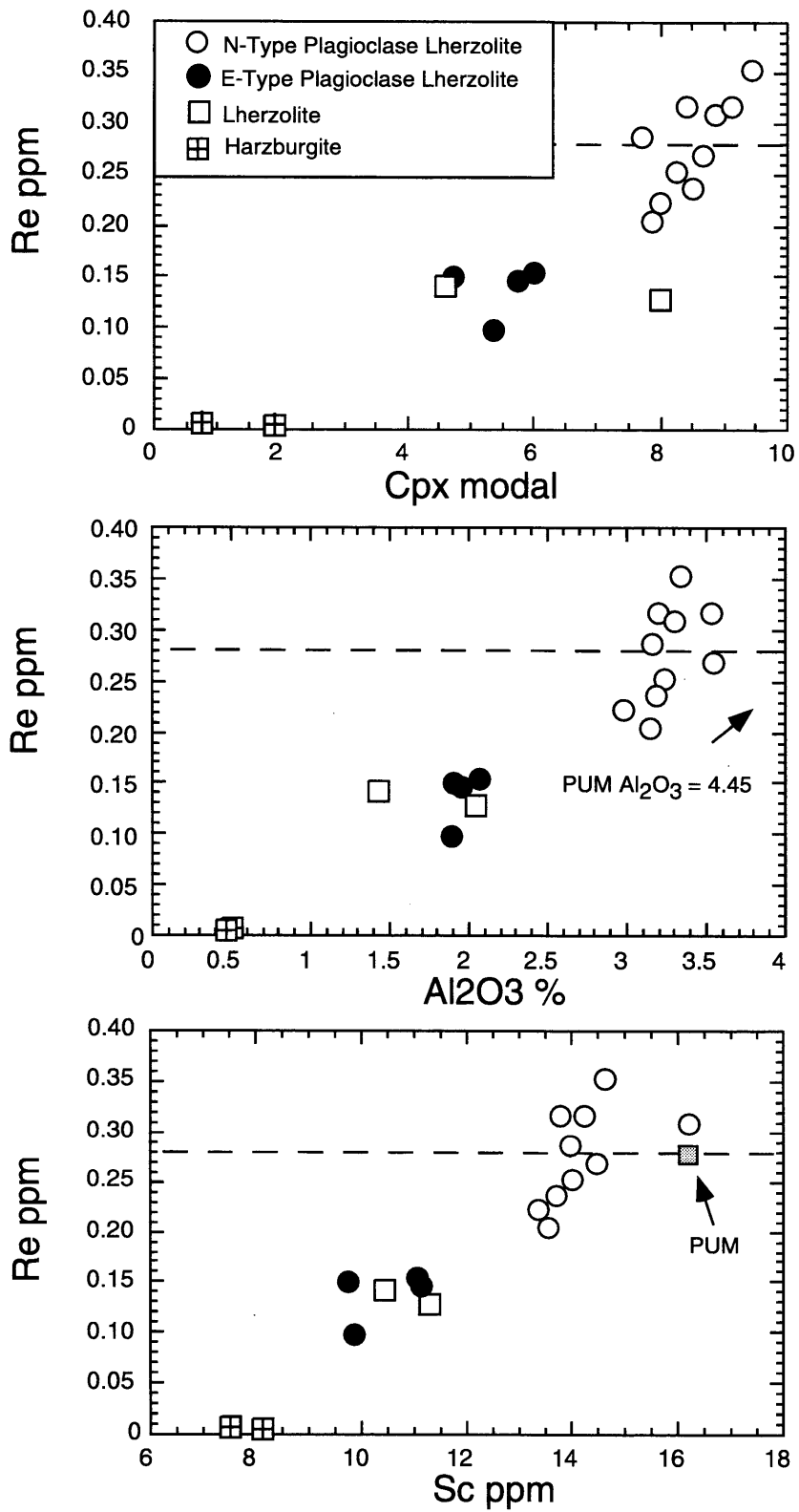


Figure 7

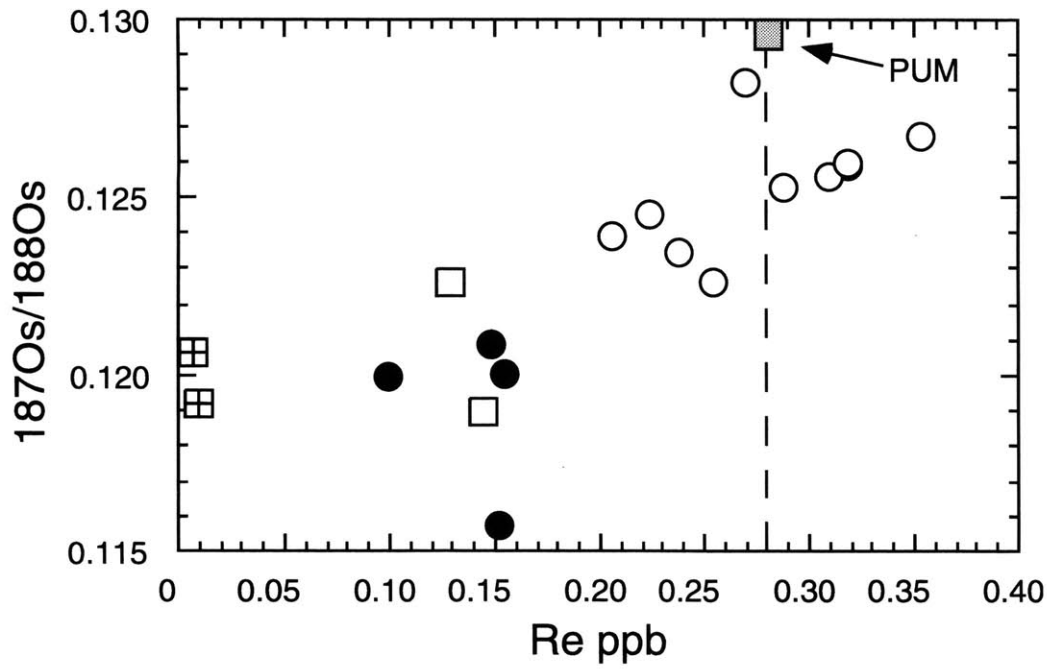
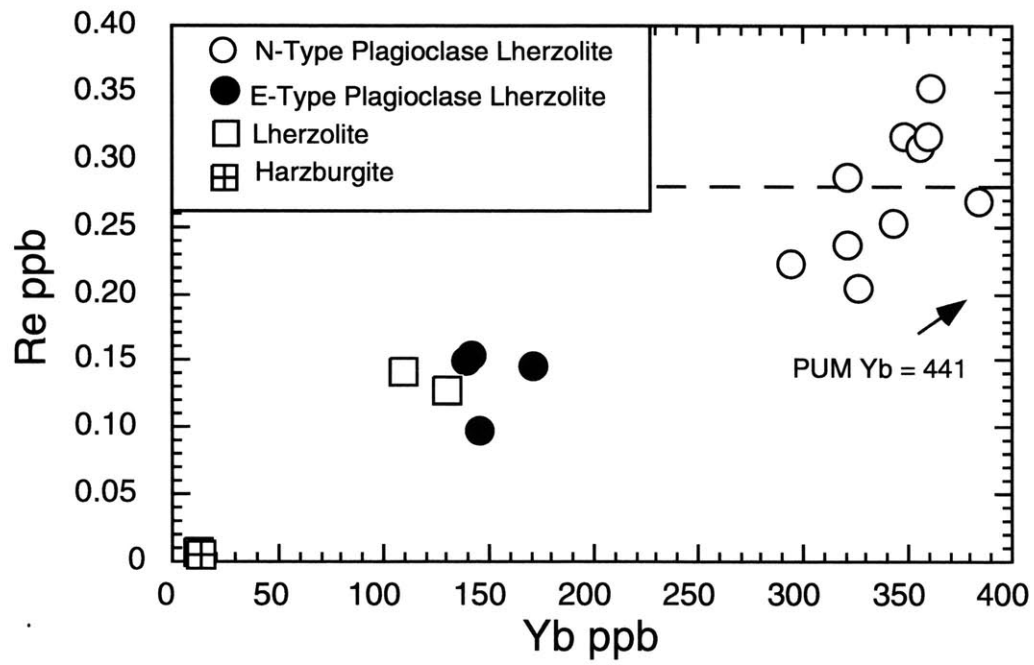


Figure 7 continuation

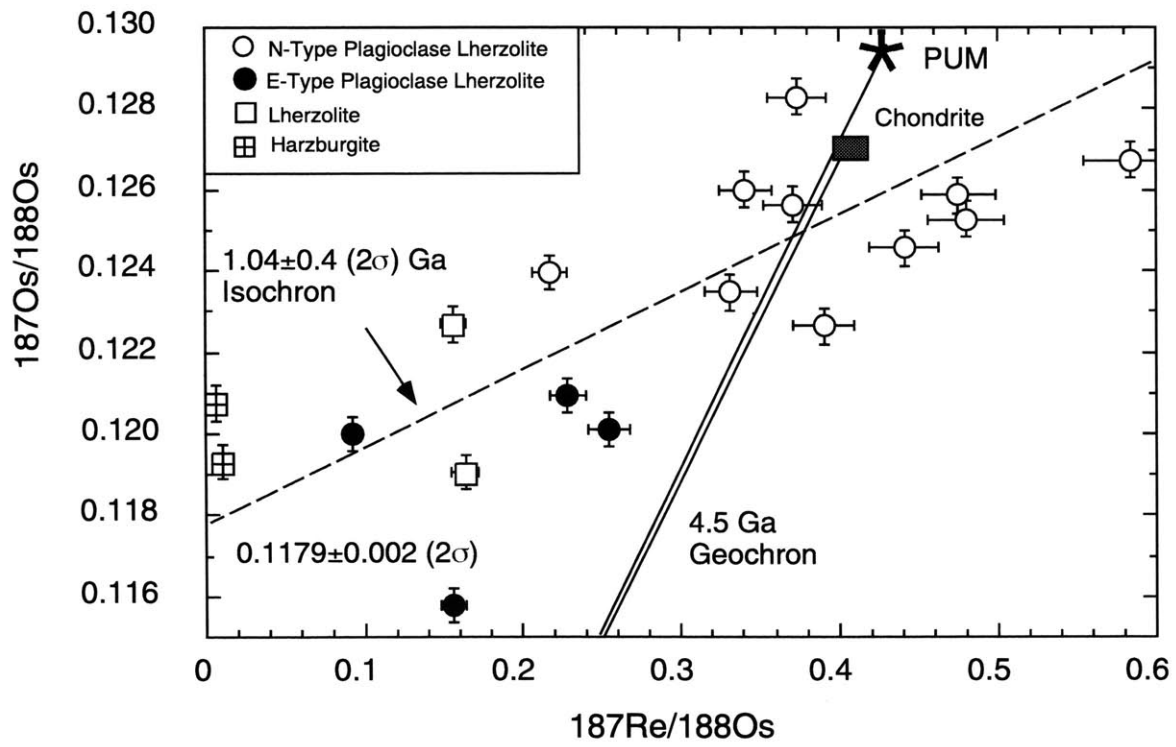


Figure 8



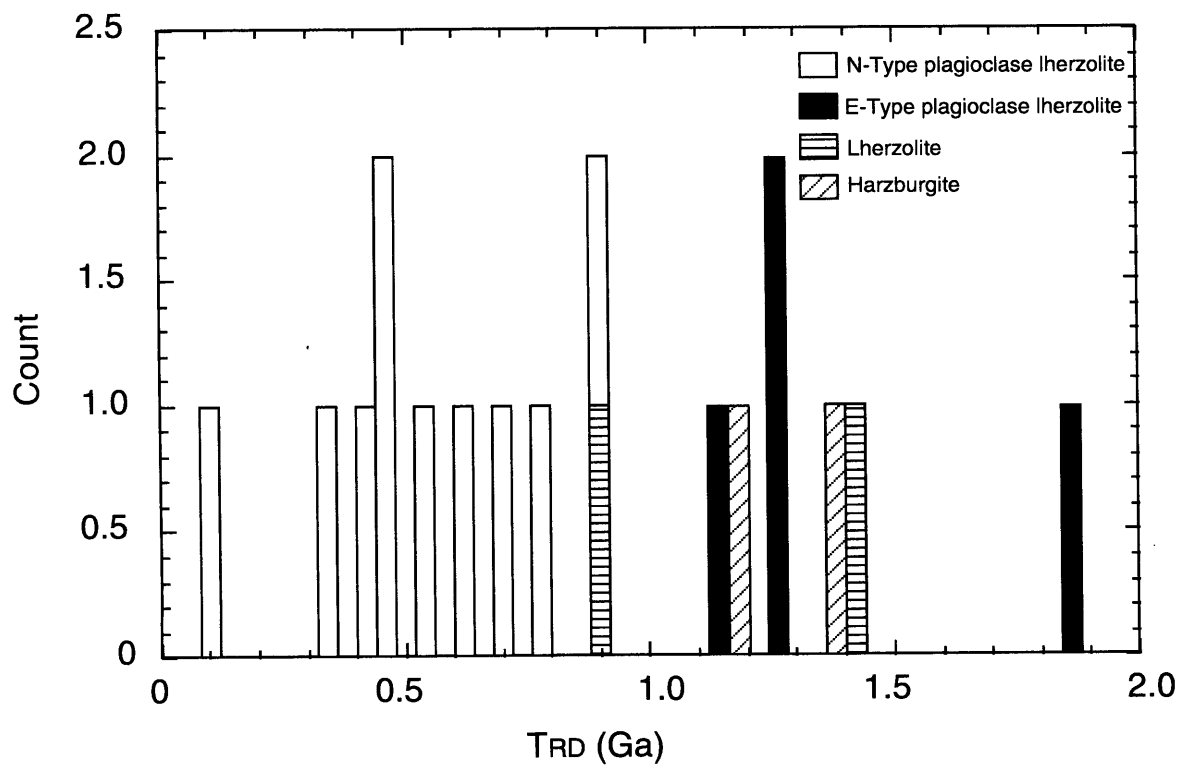


Figure 9

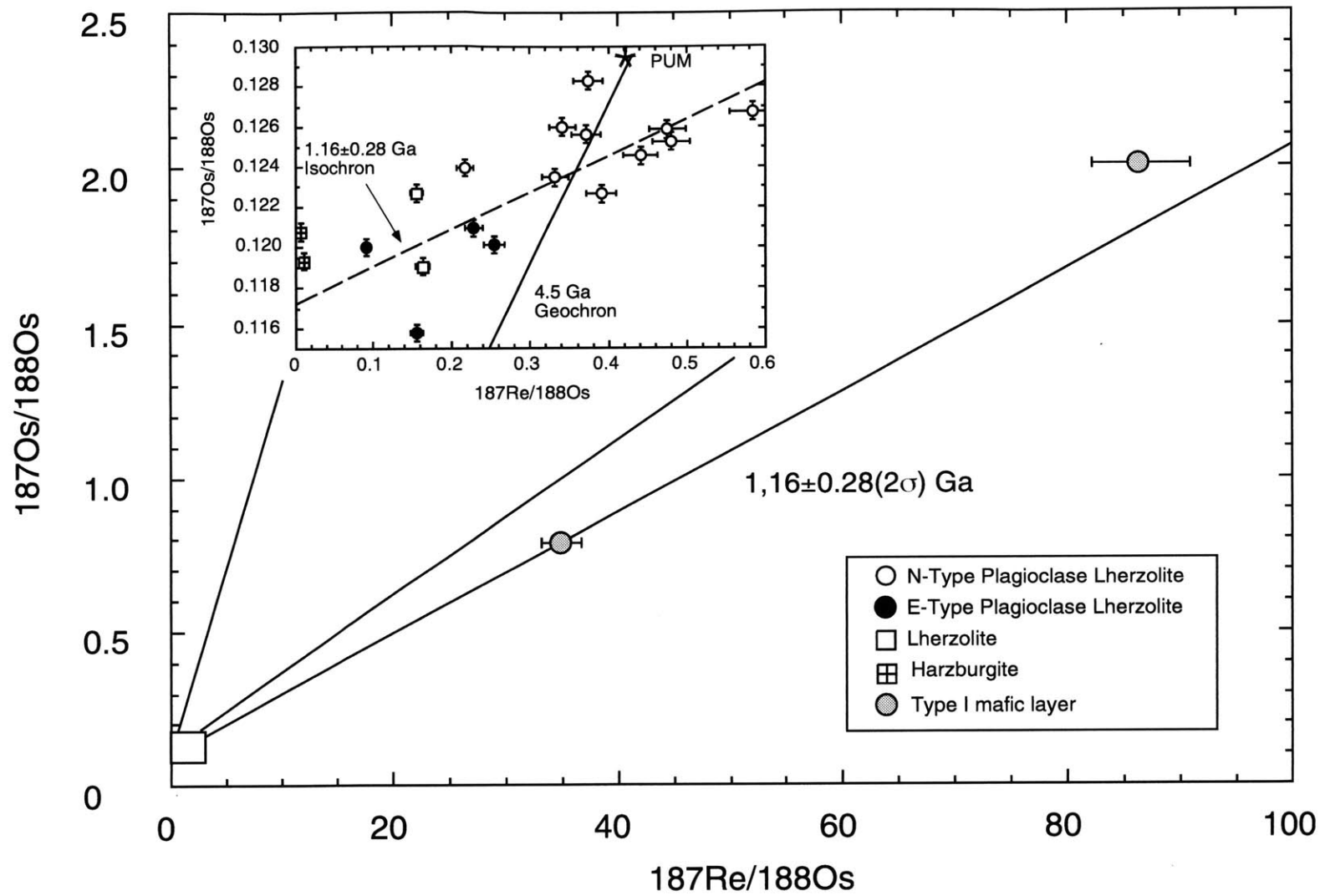


Figure 10

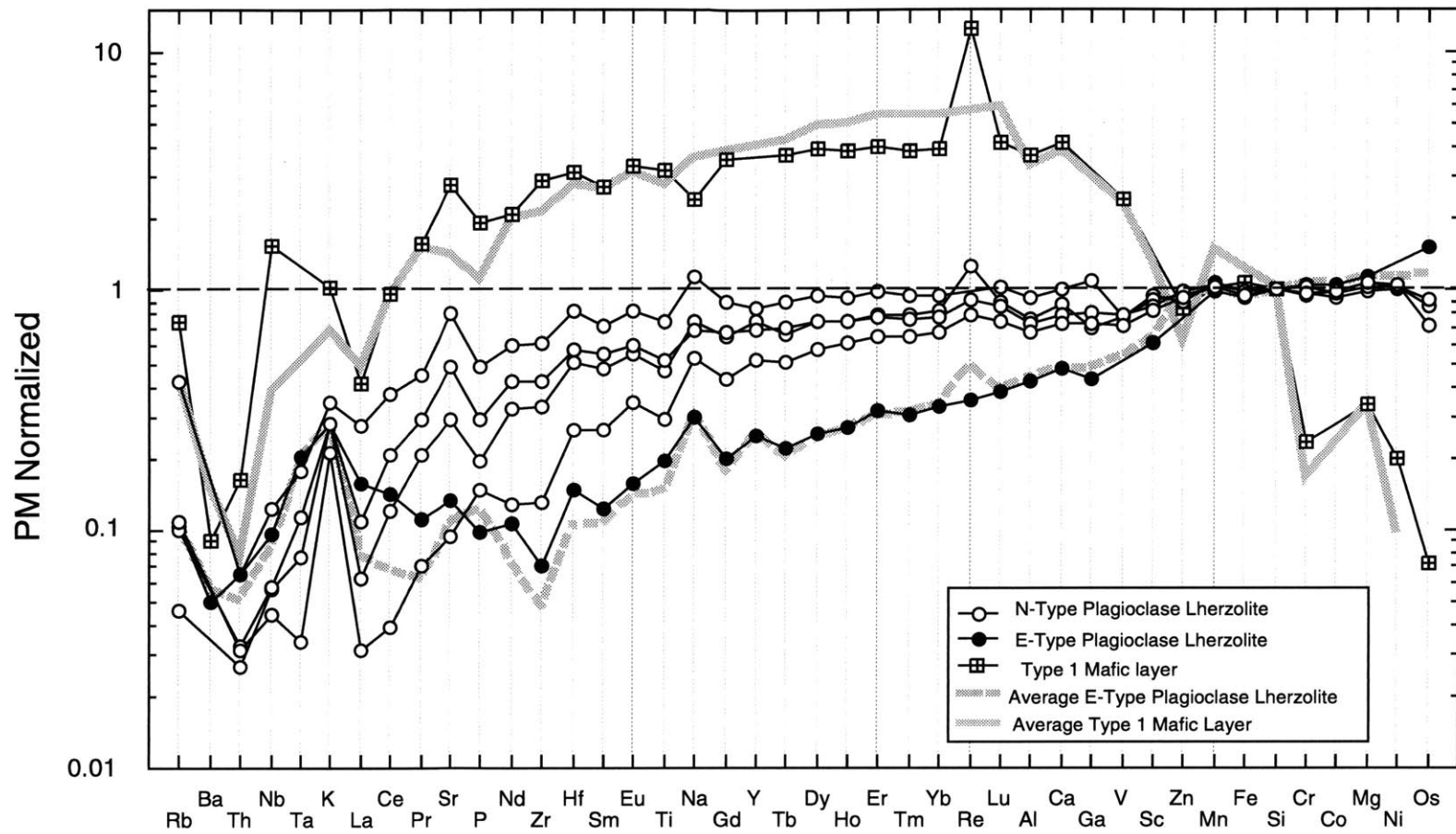


Figure 11

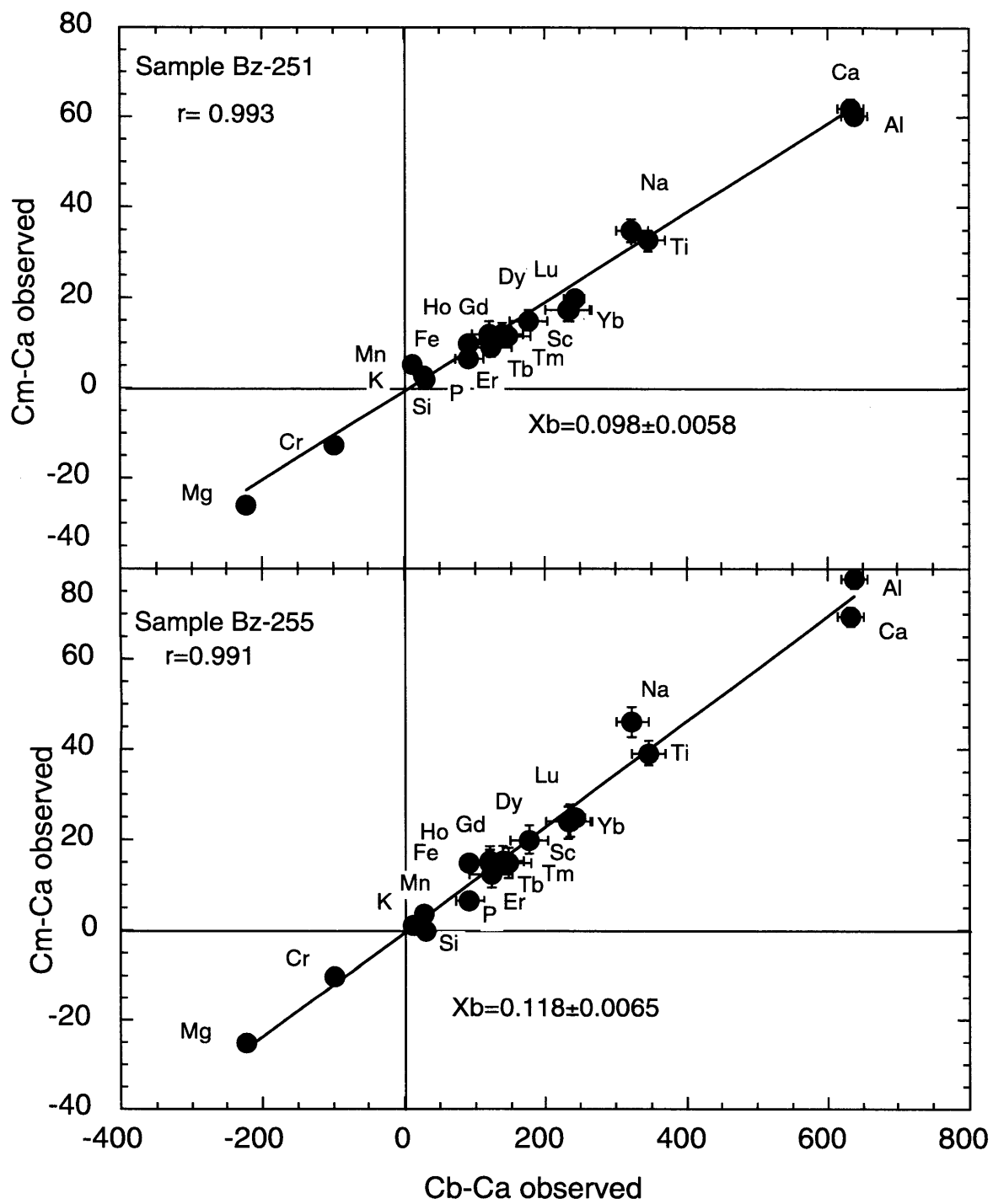


Figure 12

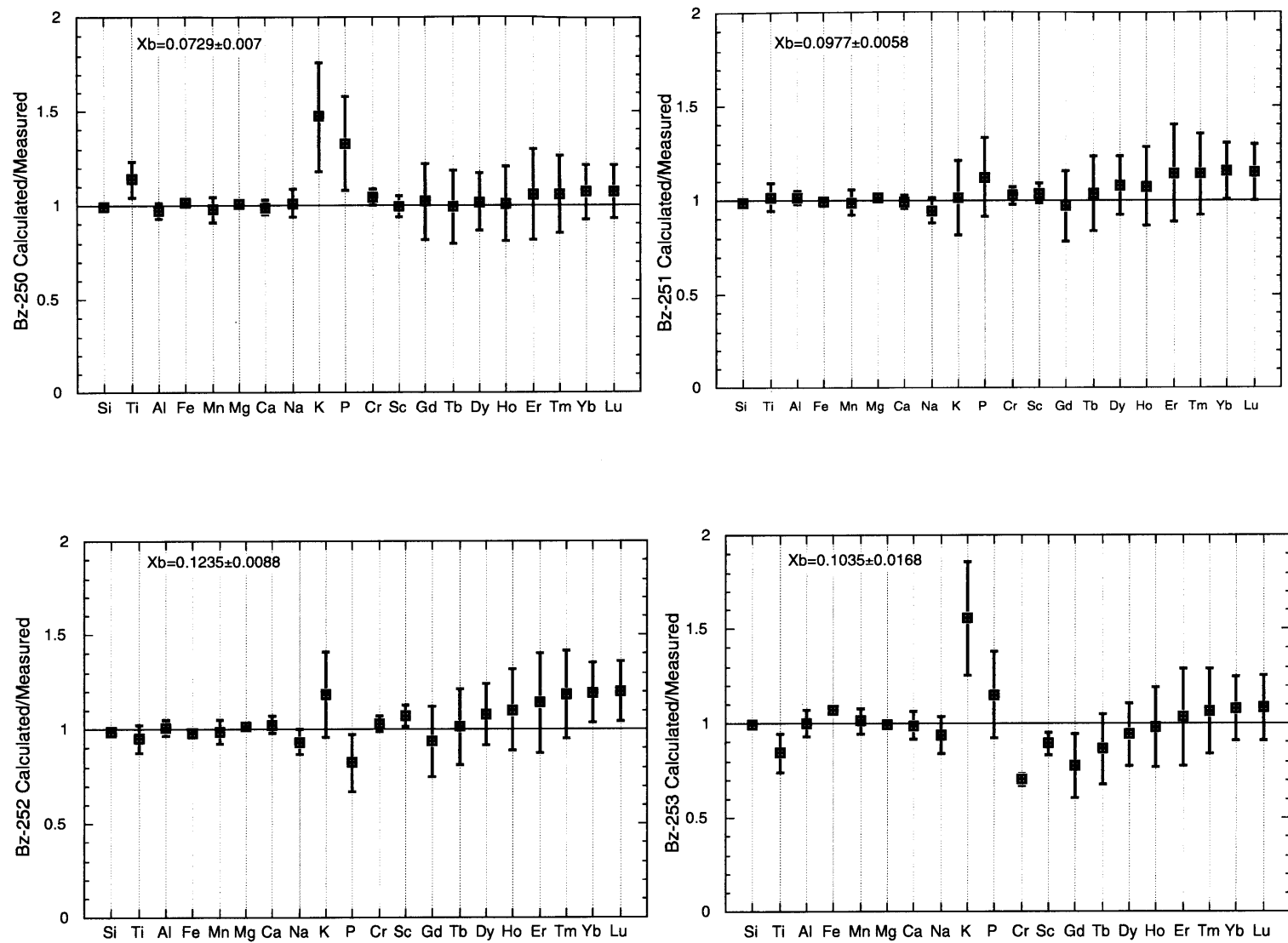


Figure 13

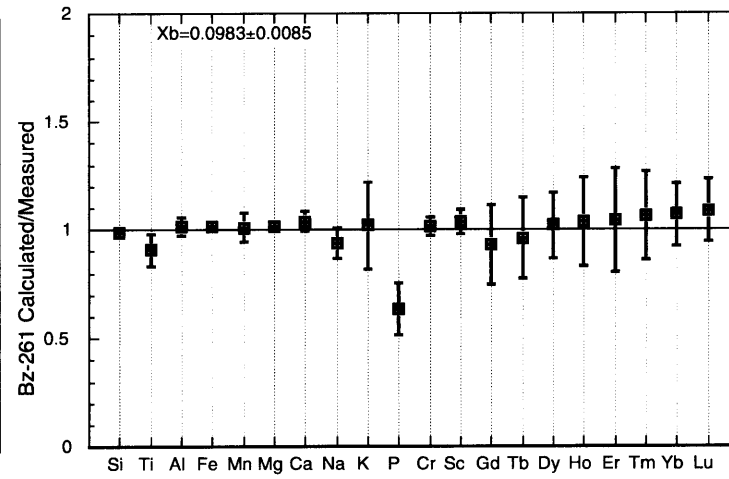
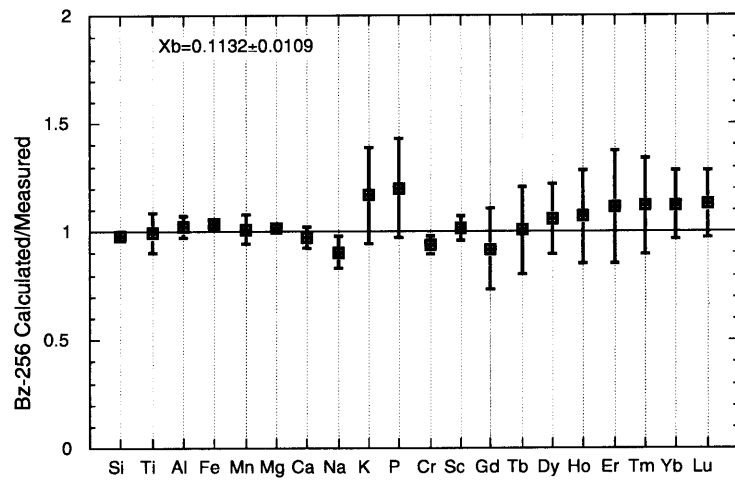
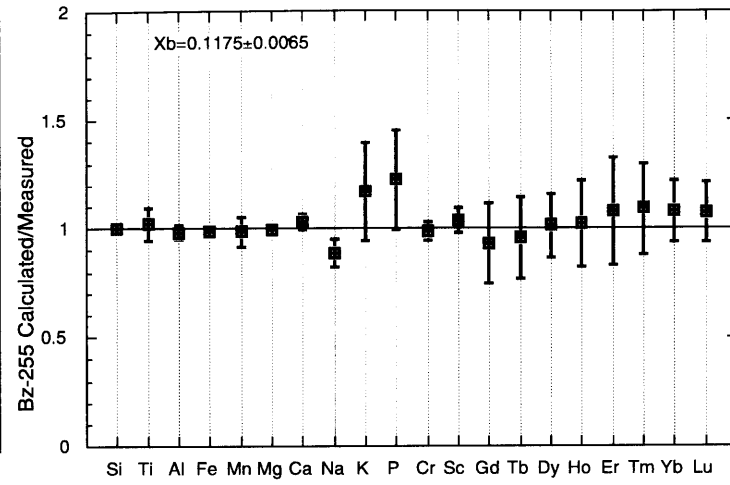
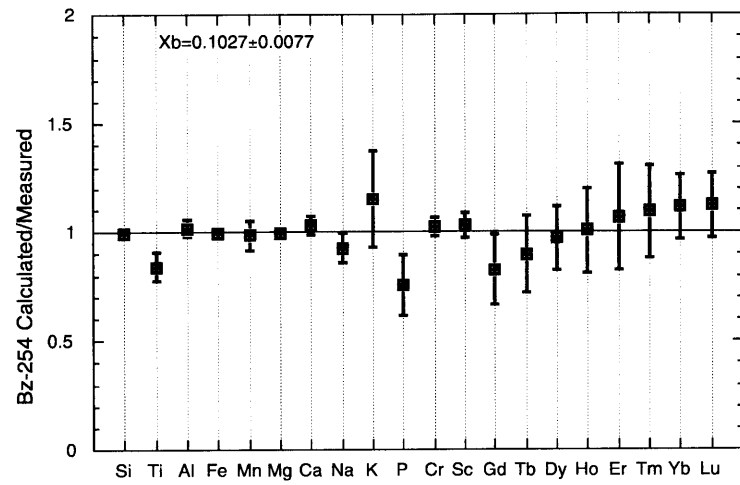


Figure 13  
Continuation

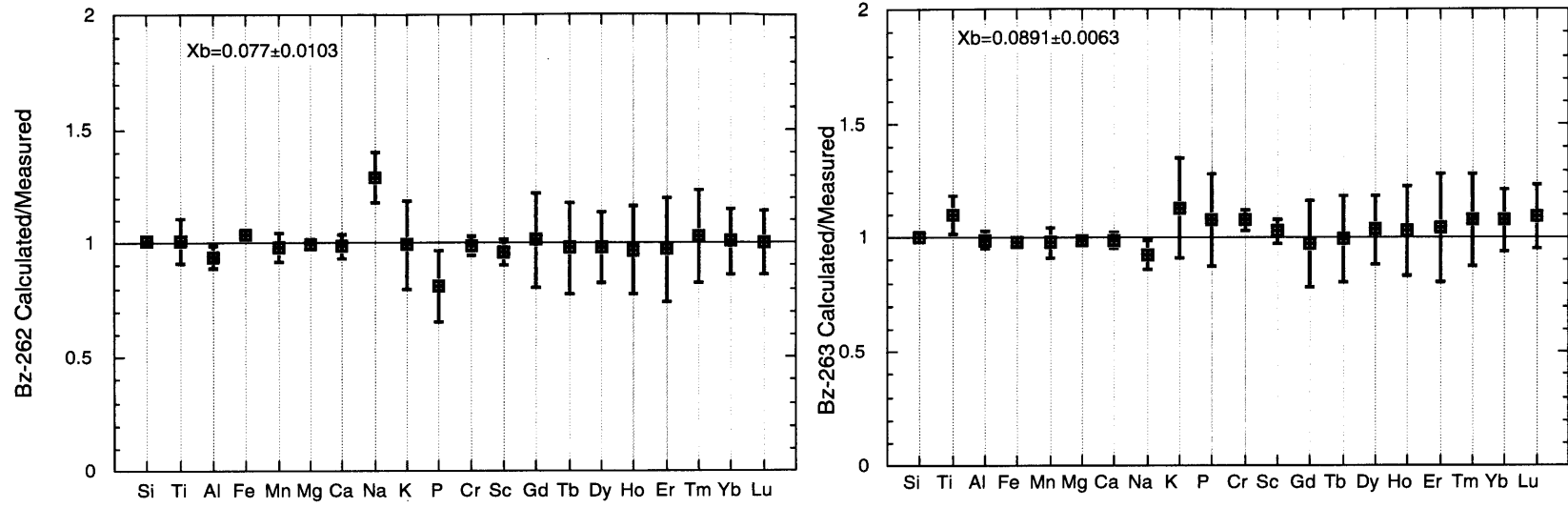


Figure 13  
Continuation

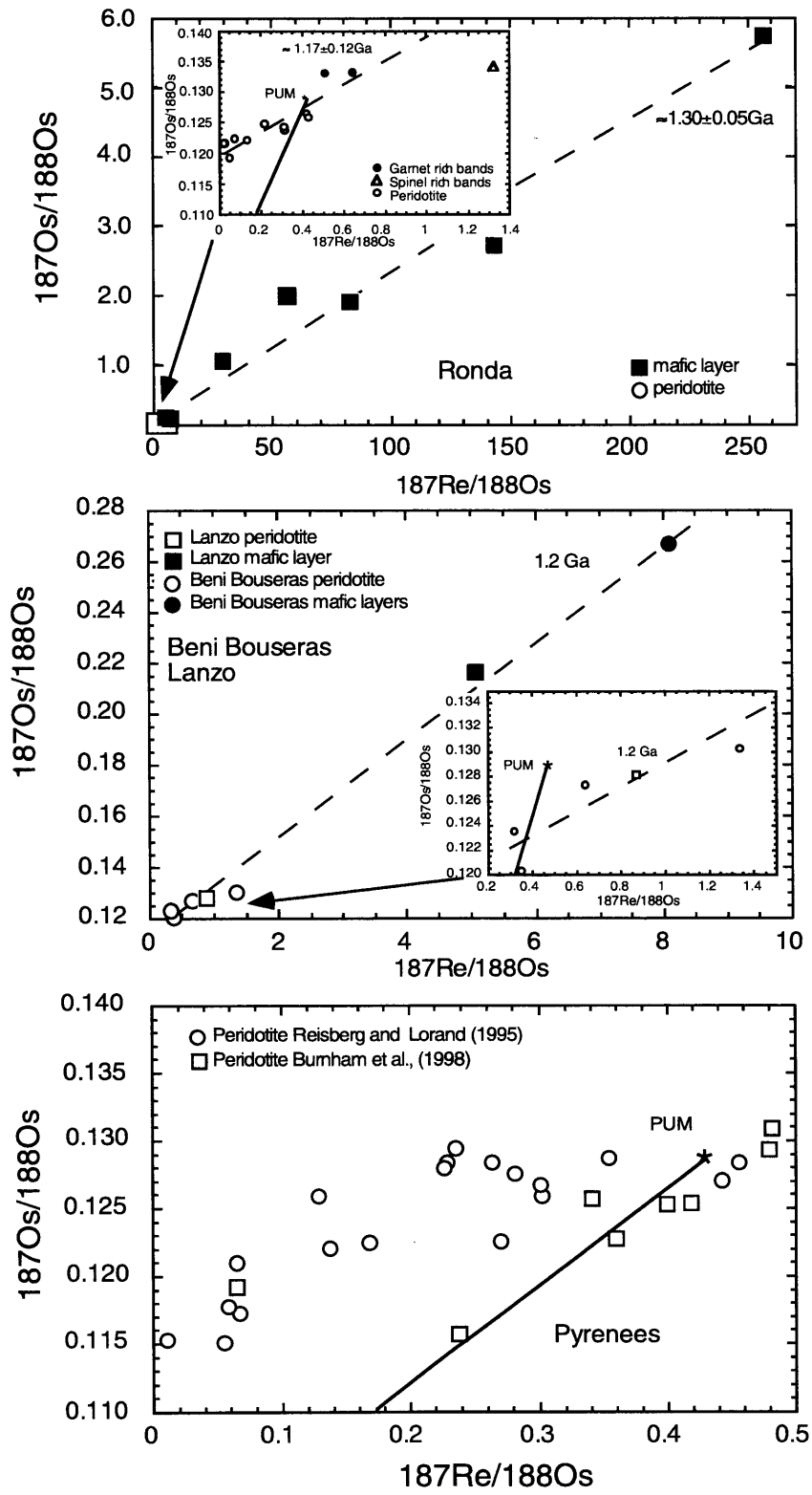


Figure 14



## **3. Chapter Three**

### **Pb Isotopic Variability in Melt Inclusions from Oceanic Island Basalts**

#### ***3.1 Abstract***

Previous studies have suggested that melting processes are responsible for the trace element variability observed in olivine-hosted basaltic melt inclusions. Melt inclusions from four individual lava samples, two from Mangaia, Cook Islands, one from Tahaa, Society chain, and one from Pitcairn, Gambier chain, have heterogeneous Pb isotopic composition, even though the erupted lavas are isotopically homogeneous. The range of Pb isotopic compositions from individual melt inclusions in a single lava flow spans 50% of the world-wide range observed for ocean island basalts (OIB). The melt inclusion data can be explained by two-component mixing for each island. Our data imply that magmas with different isotopic compositions existed in the volcanic plumbing system prior to or during melt aggregation.

#### ***3.2 Introduction***

Studies of oceanic basalts have shown that the mantle is isotopically heterogeneous (Zindler and Hart, 1986; Hart et al, 1992; Hannan and Graham, 1996, Hofmann, 1997 and references therein). However, the nature, distribution, and scale of these heterogeneities remain uncertain. The aggregation of melts on their way to the surface, and mixing in magma chambers prior to eruption, can obscure the chemical and isotopic signatures of pre-aggregated melts.

Trace and major element studies of olivine-hosted melt inclusions have been successful in defining the chemical composition of pre-aggregated melts (Sobolev and Shimizu, 1993; Gurenko and Chaussidon, 1995, Sobolev et al., 1994a, 1994b; Nielsen, 1995, Kamenetsky, 1996; Sobolev and Danyushevsky, 1994; Schiano and Clocchiatti, 1994). However, the lack of isotopic information on melt inclusions has made it difficult to distinguish whether these melt compositions represent different extents of melting from a single source or whether they originate from different source compositions.

We selected four geochemically well-characterized basalt samples, two from Mangaia, Cook Islands, one from Tahaa, Society islands, one from Pitcairn, Gambier chain, and two peridotitic xenoliths from Tubuaii, Austral chain, (Hauri et al., 1993; Hauri and Hart, 1993; 1997; Eiler, et al., 1995; Eiler, et al., 1997). The whole rock Sr, Nd and Pb isotopic compositions from Mangaia, Tahaa and Pitcairn islands are very similar to the three end-member mantle compositions known as the HIMU (high U/Pb ratios) EMII and EMI (enriched mantle) end-members respectively (Zindler and Hart, 1986; Hart et al., 1992). Previous studies reporting isotope, major and trace element compositions of these lavas have explained the observed geochemical variation by melting of single (homogeneous) mantle sources, originating through the recycling of the oceanic crust (Woodhead, 1996; Hemond et al., 1994, Eiler et al., 1997, Hauri et al., 1993; Hauri and Hart, 1993). Thus, these lavas are excellent examples on which to test the hypothesis that the geochemical variations in melt inclusions do in fact represent melt processes as opposed to source effects.

In this work, we present the first study of Pb isotopes on melt inclusions from the three mantle end-members: HIMU, EMI and EMII. The range in Pb isotopic composition observed in melt inclusions from single lava flows suggests that the geochemical variation in melt inclusions has a large component of source control, and is mainly produced by either mixing of melts or assimilation and fractional crystallization processes deep in the magma chambers.

### ***3.3 Analytical Techniques***

Approximately 100 grams of fresh interior of the basalt samples were crushed with a steel jaw crusher. The samples were sieved to separate the olivine and clinopyroxene phenocrysts (retaining grains larger than 1 mm). The grains were washed with ultrapure water in an ultrasonic bath and dried at low temperature (<80°C) in an oven. The olivines and clinopyroxenes with melt inclusions were selected under a binocular microscope, and mounted in an aluminum ring with epoxy and cured overnight. The grains were polished with silicon carbide, diamond and alumina powder until the melt inclusions were exposed. The mounts were cleaned with ethanol in ultrasonic bath for 5 minutes, dried at room temperature and placed in a vacuum chamber for at least 12 hours prior to the analysis.

Major element compositions of host crystals and daughter minerals were measured on a JEOL 733 electron microprobe at the Massachusetts Institute of Technology (Table 1 and Appendix). We applied the correction methods of Bence and Albee (1968) and Albee and Ray (1970). The typical operating conditions used were an accelerating voltage of 15 keV and currents of 10 nAmps. Counting times ranges from 20 to 40 seconds. The in-run precision is mainly governed by counting statistics, varying with the abundance of the elements. For example, for a single analysis, the one sigma uncertainties are approximately  $\leq 1\%$  for element contents  $\geq 1$  wt. %, 1 to 5% for 1-0.5 wt %, 5 to 15% for 0.5-0.1 wt % and  $> 15\%$  for contents of less than 0.1%. The melt inclusion and daughter and host mineral compositions reported are averages of two to four analyses. The spot size ranges from  $5\mu$  for the daughter minerals, to  $10\mu$  for the host minerals, and ranged from  $10\mu$  to  $30\mu$  for estimating the glass composition and the average melt inclusion composition respectively.

Trace element compositions for the melt inclusions were obtained by secondary ion mass spectrometry (SIMS) with a Cameca IMS 6f ion microprobe at the Department of Terrestrial Magnetism, Carnegie Institution of Washington and with the Cameca IMS 3f ion microprobe at the Woods Hole Oceanographic Institution, following the techniques of Shimizu and Hart, (1982) and Hauri et al., (1996) (Table 1 and Appendix). We used a primary  $^{16}\text{O}^-$  beam of 4.5 to 10 keV and current ranging from 5 to 20 nAmps. The beam was focused to a spot of 10 to  $20\mu$ . Energy filtering techniques, with secondary voltage offsets of -60 to -90, were used to eliminate molecular ion interference and to reduce the magnitude of matrix effects. Counting times ranged from 5 to 30 seconds depending on the element abundance. The intensities of the secondary ions were normalized to that of  $^{30}\text{Si}$ . The acquisition of the intensities consist of an average of five cycles of alternating measurements from light to heavy mass elements, and the  $^{30}\text{Si}$  intensity was measured at the beginning and at the end of each run. The trace element concentrations were determined using empirical working curves of intensities versus concentrations for well known standards. The analytical uncertainty for each analysis varies inversely with the intensity of the element considered, and ranged from  $<5\%$  to 30%.

Lead isotopes were determined by secondary ion mass spectrometry (SIMS) with a Cameca IMS 1270 ion microprobe at the Woods Hole Oceanographic Institution (Table 2). We used a primary ion ( $^{16}\text{O}^-$ ) beam of 22.5 keV (nominal primary accelerating voltage of 12.5 keV, and secondary accelerating voltage of 10 keV) and currents ranging from 30 to

50 nAmps focused into a spot 20 to 30  $\mu\text{m}$  in diameter. To resolve isobaric interferences, secondary ions were analyzed at a mass resolution of 3500 with no energy filtering (Figure 1). Each analysis comprised 200 to 400 cycles of alternating measurements of  $^{206}\text{Pb}$  (4 s counting),  $^{207}\text{Pb}$  (4 s counting),  $^{208}\text{Pb}$  (2 s counting) and background (1 s counting). The in-run precision inversely correlates with the Pb content of the sample; typical in-run precision for the Pb isotopic measurements ( $^{207}\text{Pb}/^{206}\text{Pb}$ ;  $^{208}\text{Pb}/^{206}\text{Pb}$ ) ranged from  $\approx 0.2\%$  ( $2\sigma$ ;  $^{208}\text{Pb}$  count rates 10,000 cps) to  $1.2\%$  ( $2\sigma$ ;  $^{208}\text{Pb}$  count rates 200 cps). When  $^{208}\text{Pb}$  counts rates were higher than 2,000 cps, we also measured  $^{204}\text{Pb}$  (8 s counting), with an in-run precision for  $^{206}\text{Pb}/^{204}\text{Pb}$  ranging from  $\approx 0.5\%$  ( $2\sigma$ ; 10,000 cps for  $^{208}\text{Pb}$ ) to  $1.6\%$  ( $2\sigma$ ; 2,000 cps for  $^{208}\text{Pb}$ ). We used Loihi glass 158-4\* previously analyzed by TIMS (Garcia et al., 1995) as a standard during the period of measurement of the melt inclusions. The Pb content of the Loihi glass is  $\approx 3$  ppm, similar to the Pb content of the melt inclusions. Typical  $^{208}\text{Pb}$  count rates were approximately 400 cps per ppm of total Pb. The average of ion probe analyses reproduced the TIMS value of the glass standard to within  $<0.1\%$  standard errors ( $2\sigma$ ). The external precision ( $2\sigma$ ) defined by the 34 runs on the Loihi standard is  $\pm 0.66\%$  ( $^{208}\text{Pb}/^{206}\text{Pb}$ ), and  $\pm 0.50\%$  ( $^{207}\text{Pb}/^{206}\text{Pb}$ ); these are essentially the same as the average internal precision of the 34 runs (Figure 2a, b and c). We carefully evaluated whether hydride or correlated errors in  $^{206}\text{Pb}$  played any role in the isotopic ratios measured; we found that the isotopic ratios are mainly controlled by counting statistics, with no effect of hydride or systematic error in  $^{206}\text{Pb}$ . The mass fractionation for SIMS analysis was  $<0.15\%$  per amu, and was smaller than the in-run precision. Comparison of TIMS and SIMS Pb isotope data for basaltic glass, feldspar, chalcopyrite, and zircon indicate that instrumental mass fractionation is generally smaller than the precision of the analyses and that there are no resolvable matrix effects (Layne and Shimizu, 1997).

### **3.4 Results**

#### **3.4.1 Major and trace element compositions**

We recovered olivine (Fo 84-89) and clinopyroxene-hosted melt inclusions from two primitive basalts from Mangaia. We selected olivines and clinopyroxenes from an ankaramite, (MGA-B-25), and olivines from a picrite, (MGA-B-47). All melt inclusions are in general ellipsoidal in shape and range from 30 to 300  $\mu\text{m}$  in the longest direction. Sometimes the inclusions have straight boundaries or rectangular habits resembling pseudomorphs after a single phase. All of the inclusions now contain “daughter minerals”

such as Ti-augite, Cr-spinel, ilmenite, Ti-magnetite, kaersutite, sphene, apatite and phlogopite that crystallized after entrapment of the melt (Table 1 and Appendix). Some of the daughter minerals, such as Cr-spinel and augite, are zoned becoming enriched in Ti toward the margin. The host olivine and clinopyroxene are also zoned close to the contact with the melt inclusions. The olivine becomes enriched in Fe and Ca and the clinopyroxene is enriched in Ti (Appendix). Sometimes crystallized daughter clinopyroxene or kaersutite forms a surface that separates the melt inclusion from the host olivine or clinopyroxene respectively (Figure 3a and b).

Some melt inclusions ( $\approx 7\%$ ) from both Mangaia samples are texturally very complex, containing co-existing silicate glasses, sulfide, and carbonate globules (Figure 3a). In many cases, the inclusions contain droplets of silicate glass within carbonate spherules, bounded by sharp, curved menisci. This texture is suggestive of processes related to carbonate-silicate liquid immiscibility (Kogarko et al., 1995) (Figure 3a). The carbonate spherules consist largely of calcium-iron-magnesium carbonates with minor alkalis. The carbonates are associated with Ca or Mg-rich volatile-bearing silicate glasses (up to 15 % volatiles) and phonolitic glasses. In addition, sulfide globules form part of the mineral association in these melt inclusions. The sulfides range in composition from chalcopyrite, pentlandite to pyrrhothite (Table 1 and Appendix). The observed mineral associations, textures, and the major element compositions of the minerals, carbonates and silicate glasses in the Mangaia melt inclusions are similar to those described for carbonate-silicate melt pockets in carbonated peridotitic xenoliths such as those reported for Spitsbergen and the Canary Islands (Ionov et al., 1996; Kogarko et al., 1995).

The silicate glasses and carbonate globules from the Mangaia melt inclusions have a large range of trace element compositions (Figure 4, Table 1 and Appendix). The silicate glasses range to very high Sr and Ba contents (up to 1.4 and 0.8 wt % respectively). Also, they have variable REE concentrations, with La and Yb ranging from 10 to 300 and 0.3 to 20 times primitive mantle respectively. Fractionation of the REEs is also large with  $(\text{La}/\text{Yb})_{\text{PM}}$  ratios reaching values up to 90, as compared to most whole rocks which typically have ratios of  $\approx 10$  (Hauri and Hart, 1997). Moreover, we found that several melt inclusions have a large fractionation between Zr and Hf, with Zr/Hf ratios ranging up to 100, similar to the Zr/Hf ratios reported for Mangaia and Tubuaii lavas (Dupuy et al., 1992).

Only two carbonate globules were analyzed for trace element contents (Table 1). From 140 melt inclusions, we found only two carbonate globule large enough to allow us to get a representative composition of the carbonate without any influence from the surrounding silicate glass. The high trace element contents of the carbonates globules, specially the high rare earth elements, compared to those of carbonates produced during secondary (hydrothermal) processes, suggests that the carbonates have had a magmatic origin. However the trace element content (i.e., La < 60 ppm) and the  $(La/Yb)_{PM}$  ratios in the carbonate globules (5 to 35) are generally lower than those reported for carbonatitic melts (Nelson et al., 1988). This is also true for their alkali contents, where the Na<sub>2</sub>O and K<sub>2</sub>O in the carbonate globules are low compared to those of carbonatitic melts experimentally produced by an immiscibility process (Lee and Wyllie, 1998). Thus, although the texture suggests that the carbonate is formed by immiscibility process, the compositions of the carbonates are not consistent with this hypothesis. The same apparent contradiction between textures and carbonate compositions was also described for carbonate-silicate melt pockets in carbonated peridotitic xenoliths from Spitsbergen (Ionov et al., 1993; 1996).

We also recovered olivine (Fo 86)-hosted melt inclusions from a primitive basalt from Tahaa (TAA-B-26). These melt inclusions are ellipsoidal in shape and range from 10 to 100  $\mu$ m in diameter. In contrast to the Mangaia melt inclusions, the Tahaa melt inclusions are much simpler. Silica-alumina-alkali-rich glass (phonolitic) records only in-situ fractional crystallization of clinopyroxene, Cr-spinel, ilmenite- and plagioclase, and segregation of sulfide globules (Table 1 and Appendix). Compared to the Mangaia inclusions the Tahaa melt inclusions have a smaller range in REE concentrations (La and Yb ranging from 50 to 250 and 3 to 10 times primitive mantle, respectively) and smaller REE fractionation ( $La/Yb_{PM}$  ranges from 9 to 33; Table 1 and Appendix).

The olivines (Fo 80) and two spinel-hosted melt inclusions recovered from Pitcairn sample (PIT 1; Eiler et al., 1995) have not yet been analyzed for major or trace element composition.

The petrography, major and trace element compositions of two mantle xenoliths from Tubuaii (TBA 1-9 and TBA 4-11), have been described in previous work (Hauri et al., 1993). TBA 1-9 is a wherlite consisting of primary olivine (Fo 90) and minor orthopyroxene. 40% by volume of the xenolith consists of melt pockets made of

clinopyroxene + silica-alumina-alkali-rich glass (phonolitic) + spinel ± euhedral apatite. These melt pockets have rectangular habits, resembling pseudomorphs after a single phase. Hauri et al., (1993) reconstructed the compositions of the melt pockets; they suggested that the composition resembles that of subcalcic pyroxene (± apatite). TBA 4-11 is a wherlitic dunite consistent of olivine (Fo 84-90), chromite and apatite. This xenolith also has silica-rich glass forming up to 10% of the total volume of the sample.

### 3.4.2 Pb isotopes

SIMS analyses of Pb isotope ratios ( $^{207}\text{Pb}/^{206}\text{Pb}$  and  $^{208}\text{Pb}/^{206}\text{Pb}$ ) of the melt inclusion populations from Mangaia, Tahaa and Pitcairn islands, and of clinopyroxene and melt pockets from Tubuaii xenoliths, are reported in Table 2. The Pb isotopic ratios of melt inclusions from the two Mangaia basalts show large and systematic variations that overlap with the Pb isotope compositions of whole rocks from the entire Cook-Austral Islands chain between Mangaia island and Macdonald seamount (Mangaia, Tubuaii, Rimatara, Rurutu, Raivavae, Rapa, Morotiri and Macdonald) (Figure 5 and 6). In the Mangaia melt inclusions,  $^{207}\text{Pb}/^{206}\text{Pb}$  varies from 0.706 to 0.815, and  $^{208}\text{Pb}/^{206}\text{Pb}$  ratios range from 1.840 to 2.02. These isotopic variations in the Mangaia melt inclusions are much larger than those observed for Mangaia whole rock lavas (0.726 to 0.735 and 1.869 to 1.883 respectively), and cover 50% of the entire range defined by the world-wide population of oceanic island basalts (OIB, Figure 5). In a similar way, the range in Pb isotopes of the Tahaa melt inclusions is larger than that observed for Tahaa whole rock lavas, and overlaps with the Pb isotopic composition of whole rock lavas from the entire Society Island chain (Cyana, Mehetia, Moua Piha, Rocard, Tahaa, Tahiti, and Teahitia) (Figure 7). For Tahaa melt inclusions,  $^{207}\text{Pb}/^{206}\text{Pb}$  varies from 0.809 to 0.824, and  $^{208}\text{Pb}/^{206}\text{Pb}$  ranges from 2.009 to 2.048, while the Pb isotopes for Tahaa whole rock lavas vary from 0.811 to 0.816 and from 2.027 to 2.032 respectively. The Pb isotopic composition of Pitcairn melt inclusions have a similar range in  $^{207}\text{Pb}/^{206}\text{Pb}$  ratios, but a larger range in  $^{208}\text{Pb}/^{206}\text{Pb}$  ratios than those reported for Pitcairn whole rock lavas (Figure 7). Thus, for Pitcairn melt inclusions,  $^{207}\text{Pb}/^{206}\text{Pb}$  and  $^{208}\text{Pb}/^{206}\text{Pb}$  vary from 0.841 to 0.875 and from 2.054 to 2.196, meanwhile the whole rock lavas range from 0.838 to 0.878 and from 2.107 to 2.210 respectively. Finally, the Pb isotopic composition of the melt pockets and clinopyroxene from the Tubuaii xenoliths reproduced exactly the Pb isotopic measurements reported by Hauri et al., (1993) (Figure 7). Thus, the range in  $^{207}\text{Pb}/^{206}\text{Pb}$  (0.739-0.751) and

$^{208}\text{Pb}/^{206}\text{Pb}$  (1.905-1.921) obtained by TIMS compare very well with the variation in Pb isotopes (0.746-0.750 and 1.918-1.925) obtained in this work by SIMS.

For samples with high Pb concentrations (measured  $^{208}\text{Pb}$  counts rates higher than 2,000 cps;  $\approx$  5 ppm Pb content), we measured  $^{206}\text{Pb}/^{204}\text{Pb}$  ratios on subsets of melt inclusions from Mangaia, Tahaa and melt pockets from Tubuaii xenoliths (Table 2). The range in  $^{206}\text{Pb}/^{204}\text{Pb}$  ratios for Mangaia and Tahaa melt inclusions is larger than that for whole rock lavas, while the  $^{206}\text{Pb}/^{204}\text{Pb}$  ratios measured on Tubuaii melt pockets by SIMS reproduce those measured by TIMS. Thus, the  $^{206}\text{Pb}/^{204}\text{Pb}$  ratios measured by SIMS substantiate the variation observed in  $^{207}\text{Pb}/^{206}\text{Pb}$  and  $^{208}\text{Pb}/^{206}\text{Pb}$  ratios. SIMS  $^{206}\text{Pb}/^{204}\text{Pb}$  ratios measured in melt inclusions and melt pockets range from 22.01 to 19.63 in Mangaia, from 18.94 to 19.73 in Tahaa, and from 21.07 to 21.16 in Tubuaii. These variation in  $^{206}\text{Pb}/^{204}\text{Pb}$  ratios in melt inclusions are very different from the ranges measured by TIMS on Mangaia and Tahaa whole rock lavas (from 21.480 to 21.932, and from 19.16 to 19.29 respectively), and similar to those measured in melt pockets of the Tubuaii xenolith (from 20.97 to 21.31), reproducing the variations observed in  $^{207}\text{Pb}/^{206}\text{Pb}$  and  $^{208}\text{Pb}/^{206}\text{Pb}$  ratios (Hart et al., 1992; Hauri and Hart, 1993; Hauri et al., 1993; Woodhead, 1996).

### **3.5 Discussion**

#### **3.5.1 Petrography, major and trace element composition of the melt inclusions**

The compositional zoning observed in daughter and host minerals near the inclusion is explained simply by post-entrapment fractional crystallization of the melt inclusion. Sometimes, euhedral phases inside the melt inclusions are too large to be daughter phases crystallized from the trapped melt. These are crystals trapped along with the melt at the same time, by the olivine or clinopyroxene host. Common trapped phases are Cr-spinel or clinopyroxene. They are strongly zoned, with the core composition similar to that of groundmas Cr-Spinel and the clinopyroxene phenocrysts respectively; meanwhile, the margins become more enriched in Ti produced during post-entrapment crystallization of the melt inclusion (see Appendix). As a cautionary note, trapped crystals are an important reason why melt inclusions must be studied first before they are homogenized.

Previous works suggested the presence of a carbonated mantle source for the magmas from the Cook-Austral islands, largely based on indirect evidence :



1) the compositional and mineralogical characteristics of peridotitic xenoliths from Tubuai island (Hauri et al., 1993);

2) the observed supra-chondritic Zr/Hf ratios in basaltic lavas from Tubuai and Mangaia islands (Dupuy et al., 1992); and

3) the similar geochemical and isotopic characteristics between carbonatitic magmas and some oceanic island basalts (Nelson et al., 1988). However, direct evidence for primary carbonate melts in peridotite xenoliths or in basaltic rocks has never been reported for the Cook-Austral chain.

The association of the carbonate globules with Ca or Mg volatile-rich silicate glasses (up to 15% volatiles), and with silica-alumina-alkali rich glasses in Mangaia melt inclusions, are essentially the same as those described for carbonate-silicate melt pockets in carbonated peridotitic xenoliths such as Spitsbergen, Canary Island, Tubuaii, Kerguelen and Northern Tanzania (Ionov et al., 1993; 1996; ; Kogarko et al., 1995; Hauri et al., 1993; Schiano et al., 1994; Rudnick et al., 1993; among others). The observed mineral associations, mineral and glass compositions, and the ocelli textures between carbonate globules and silicate glasses in the melt inclusions from both Mangaia lavas provide the first direct evidence for a carbonated mantle source for HIMU basalts (Table 1 and Appendix).

As mentioned previously, the ocelli texture between carbonate globules and silicate glass, and the presence of droplets of silicate glass in carbonate spherules, bounded by sharp, curved menisci (Figure 3a) has been considered as clear evidence of liquid immiscibility (Ionov et al., 1993; 1996; ; Kogarko et al., 1995). However, the REE content and the  $(La/Yb)_{PM}$  ratios in carbonate globules from the melt inclusions or melt pockets in peridotites, are generally lower than those reported for carbonatitic rocks that have been formed by immiscibility processes (Nelson, et al., 1988). To evaluate whether liquid immiscibility is the process responsible for the formation of the carbonate globules in the Mangaia melt inclusions, we used the phase diagram for the pseudo-quaternary system  $CaO-(MgO+FeO^*)-(Na_2O+K_2O)-(SiO_2+Al_2O_3+TiO_2)$  with  $CO_2$  at 1.0 and 2.5 GPa pressure (Lee and Wyllie, 1998). This diagram provides a petrogenetic framework for magmatic processes occurring within the mantle and deep crust (Figure 8 a and b). In Figure 8, the composition of the proposed immiscible carbonate globules and associated silicate glasses are compared with the 1.0 GPa and 2.5 GPa phase diagram. Moreover in

both Figures, 8a and b, we have also plotted the composition of carbonate globules and silicate glasses from Spitsbergen xenoliths (Ionov et al. 1993). Carbonate globules and silicate glasses from both, Mangaia melt inclusions and melt pockets from Spitsbergen xenoliths, have similar compositions, plotting well outside the miscibility gap. Therefore, the phase relations shown in Figure 8 indicates that none of the Mangaia silicate glass and carbonate globule compositions from the melt inclusions represent equilibrium immiscible liquids.

Carbonate globules and silicate glasses are too far removed from the miscibility gap, well inside the forbidden volumes of the primary silicate and carbonate liquidus respectively (Lee and Wyllie, 1998). Thus, although the ocelli texture is suggestive of immiscibility, the major and trace element composition of the carbonate globules and silicate glasses do not support this hypothesis. Lee et al. (1994) and Lee and Wyllie (1996 and 1997) were able to reconcile the textural with the chemical evidences. They demonstrated that, through a wide range in compositions, rounded calcites coexisting with silicate or carbonate liquids were in fact rounded single crystal phases. Therefore, some of the round calcites can be explained as crystal phases crystallizing from a silicate or silicate-carbonate melt.

Ionov et al., (1996) suggested that dolomite and magnesian calcite represent primary melts and that amorphous magnesite-ankerite phases are formed as a result of melting and breakdown of primary carbonate minerals in carbonate-bearing mantle xenoliths from Spitsbergen. Moreover, they consider that the Mg-rich silicate glass reported for Spitsbergen xenoliths (Amudsen, 1987 and Ionov et al., 1993) may have been formed by decompression-induced dissolution of olivine into a preexisting carbonate-rich melt. Due to the similarity in composition between the melt pockets in the Spitsbergen xenoliths and the Mangaia melt inclusions, we will follow the Ionov et al. (1996), and Lee and Wyllie (1998) interpretation. Thus, our results would suggest that Mangaia silicate glasses and carbonate globules record a complex history of crystal fractionation, and probably in some cases (Mg-rich volatile bearing glasses), decompression-induced dissolution or reaction of the olivine host in contact with the volatile-rich melt inclusions. Moreover, the round carbonate globules observed in the Mangaia melt inclusions are probably a phase crystallized from a primitive CO<sub>2</sub>-rich magma at pressures lower than 2.5 GPa.

Tahaa melt inclusions have glasses with simple compositions, recording fractional crystallization of Ti-augite, ilmenite, Cr-spinel and plagioclase and segregation of sulfide globules (Table 1 and Appendix).

The origin of the silica-alumina-alkali rich (phonolitic) glasses in mantle xenoliths and in Tahaa and Mangaia melt inclusions is controversial and several hypotheses have been proposed:

- 1) silicate-carbonate melt immiscibility (Ionov et al., 1993; 1996; Kogarko et al., 1995).
- 2) break-down of a crystal phase such as clinopyroxene during decompression-induced melting (Hauri et al., 1993)
- 3) low extent of partial melting of an amphibole peridotite (Schiano and Colocciatti, 1994)
- 4) partial melting of an eclogite created through subduction of oceanic crust, (Schiano et al., 1995)
- 5) extreme fractional crystallization of a small volume of melt percolating through the lithosphere (Schiano and Colocciatti, 1994)

The silica-rich glasses in the melt inclusions from Tahaa and Mangaia suggest that they are not necessarily related to silicate-carbonate immiscibility because:

- a) the Tahaa melt inclusions do not have carbonate globules; and
- b) the major element composition of the carbonate globules and silicate melts are not consistent with immiscibility.

Most of the melt inclusions are rounded and do not have sharp rectangular borders resembling pseudomorphs of a crystal phase (clinopyroxene) undergoing decompression melting. Therefore, we discard the hypothesis of a break-down of crystal phases by decompression-induced melting as the process responsible for the formation of the silica-rich glass. The silica-rich glasses are usually associated with many different daughter crystals (Table 1 and Appendix). Thus, when the composition of the melt inclusion or melt pocket is averaged over the whole inclusion, the composition is no longer phonolitic, but basaltic [see Appendix and compare the phonolitic compositions obtained by analyses done

with small beams (5 to 10  $\mu$ ) with the basaltic ones obtained with large beams (20 to 30 $\mu$ ) for the same melt inclusions]. Thus, the melt inclusions cannot be low extent melts from eclogites or amphibole peridotites because they do not have an averaged composition similar to that of the silica-alumina-alkali rich glasses. Therefore, the high silica glasses seem to be simple residual melts after post-entrapment fractional crystallization of the melt inclusions or after extensive fractional crystallization during melt percolation in the lithosphere prior to entrapment.

In the Mangaia and Tahaa melt inclusions, some of the trace element variability is likely a result of partial melting processes, but much of the trace element signal is complicated by extensive in-situ fractional crystallization. In the Mangaia melt inclusions, additional complexity is added by silicate-carbonate fractionation and probably by decompression-induced dissolution or reaction of crystal phases in the volatile-rich melt inclusions. Thus, while the presence of carbonate melts shows that the sub-Mangaia mantle is at least locally carbon-rich, the trace element data indicate little else about the Mangaia and Tahaa mantle source compositions, and it is difficult to deduce information regarding mantle source variability from the trace element compositions observed in these melt inclusions. Thus, the distinct mantle components beneath Mangaia and Tahaa are best resolved by measuring radiogenic isotopes on individual melt inclusions.

### 3.5.2 Pb isotopes

Figure 5, 6 and 7 show the large and systematic variations in Pb isotopic composition ( $^{207}\text{Pb}/^{206}\text{Pb}$  and  $^{208}\text{Pb}/^{206}\text{Pb}$ ) of melt inclusions from Mangaia, Tahaa and Pitcairn basalts. These variations are larger than the range in Pb isotopes reported for whole rock lavas for each island and overlap with the entire isotopic variation reported for each one of the islands chains: Cook-Austral chain, Society islands and Pitcairn island and seamounts.

The Pb isotope data for Mangaia, Tahaa and Pitcairn forms quasi-linear arrays that can be described by simple mixing of two end-members. For Mangaia, one component (end-member) is rich in radiogenic Pb and is identical to the HIMU component observed in erupted Mangaia lavas (Hauri and Hart, 1993; Woodhead, 1996). The Pb isotopic compositions of sulfide- and carbonate-bearing melt inclusions invariably reflect this end-member HIMU (Table 2, Figure 6). Whether the carbonation is responsible for the HIMU signature or merely accompanies it is still an unresolved issue. The second component is

defined by a less-radiogenic Pb composition (hereafter called the "less radiogenic component"). The exact identity of this component is difficult to establish, partly because of the analytical errors involved, as can be seen from Figure 6 and 7. It could be either the enriched mantle component observed in erupted lavas from other islands in the Cook-Austral chain, such as Macdonald seamount (Nakamura and Tatsumoto 1988; Chauvel et al., 1992; Chauvel et al., 1997; Hemond et al., 1994) or a depleted upper mantle component, such as the Pacific MORB (Zindler and Hart, 1986 and references therein). For the Tahaa melt inclusions, the Pb isotopic composition can also be explained by mixing between a component represented by the EMII mantle end-member, and a second component that is identical to the less radiogenic component observed in the Mangaia melt inclusions. The Pb isotopic compositions of sulfide from Tahaa melt inclusions are end-member EMII (Figure 7). The Pb data for the Pitcairn sample can be explained by a similar mixing trend between the EM I mantle end-member, and probably the same less-radiogenic component as in Mangaia and Tahaa (Figure 7). The less radiogenic component present in melt inclusions from the three end-member oceanic island basalts is similar to ubiquitous mantle components such as "FOZO" (Hart et al., 1992) or C (Hannan and Graham, 1996), although the geologic significance of this similarity is not clear at this point.

The Pb isotopic compositions in melt inclusions from sample MAG-B-47 (Mangaia) correlate negatively with the Cr content of daughter clinopyroxenes (Figure 9). The melt inclusions from Tahaa plot at one extreme of this correlation, corresponding to the less radiogenic component, and the lowest Cr content of daughter clinopyroxenes (Figure 9). The Cr variation in the daughter clinopyroxenes could be considered to reflect a different extent of in-situ fractionation resulting from olivine and Cr-spinel crystallization in the melt inclusion. However, the correlation between the Pb isotopes and the Cr content in the clinopyroxene cannot be explained by a simple crystal fractionation model and instead requires a process that couples the change of Cr in clinopyroxene with the Pb isotope composition of the melt inclusions. This negative correlation suggests that the melt corresponding to the less radiogenic component may have suffered differentiation before its entrapment, or represents a source material present in the mantle which is already differentiated (eclogite?). We can explain this correlation in several ways:

- 1) A HIMU melt underwent coupled assimilation and fractional crystallization processes (AFC); as the primitive HIMU melt differentiates, it assimilates lithospheric material (lower oceanic crust?) with a less radiogenic Pb isotope composition.

2) Mixing of two melts: one, a primitive HIMU melt and one melt having a less radiogenic signature that has undergone crystal fractionation (low Cr) within the oceanic lithosphere prior to the mixing.

3) The less radiogenic component could also be a melt originating from a previously differentiated source. For example, partial melts of eclogites produced during recycling of the oceanic crust.

4) It is entirely possible that, during basalt-lithospheric interaction, many of the olivines observed in the Mangaia samples could be xenocrysts coming from the lithosphere, carrying inclusions with compositions not consistent with the host lava. MAG-B-47 olivines have a range in Fo content from 86 to 89. Therefore, if some of the olivines come from the lithosphere (Fo 91-92), the lower range in Fo content of the olivine phenocrysts in Mangaia lavas may indicate a tendency for the xenocrysts to equilibrate with the host lava.

To test whether there is any indication that the less radiogenic Pb component is situated in the lithosphere of the Cook-Austral islands, we analyzed the melt pockets and clinopyroxene from two peridotitic xenoliths from Tubuaii (Hauri et al., 1993). From seven analyses (six on melt pockets and one on clinopyroxene), none gave an isotopic composition similar to the unradiogenic component, and all of the isotopic compositions measured are similar to those of Tubuaii whole rock lavas.

The SIMS Pb data, and the range of  $\delta^{18}\text{O}$  (from 4.98 to 5.03 for Mangaia, 5.71 for Tahaa and 5.22 for Pitcairn) previously measured in olivine phenocrysts from the same samples (Eiler et al., 1995; 1997), suggests that the less radiogenic component could be derived from either the mantle or oceanic lithosphere.

There exists no evidence from data on erupted Mangaia lavas for the second, less radiogenic, component that we observe in the melt inclusions. This lack of evidence in the erupted lavas for the less radiogenic Pb isotopic compositions suggests that the proportion of this component in the erupted lavas must be small. In the case of Pitcairn, we see a trend toward a less radiogenic component in the isotopic composition of the whole rock lavas (Figure 5) although it is not as extreme as the range shown by the melt inclusions. In contrast to Mangaia and Pitcairn, the Pb isotopic composition of whole rock lavas from Tahaa is within errors the same as the less radiogenic component observed in the melt

inclusions, suggesting a high proportion of this end-member and a dilution of the EMII end member in the Society chain. It is apparent from the Pb isotope data that shallow-level storage and mixing of distinct magma batches resulted in homogenization and dilution of the less radiogenic component within Mangaia and Pitcairn magmas, and dilution of the EMII component in the Tahaa lavas, prior to eruption. The melt inclusions are thus valuable samples for revealing magmatic diversity within the volcanic plumbing system beneath Mangaia, Pitcairn and Tahaa, and for providing information that is invisible in the erupted lavas. Moreover, we did not find melt inclusions with more extreme end-member isotopic compositions than those defined by the whole rocks. Therefore, the end-member Pb isotopic compositions remain as previously defined (Zindler and Hart, 1986; Hart et al., 1992).

In synthesis, the broad linear trend in  $^{207}\text{Pb}/^{206}\text{Pb}$  -  $^{208}\text{Pb}/^{206}\text{Pb}$  space produced by the melt inclusions from Mangaia, Tahaa and the Pitcairn samples reproduces the entire trend defined by the Austral and Society chains and the Pitcairn island and seamount groups. The inclusions preserve a record of melt composition of far greater isotopic diversity than that sampled in whole rock basalts. The results suggest that the entire Cook-Austral chain, the Pitcairn island group and the Society island array can be explained by mixing of the HIMU, EMI and EMII end-members with a common component (less radiogenic component), respectively. The HIMU Pb isotopic composition measured for the carbonated globules, in conjunction with their textures and compositions, provide direct evidence for a carbonated mantle source beneath the Austral Islands.

### **3.6 Conclusions**

The major and trace element composition of carbonate globules and silicate glasses from melt inclusions from phenocrysts from Mangaia basalts are not consistent with an immiscibility processes, but suggest simple crystallization of carbonates from a  $\text{CO}_2$ -rich silicate melt. Moreover, the silica-alumina-alkali-rich glasses in melt inclusions from Mangaia and Tahaa lavas, and in melt pockets in mantle xenoliths, seem to be the result of post-entrapment crystallization of the melt inclusions or extensive fractional crystallization during melt percolation in the lithosphere.

The observed HIMU Pb isotopic composition of carbonate globules, their ocelli texture, and mineral associations, and the mineral and glass compositions in melt inclusions

from two Mangaia basalts provide the first direct evidence for a carbonated mantle source for HIMU basalts.

The melt inclusion Pb isotope data for Mangaia, Tahaa and Pitcairn forms quasi-linear arrays that can be described by simple binary mixing between the HIMU, EMI and EMII end-member compositions and an unradiogenic component common to all three isotopic arrays. Therefore, the geochemical variation in trace and major elements from melt inclusions does not always represent variable melting processes of a single mantle source, but clearly reflect geochemical source heterogeneities as well.

The negative correlation between the Cr content in daughter clinopyroxene and the Pb isotope composition of the melt inclusions from Mangaia suggests that the melt corresponding to the less radiogenic component may have suffered differentiation before its entrapment, or represents a source material in the mantle that has already been differentiated (eclogite?).



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### **3.8 Table Caption**

**Table 1:** Major and trace element compositions for representative glasses, carbonates, sulfides and minerals from melt inclusions and their mineral hosts. Major elements (in wt%) were determined on the Jeol-733 electron microprobe at MIT, and trace elements (in ppm) were measured with a Cameca 6f ion probe at the Department of Terrestrial Magnetism. Glass 1 is common to both olivine and clinopyroxene hosts; glass 2 is commonly found in clinopyroxene host, and glass 3 is found in olivine host. The particular mineral phases listed in the table are all hosted in olivine, and the reported carbonates are all hosted in clinopyroxene, but similar mineral and carbonate compositions were found in clinopyroxene and olivine phenocrysts respectively. The composition of Glass 4 and plagioclase are from Tahaa melt inclusions, the other compositions are from Mangaia melt inclusions. Plag = plagioclase; Phl = phlogopite; Kaer = kaersutite; Sp = spinel; Ilm = ilmenite; Sph = sphene; Ti-Au = Ti-augite; H-Cpx = host clinopyroxene; H-Ol = host olivine; Ap = apatite, Cpy = chalcopyrite, Pn = pentlandite, Po = pyrrhothite, Carb = carbonate. Table 1 lists only representative analyses. A full listing of all of the analyses is reported in the Appendix.

**Table 2:** Pb isotopic composition for single melt inclusions from two basalts from Mangaia and one basalt each from Tahaa and Pitcairn island. Analyses of melt pockets and one clinopyroxene from two mantle xenoliths from Tubuaii are also reported. All samples analyzed were silicates except where designated. Two spinel-hosted melt inclusions from Pitcairn were analyzed also.

**Appendix:** We report major element compositions of representative glasses, daughter mineral phases, carbonates and sulfides from melt inclusions from Mangaia and Tahaa basalts. Each reported analysis is the average of two to four spot analyses. We also report trace element compositions for melt inclusions from both islands. Trace elements (in ppm) were measured with a Cameca 6f ion probe at the Department of Terrestrial Magnetism (for Mangaia melt inclusions) and with a Cameca 3f ion probe (for Tahaa melt inclusions). In each table, the number corresponds to a number assigned to each melt inclusion. The size of the beam used during the analysis is denoted between parenthesis. Note the different compositions obtained with a small beam ( $10\mu$ ) as compared with those of a larger beam ( $30\mu$ ) for the same inclusion. The table for olivine compositions reports analyses for olivine hosts. No daughter olivine has been found. The number for each olivine analysis

represents the melt inclusion number hosted by the olivine. The letters “NI” indicates a second analysis for the same olivine host near the melt inclusion. In the Appendix we report analyses for daughter and host-clinopyroxene. The analyses for the hosts are labeled with an “H”. The number is the melt inclusion number the clinopyroxene either hosts or crystallized from. We report analyses for daughter crystals such as amphibole (AMPH), ilmenite (ILM), Ti-magnetite (TI-MNT), plagioclase (PLAG), spinel, sphene, and biotite. Several spinels analyzed are included in olivine phenocrysts (“in OL”). When clinopyroxene or spinel are zoned, the analyses for core and margin are denoted with a “c” or a “m” respectively. Carbonate and sulfide major element compositions are also included in this appendix as are trace element compositions for melt inclusions, daughter clinopyroxene (“cpx”) and carbonates.

### 3.9 Figure Caption

**Figure 1:** Mass spectrum for  $^{206}\text{Pb}$ ,  $^{207}\text{Pb}$  and  $^{204}\text{Pb}$  obtained with the IMS Cameca 1270 ion probe. To resolve isobaric interferences, secondary ions were analyzed at a mass resolution of 3500 with no energy filtering. Note that we are able to separate isobaric interferences such as major element and REE oxides. The spectrum was taken on Loihi standard 158-4.

**Figure 2:** a and b) Reproducibility of  $^{207}\text{Pb}/^{206}\text{Pb}$  and  $^{206}\text{Pb}/^{204}\text{Pb}$  for 34 and 11 analyses of the Loihi 158-4 standard glass used during the course of the melt inclusion analyses respectively. The horizontal lines define the TIMS value reported for the standard and the 2 sigma standard error limits (Garcia et al., 1995). The grey circle define the mean of the 34 and 11 analyses c)  $^{208}\text{Pb}/^{206}\text{Pb}$  versus  $^{207}\text{Pb}/^{206}\text{Pb}$  for Loihi 158-4. The error bars are the 2 sigma in-run (internal) precision. We also add the trends produced by mass fractionation, correlated error in  $^{206}\text{Pb}$ , hydride and counting statistics. There is perhaps a small trend in the data in the direction defined by mass fractionation and counting statistics.

**Figure 3:** BSE microphotograph of two representative melt inclusions. a) Inclusion #25 from sample MAGB25. The inclusion contains droplets of silicate glass in carbonate spherules, bounded by sharp, curved menisci. Surprisingly, this one melt inclusion contains three different types of silicate glasses (G1; G2; G3) and two different types of carbonates (C1; C2) as reported in Table 1 and Figure 4 and 8. The host is clinopyroxene; note the crystallization of kaersutite along the contact between the host and the melt

inclusion. b) Typical and petrographically simple melt inclusion hosted by olivine with crystallization of daughter clinopyroxene; note how the clinopyroxene crystallizes along the contact between the olivine host and the inclusion.

**Figure 4:** Primitive mantle normalized trace element data for representative melt inclusions from Mangaia. (a) carbonates, (b) silicate glasses contained in the melt inclusions, and (c) whole rock composition for the two Mangaia basalts (Hauri and Hart, 1997). The shaded area defines the compositional range of 46 analyses of melt inclusions. Primitive mantle values are from Sun and McDonough, (1989).

**Figure 5:**  $^{208}\text{Pb}/^{206}\text{Pb}$  versus  $^{207}\text{Pb}/^{206}\text{Pb}$  ratios for whole rock basalts from Pacific MORB, Pitcairn, Marquesas, Society, St. Helena and the Cook-Austral islands. The shaded fields represent Pb isotopic compositions for the islands and seamounts from the Cook-Austral chain between Mangaia island and Macdonald seamount. The dashed fields represent the Society islands, where the field for Tahaa is highlighted. The field for Pitcairn island and adjacent seamounts is highlighted by vertical lines. The data used are from Zindler and Hart (1986); Hauri and Hart, (1993); Hart et al., (1992); Woodhead (1996); Nakamura and Tatsumoto, (1988); Chauvel et al. (1997); Hemond et al., (1994); Eiler et al., (1995); Chauvel et al., (1992); Lassiter, et al. (1998). Large open circles are the four end-members of the mantle tetrahedron (Hart et al., 1992) projected onto the  $^{208}\text{Pb}/^{206}\text{Pb}$ - $^{207}\text{Pb}/^{206}\text{Pb}$  plane.

**Figure 6:**  $^{208}\text{Pb}/^{206}\text{Pb}$  versus  $^{207}\text{Pb}/^{206}\text{Pb}$  ratios (Table 2) for individual olivine and clinopyroxene-hosted melt inclusions from Mangaia basalts MAG-B-47 (open circle) and MAG-B-25 (open square). Filled diamonds: sulfide globule and filled square: carbonate globule (both from MAG-B-25); filled triangle: carbonate globule from MAG-B-47. The error bars are in-run precision ( $2\sigma$ , standard error)

**Figure 7:**  $^{208}\text{Pb}/^{206}\text{Pb}$  versus  $^{207}\text{Pb}/^{206}\text{Pb}$  ratios (Table 2) for individual olivine-hosted melt inclusions from Tahaa TAA B 26 (open triangle), Pitcairn PIT 1 (gray filled square) and melt pockets and clinopyroxene from Tubuaii xenoliths TBA 4-11 and TBA 9-11 (open diamond). Inset shows a magnified view of the Pb isotopic composition for individual olivine-hosted melt inclusions from Tahaa basalt TAA-B-26 (open triangle); filled circle: sulfide globule composition (Table 2). The error bars are in-run precision ( $2\sigma$ , standard error)

**Figure 8:**  $(\text{SiO}_2+\text{Al}_2\text{O}_3+\text{TiO}_2)$ - $(\text{MgO}+\text{FeO}^*)$ - $(\text{CaO})$ - $(\text{Na}_2\text{O}+\text{K}_2\text{O})$  generalized pseudoquaternary phase diagram (in weight percent) a) at 1.0 GPa and b) at 2.5 GPa showing the major element compositions for silicate glasses and carbonate globules from Mangaia melt inclusions (Table 1) and from Spitsbergen mantle xenoliths (Ionov et al., 1996). The figures show the three major liquidus volumes (the miscibility gap, silicate and carbonate liquidus fields), and the liquidus surfaces between them. Contours and values for  $(\text{MgO}+\text{FeO}^*)$  of the surfaces are also shown. Glass 1 - filled triangles; glass 2 - filled diamonds; glass 3 - dotted circle; carbonate globules - black circle, black square, grey diamond and open triangles; silicate glasses and carbonate globules from Spitsbergen xenoliths - open diamonds (Ionov et al., 1996); whole rock lavas from Mangaia (MAG-B-47 and MAG-B-25 respectively; Hauri and Hart, 1997)) - open circle and open square. The small circles are the projection of each point from the  $(\text{MgO}+\text{FeO}^*)$  vertex onto the basal plane of the tetrahedron. The composition of the Mangaia silicate glasses and carbonate globules are very similar to the composition of those found in Spitsbergen xenoliths. The phase relationships at 1.0 GPa and at 2.5 GPa indicate that none of these carbonate and glass compositions represent equilibrium immiscible liquids.

**Figure 9:**  $^{207}\text{Pb}/^{206}\text{Pb}$  ratios for single melt inclusions versus the Cr content of daughter clinopyroxenes from Mangaia (MAG-B-47) and Tahaa (TAA-B-26). The negative correlation suggests that the “less radiogenic component” has suffered differentiation prior to entrapment by the olivine phenocrysts, or that it originated from partial melting of a differentiated source (recycled oceanic crust, eclogite). The error bars for the  $^{207}\text{Pb}/^{206}\text{Pb}$  ratios are in-run precision (2s standard errors); for Cr content the error bars are standard deviation (2s) from replicated analyses. Symbols as in Figure 6 and 7.



**Table 1: Representative mineral, sulfide, glass and carbonate analyses from melt inclusions from samples MAG-B-47, MAG-B-25 and TAA-B-26**

	Phl	Ap	Sph	Ilm	Kaer	Sp	Sp	Ti-Au	Plag	H- Cpx	H-Ol	H- Ol
SiO <sub>2</sub>	43.86	2.67	30.50	0.48	38.99	0.14	0.11	40.99	54.73	51.39	39.19	39.38
TiO <sub>2</sub>	4.11	0.00	36.95	54.02	8.45	7.54	17.99	6.08	0.00	0.85	0.00	0.00
Al <sub>2</sub> O <sub>3</sub>	13.11	0.51	1.89	0.27	14.35	17.97	5.97	12.66	28.27	3.00	0.03	0.03
Cr <sub>2</sub> O <sub>3</sub>	0.10	0.00	0.00	0.34	0.00	33.81	17.72	0.00	0.00	0.51	0.06	0.04
FeO	9.81	1.04	0.85	37.81	10.13	33.62	54.75	10.81	0.93	5.10	16.37	13.15
MnO	0.12	0.00	0.00	0.73	0.11	0.44	0.75	0.21	0.00	0.20	0.05	0.22
MgO	17.30	0.78	0.00	5.53	11.45	7.09	2.95	7.02	0.11	16.00	44.09	46.95
CaO	0.06	51.53	29.16	0.58	12.66	0.09	0.04	22.33	10.46	22.21	0.37	0.39
Na <sub>2</sub> O	0.80	0.00	0.02	0.00	2.91	0.00	0.00	0.67	2.29	0.45	0.00	0.00
K <sub>2</sub> O	8.46	0.00	0.00	0.00	0.41	0.00	0.00	0.00	3.23	0.00	0.00	0.00
P <sub>2</sub> O <sub>5</sub>	0.11	36.57	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.07	0.19	0.00	0.00	0.04	0.03	0.17
Total	97.84	93.09	99.37	99.76	99.45	100.77	100.46	100.78	100.02	99.75	100.18	100.33
Mg#	0.80	0.62	0.00	0.24	0.71	0.32	0.11	0.59	0.20	0.87	0.86	0.89
											Fo 84- 86	Fo 86- 89
	Cpy	Pn	Pn-Po	Po								
Fe	39.98	33.57	58.28	60.01								
Co	0.13	0.68	0.18	0.13								
Ni	0.33	31.59	2.16	0.32								
Cu	22.07	0.27	0.22	0.16								
Zn	0.15	0.21	0.12	0.18								
Pb	0.18	0.00	0.00	0.00								
S	35.58	33.03	37.95	39.39								
Total	98.42	99.34	98.90	100.20								

**Table 1: Representative glass and carbonate analyses from melt inclusions from samples MAG-B-47, MAG-B-25 and TAA-B-26 (continuation)**

	Glass 1	Glass 1	Glass 2	Glass 2	Glass 3	Glass 4	Carb	Carb	Carb
SiO <sub>2</sub>	57.39	58.61	55.07	52.95	43.69	58.53	0.02	0.71	0.19
TiO <sub>2</sub>	0.30	0.31	0.00	0.14	0.03	0.45			
Al <sub>2</sub> O <sub>3</sub>	21.54	21.74	20.78	18.84	0.99	22.04			
Cr <sub>2</sub> O <sub>3</sub>	0.00	0.00	0.00	0.03	0.02	0.09			
FeO	1.86	1.62	0.51	1.83	9.41	1.08	16.76	26.86	14.73
MnO	0.04	0.00	0.00	0.00	0.26	0.00	0.27	0.46	1.99
MgO	0.10	0.06	0.90	1.19	30.27	0.27	30.18	11.09	4.05
CaO	3.42	3.09	7.92	10.44	0.16	1.44	4.26	18.20	36.10
Na <sub>2</sub> O	7.81	8.00	0.31	0.37	0.00	6.43	0.38	0.15	0.08
K <sub>2</sub> O	2.14	2.12	0.77	0.72	0.00	6.83			
P <sub>2</sub> O <sub>5</sub>	1.08	0.60	0.12	2.23	0.00	0.88			
CO <sub>2</sub>							47.05	44.31	43.40
NiO	0.00	0.00	0.00	0.00	0.00	0.00			
Total	95.68	96.15	86.39	88.75	84.84	98.02	98.94	101.77	100.54
Mg#	0.10	0.08	0.80	0.59	0.88	0.36			
Rb	26.6			2.76	1.78		0.84		4.20
Ba	524			145	1.06	1077	33.3		100
Nb	74.9			7.94	8.22	42.6	29.5		90.7
La	131			6.31	4.59	51	29.9		60.2
Ce	232			10.9	8.82	76.4	54.7		137
Sr	972			14454	25.7	1022	313		10281
Nd	78.6			4.64	4.73	20.5	19.4		57.1
Sm	14.9			0.89	0.96	2.61	2.76		10.9
Zr	272			33.8	48.7	327	112		554
Hf	4.90			0.38	0.74		1.17		5.31
Eu	4.47			0.18	0.27		0.68		3.16
Ti	1528			96.5	3087	5113	805		16049
Gd	11.0			0.61	0.74		1.35		7.25
Dy	7.14			0.45	0.56	1.11	1.24		6.75
Y	33.9			2.70	3.42	10.5	7.67		50
Er	2.41			0.21	0.22		0.56		5.29
Yb	1.13			0.14	0.18	0.41	0.59		8.93

**Table 2: Pb isotopic composition of single melt inclusions from samples MAG-B-47, MAG-B-25, TAA-B-26 and PIT 1, and from melt pockets and clinopyroxene TBA 4 11 and TBA 9-11**

MAG B 47 inclusion #	207/206	in-run error (2s)	208/206	in-run error (2s)	206/204	in-run error (2s)
1	0.754	0.006	1.905	0.017		
2	0.743	0.005	1.887	0.013		
5	0.765	0.006	1.924	0.017		
6	0.731	0.009	1.890	0.022		
8 carbonate	0.728	0.010	1.890	0.026		
8	0.745	0.008	1.894	0.023		
9	0.768	0.012	1.940	0.032		
10	0.793	0.003	1.969	0.009		
11	0.793	0.004	1.974	0.014		
duplicate	0.804	0.009	1.979	0.024		
13	0.729	0.007	1.861	0.021		
16	0.735	0.006	1.868	0.017		
19	0.812	0.004	2.020	0.011		
23	0.738	0.003	1.892	0.011		
25	0.745	0.007	1.901	0.017		
33A	0.727	0.008	1.864	0.022		
36	0.813	0.005	2.022	0.014		
39	0.792	0.004	1.967	0.011		
42	0.744	0.004	1.883	0.010		
43	0.727	0.003	1.844	0.009		
44	0.735	0.004	1.860	0.013		
45A	0.721	0.008	1.854	0.023		
46	0.806	0.004	1.995	0.013		
48	0.805	0.003	1.987	0.008		
49	0.793	0.004	1.978	0.010		
50	0.748	0.008	1.911	0.027		
52	0.706	0.009	1.840	0.027		
55	0.776	0.006	1.957	0.019		
56	0.753	0.006	1.906	0.016		
57	0.803	0.003	1.986	0.007	19.64	0.31
duplicate	0.800	0.005	2.000	0.013		
58	0.783	0.011	1.964	0.026		
60	0.747	0.007	1.897	0.016		
61	0.813	0.009	2.003	0.006		
duplicate	0.796	0.008	1.986	0.019		
67	0.727	0.009	1.866	0.025		

**Table 2: Pb isotopic composition of single melt inclusions from samples MAG-B-47, MAG-B-25, TAA-B-26 and PIT 1, and from melt pockets and clinopyroxene TBA 4 11 and TBA 9-11 (continuation)**

MAG B 25 inclusion #	207/206	in-run error (2s)	208/206	in-run error (2s)	206/204	in-run error (2s)
1	0.731	0.007	1.866	0.019		
2	0.750	0.008	1.916	0.022		
12	0.748	0.009	1.899	0.025		
24	0.760	0.004	1.923	0.012		
25 carbonate	0.737	0.010	1.871	0.027		
25	0.731	0.002	1.866	0.005	21.46	0.12
28	0.732	0.005	1.865	0.013		
31E	0.740	0.003	1.891	0.009		
32 sulfide	0.727	0.002	1.865	0.007	22.01	0.16
35	0.808	0.009	2.003	0.022		
44	0.750	0.005	1.909	0.012	21.28	0.30
48	0.736	0.007	1.884	0.017		
54	0.739	0.003	1.885	0.009	21.61	0.21
TAA B 26 inclusion #	207/206	in-run error (2s)	208/206	in-run error (2s)	206/204	in-run error (2s)
1	0.814	0.003	2.027	0.007	19.73	0.20
E1	0.819	0.005	2.045	0.014		
E2	0.815	0.004	2.033	0.010		
3 sulfide	0.821	0.005	2.043	0.016		
4 sulfide	0.822	0.005	2.039	0.012	19.25	0.29
4	0.814	0.002	2.028	0.004	19.17	0.10
6	0.819	0.004	2.032	0.011		
7	0.823	0.004	2.036	0.011	19.21	0.26
14	0.823	0.006	2.044	0.015		
20	0.816	0.003	2.022	0.008		
21	0.810	0.003	2.024	0.007	19.33	0.17
duplicate	0.818	0.004	2.028	0.009		
22	0.814	0.004	2.015	0.009		
24	0.817	0.005	2.021	0.011		
28	0.809	0.003	2.023	0.007		
35	0.817	0.003	2.035	0.008		
38	0.816	0.007	2.040	0.023	19.54	0.24
duplicate	0.814	0.004	2.034	0.011		
40	0.816	0.006	2.033	0.013		
50	0.824	0.005	2.041	0.013		
51	0.809	0.008	2.048	0.020		
52	0.816	0.003	2.028	0.006	18.94	0.12

**Table 2: Pb isotopic composition of single melt inclusions from samples MAG-B-47, MAG-B-25, TAA-B-26 and PIT 1, and from melt pockets and clinopyroxene TBA 4 11 and TBA 9-11 (continuation)**

TAA B 26 inclusion #	207/206	in-run error (2s)	208/206	in-run error (2s)	206/204	in-run error (2s)
54	0.811	0.007	2.027	0.022		
59	0.816	0.003	2.026	0.008		
60	0.817	0.003	2.033	0.008		
61	0.810	0.003	2.013	0.009		
62E	0.815	0.003	2.030	0.007	19.26	0.16
62	0.820	0.006	2.036	0.014		
63	0.818	0.003	2.032	0.008		
65	0.812	0.004	2.009	0.013	19.27	0.21
68	0.814	0.003	2.023	0.009		
69	0.815	0.007	2.023	0.017		
PIT 1 inclusion #	207/206	in-run error (2s)	208/206	in-run error (2s)		
5A	0.850	0.002	2.118	0.006		
22	0.864	0.008	2.159	0.019		
23A spinel	0.868	0.004	2.154	0.009		
25B	0.865	0.003	2.165	0.008		
27A	0.870	0.004	2.177	0.010		
32A	0.874	0.002	2.196	0.006		
46A	0.841	0.008	2.054	0.022		
48A	0.875	0.004	2.190	0.009		
52A	0.849	0.003	2.116	0.008		
54 spinel	0.849	0.006	2.089	0.017		
59	0.863	0.004	2.173	0.012		
73A	0.874	0.003	2.195	0.008		
81	0.865	0.007	2.170	0.019		
81B	0.865	0.003	2.177	0.008		
83A	0.870	0.006	2.185	0.015		
89A	0.858	0.003	2.154	0.007		
91B	0.874	0.011	2.160	0.026		
TBA inclusion #	207/206	in-run error (2s)	208/206	in-run error (2s)	206/204	in-run error (2s)
9-1 melt 1	0.748	0.002	1.922	0.005	21.07	0.08
9-1 melt 2	0.750	0.001	1.925	0.003	21.12	0.13
9-1 melt 3	0.746	0.002	1.918	0.007		
9-1 cpx	0.756	0.005	1.914	0.014	21.32	0.06
4-11 melt 1	0.749	0.003	1.920	0.008	21.16	0.20
4-11 melt 2	0.749	0.003	1.916	0.008	21.16	0.17

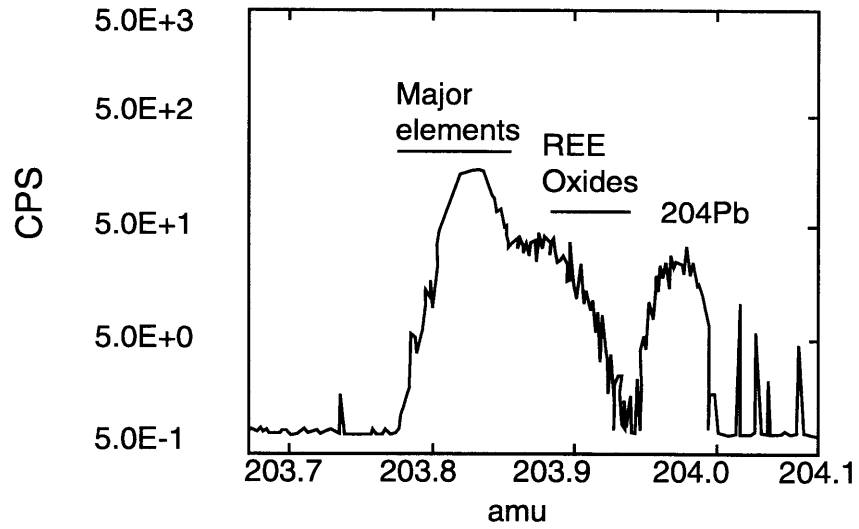
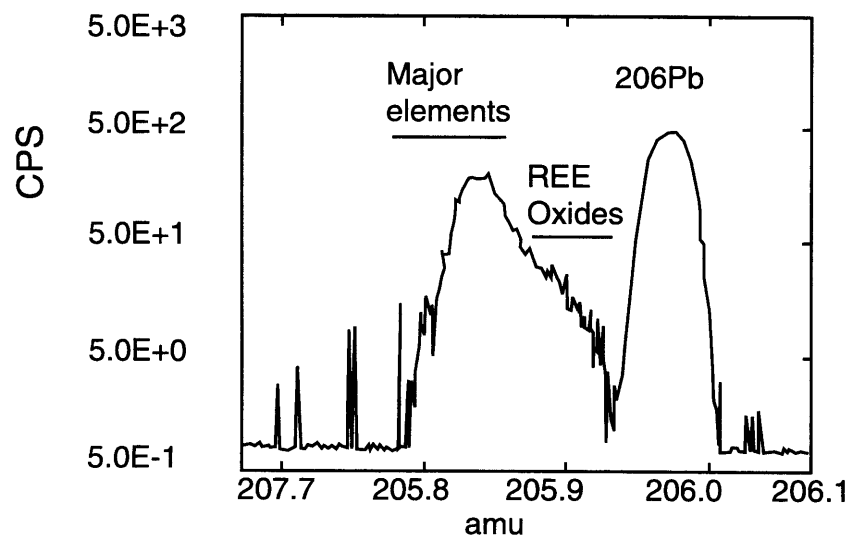
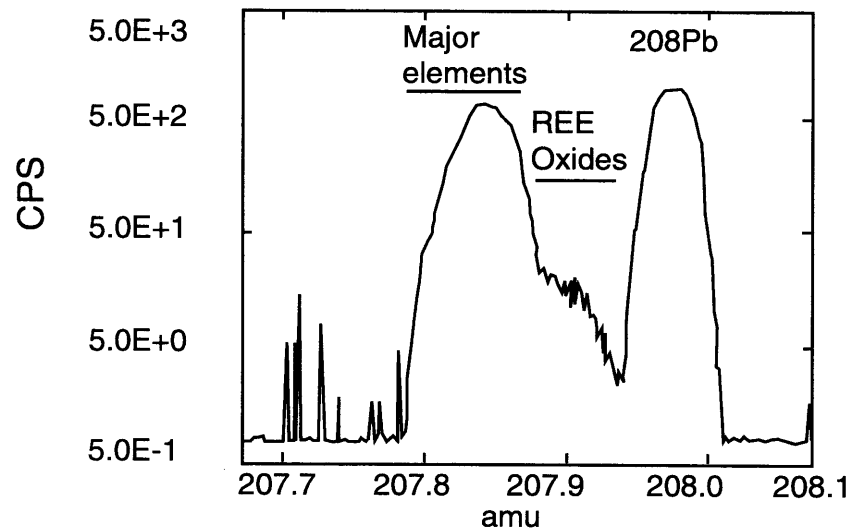


Figure 1

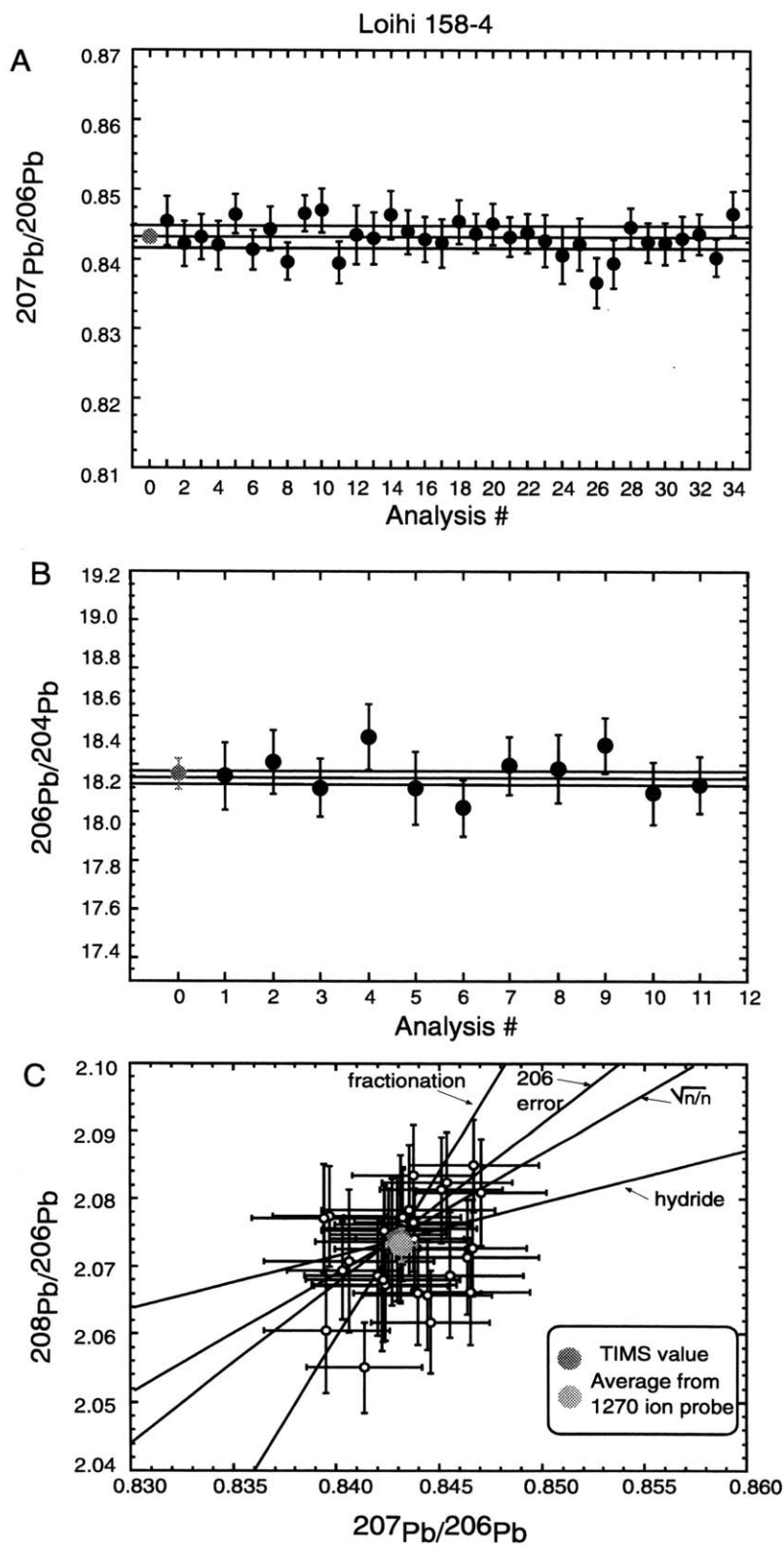


Figure 2

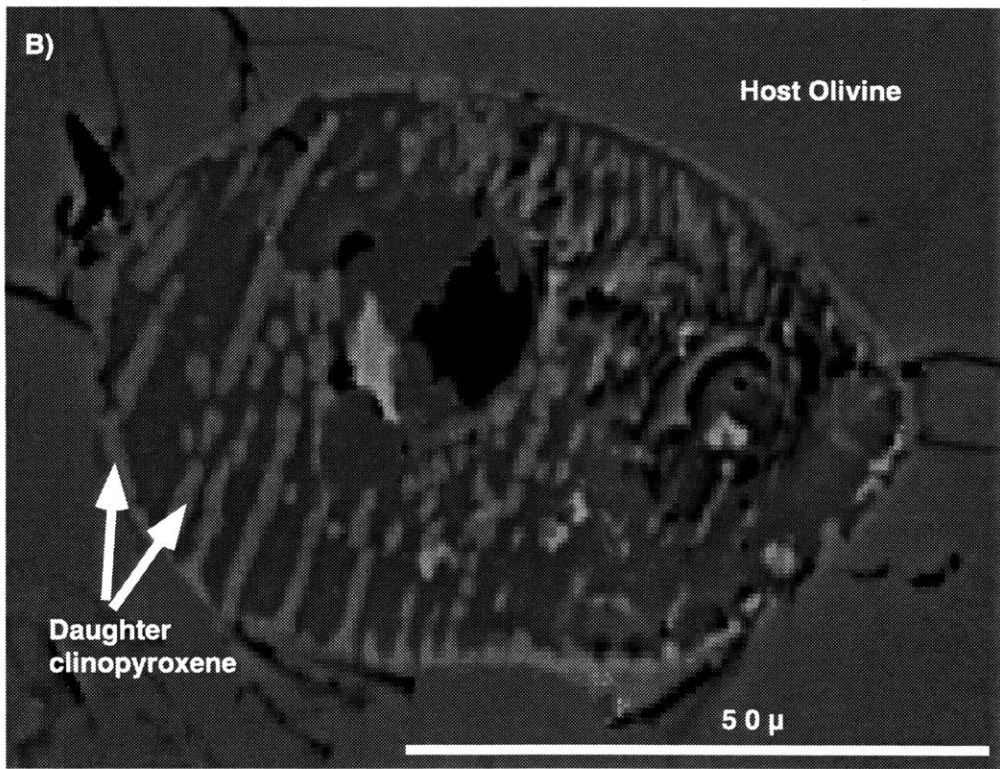
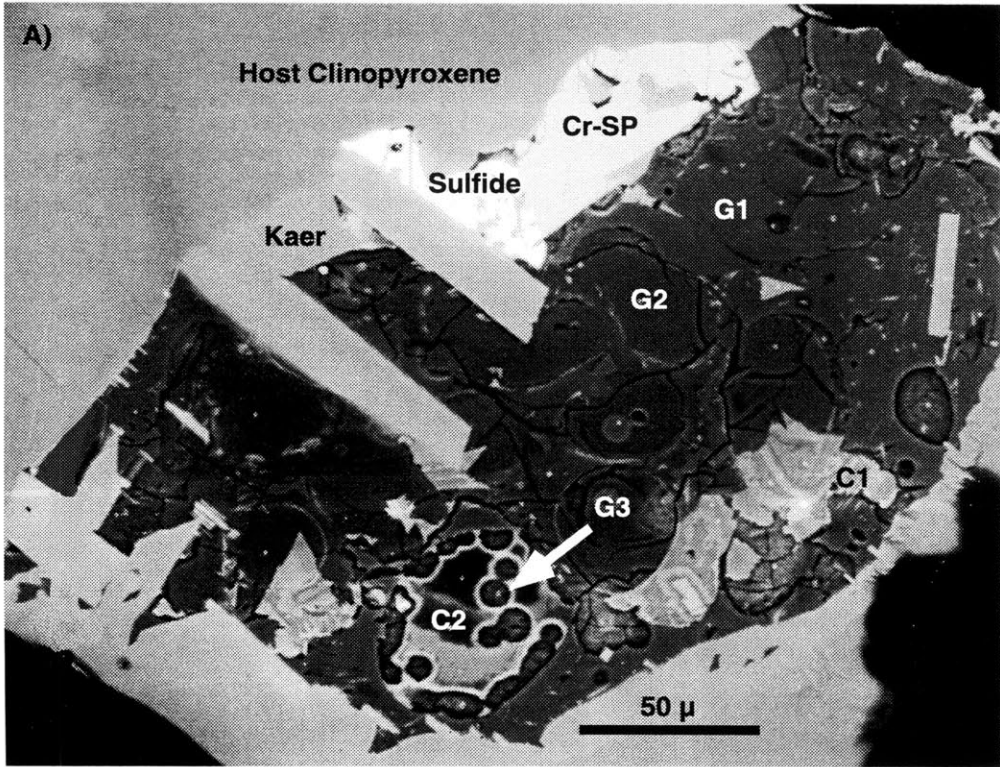


Figure 3



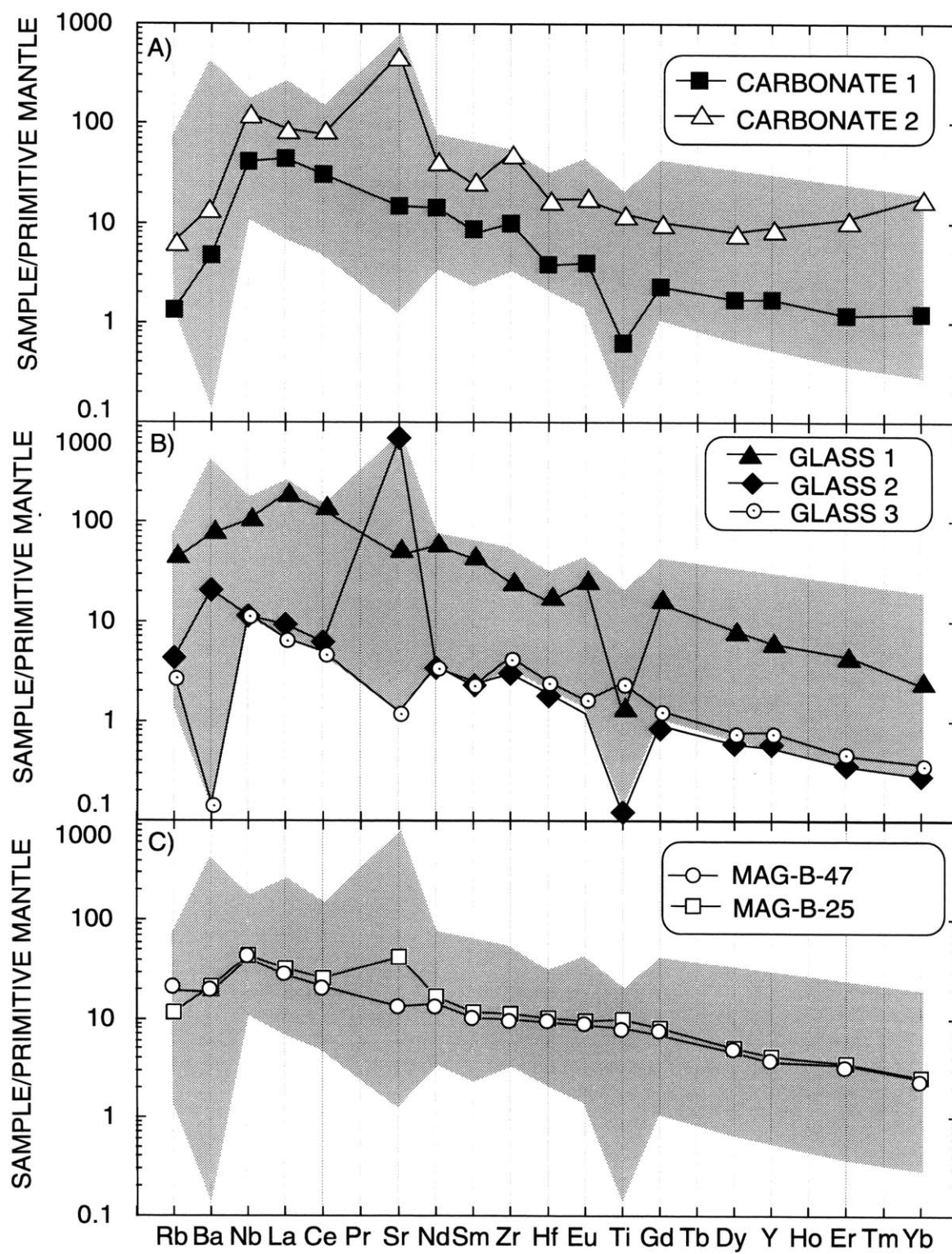


Figure 4

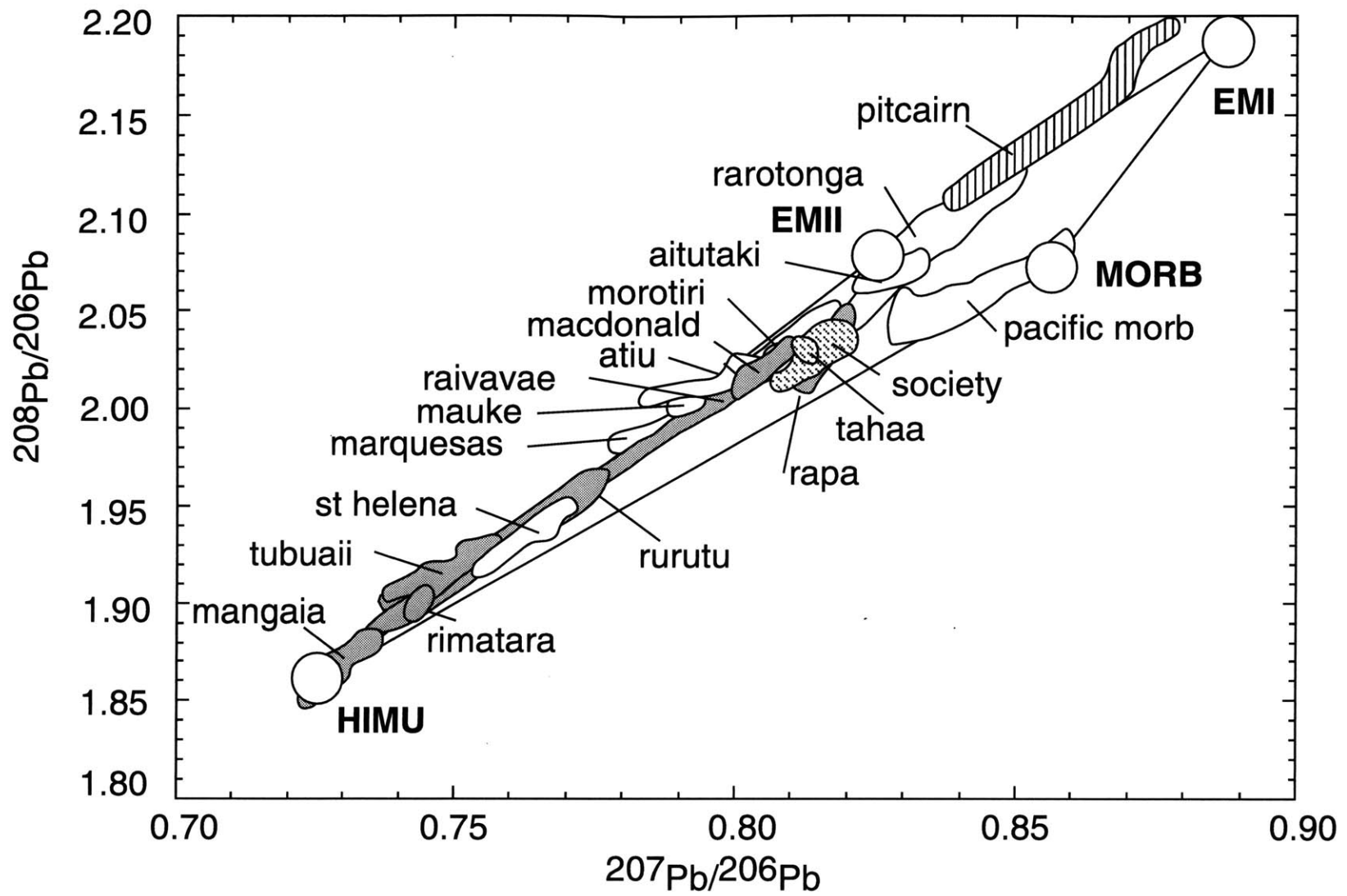


Figure 5

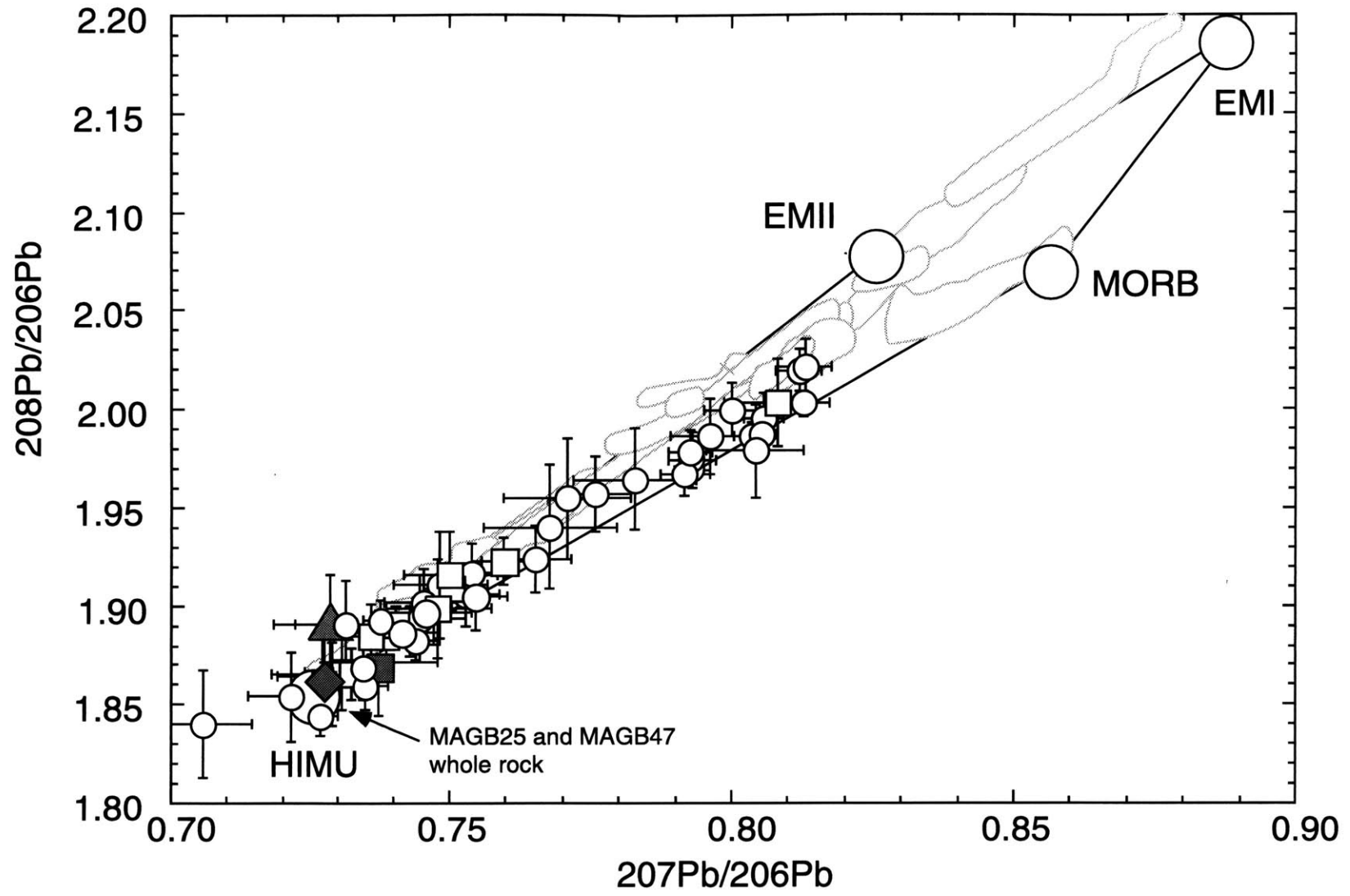


Figure 6

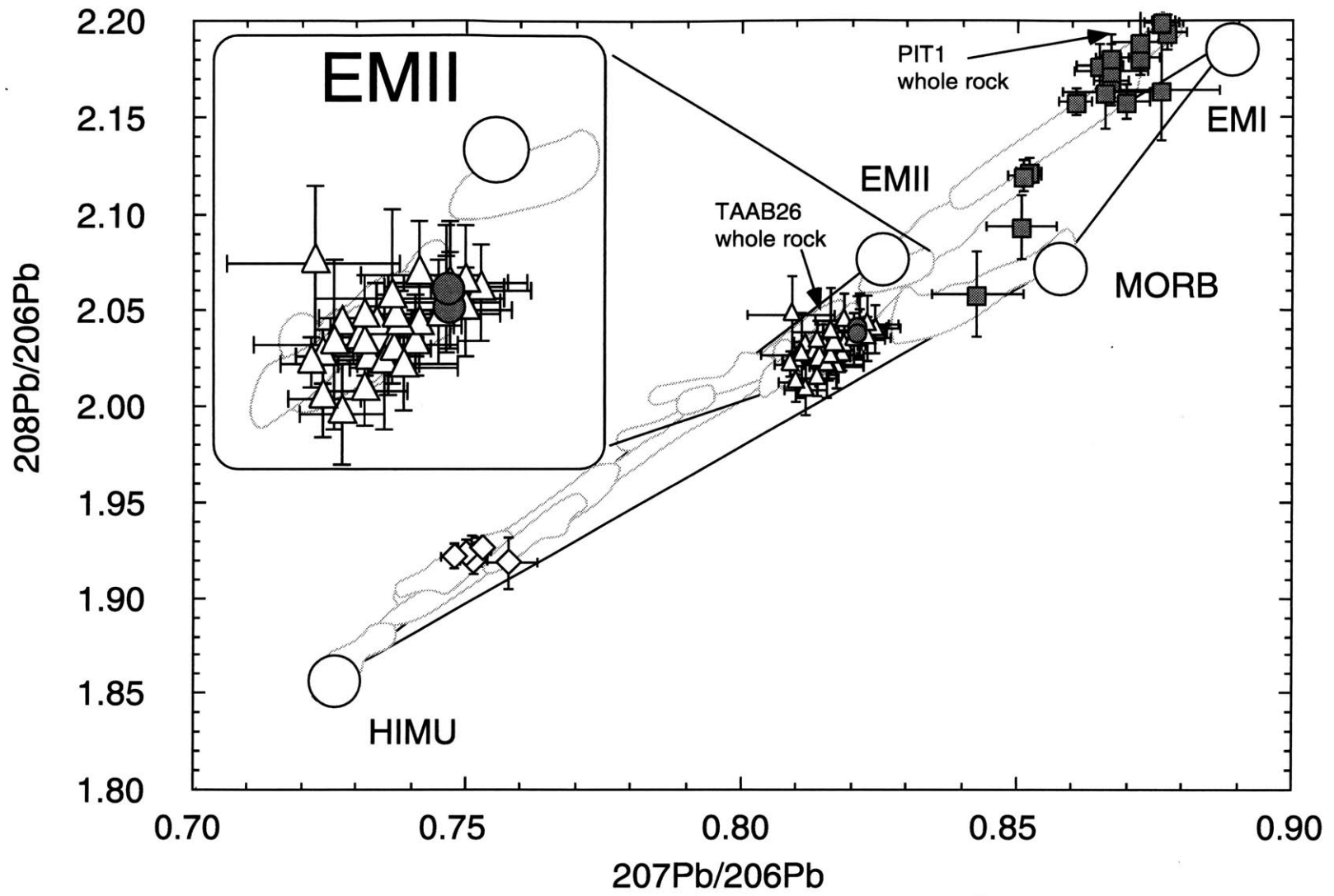


Figure 7

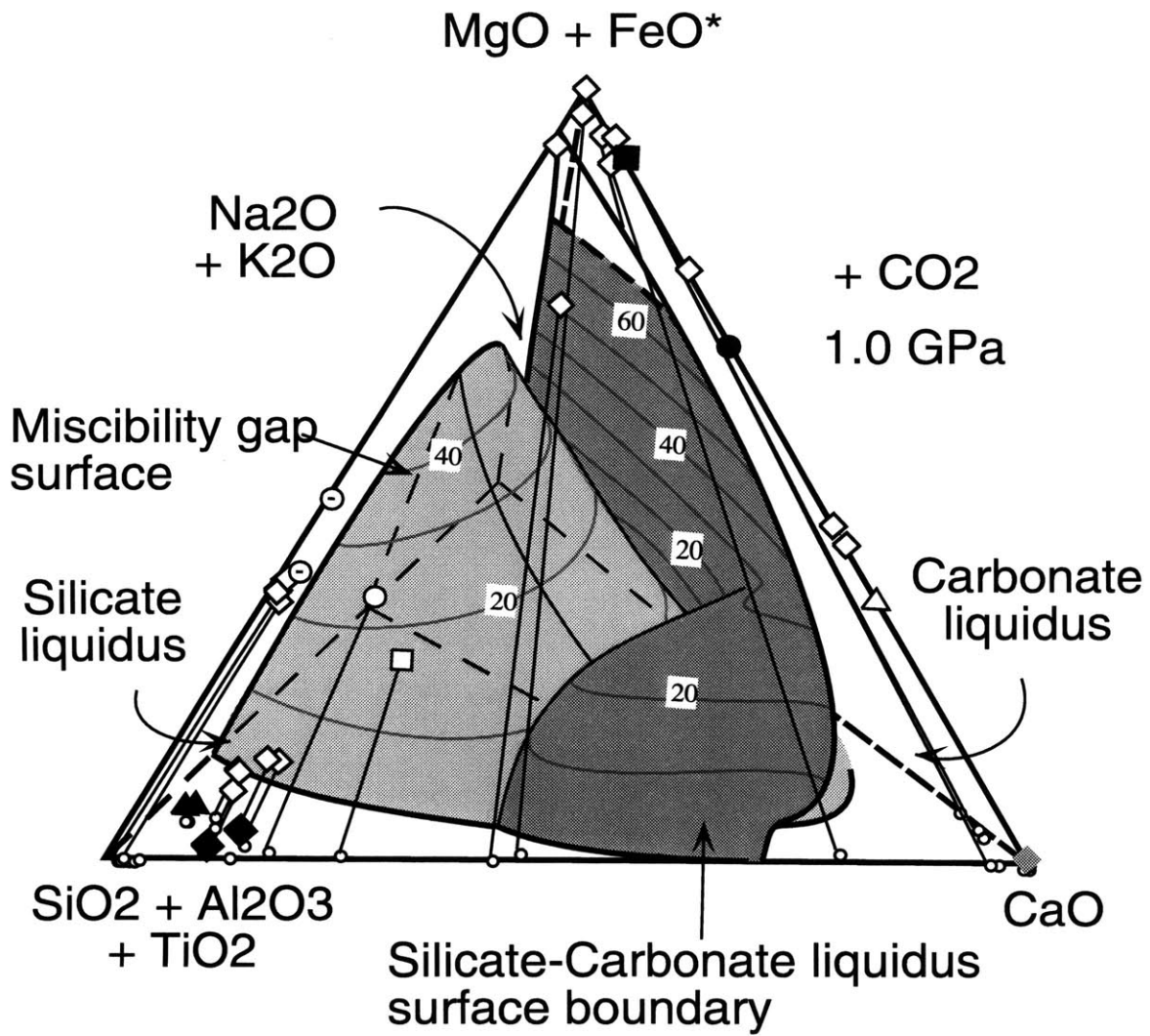


Figure 8a

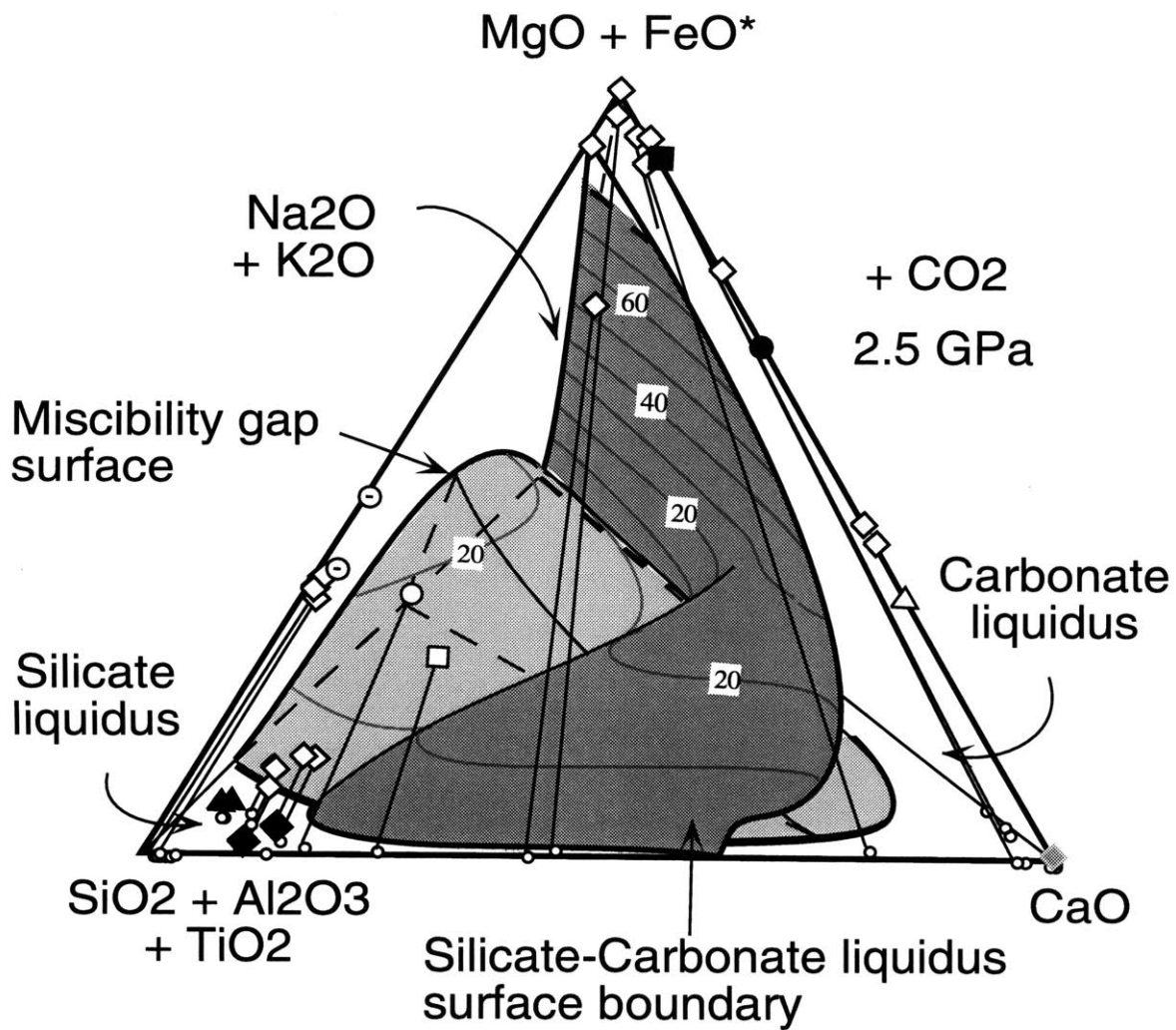


Figure 8b

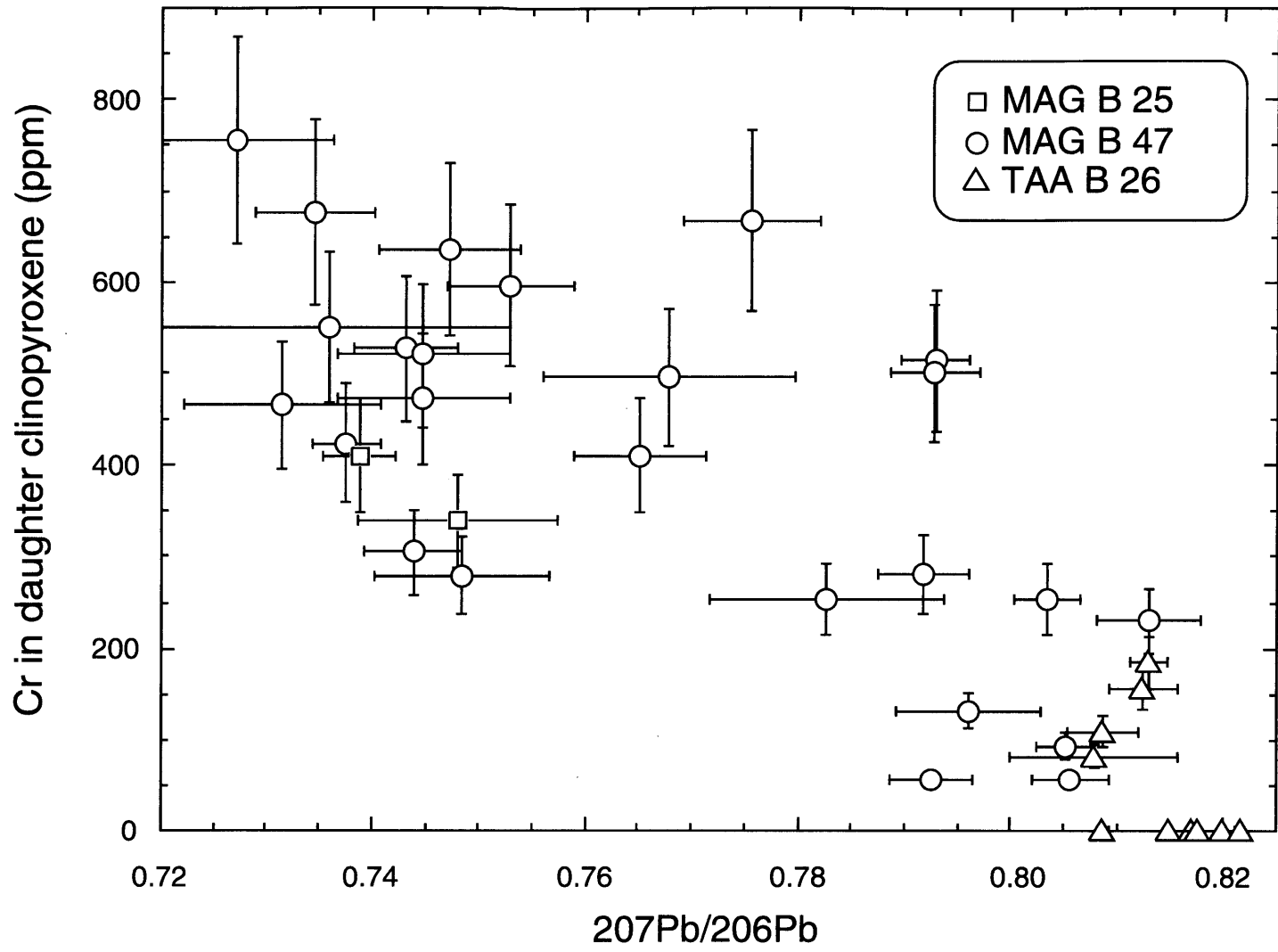


Figure 9

APPENDIX TABLE 1: Major element composition of melt inclusions

MAG B 47	1 (30 $\mu$ )	2 (10 $\mu$ )	2 (30 $\mu$ )	3 (10 $\mu$ )	3 (30 $\mu$ )	5 (10 $\mu$ )	5 (30 $\mu$ )	6 (10 $\mu$ )	6 (30 $\mu$ )	7 (10 $\mu$ )	7 (30 $\mu$ )	8 (10 $\mu$ )	8 (30 $\mu$ )	9 (10 $\mu$ )
SiO2	53.54	58.43	51.60	62.28	52.79	51.47	49.98	61.60	52.79	59.65	55.00	40.78	41.62	61.02
TiO2	2.16	0.17	2.75	0.29	2.15	0.24	2.06	0.05	2.09	0.23	1.80	0.08	2.50	0.07
Al2O3	20.07	29.24	18.61	25.74	19.24	28.27	19.67	25.25	18.91	24.38	20.54	0.06	4.71	25.51
Cr2O3	0.03	0.02	0.12	0.04	0.10	0.04	0.02	0.01	0.02	0.01	0.01	0.01	0.04	0.02
FeO	3.63	0.93	3.90	1.04	3.97	1.00	5.60	0.97	3.67	1.11	3.18	10.07	10.01	0.53
MnO	0.09	0.01	0.12	0.03	0.07	0.03	0.05	0.01	0.02	0.01	0.02	0.20	0.18	0.00
MgO	3.74	0.05	4.71	0.26	2.89	0.26	2.58	0.01	4.09	0.36	3.43	33.02	22.01	0.01
CaO	8.61	0.46	11.95	1.40	9.23	9.70	11.32	0.77	10.02	1.35	7.79	0.30	9.37	1.02
Na2O	6.57	8.50	5.13	7.32	6.97	5.17	5.89	8.92	6.71	9.16	5.46	0.00	0.35	10.17
K2O	1.55	2.98	1.51	2.82	1.75	3.39	2.83	3.96	2.25	3.73	2.51	0.00	0.04	2.70
P2O5	0.25	0.07	0.29	0.15	0.60	0.22	0.39	0.29	0.21	0.14	0.32	0.03	0.35	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	100.24	100.86	100.69	101.37	99.76	99.79	100.38	101.84	100.77	100.14	100.06	84.55	91.17	101.04
Mg#	0.70	0.11	0.73	0.31	0.62	0.36	0.50	0.03	0.71	0.40	0.70	0.88	0.83	0.03

MAG B 25	1 (15 $\mu$ )	1 (5 $\mu$ )	2 (5 $\mu$ )	2 (10 $\mu$ )	3 (10 $\mu$ )	3 (15 $\mu$ )	4 (10 $\mu$ )	5 (20 $\mu$ )	7 (10 $\mu$ )	8 (5 $\mu$ )	9 (5 $\mu$ )	11A (5 $\mu$ )	12 (5 $\mu$ )	13 (5 $\mu$ )
SiO2	59.42	60.73	61.83	48.93	57.39	43.03	51.36	51.94	58.11	54.25	56.40	60.61	61.44	61.05
TiO2	0.14	0.08	0.12	2.34	0.30	4.10	0.00	0.02	0.02	0.02	0.00	0.07	0.09	0.12
Al2O3	21.45	23.43	20.81	17.17	21.54	11.56	21.22	21.09	18.96	21.67	20.77	22.43	23.91	21.31
Cr2O3	0.00	0.03	0.00	0.00	0.00	0.04	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
FeO	1.49	1.15	1.11	7.91	1.86	10.74	0.61	0.91	2.69	0.02	0.00	0.00	1.24	0.00
MnO	0.37	0.06	0.00	0.14	0.04	0.67	0.23	0.10	0.09	0.00	0.00	0.00	0.04	0.00
MgO	0.00	0.02	0.11	4.89	0.10	8.37	0.40	0.93	3.08	0.65	0.81	0.00	0.11	0.03
CaO	2.89	2.83	1.30	6.72	3.42	17.05	8.50	8.17	2.11	9.46	8.16	2.45	1.27	1.20
Na2O	8.46	7.54	9.15	6.59	7.81	1.76	0.00	0.22	2.48	1.05	0.33	9.11	7.01	10.22
K2O	1.97	2.11	2.96	1.74	2.14	0.42	0.03	0.18	2.08	2.23	0.51	2.34	3.08	2.68
P2O5	0.41	0.81	0.19	0.40	1.08	0.44	0.07	0.04	1.31	0.70	0.08	0.30	0.68	0.62
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	96.59	98.78	97.59	96.83	95.68	98.15	82.42	83.59	90.92	90.04	87.06	97.31	98.87	97.23
Mg#	0.00	0.04	0.18	0.58	0.10	0.63	0.60	0.70	0.72	0.99	1.00	0.00	0.16	1.00

TAA B 26	1 (20 $\mu$ )	2 (20 $\mu$ )	3 (20 $\mu$ )	4 (20 $\mu$ )	5 (20 $\mu$ )	6 (20 $\mu$ )	6 (5 $\mu$ )	8 (20 $\mu$ )	9 (20 $\mu$ )	10 (5 $\mu$ )	10 (10 $\mu$ )	12 (20 $\mu$ )	14 (10 $\mu$ )	15 (10 $\mu$ )
SiO2	58.53	57.61	59.08	58.96	61.25	58.04	62.11	56.53	56.72	60.62	50.37	58.09	59.56	58.42
TiO2	0.45	0.49	0.41	0.39	0.43	0.41	0.28	1.48	0.44	0.44	5.62	1.50	0.42	0.39
Al2O3	22.04	22.66	21.90	22.28	22.37	21.40	21.96	18.38	23.95	21.44	17.34	18.83	21.68	22.71
Cr2O3	0.09	0.00	0.00	0.00	0.00	0.00	0.00	0.02	0.00	0.00	0.01	0.00	0.00	0.00
FeO	1.08	1.16	1.44	1.33	1.06	1.19	1.08	3.79	1.32	1.06	6.33	1.51	1.65	1.38
MnO	0.00	0.03	0.07	0.03	0.06	0.04	0.05	0.09	0.04	0.02	0.14	0.06	0.03	0.06
MgO	0.27	0.25	0.13	0.10	0.24	0.21	0.19	1.42	0.34	0.26	5.07	0.14	0.54	0.46
CaO	1.44	1.73	1.78	1.60	1.17	2.49	1.38	9.10	2.13	0.99	6.69	7.97	2.42	1.83
Na2O	6.43	6.78	6.80	7.25	6.66	7.22	3.52	4.73	8.68	6.69	4.33	6.52	6.87	7.48
K2O	6.83	6.45	7.17	6.08	6.80	6.51	4.94	3.85	4.54	6.95	4.35	3.45	5.82	4.64
P2O5	0.88	1.16	0.76	0.50	0.94	1.37	0.59	0.96	1.41	0.98	0.79	1.71	0.77	0.98
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	98.02	98.30	99.53	98.52	100.97	98.88	96.09	100.34	99.55	99.46	101.05	99.77	99.76	98.34
Mg#	0.36	0.32	0.17	0.15	0.34	0.28	0.28	0.45	0.36	0.35	0.64	0.17	0.42	0.42



APPENDIX TABLE 1: Major element composition of melt inclusions

MAG B 47	9 (30μ)	12 (20μ)	13 (10μ)	13 (30μ)	14 (10μ)	14 (30μ)	15 (10μ)	15 (30μ)	16 (30μ)	17 (10μ)	17 (30μ)	18 (10μ)	18 (30μ)	21 (20μ)
SiO <sub>2</sub>	53.98	39.39	40.99	42.55	55.27	48.03	35.40	36.44	54.08	52.51	49.36	61.79	50.65	44.75
TiO <sub>2</sub>	1.34	0.04	0.05	0.42	0.76	2.97	0.06	1.63	0.71	0.35	2.24	0.19	2.72	0.03
Al <sub>2</sub> O <sub>3</sub>	20.05	0.00	0.20	0.84	24.93	15.85	6.74	8.58	20.95	25.66	19.48	23.19	19.36	22.37
Cr <sub>2</sub> O <sub>3</sub>	0.03	0.04	0.05	0.05	0.09	0.12	0.28	0.38	0.03	0.02	0.03	0.02	0.03	0.02
FeO	3.37	26.70	11.79	9.43	1.17	4.31	24.32	11.36	2.23	1.04	3.79	1.03	4.16	2.87
MnO	0.02	0.35	0.19	0.21	0.00	0.09	0.26	0.06	0.05	0.00	0.04	0.00	0.06	0.27
MgO	2.07	32.24	32.31	32.33	1.54	7.78	19.14	18.40	1.00	0.70	5.25	0.27	4.10	1.46
CaO	8.91	0.41	0.07	0.10	3.78	15.36	0.43	10.30	13.21	3.60	9.01	1.39	10.37	7.43
Na <sub>2</sub> O	8.05	0.00	0.01	0.00	10.88	4.52	0.08	0.63	7.18	14.18	9.76	10.91	7.11	0.06
K <sub>2</sub> O	1.83	0.00	0.01	0.01	2.83	0.90	0.05	0.12	0.92	1.57	1.30	2.82	1.81	2.10
P <sub>2</sub> O <sub>5</sub>	1.21	0.14	0.00	0.00	0.38	0.35	0.00	0.94	0.22	1.41	0.28	0.32	0.69	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	100.85	99.32	85.67	85.94	101.64	100.27	86.75	88.83	100.58	101.02	100.54	101.92	101.06	81.36
Mg#	0.57	0.73	0.86	0.88	0.74	0.80	0.63	0.78	0.34	0.60	0.75	0.37	0.69	0.53
MAG B 25	14 (15μ)	15 (25μ)	15 (5μ)	24 (5μ)	24 (5μ)	25 (5μ)	25 (5μ)	25 (5μ)	27 (10μ)	28 (5μ)	29 (5μ)	30 (5μ)	31 (5μ)	31A (5μ)
SiO <sub>2</sub>	40.15	48.23	62.15	51.93	58.28	54.01	48.26	58.87	51.08	57.77	47.30	50.14	59.09	56.07
TiO <sub>2</sub>	6.38	1.88	0.19	0.12	0.23	0.07	0.21	0.20	0.09	0.21	0.04	0.05	0.10	0.27
Al <sub>2</sub> O <sub>3</sub>	13.26	16.06	20.89	20.52	22.56	20.03	9.02	21.48	19.75	21.40	18.92	0.85	20.74	23.77
Cr <sub>2</sub> O <sub>3</sub>	0.02	0.04	0.02	0.03	0.00	0.01	0.00	0.02	0.02	0.01	0.00	0.08	0.02	0.00
FeO	10.82	7.37	1.05	1.11	1.31	0.81	12.96	1.53	2.10	0.94	1.45	23.00	0.73	1.44
MnO	0.00	0.14	0.04	0.03	0.10	0.01	0.07	0.00	0.00	0.00	0.00	0.00	0.00	0.00
MgO	4.64	5.12	0.06	0.60	0.06	0.95	11.45	0.12	0.85	0.04	0.68	11.23	0.04	0.23
CaO	17.64	9.90	2.19	7.62	3.95	8.39	1.97	3.19	9.55	3.36	8.32	1.50	4.03	7.65
Na <sub>2</sub> O	1.52	5.43	8.37	0.50	6.45	0.29	0.22	7.90	0.61	8.64	0.11	0.16	8.14	7.39
K <sub>2</sub> O	0.52	1.36	2.82	0.58	2.25	0.66	0.36	2.17	1.10	2.33	0.20	0.41	2.20	0.41
P <sub>2</sub> O <sub>5</sub>	0.82	0.77	0.90	0.46	1.06	0.50	0.26	0.84	1.11	1.05	0.13	0.20	1.65	1.21
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	95.77	96.28	98.67	83.50	96.24	85.73	84.80	96.30	86.25	95.75	77.14	87.62	96.73	98.44
Mg#	0.49	0.61	0.11	0.57	0.09	0.75	0.66	0.14	0.54	0.07	0.64	0.52	0.10	0.22
TAA B 26	16 (10μ)	17 (10μ)	18 (20μ)	20 (20μ)	21 (20μ)	22 (30μ)	23 (30μ)	24 (30μ)	26 (10μ)	28 (30μ)	30 (30μ)	30 (10 μ)	32 (30μ)	33 (20μ)
SiO <sub>2</sub>	59.02	57.58	54.23	57.61	57.88	58.63	55.85	58.35	57.61	57.87	53.55	60.72	58.94	58.95
TiO <sub>2</sub>	0.37	0.23	1.81	0.32	0.23	0.31	1.01	0.51	0.96	0.81	1.84	0.63	0.38	0.47
Al <sub>2</sub> O <sub>3</sub>	23.82	23.20	20.45	22.99	20.20	22.21	19.53	22.47	21.05	20.37	17.58	20.59	22.50	22.99
Cr <sub>2</sub> O <sub>3</sub>	0.00	0.00	0.05	0.00	0.00	0.00	0.04	0.00	0.00	0.00	0.02	0.00	0.00	0.00
FeO	1.51	1.31	2.58	1.21	0.61	1.15	2.48	1.14	1.84	1.59	3.30	1.34	1.27	1.33
MnO	0.03	0.04	0.08	0.04	0.06	0.09	0.04	0.04	0.05	0.05	0.12	0.06	0.06	0.08
MgO	0.31	0.35	2.87	0.32	0.00	0.28	2.73	0.48	1.33	1.30	4.50	0.94	0.34	0.28
CaO	1.74	1.73	5.00	1.83	5.25	1.58	6.00	1.39	3.17	3.37	8.95	2.46	1.67	1.68
Na <sub>2</sub> O	7.59	6.82	6.98	7.33	5.73	6.71	5.88	7.40	6.68	5.85	5.01	6.80	6.99	7.37
K <sub>2</sub> O	5.16	5.23	4.35	6.37	6.44	6.78	4.29	7.02	5.13	6.58	3.44	5.16	6.02	5.66
P <sub>2</sub> O <sub>5</sub>	1.37	1.21	0.74	1.32	1.90	1.32	0.99	0.72	0.94	0.80	0.84	0.95	1.54	0.91
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	100.93	97.72	99.13	99.34	98.30	99.06	98.84	99.51	98.74	98.59	99.15	99.66	99.70	99.72
Mg#	0.31	0.37	0.71	0.37	0.00	0.35	0.71	0.48	0.62	0.64	0.75	0.61	0.37	0.32

APPENDIX TABLE 1: Major element composition of melt inclusions

MAG B 47	22 (10 $\mu$ )	24 (20 $\mu$ )	24 (10 $\mu$ )	25 (30 $\mu$ )	25 (10 $\mu$ )	29 (10 $\mu$ )	30 (10 $\mu$ )	31 (20 $\mu$ )	I31 (30 $\mu$ )	34 (10 $\mu$ )	36 (10 $\mu$ )	37 (30 $\mu$ )	41 (20 $\mu$ )	42 (10 $\mu$ )
SiO <sub>2</sub>	40.82	47.71	60.26	43.87	48.16	53.44	43.98	50.29	50.52	57.97	51.01	35.85	43.64	62.15
TiO <sub>2</sub>	0.06	0.38	0.36	0.00	0.02	1.17	0.00	2.09	2.85	0.11	0.23	0.72	1.52	0.09
Al <sub>2</sub> O <sub>3</sub>	2.79	28.65	19.61	0.76	1.10	21.14	0.78	18.79	19.48	24.35	29.40	6.18	13.10	23.27
Cr <sub>2</sub> O <sub>3</sub>	0.00	0.00	0.00	0.00	0.01	0.01	0.00	0.06	0.05	0.00	0.00	0.02	0.06	0.04
FeO	3.96	1.20	2.47	5.58	16.32	2.50	4.79	4.60	5.33	0.95	0.61	5.47	7.52	0.78
MnO	0.27	0.03	0.03	0.08	0.14	0.03	0.07	0.08	0.13	0.00	0.00	0.17	0.19	0.00
MgO	36.53	0.37	1.94	30.55	21.23	2.91	34.26	2.06	2.83	0.00	0.00	23.79	8.66	0.00
CaO	1.05	6.62	5.47	0.13	0.47	7.47	0.07	12.15	14.82	0.62	0.75	4.44	14.88	0.94
Na <sub>2</sub> O	0.00	13.15	3.00	0.00	0.07	11.00	0.00	4.75	4.90	10.25	15.77	0.00	0.43	11.22
K <sub>2</sub> O	0.00	1.70	7.45	0.00	0.12	0.95	0.00	0.78	0.45	4.68	2.34	0.00	1.42	2.48
P <sub>2</sub> O <sub>5</sub>	1.03	2.97	0.41	0.33	0.40	1.02	0.19	0.25	0.83	0.06	0.00	0.57	0.40	0.01
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	86.51	102.76	100.99	81.31	88.05	101.64	84.14	95.89	102.18	98.99	100.10	77.19	91.82	100.98
Mg#	0.95	0.41	0.63	0.92	0.74	0.66	0.94	0.50	0.54	0.00	0.00	0.91	0.72	0.00

MAG B 25	32 (5 $\mu$ )	33 (5 $\mu$ )	34 (5 $\mu$ )	34 (15 $\mu$ )	35 (5 $\mu$ )	36 (5 $\mu$ )	36 (5 $\mu$ )	37 (5 $\mu$ )	38 (5 $\mu$ )	39 (5 $\mu$ )	40 (20 $\mu$ )	40A (20 $\mu$ )	41 (5 $\mu$ )	42 (5 $\mu$ )
SiO <sub>2</sub>	46.93	41.58	62.71	52.19	50.63	50.81	50.16	60.90	51.85	48.62	51.60	50.55	59.35	60.06
TiO <sub>2</sub>	0.25	0.02	0.17	1.85	0.04	0.14	0.04	0.24	0.03	0.10	1.99	1.90	0.11	0.13
Al <sub>2</sub> O <sub>3</sub>	6.72	10.67	23.78	17.88	19.69	19.53	20.61	19.25	18.73	19.43	16.87	17.25	21.79	21.15
Cr <sub>2</sub> O <sub>3</sub>	0.03	0.00	0.00	0.00	0.01	0.03	0.01	0.00	0.02	0.00	0.06	0.02	0.01	0.00
FeO	18.28	15.09	0.77	4.41	1.62	2.89	0.34	1.83	1.84	0.36	5.08	6.71	0.84	0.70
MnO	0.00	0.00	0.05	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.16	0.00	0.00	0.00
MgO	10.97	15.02	0.09	2.50	1.08	1.93	0.65	1.10	1.31	0.83	3.01	3.48	0.04	0.03
CaO	3.39	0.99	1.80	9.16	8.22	3.69	10.07	3.12	7.58	10.89	8.99	6.83	3.66	3.69
Na <sub>2</sub> O	0.08	0.06	3.79	4.96	0.19	6.28	0.06	7.80	1.44	0.11	6.67	5.36	8.33	7.99
K <sub>2</sub> O	0.65	0.15	2.30	1.71	0.43	2.52	0.11	2.59	2.26	0.65	1.72	2.96	2.13	2.22
P <sub>2</sub> O <sub>5</sub>	1.70	0.09	0.02	0.82	0.36	0.38	1.07	0.55	0.32	2.40	0.85	0.79	1.35	1.18
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	89.00	83.66	95.48	95.48	82.26	88.20	83.13	97.38	85.37	83.39	96.99	95.84	97.60	97.15
Mg#	0.57	0.69	0.21	0.56	0.69	0.60	0.81	0.57	0.61	0.83	0.57	0.53	0.09	0.09

TAA B 26	34 (20 $\mu$ )	35 (30 $\mu$ )	36 (30 $\mu$ )	37 (20 $\mu$ )	38 (20 $\mu$ )	40 (30 $\mu$ )	41 (30 $\mu$ )	43 (10 $\mu$ )	48 (10 $\mu$ )	48 (10 $\mu$ )	49 (30 $\mu$ )	50 (30 $\mu$ )	51 (30 $\mu$ )	52 (10 $\mu$ )
SiO <sub>2</sub>	58.19	57.37	52.61	58.64	53.94	55.23	50.03	58.10	54.56	53.80	57.66	53.88	57.47	58.26
TiO <sub>2</sub>	0.48	1.00	1.93	0.54	1.91	1.86	3.34	0.28	2.12	2.42	1.74	1.51	0.47	0.39
Al <sub>2</sub> O <sub>3</sub>	22.35	21.21	19.66	21.17	20.20	19.29	19.04	20.86	20.34	20.71	19.11	18.48	22.41	21.77
Cr <sub>2</sub> O <sub>3</sub>	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
FeO	1.09	2.21	3.05	1.39	3.16	3.13	4.18	1.68	2.78	3.31	1.98	3.17	1.05	1.29
MnO	0.06	0.07	0.06	0.03	0.12	0.10	0.06	0.04	0.09	0.04	0.04	0.09	0.05	0.05
MgO	0.31	0.81	3.68	0.69	3.18	1.88	4.28	0.03	4.20	3.65	0.38	1.21	0.28	0.08
CaO	2.38	4.07	6.15	2.10	5.31	6.58	7.12	1.95	4.98	5.05	6.99	8.82	2.08	2.37
Na <sub>2</sub> O	7.15	6.85	5.16	6.14	5.46	6.02	4.54	7.13	2.49	4.13	5.07	6.07	6.35	6.85
K <sub>2</sub> O	5.58	5.03	4.84	6.68	5.28	4.07	4.87	7.38	2.30	3.64	5.41	3.95	6.91	6.62
P <sub>2</sub> O <sub>5</sub>	1.12	0.38	0.52	0.94	0.77	1.10	1.14	1.22	0.83	0.86	1.15	0.87	1.24	0.85
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	98.72	98.99	97.65	98.32	99.33	99.26	98.58	98.67	94.70	97.62	99.51	98.05	98.32	98.52
Mg#	0.38	0.45	0.73	0.52	0.69	0.57	0.69	0.03	0.77	0.71	0.30	0.46	0.37	0.12

APPENDIX TABLE 1: Major element composition of melt inclusions

MAG B 47	42 (30μ)	45A (10μ)	48 (10μ)	51 (10μ)	51A (10μ)	53 (10μ)	54 (10μ)	56 (10μ)	57 (10μ)	58 (10μ)	59 (5μ)	60 (10μ)	61 (10μ)	62 (15μ)
SiO2	52.96	51.75	45.44	49.76	43.86	36.88	37.32	47.70	40.46	50.58	46.58	43.04	52.24	39.97
TiO2	1.40	0.07	0.10	0.10	4.11	0.11	0.06	0.37	2.45	1.79	0.08	1.76	0.22	3.26
Al2O3	20.36	29.79	12.56	27.69	13.11	9.99	9.39	7.30	9.89	18.74	10.47	15.26	1.56	8.24
Cr2O3	0.06	0.00	0.00	0.01	0.10	0.02	0.01	0.00	0.00	0.03	0.02	0.06	0.00	0.03
FeO	2.96	0.69	9.73	1.74	9.81	19.03	12.57	3.22	7.83	2.22	15.97	6.78	17.17	8.49
MnO	0.10	0.00	0.16	0.01	0.12	0.23	0.14	0.05	0.08	0.00	0.56	0.04	0.17	0.36
MgO	1.48	0.06	14.76	1.28	17.30	21.24	24.25	23.14	12.14	4.67	9.86	5.69	12.99	14.09
CaO	7.24	0.09	2.43	1.40	0.06	0.71	0.41	1.86	11.13	9.09	0.56	9.83	1.07	15.36
Na2O	8.60	16.94	0.07	15.34	0.80	0.16	0.15	0.34	0.85	8.02	0.19	0.59	0.10	0.60
K2O	1.95	2.04	0.53	2.01	8.46	0.06	0.14	0.32	0.22	1.48	0.70	0.16	0.05	0.01
P2O5	1.17	0.05	0.00	0.98	0.11	0.06	0.00	0.11	0.27	0.36	0.09	0.11	0.05	0.24
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	98.28	101.48	85.78	100.32	97.84	88.48	84.44	84.40	85.33	96.97	85.07	83.31	85.63	90.66
Mg#	0.53	0.16	0.76	0.62	0.80	0.71	0.80	0.94	0.77	0.82	0.58	0.65	0.62	0.79

MAG B 25	43 (20μ)	43 (10μ)	44 (10μ)	45 (5μ)	46 (20μ)	47 (15μ)	48 (20μ)	49 (20μ)	50 (10μ)	52 (10μ)	53 (10μ)	53 (10μ)	54 (10μ)	55 (10μ)
SiO2	47.84	46.91	59.22	58.24	42.66	45.96	46.89	42.88	57.90	58.35	52.18	49.15	48.44	26.98
TiO2	2.82	0.03	0.13	0.13	2.38	2.96	2.91	2.07	0.06	0.56	0.07	0.07	0.10	0.00
Al2O3	14.23	10.45	21.01	21.83	18.12	16.72	14.97	11.45	5.11	21.16	8.19	2.52	0.91	2.94
Cr2O3	0.03	0.03	0.00	0.00	0.02	0.06	0.09	0.04	0.00	0.00	0.03	0.06	0.05	0.09
FeO	7.64	12.13	0.59	0.83	7.30	8.18	6.56	7.91	11.23	1.54	10.03	21.31	16.20	11.34
MnO	0.00	0.01	0.00	0.00	0.00	0.00	0.00	0.04	0.00	0.00	0.00	0.00	0.00	0.01
MgO	6.70	20.72	0.05	0.02	4.54	5.36	3.37	6.87	13.18	0.42	6.81	11.57	16.15	11.84
CaO	12.73	0.66	2.99	3.77	9.30	10.07	12.80	11.84	1.44	4.07	5.73	1.30	2.20	6.45
Na2O	4.25	4.46	7.96	8.02	4.68	4.91	5.05	4.27	0.18	7.42	1.04	0.07	0.12	0.09
K2O	1.11	1.40	2.18	2.13	1.05	1.23	1.14	0.84	0.64	2.19	0.85	0.31	0.24	0.13
P2O5	0.56	0.38	0.49	1.22	0.38	0.65	1.30	0.32	0.10	0.28	3.20	0.09	0.11	0.16
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	97.91	97.17	94.63	96.19	90.43	96.08	95.07	88.52	89.85	95.98	88.13	86.45	84.52	60.03
Mg#	0.66	0.79	0.15	0.06	0.58	0.59	0.53	0.66	0.72	0.26	0.60	0.55	0.68	0.70

TAA B 26	52 (20μ)	53 (20μ)	54 (30μ)	55 (20μ)	59 (10μ)	59 (30μ)	60 (30μ)	61 (30μ)	62A (30μ)	63 (30μ)	66 (20μ)	67 (10μ)	68 (30μ)
SiO2	52.65	59.14	52.34	57.79	60.88	58.21	57.44	58.14	58.66	57.51	56.61	56.87	57.55
TiO2	1.83	0.40	2.48	1.78	0.22	1.06	0.45	0.47	0.27	0.48	0.39	0.69	0.77
Al2O3	20.78	22.15	19.05	19.51	22.10	21.14	21.94	21.55	21.31	21.53	23.11	22.76	20.41
Cr2O3	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
FeO	3.74	1.14	3.75	1.34	1.16	2.00	1.27	0.93	1.91	1.01	1.24	1.28	1.72
MnO	0.07	0.03	0.10	0.03	0.03	0.08	0.07	0.08	0.08	0.05	0.04	0.05	0.04
MgO	1.84	0.38	4.09	0.32	0.20	1.02	0.35	0.45	0.30	0.20	0.43	0.55	0.45
CaO	8.84	1.45	6.58	4.22	1.67	2.55	1.68	1.87	1.48	2.97	2.35	2.87	5.47
Na2O	4.47	6.58	4.93	8.00	7.31	6.80	7.40	6.31	6.36	7.12	6.69	7.39	6.31
K2O	4.63	6.65	4.88	5.01	5.90	5.39	7.39	6.97	7.06	6.20	6.23	4.74	5.25
P2O5	0.83	0.70	0.85	0.75	0.69	0.55	0.67	1.22	1.06	1.79	1.55	1.50	0.86
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	99.68	98.61	99.04	98.74	100.15	98.79	98.65	98.00	98.48	98.85	98.65	98.69	98.81
Mg#	0.52	0.43	0.71	0.35	0.27	0.53	0.38	0.52	0.25	0.31	0.44	0.49	0.36

**APPENDIX TABLE 1: Major element composition of melt inclusions**

MAG B 47	64 (5 $\mu$ )	65 (15 $\mu$ )	66 (15 $\mu$ )	67 (5 $\mu$ )	67 (15 $\mu$ )	68 (10 $\mu$ )	71 (5 $\mu$ )	72 (10 $\mu$ )	73 (5 $\mu$ )
SiO <sub>2</sub>	41.95	55.46	52.22	42.17	35.17	51.61	37.58	41.82	43.62
TiO <sub>2</sub>	1.46	1.82	2.52	0.04	2.82	2.27	0.08	0.04	0.02
Al <sub>2</sub> O <sub>3</sub>	3.04	22.46	20.64	0.80	7.41	18.17	7.64	0.79	0.00
Cr <sub>2</sub> O <sub>3</sub>	0.00	0.00	0.08	0.01	0.12	0.02	0.03	0.00	0.03
FeO	10.06	3.18	3.85	9.93	7.98	3.13	16.02	7.33	12.45
MnO	0.09	0.32	0.06	0.26	0.19	0.00	0.21	0.19	0.16
MgO	28.06	2.56	3.78	28.49	17.97	6.22	20.91	33.85	27.44
CaO	0.11	4.72	7.38	1.42	12.88	5.53	0.42	0.22	0.19
Na <sub>2</sub> O	0.16	4.27	4.50	0.00	0.38	4.70	0.06	0.00	0.03
K <sub>2</sub> O	1.40	2.22	2.07	0.00	0.32	5.67	0.15	0.00	0.02
P <sub>2</sub> O <sub>5</sub>	0.07	0.26	0.65	0.88	0.09	0.31	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	86.40	97.26	97.75	84.00	85.33	97.62	83.08	84.25	83.96
Mg#	0.86	0.64	0.68	0.86	0.83	0.81	0.74	0.91	0.83

MAG B 25	56 (10 $\mu$ )	56 (10 $\mu$ )	56 (10 $\mu$ )	57 (10 $\mu$ )
SiO <sub>2</sub>	49.71	52.31	59.30	49.84
TiO <sub>2</sub>	0.43	0.07	0.12	0.26
Al <sub>2</sub> O <sub>3</sub>	8.86	20.57	21.03	7.11
Cr <sub>2</sub> O <sub>3</sub>	0.04	0.00	0.00	0.00
FeO	16.33	1.24	0.87	15.18
MnO	0.00	0.00	0.00	0.00
MgO	10.81	0.86	0.04	9.86
CaO	2.22	7.79	2.38	1.79
Na <sub>2</sub> O	0.13	0.37	7.98	0.98
K <sub>2</sub> O	0.63	0.26	2.47	0.63
P <sub>2</sub> O <sub>5</sub>	0.17	0.22	0.65	0.29
NiO	0.00	0.00	0.00	0.00
Total	89.32	83.68	94.83	85.93
Mg#	0.59	0.61	0.10	0.59

APPENDIX TABLE 2: Major element compositions of host Olivines

MAG B 47	OLIVINE 1	OLIVINE 2	OLIVINE 3	OLIVINE 4	OLIVINE 5	OLIVINE 6-7	OLIVINE 8	OLIVINE 9	OLIVINE 10	OLIVINE 11	OLIVINE 12-13	OLIVINE 12 NI	OLIVINE 14
SiO2	39.85	39.51	39.13	39.68	38.82	39.83	39.93	39.55	39.99	38.94	38.69	38.37	40.07
TiO2	0.01	0.01	0.00	0.00	0.01	0.00	0.00	0.00	0.03	0.00	0.00	0.01	0.00
Al2O3	0.02	0.03	0.03	0.00	0.02	0.00	0.00	0.00	0.05	0.03	0.00	0.08	0.02
Cr2O3	0.03	0.02	0.04	0.00	0.02	0.01	0.00	0.04	0.05	0.03	0.04	0.03	0.07
FeO	13.11	13.01	15.02	14.03	17.32	15.44	14.58	13.16	13.58	13.09	17.21	23.52	12.83
MnO	0.20	0.18	0.24	0.21	0.27	0.22	0.24	0.18	0.22	0.20	0.28	0.17	0.17
MgO	47.24	46.65	45.30	46.49	44.45	45.45	45.98	47.22	46.03	47.09	43.55	10.81	47.33
CaO	0.38	0.43	0.42	0.40	0.40	0.40	0.40	0.35	0.43	0.32	0.38	0.45	0.46
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.23	0.21	0.17	0.26	0.19	0.22	0.20	0.26	0.25	0.22	0.19	0.10	0.24
Total	101.07	100.04	100.35	101.05	101.49	101.57	101.32	100.76	100.62	99.90	100.33	73.53	101.18
Mg#	0.89	0.89	0.87	0.88	0.85	0.87	0.87	0.89	0.88	0.89	0.85	0.50	0.89

MAG B 25	OLIVINE NI 2	OLIVINE 2	OLIVINE NI 3	OLIVINE 3	OLIVINE NI 6	OLIVINE 6-7	OLIVINE NI13	OLIVINE 12-13	OLIVINE NI14	OLIVINE 14	OLIVINE NI15	OLIVINE 15	OLIVINE NI17
SiO2	38.55	38.21	38.30	38.58	38.75	38.41	38.39	38.56	38.73	38.92	38.04	38.61	39.02
TiO2	0.00	0.02	0.00	0.05	0.00	0.01	0.00	0.00	0.02	0.02	0.03	0.00	0.04
Al2O3	0.05	0.02	0.04	0.04	0.04	0.04	0.06	0.04	0.03	0.00	0.04	0.03	0.05
Cr2O3	0.00	0.00	0.00	0.04	0.01	0.04	0.02	0.09	0.05	0.04	0.06	0.03	0.05
FeO	17.88	18.19	17.76	18.42	17.92	18.15	17.18	17.43	16.50	16.46	18.43	18.67	16.51
MnO	0.25	0.26	0.30	0.26	0.27	0.09	0.08	0.09	0.04	0.07	0.14	0.09	0.06
MgO	43.20	43.10	43.08	43.01	42.84	42.88	43.07	43.04	43.63	43.96	42.80	42.22	44.72
CaO	0.38	0.44	0.39	0.41	0.40	0.40	0.42	0.38	0.43	0.44	0.43	0.36	0.48
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.15	0.16	0.17	0.17	0.18	0.03	0.04	0.02	0.04	0.05	0.03	0.04	0.05
Total	100.45	100.39	100.03	100.98	100.41	100.07	99.27	99.65	99.45	99.97	100.00	100.03	100.99
Mg#	0.84	0.84	0.84	0.84	0.84	0.84	0.85	0.85	0.85	0.86	0.84	0.83	0.86

TAA B 26	OLIVINE 1	OLIVINE 3	OLIVINE 4	OLIVINE 5-6	OLIVINE 7	OLIVINE 8	OLIVINE 9	OLIVINE 10	OLIVINE 12	OLIVINE 14	OLIVINE 17	OLIVINE 18	OLIVINE 20
SiO2	38.56	38.63	38.67	38.92	38.85	39.16	38.55	38.88	39.10	38.49	38.87	38.75	38.81
TiO2	0.05	0.04	0.06	0.04	0.00	0.00	0.03	0.00	0.03	0.02	0.03	0.00	0.03
Al2O3	0.05	0.04	0.04	0.05	0.05	0.06	0.06	0.04	0.07	0.06	0.05	0.05	0.04
Cr2O3	0.00	0.00	0.00	0.03	0.05	0.05	0.05	0.02	0.02	0.01	0.04	0.00	0.01
FeO	19.75	19.97	19.23	19.85	19.57	19.47	19.56	19.06	18.96	19.83	20.29	19.46	19.63
MnO	0.25	0.27	0.25	0.27	0.25	0.26	0.26	0.25	0.22	0.24	0.26	0.25	0.25
MgO	41.19	41.97	42.14	42.15	41.16	41.48	41.51	43.07	41.65	41.03	41.01	41.57	41.60
CaO	0.34	0.32	0.29	0.32	0.28	0.25	0.31	0.32	0.30	0.27	0.32	0.29	0.30
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.18	0.13	0.19	0.17	0.22	0.22	0.20	0.25	0.25	0.20	0.18	0.21	0.23
Total	100.37	101.38	100.86	101.79	100.43	100.94	100.53	101.87	100.58	100.16	101.05	100.58	100.90
Mg#	0.82	0.82	0.83	0.82	0.82	0.82	0.82	0.83	0.83	0.82	0.82	0.83	0.82

APPENDIX TABLE 2: Major element compositions of host Olivines

MAG B 47	OLIVINE 15	OLIVINE 16	OLIVINE 17	OLIVINE 18	OLIVINE 19-20	OLIVINE 21	OLIVINE 22	OLIVINE 23	OLIVINE 24	OLIVINE 25	OLIVINE 26-27	OLIVINE 28	OLIVINE 29-30
SiO2	39.25	39.95	39.36	39.45	39.43	38.69	39.24	38.92	38.55	39.46	39.48	40.10	39.95
TiO2	0.02	0.02	0.00	0.01	0.01	0.07	0.00	0.02	0.03	0.02	0.04	0.00	0.00
Al2O3	0.03	0.06	0.04	0.02	0.04	0.10	0.02	0.03	0.02	0.06	0.03	0.00	0.03
Cr2O3	0.04	0.07	0.02	0.00	0.06	0.06	0.03	0.05	0.04	0.04	0.04	0.02	0.05
FeO	15.36	13.13	13.15	17.93	14.50	22.60	15.16	18.09	17.82	14.90	15.12	14.62	13.32
MnO	0.23	0.17	0.21	0.27	0.22	0.38	0.25	0.25	0.30	0.23	0.27	0.22	0.21
MgO	45.86	47.01	48.58	42.90	45.63	38.25	45.29	43.62	44.11	46.66	46.45	45.57	47.15
CaO	0.47	0.35	0.38	0.39	0.41	0.48	0.37	0.39	0.38	0.31	0.38	0.43	0.42
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.23	0.21	0.27	0.20	0.21	0.18	0.24	0.23	0.21	0.19	0.24	0.22	0.24
Total	101.50	100.96	102.01	101.18	100.51	100.81	100.59	101.59	101.46	101.85	102.04	101.18	101.36
Mg#	0.87	0.89	0.89	0.84	0.87	0.79	0.87	0.84	0.85	0.87	0.87	0.87	0.89

MAG B 25	OLIVINE 17	OLIVINE NI 34	OLIVINE 34	OLIVINE NI37	OLIVINE 37	OLIVINE NI40	OLIVINE 40	OLIVINE NI43	OLIVINE 43	OLIVINE NI 47	OLIVINE 46-47	OLIVINE NI48	OLIVINE 48
SiO2	39.19	39.17	38.74	39.44	38.91	38.90	38.63	38.94	38.00	37.96	38.35	38.82	38.95
TiO2	0.00	0.02	0.03	0.03	0.02	0.05	0.02	0.05	0.00	0.03	0.03	0.00	0.00
Al2O3	0.03	0.04	0.02	0.03	0.02	0.03	0.00	0.03	1.01	0.03	0.05	0.04	0.03
Cr2O3	0.06	0.08	0.04	0.04	0.03	0.00	0.04	0.02	0.04	0.04	0.03	0.06	0.02
FeO	16.37	16.96	16.80	17.76	18.01	17.90	17.90	17.80	17.08	18.13	18.10	16.98	17.20
MnO	0.05	0.25	0.26	0.25	0.23	0.20	0.26	0.29	0.24	0.22	0.26	0.23	0.22
MgO	44.09	42.86	43.74	41.97	42.74	42.95	43.28	42.16	42.16	43.36	43.11	43.42	43.61
CaO	0.37	0.41	0.36	0.44	0.41	0.46	0.43	0.51	0.40	0.41	0.40	0.47	0.38
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.03	0.20	0.19	0.15	0.15	0.15	0.18	0.17	0.15	0.15	0.14	0.17	0.19
Total	100.18	99.99	100.18	100.10	100.52	100.64	100.74	99.95	99.07	100.32	100.47	100.19	100.60
Mg#	0.86	0.85	0.85	0.84	0.84	0.84	0.84	0.84	0.85	0.84	0.84	0.85	0.85

TAA B 26	OLIVINE 21	OLIVINE 22	OLIVINE 23	OLIVINE 24	OLIVINE 25	OLIVINE 28	OLIVINE 30	OLIVINE 32	OLIVINE 35	OLIVINE 37	OLIVINE 40	OLIVINE NI 41	OLIVINE 41
SiO2	38.46	38.64	38.83	38.59	38.56	38.43	38.44	38.60	38.44	38.83	38.61	38.26	38.70
TiO2	0.00	0.03	0.02	0.05	0.00	0.03	0.00	0.02	0.00	0.00	0.01	0.02	0.01
Al2O3	0.05	0.06	0.06	0.05	0.06	0.05	0.04	0.05	0.04	0.05	0.05	0.07	0.02
Cr2O3	0.07	0.00	0.02	0.02	0.00	0.04	0.04	0.03	0.02	0.01	0.00	0.04	0.00
FeO	19.45	19.46	19.65	19.43	19.83	19.71	19.66	19.46	18.94	19.60	19.60	19.83	19.48
MnO	0.28	0.25	0.25	0.26	0.26	0.26	0.29	0.25	0.22	0.23	0.25	0.25	0.30
MgO	41.49	41.94	41.49	41.74	41.18	41.51	41.95	41.63	42.79	41.75	41.40	41.45	40.94
CaO	0.30	0.28	0.35	0.32	0.29	0.31	0.29	0.29	0.28	0.30	0.29	0.30	0.30
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.23	0.21	0.18	0.19	0.18	0.18	0.17	0.21	0.23	0.18	0.20	0.19	0.19
Total	100.33	100.85	100.86	100.64	100.36	100.52	100.88	100.55	100.95	100.94	100.41	100.41	99.93
Mg#	0.83	0.83	0.82	0.83	0.82	0.82	0.83	0.83	0.83	0.82	0.82	0.82	0.82

**APPENDIX TABLE 2: Major element compositions of host Olivines**

MAG B 47	OLIVINE 31	OLIVINE 32	OLIVINE 33	OLIVINE 33	OLIVINE 34	OLIVINE 35-36	OLIVINE 37	OLIVINE 38	OLIVINE 39	OLIVINE 40	OLIVINE 41	OLIVINE 42-43	OLIVINE 44-45
SiO2	39.19	39.21	38.98	36.62	39.84	39.45	39.62	39.77	39.93	39.57	39.65	39.27	38.24
TiO2	0.00	0.00	0.01	0.06	0.00	0.03	0.00	0.03	0.00	0.00	0.00	0.02	0.05
Al2O3	0.04	0.03	0.03	0.18	0.00	0.03	0.02	0.04	0.00	0.02	0.00	0.03	0.03
Cr2O3	0.04	0.06	0.03	0.06	0.06	0.05	0.03	0.05	0.03	0.04	0.00	0.00	0.00
FeO	15.40	14.88	16.09	15.24	12.94	14.97	15.05	13.39	13.13	14.90	17.38	15.79	18.80
MnO	0.24	0.23	0.23	0.28	0.19	0.26	0.22	0.22	0.20	0.21	0.26	0.22	0.29
MgO	46.36	45.61	43.35	41.33	47.04	45.28	45.92	47.66	47.33	45.49	43.53	44.40	41.03
CaO	0.41	0.42	0.37	0.54	0.39	0.43	0.39	0.42	0.43	0.44	0.35	0.39	0.44
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.23	0.17	0.22	0.21	0.24	0.21	0.22	0.27	0.22	0.21	0.20	0.20	0.19
Total	101.91	100.62	99.30	94.52	100.71	100.71	101.47	101.84	101.25	100.87	101.36	100.33	99.07
Mg#	0.87	0.87	0.86	0.86	0.89	0.87	0.87	0.89	0.89	0.87	0.85	0.86	0.83

MAG B 25	OLIVINE NI49	OLIVINE 49	OLIVINE NI52	OLIVINE 52	OLIVINE NI53	OLIVINE 53	OLIVINE NI54
SiO2	38.69	38.00	38.64	38.59	39.68	38.94	37.94
TiO2	0.03	0.00	0.00	0.00	0.00	0.00	0.00
Al2O3	0.11	0.25	0.05	0.06	0.04	0.01	0.03
Cr2O3	0.03	0.02	0.06	0.04	0.05	0.03	0.04
FeO	16.68	16.74	17.93	17.78	18.10	18.02	18.16
MnO	0.21	0.26	0.24	0.25	0.25	0.33	0.31
MgO	43.37	43.80	43.19	42.67	40.76	42.69	44.01
CaO	0.49	0.40	0.45	0.38	0.40	0.41	0.45
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.17	0.13	0.13	0.17	0.18	0.26	0.18
Total	99.77	99.60	100.68	99.94	99.46	100.69	101.11
Mg#	0.85	0.85	0.84	0.84	0.83	0.84	0.84

TAA B 26	OLIVINE 43	OLIVINE 48	OLIVINE 49	OLIVINE 50	OLIVINE 51	OLIVINE 52	OLIVINE 53	OLIVINE 58	OLIVINE 59	OLIVINE 61	OLIVINE 62	OLIVINE 63	OLIVINE 66
SiO2	38.58	38.47	38.57	38.67	38.22	38.61	38.56	38.79	38.53	38.64	39.35	38.04	38.04
TiO2	0.02	0.02	0.02	0.00	0.03	0.00	0.02	0.01	0.02	0.01	0.00	0.01	0.02
Al2O3	0.06	0.04	0.04	0.06	0.05	0.07	0.02	0.08	0.06	0.04	0.06	0.05	0.06
Cr2O3	0.04	0.05	0.02	0.04	0.03	0.02	0.04	0.00	0.00	0.00	0.00	0.05	0.01
FeO	19.47	19.85	19.71	19.81	18.69	18.78	19.69	19.69	19.75	19.99	19.96	19.06	19.33
MnO	0.26	0.22	0.30	0.29	0.24	0.25	0.26	0.32	0.26	0.27	0.27	0.23	0.27
MgO	41.62	40.76	41.89	41.88	41.76	42.12	41.87	41.52	41.78	41.69	41.27	42.20	41.96
CaO	0.33	0.26	0.31	0.31	0.31	0.29	0.28	0.31	0.28	0.27	0.28	0.26	0.26
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.20	0.20	0.16	0.19	0.21	0.20	0.19	0.19	0.21	0.23	0.21	0.23	0.22
Total	100.56	99.87	101.00	101.25	99.55	100.35	100.93	100.91	100.88	101.13	101.42	100.13	100.17
Mg#	0.83	0.82	0.82	0.82	0.83	0.83	0.82	0.82	0.82	0.82	0.82	0.83	0.83

**APPENDIX TABLE 2: Major element compositions of host Olivines**

MAG B 47	OLIVINE 45 NI45A	OLIVINE 46-47	OLIVINE 48	OLIVINE 49	OLIVINE 50	OLIVINE 51	OLIVINE 51	OLIVINE 51	OLIVINE 51	OLIVINE 51	OLIVINE 51	OLIVINE 51	OLIVINE 51
SiO2	34.36	38.88	37.84	39.38	39.97	39.16	35.73	39.83	40.13	39.18	39.11	39.44	
TiO2	0.10	0.01	0.02	0.00	0.00	0.00	0.08	0.01	0.00	0.00	0.00	0.01	
Al2O3	0.03	0.02	0.04	0.03	0.00	0.04	0.16	0.02	0.00	0.00	0.01	0.00	
Cr2O3	0.03	0.03	0.01	0.04	0.05	0.02	0.16	0.04	0.04	0.07	0.04	0.03	
FeO	39.06	17.26	18.91	13.15	13.33	17.86	34.30	14.69	14.23	12.94	13.08	14.37	
MnO	0.71	0.29	0.29	0.22	0.22	0.27	0.65	0.20	0.21	0.17	0.19	0.21	
MgO	23.05	43.53	41.68	46.95	47.07	43.43	27.55	44.94	46.13	47.52	43.88	44.88	
CaO	0.81	0.41	0.42	0.39	0.34	0.57	0.71	0.44	0.36	0.41	0.40	0.39	
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	
NiO	0.05	0.20	0.17	0.17	0.25	0.17	0.12	0.23	0.23	0.20	0.24	0.24	
Total	98.20	100.62	99.37	100.33	101.22	101.52	99.45	100.40	101.32	100.49	96.96	99.59	
Mg#	0.57	0.85	0.83	0.89	0.89	0.84	0.64	0.87	0.88	0.89	0.88	0.87	

MAG B 25
SiO2
TiO2
Al2O3
Cr2O3
FeO
MnO
MgO
CaO
Na2O
K2O
P2O5
NiO
Total
Mg#

TAA B 26	OLIVINE 67	OLIVINE 68
SiO2	38.25	38.79
TiO2	0.00	0.00
Al2O3	0.06	0.04
Cr2O3	0.00	0.01
FeO	19.61	19.56
MnO	0.25	0.26
MgO	41.84	41.55
CaO	0.32	0.29
Na2O	0.00	0.00
K2O	0.00	0.00
P2O5	0.00	0.00
NiO	0.20	0.20
Total	100.52	100.70
Mg#	0.83	0.82



APPENDIX TABLE 2: Major element compositions of host Olivines

MAG B 47	OLIVINE 58	OLIVINE 59	OLIVINE 60	OLIVINE 61	OLIVINE 61A	OLIVINE NI 62	OLIVINE 62	OLIVINE NI 63	OLIVINE 63	OLIVINE 64	OLIVINE NI 64	OLIVINE NI 66	OLIVINE 65-66
SiO2	39.39	38.32	39.13	38.36	38.66	38.10	39.06	39.12	39.40	39.62	40.36	45.88	38.83
TiO2	0.01	0.01	0.00	0.01	0.01	0.02	0.01	0.00	0.00	0.02	0.00	1.17	0.03
Al2O3	0.04	0.18	0.02	0.05	0.06	0.02	0.03	0.02	0.02	0.03	0.04	13.48	0.01
Cr2O3	0.06	0.01	0.07	0.02	0.02	0.02	0.05	0.03	0.03	0.03	0.00	0.02	0.02
FeO	13.10	17.62	13.17	18.11	18.14	15.31	14.67	15.21	14.97	13.12	15.22	8.22	17.69
MnO	0.21	0.28	0.19	0.29	0.30	0.22	0.24	0.26	0.24	0.21	0.22	0.13	0.29
MgO	45.47	41.98	45.46	42.15	42.18	46.19	45.23	45.26	45.34	45.92	44.81	7.34	42.55
CaO	0.41	0.40	0.40	0.39	0.42	0.55	0.43	0.42	0.39	0.39	0.45	6.54	0.41
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.25	0.19	0.24	0.19	0.15	0.18	0.21	0.21	0.21	0.22	0.20	0.00	0.21
Total	98.91	98.99	98.67	99.57	99.93	100.59	99.92	100.53	100.60	99.57	101.30	82.79	100.05
Mg#	0.88	0.84	0.88	0.84	0.84	0.87	0.87	0.87	0.87	0.89	0.87	0.66	0.84

MAG B 47	OLIVINE NI 65	OLIVINE 67	OLIVINE NI 67	OLIVINE 68	OLIVINE 69	OLIVINE 70	OLIVINE 71	OLIVINE 72	OLIVINE I73
SiO2	38.90	39.10	38.21	39.45	39.11	39.04	39.31	38.91	38.43
TiO2	0.05	0.00	0.03	0.00	0.00	0.01	0.01	0.00	0.02
Al2O3	0.05	0.02	0.00	0.02	0.04	0.02	0.01	0.02	0.03
Cr2O3	0.03	0.03	0.04	0.05	0.03	0.03	0.06	0.05	0.03
FeO	17.83	15.91	17.43	13.32	13.31	15.00	13.00	15.10	17.79
MnO	0.28	0.22	0.26	0.21	0.17	0.23	0.21	0.21	0.28
MgO	42.02	43.85	43.39	45.49	44.59	44.06	45.78	44.26	41.47
CaO	0.41	0.41	0.50	0.40	0.42	0.40	0.36	0.39	0.39
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.20	0.19	0.14	0.24	0.26	0.21	0.24	0.22	0.20
Total	99.78	99.73	99.99	99.18	97.92	99.02	98.98	99.16	98.63
Mg#	0.84	0.86	0.85	0.88	0.88	0.87	0.89	0.87	0.84

**APPENDIX TABLE 3: Major element compositions of host and daughter Clinopyroxenes compositions**

MAG B 47	PYROXEN1	PYROXEN 6	PYROXEN 7	PYROXEN 8	PYROXEN 8	PYROXEN 9	PYROXEN 10	PYROXEN 11	PYROXEN 14	PYROXEN 15
SiO2	42.81	44.57	44.03	42.03	42.15	42.51	44.46	44.71	44.22	44.87
TiO2	5.87	4.54	4.64	5.24	5.56	4.59	4.08	3.72	4.47	2.84
Al2O3	9.80	12.17	10.65	10.65	12.28	10.40	10.04	9.95	9.54	8.76
Cr2O3	0.16	0.07	0.02	0.08	0.07	0.07	0.08	0.07	0.12	0.66
FeO	6.29	7.07	8.29	8.96	7.95	8.41	7.99	8.22	6.09	7.67
MnO	0.11	0.13	0.14	0.19	0.14	0.18	0.15	0.15	0.08	0.11
MgO	10.50	9.54	9.46	10.48	9.02	9.09	9.83	10.41	11.44	10.93
CaO	22.90	21.89	23.02	20.61	22.72	23.34	22.97	22.39	23.08	22.91
Na2O	0.65	0.80	0.69	0.54	0.68	0.89	0.63	0.53	0.72	0.59
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.02	0.04
Total	99.07	100.77	100.93	98.77	100.57	99.48	100.22	100.16	99.78	99.39
Mg#	0.79	0.75	0.72	0.71	0.71	0.71	0.73	0.73	0.81	0.76

MAG B 25	PYROXEN H1	PYROXEN H 4	PYROXEN H 5	PYROXEN H8-9	PYROXEN H10-11	PYROXEN H 18	PYROXEN H19-23	PYROXEN H24	PYROXEN H 25	PYROXEN H26
SiO2	50.56	49.62	50.12	50.84	50.06	51.29	51.29	50.90	50.92	50.53
TiO2	1.01	1.36	1.12	0.93	1.00	0.84	0.72	0.91	1.12	1.07
Al2O3	3.35	4.05	4.09	3.26	4.01	3.07	2.63	3.16	3.49	3.71
Cr2O3	0.66	0.64	0.97	0.78	0.91	0.64	0.90	0.60	0.48	0.81
FeO	5.16	5.37	4.95	4.91	4.99	4.89	4.31	4.83	5.29	4.74
MnO	0.12	0.09	0.00	0.00	0.00	0.00	0.00	0.00	0.17	0.00
MgO	15.65	15.15	15.03	15.97	15.08	15.37	16.36	15.59	15.43	15.17
CaO	22.57	22.71	22.40	21.86	22.24	22.36	22.15	22.13	22.34	22.29
Na2O	0.33	0.45	0.47	0.42	0.47	0.38	0.41	0.38	0.38	0.36
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.01	0.06	0.00	0.00	0.00	0.00	0.00	0.03	0.19	0.01
Total	99.43	99.50	99.15	98.96	98.75	98.84	98.78	98.52	99.80	98.69
Mg#	0.87	0.86	0.87	0.88	0.87	0.87	0.89	0.88	0.87	0.88

TAA B 26	PYROXEN 1a	PYROXEN 1b	PYROXEN 2	PYROXEN 3	PYROXEN core 4	PYROXEN margin	PYROXEN 5	PYROXEN 6a	PYROXEN 6b	PYROXEN 8
SiO2	40.66	39.85	41.59	39.17	42.38	44.18	39.40	40.75	37.22	47.62
TiO2	5.91	5.98	6.90	6.05	3.73	3.32	7.54	2.36	7.03	2.44
Al2O3	13.34	13.00	12.24	13.75	14.57	18.02	13.47	13.83	17.28	12.47
Cr2O3	0.02	0.01	0.05	0.00	0.00	0.03	0.03	0.00	0.00	0.00
FeO	8.00	7.97	6.66	9.55	7.52	6.50	8.81	7.21	8.84	9.66
MnO	0.02	0.13	0.08	0.05	0.05	0.06	0.12	0.07	0.09	0.16
MgO	8.93	8.99	10.79	7.62	8.68	2.95	9.16	10.18	6.74	6.19
CaO	21.03	22.18	22.42	22.76	23.39	16.03	21.78	23.54	21.95	19.78
Na2O	0.50	0.44	0.61	0.53	0.49	2.77	0.65	0.42	1.61	1.12
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	98.42	98.54	101.33	99.48	100.81	93.86	100.95	98.36	100.74	99.44
Mg#	0.71	0.71	0.78	0.64	0.72	0.50	0.70	0.76	0.63	0.59

**APPENDIX TABLE 3: Major element compositions of host and daughter Clinopyroxenes compositions**

MAG B 47	PYROXEN 16	PYROXEN 17	PYROXEN19	PYROXEN 20	PYROXEN 21	PYROXEN 21	PYROXEN 22	PYROXEN 23	PYROXEN 24	PYROXEN 27
SiO2	43.95	45.07	44.80	40.52	52.19	47.83	42.98	41.21	47.18	45.42
TiO2	4.87	4.01	6.07	7.47	0.79	2.72	5.26	5.84	2.58	3.05
Al2O3	10.60	8.24	10.22	12.48	3.03	6.49	12.98	14.70	11.61	9.07
Cr2O3	0.10	0.05	0.33	0.12	1.20	0.33	0.06	0.06	0.01	0.28
FeO	7.69	8.49	6.32	6.63	4.00	7.34	6.44	7.63	9.23	7.00
MnO	0.09	0.13	0.08	0.07	0.04	0.11	0.08	0.11	0.15	0.11
MgO	8.92	9.97	10.78	10.54	15.83	12.65	11.10	8.05	8.02	11.99
CaO	23.37	23.53	22.69	19.63	23.49	23.01	20.35	22.81	21.92	22.39
Na2O	0.62	0.74	0.69	0.46	0.32	0.44	1.09	0.61	0.81	0.72
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.02	0.00	0.06	0.02	0.01	0.02	0.02	0.02
Total	100.20	100.23	102.00	97.91	100.95	100.94	100.35	101.03	101.54	100.05
Mg#	0.72	0.72	0.79	0.78	0.90	0.79	0.79	0.70	0.66	0.79

MAG B 25	PYROXEN H27-28	PYROXEN H 29	PYROXEN H30-32	PYROXEN H35	PYROXEN H36	PYROXEN H38-39	PYROXEN H 41-42	PYROXEN H 44	PYROXEN H 45	PYROXEN H N150
SiO2	50.52	50.65	50.99	50.86	48.66	50.94	50.43	50.36	50.80	49.64
TiO2	1.06	1.01	0.88	0.93	0.77	0.94	1.12	1.32	0.94	1.29
Al2O3	3.65	3.65	3.32	3.41	2.77	3.44	3.79	4.09	3.17	4.08
Cr2O3	0.68	0.65	0.64	0.60	0.63	0.77	0.51	0.26	0.37	0.63
FeO	5.12	5.17	4.86	4.69	4.27	4.80	4.84	5.50	4.86	5.67
MnO	0.37	0.22	0.23	0.08	0.00	0.00	0.00	0.00	0.00	0.05
MgO	15.44	15.59	15.53	15.65	16.84	15.41	15.25	14.79	15.51	14.27
CaO	22.52	22.08	21.80	22.20	22.30	22.01	22.19	22.36	22.37	22.83
Na2O	0.41	0.38	0.39	0.40	0.39	0.45	0.36	0.39	0.34	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.25	0.19	0.29	0.00	0.04	0.00	0.00	0.00	0.00	0.00
Total	100.01	99.56	98.92	98.81	96.67	98.75	98.47	99.06	98.35	98.47
Mg#	0.87	0.87	0.88	0.88	0.90	0.88	0.87	0.86	0.88	0.85

TAA B 26	PYROXEN 9	PYROXEN 10a	PYROXEN10b	PYROXEN 12	PYROXEN 14	PYROXEN 15	PYROXEN 16a	PYROXEN 16b	PYROXEN 16c	PYROXEN 17a
SiO2	43.29	39.07	43.34	44.02	39.98	45.28	49.30	50.28	43.31	49.02
TiO2	4.77	7.83	5.34	3.96	6.97	3.93	1.69	1.55	5.11	1.97
Al2O3	10.96	13.40	11.51	12.63	15.11	9.52	4.98	4.81	11.85	5.30
Cr2O3	0.15	0.04	0.00	0.00	0.00	0.06	0.90	0.86	0.12	0.62
FeO	7.61	7.35	8.06	9.51	9.15	7.48	6.44	6.35	5.99	6.88
MnO	0.10	0.12	0.15	0.13	0.12	0.15	0.12	0.12	0.15	0.13
MgO	11.60	9.61	11.52	6.68	10.47	12.44	14.93	15.03	11.58	14.77
CaO	21.79	22.29	20.61	21.93	17.33	21.74	21.90	21.89	21.93	21.86
Na2O	0.59	0.72	0.74	1.25	1.47	0.57	0.42	0.44	0.59	0.42
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	100.85	100.42	101.27	100.11	100.62	101.18	100.67	101.32	100.63	100.97
Mg#	0.77	0.74	0.76	0.61	0.72	0.79	0.84	0.84	0.81	0.83

**APPENDIX TABLE 3: Major element compositions of host and daughter Clinopyroxenes compositions**

MAG B 47	PYROXEN 29	PYROXEN 30	PYROXEN 31	PYROXEN 32	PYROXEN 33	PYROXEN 33A	PYROXEN 33B	PYROXEN 33B	PYROXEN 36	PYROXEN 37	PYROXEN 38
SiO2	44.50	45.59	41.77	43.61	43.02	46.22	43.26	48.20	40.93	41.54	42.31
TiO2	4.81	4.15	8.05	4.85	4.34	4.24	4.45	3.03	6.18	5.66	5.09
Al2O3	10.47	9.79	12.78	10.25	12.75	11.33	11.39	9.87	12.03	12.69	11.14
Cr2O3	0.12	0.10	0.08	0.18	0.01	0.11	0.06	0.07	0.03	0.08	0.04
FeO	5.79	6.36	7.35	6.94	7.49	7.33	7.68	6.76	7.05	8.01	7.95
MnO	0.07	0.10	0.09	0.09	0.11	0.14	0.15	0.14	0.09	0.10	0.09
MgO	10.79	11.19	7.71	11.12	9.11	10.22	9.93	12.66	9.89	11.65	8.64
CaO	23.33	23.10	23.02	22.84	22.52	21.04	22.24	21.52	22.78	18.97	23.08
Na2O	0.61	0.57	0.70	0.66	0.77	0.81	0.69	0.74	0.44	1.08	0.56
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.06	0.02	0.02	0.02	0.03	0.02	0.00	0.00	0.00	0.00	0.00
Total	100.54	100.99	101.56	100.56	100.15	101.45	99.84	102.99	99.41	99.80	98.91
Mg#	0.80	0.79	0.70	0.78	0.73	0.76	0.74	0.80	0.76	0.76	0.71

MAG B 25	PYROXEN H 50	PYROXEN H 51	PYROXEN H 54	PYROXEN H 56	PYROXEN 3	PYROXEN 6	PYROXEN 8	PYROXEN 12	PYROXEN 13	PYROXEN 14	PYROXEN 16
SiO2	49.55	50.91	51.01	50.96	40.17	40.81	51.87	41.71	40.63	40.03	40.10
TiO2	1.24	0.99	1.07	1.18	5.60	4.99	0.78	4.55	5.69	4.73	5.24
Al2O3	3.72	3.39	3.34	4.23	10.22	10.78	2.94	11.09	10.40	11.59	11.21
Cr2O3	0.37	0.57	0.56	0.86	0.07	0.04	0.98	0.03	0.02	0.06	0.00
FeO	5.64	4.49	5.37	4.87	11.74	11.00	4.41	10.36	10.86	10.91	11.38
MnO	0.05	0.00	0.12	0.00	0.51	0.37	0.00	0.20	0.00	0.00	0.33
MgO	14.79	15.43	15.98	14.90	8.39	9.26	16.55	9.61	9.30	6.02	7.37
CaO	22.77	22.47	23.09	22.41	21.50	22.00	22.58	21.98	21.70	22.77	22.34
Na2O	0.00	0.36	0.32	0.41	0.78	0.65	0.36	0.66	0.76	0.66	0.63
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.06	0.05
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.35	0.27	0.00	0.00	0.00	0.00	0.00
Total	98.13	98.61	100.85	99.82	99.33	100.18	100.47	100.19	99.36	96.83	98.63
Mg#	0.85	0.88	0.87	0.87	0.61	0.65	0.89	0.67	0.65	0.55	0.59

TAA B 26	PYROXEN 17b	PYROXEN 18	PYROXEN 20	PYROXEN 21	PYROXEN 22	PYROXEN 30	PYROXEN 32	PYROXEN 33	PYROXEN 34	PYROXEN 40	PYROXEN 48a
SiO2	43.71	43.87	48.85	42.26	41.86	45.54	47.22	44.36	44.59	39.36	45.72
TiO2	4.33	4.98	2.39	3.04	5.39	3.61	2.93	4.26	4.33	6.76	3.87
Al2O3	10.24	12.26	5.03	13.86	12.07	8.83	7.08	8.39	10.58	14.76	9.16
Cr2O3	0.03	0.08	0.10	0.00	0.07	0.02	0.20	0.05	0.21	0.00	0.04
FeO	7.76	7.44	7.50	8.18	7.03	7.00	7.44	8.96	6.17	8.14	7.75
MnO	0.10	0.18	0.12	0.11	0.14	0.12	0.13	0.15	0.12	0.09	0.13
MgO	11.99	11.28	14.82	8.58	11.06	12.59	13.63	11.79	12.29	8.12	12.40
CaO	22.01	20.04	22.06	23.40	21.83	21.80	21.76	21.75	21.92	22.73	21.67
Na2O	0.51	0.95	0.35	0.45	0.62	0.45	0.46	0.52	0.58	0.61	0.53
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	100.68	101.07	101.22	99.85	100.08	99.97	100.86	100.22	100.79	100.57	101.28
Mg#	0.77	0.77	0.81	0.70	0.78	0.80	0.80	0.74	0.81	0.69	0.78

**APPENDIX TABLE 3: Major element compositions of host and daughter Clinopyroxenes compositions**

MAG B 47	PYROXEN 39	PYROXEN 39B	PYROXEN 40	PYROXEN 41	PYROXEN 42	PYROXEN 44	PYROXEN 45	PYROXEN 45A	PYROXEN 45A	PYROXEN 46	PYROXEN 48
SiO2	43.35	43.44	42.12	46.42	41.05	42.22	43.89	49.69	47.70	40.01	41.37
TiO2	4.14	4.72	5.00	2.08	5.24	5.43	2.91	1.50	2.91	5.51	5.81
Al2O3	9.70	10.28	10.46	6.71	13.70	13.38	10.24	1.52	4.35	13.06	13.35
Cr2O3	0.04	0.06	0.04	0.02	0.04	0.00	0.00	0.00	0.00	0.01	0.01
FeO	7.17	7.42	6.81	12.46	6.71	9.25	8.84	14.92	13.80	8.86	10.13
MnO	0.09	0.11	0.10	0.20	0.11	0.16	0.15	0.32	0.25	0.11	0.13
MgO	10.05	9.89	10.23	8.78	7.92	7.51	9.15	8.92	8.38	9.19	8.42
CaO	23.18	23.62	23.09	22.39	22.43	21.34	21.96	21.18	21.45	21.62	19.90
Na2O	0.41	0.58	0.50	0.63	0.66	0.96	0.42	1.12	0.99	0.92	1.02
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	98.13	100.10	98.34	99.67	97.85	100.25	97.56	99.15	99.83	99.29	100.15
Mg#	0.76	0.75	0.77	0.61	0.72	0.64	0.68	0.57	0.57	0.69	0.64

MAG B 25	PYROXEN 17	PYROXEN 18	PYROXEN 30	PYROXEN 34	PYROXEN 37	PYROXEN 40	PYROXEN 40A	PYROXEN 43	PYROXEN 46	PYROXEN 49	PYROXEN 54
SiO2	41.37	45.62	42.86	41.49	47.91	43.96	41.16	42.68	41.83	42.90	45.17
TiO2	5.38	2.62	4.60	3.79	3.45	3.87	5.52	4.59	4.78	4.58	3.33
Al2O3	10.76	7.47	9.99	12.57	10.57	7.63	11.63	9.75	9.70	8.10	7.54
Cr2O3	0.07	0.41	0.02	0.00	0.02	0.00	0.04	0.09	0.04	0.08	0.04
FeO	9.80	7.54	9.96	10.57	9.79	11.22	10.59	10.25	11.01	10.55	10.24
MnO	0.36	0.07	0.15	0.12	0.15	0.19	0.14	0.15	0.19	0.17	0.20
MgO	9.34	11.98	9.84	8.05	8.35	9.62	7.65	9.15	8.34	9.42	10.88
CaO	21.99	22.30	21.50	23.44	20.15	21.82	21.07	21.44	21.55	21.65	21.77
Na2O	0.60	0.00	0.73	0.57	0.88	0.53	0.79	0.70	0.70	0.67	0.56
K2O	0.03	0.00	0.02	0.00	0.26	0.02	0.32	0.05	0.04	0.02	0.02
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.03	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	99.69	98.03	99.66	100.59	101.52	98.86	98.89	98.84	98.18	98.13	99.74
Mg#	0.68	0.77	0.69	0.63	0.65	0.65	0.61	0.66	0.62	0.66	0.70

TAA B 26	PYROXEN 48b	PYROXEN 51a	PYROXEN 51b	PYROXEN 59a	PYROXEN 59b	PYROXEN 61	PYROXEN 66	PYROXEN 67
SiO2	39.79	40.44	24.80	37.77	41.28	45.56	43.54	41.97
TiO2	7.32	6.37	11.87	8.99	4.76	3.70	4.25	6.01
Al2O3	16.42	12.53	18.78	17.31	14.88	9.25	9.51	10.74
Cr2O3	0.00	0.01	0.14	0.00	0.00	0.02	0.06	0.03
FeO	8.25	6.63	18.72	10.74	7.60	6.53	7.53	7.37
MnO	0.12	0.14	0.15	0.14	0.09	0.14	0.10	0.11
MgO	12.95	10.44	12.83	9.99	8.52	13.13	12.19	11.67
CaO	11.70	23.52	12.47	12.16	22.97	22.36	22.28	21.94
Na2O	2.63	0.53	0.88	2.30	0.67	0.45	0.46	0.54
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	99.16	100.60	100.63	99.38	100.77	101.14	99.91	100.37
Mg#	0.78	0.78	0.60	0.67	0.71	0.82	0.78	0.78

**APPENDIX TABLE 3: Major element compositions of host and daughter Clinopyroxenes compositions**

MAG B 47	PYROXEN 49	PYROXEN 50	PYROXEN 51	PYROXEN 52	PYROXEN 53	PYROXEN 54	PYROXEN 54	PYROXEN 55	PYROXEN 56	PYROXEN 57	PYROXEN 58
SiO2	41.41	44.39	44.37	43.25	42.92	45.37	40.65	42.29	44.38	40.20	44.83
TiO2	8.04	4.25	4.07	4.51	4.86	3.65	5.79	5.52	4.22	6.92	4.75
Al2O3	11.80	10.61	9.58	11.64	10.17	9.48	12.80	12.03	9.98	12.83	10.29
Cr2O3	0.01	0.04	0.00	0.04	0.06	0.06	0.05	0.10	0.09	0.06	0.10
FeO	8.20	6.81	6.93	6.44	9.97	5.99	5.95	6.93	8.70	8.93	6.84
MnO	0.10	0.10	0.12	0.06	0.15	0.13	0.08	0.11	0.13	0.34	0.54
MgO	8.96	10.41	10.96	10.26	8.02	12.45	10.69	9.72	8.77	10.29	9.63
CaO	22.63	23.00	22.68	23.36	22.85	22.55	22.51	22.97	23.16	17.01	22.67
Na2O	0.64	0.52	0.50	0.49	0.57	0.45	0.63	0.59	0.59	1.81	0.56
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.01	0.00	0.00	0.17	0.36
Total	101.77	100.13	99.20	100.02	99.56	100.12	99.17	100.25	100.00	98.55	100.57
Mg#	0.71	0.77	0.78	0.78	0.64	0.82	0.80	0.76	0.69	0.72	0.76

MAG B 25	PYROXEN 56	PYROXEN 57
SiO2	39.75	48.10
TiO2	6.17	1.91
Al2O3	11.61	5.45
Cr2O3	0.00	0.28
FeO	10.61	6.14
MnO	0.14	0.06
MgO	8.25	14.36
CaO	21.40	22.72
Na2O	0.69	0.40
K2O	0.06	0.00
P2O5	0.00	0.00
NiO	0.00	0.00
Total	98.67	99.43
Mg#	0.63	0.84

**APPENDIX TABLE 3: Major element compositions of host and daughter Clinopyroxenes compositions**

MAG B 47	PYROXEN 59	PYROXEN 60	PYROXEN 61	PYROXEN 62	PYROXEN 63	PYROXEN 64	PYROXEN 65	PYROXEN 66	PYROXEN 67	PYROXEN 69	PYROXEN 70
SiO2	41.72	43.65	42.78	43.43	44.78	42.96	39.86	41.62	42.52	43.47	45.34
TiO2	5.68	5.11	5.18	3.66	3.37	4.21	5.62	6.47	4.49	4.16	4.39
Al2O3	13.03	10.06	12.05	10.32	10.32	11.01	13.07	13.26	10.97	9.68	12.25
Cr2O3	0.06	0.09	0.02	0.06	0.11	0.03	0.09	0.06	0.11	0.06	0.14
FeO	9.33	7.38	7.64	7.35	7.28	6.40	7.16	8.19	7.28	8.23	5.16
MnO	0.59	0.61	0.12	0.33	0.34	0.53	0.18	0.14	0.09	0.15	0.08
MgO	7.14	10.08	9.91	10.47	10.94	11.47	9.41	7.70	10.27	9.06	10.09
CaO	22.74	23.16	22.36	23.05	22.51	22.83	22.55	22.09	22.37	22.20	19.19
Na2O	0.67	0.66	0.64	0.58	0.57	0.54	0.71	0.63	0.59	0.92	1.25
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.40	0.40	0.00	0.18	0.17	0.33	0.06	0.00	0.00	0.00	0.00
Total	101.37	101.21	100.71	99.42	100.38	100.32	98.70	100.15	98.68	97.91	97.90
Mg#	0.63	0.75	0.74	0.76	0.77	0.80	0.74	0.68	0.76	0.71	0.81

MAG B 47	PYROXEN 71	PYROXEN 72
SiO2	49.32	42.00
TiO2	1.91	5.26
Al2O3	5.10	11.74
Cr2O3	0.31	0.07
FeO	5.38	7.69
MnO	0.05	0.14
MgO	14.53	10.84
CaO	22.56	19.34
Na2O	0.46	1.50
K2O	0.00	0.00
P2O5	0.00	0.00
NiO	0.05	0.00
Total	99.66	98.58
Mg#	0.86	0.76

APPENDIX TABLE 4: Major element compositions of daughter mineral

MAG B 47	AMPH 13	AMPH 18	AMPH 23	AMPH 27	AMPH 34	AMPH 44	AMPH 54	AMPH 61	AMPH 65	AMPH 66	AMPH 68	AMPH 70
SiO2	38.18	39.53	37.76	36.92	40.10	37.71	37.75	38.93	36.95	38.37	40.84	40.93
TiO2	9.80	8.16	8.72	8.00	5.04	8.28	6.24	8.37	8.57	8.72	7.71	5.74
Al2O3	14.71	14.48	15.76	12.14	14.13	14.57	16.40	14.02	14.38	15.39	13.50	10.47
Cr2O3	0.10	0.03	0.02	0.13	0.02	0.00	0.00	0.05	0.04	0.01	0.06	0.07
FeO	8.84	10.01	9.99	7.45	8.12	11.21	8.09	9.35	10.06	9.89	6.54	6.62
MnO	0.14	0.11	0.12	0.05	0.11	0.17	0.09	0.12	0.16	0.14	0.07	0.10
MgO	11.77	10.81	10.93	13.49	13.06	10.10	12.80	10.53	11.64	10.41	12.45	10.66
CaO	11.96	12.82	12.31	11.38	11.85	11.79	12.04	14.03	12.18	12.38	12.89	22.64
Na2O	2.81	2.80	2.68	3.04	2.69	2.49	3.12	2.07	3.06	2.68	2.27	0.68
K2O	0.07	0.10	0.31	0.40	0.47	0.00	0.00	0.05	0.00	0.24	0.54	0.06
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.01	0.02	0.00	0.00	0.00	0.00	0.00	0.00	0.01	0.00	0.00
Total	98.38	98.85	98.62	92.99	95.59	96.32	96.52	97.52	97.04	98.24	96.86	97.96

MAG B 25	AMPH 1	AMPH 2	AMPH 3	AMPH 4	AMPH 5	AMPH 7	AMPH 8	AMPH 9	AMPH 10	AMPH 11A	AMPH 12	AMPH 15
SiO2	36.67	38.60	42.39	37.43	38.09	38.44	38.00	38.36	37.29	36.48	38.68	37.43
TiO2	2.58	5.45	1.66	5.08	5.06	5.44	5.09	4.50	3.44	3.76	5.20	6.05
Al2O3	16.66	13.76	14.39	17.19	15.41	14.44	16.12	14.45	16.21	17.37	14.70	13.93
Cr2O3	0.00	0.00	0.00	0.00	0.06	0.00	0.03	0.04	0.06	0.09	0.02	0.04
FeO	18.31	12.63	13.09	14.75	12.93	10.84	12.84	12.66	15.12	13.37	11.52	12.58
MnO	0.23	0.06	0.20	0.27	0.12	0.11	0.00	0.00	0.00	0.00	0.00	0.19
MgO	6.35	10.45	12.76	8.84	10.42	13.11	10.12	10.82	8.79	9.82	11.64	11.49
CaO	11.53	11.88	11.15	11.73	11.51	11.67	11.74	11.52	11.41	11.40	11.80	11.70
Na2O	2.99	2.87	3.12	2.71	3.01	3.00	2.88	3.00	3.03	3.17	2.85	2.88
K2O	0.42	0.56	0.62	0.63	0.39	0.59	0.39	0.42	0.45	0.43	0.60	0.64
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	95.74	96.26	99.37	98.64	96.98	97.63	97.21	95.77	95.79	95.90	97.00	96.92

TAA B 26	ILM 3	ILM 4	ILM 8	ILM 12	ILM 17	ILM 22	ILM 30	ILM 49	ILM 61	SPINEL 2a	SPINEL 2b	SPINEL 19
SiO2	0.00	0.00	0.00	1.11	0.07	0.06	0.00	0.00	0.00	6.06	3.65	0.00
TiO2	51.87	53.82	53.57	53.58	52.67	51.97	53.89	53.55	53.93	17.70	18.86	9.89
Al2O3	0.26	0.18	0.07	0.80	0.31	0.19	0.14	0.16	0.14	10.67	10.75	11.96
Cr2O3	0.04	0.08	0.04	0.13	0.00	0.00	0.00	0.02	0.00	11.16	12.01	12.89
FeO	42.93	42.61	36.94	37.11	40.39	42.52	41.04	38.71	40.34	43.71	46.66	55.93
MnO	0.52	0.53	0.48	0.46	0.46	0.56	0.58	0.59	0.47	0.29	0.46	0.43
MgO	4.62	4.67	7.76	8.03	6.03	4.86	5.94	7.31	6.27	7.92	8.06	7.58
CaO	0.32	0.03	0.20	0.55	0.25	0.09	0.35	0.31	0.31	3.29	2.06	0.08
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.01	0.03	0.00	0.00	0.02	0.00	0.00	0.00	0.03	0.15	0.20	0.44
Total	100.56	101.95	99.06	101.76	100.20	100.24	101.94	100.64	101.48	100.94	102.69	99.20



APPENDIX TABLE 4: Major element compositions of daughter mineral

MAG B 47	Ti-MNT 21	ILM 45	ILM 64	SPINEL 4	SPINEL 6	SPINEL 8	SPINEL 9	SPINEL 11	SPINEL 11 in OI	SPINEL 13	SPINEL 15	SPINEL c 19
SiO2	0.13	0.48	0.08	0.09	0.00	0.00	0.05	0.00	0.00	0.00	0.00	0.14
TiO2	29.83	54.02	52.90	3.33	3.80	4.57	3.12	3.37	2.27	7.14	2.93	7.54
Al2O3	2.40	0.27	0.02	18.96	15.99	18.35	18.16	15.48	14.71	16.27	14.85	17.97
Cr2O3	0.07	0.34	0.00	40.33	40.57	36.67	42.65	42.78	44.78	33.00	40.14	33.81
FeO	62.74	37.81	40.27	26.16	31.66	32.57	26.62	30.84	27.72	33.09	32.27	33.62
MnO	0.88	0.73	0.72	0.34	0.46	0.54	0.36	0.58	0.37	0.52	0.47	0.44
MgO	1.15	5.53	5.05	9.72	8.01	7.54	9.44	7.37	12.24	8.39	8.24	7.09
CaO	0.09	0.58	0.07	0.11	0.25	0.29	0.21	0.24	0.08	0.20	0.18	0.09
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.06	0.00	0.05	0.07	0.27	0.28	0.04	0.30	0.38	0.28	0.29	0.07
Total	97.36	99.76	99.14	99.12	101.01	100.80	100.65	100.96	102.55	98.89	99.36	100.77

MAG B 25	AMPH 24	AMPH 25	AMPH 27	AMPH 28	AMPH 30	AMPH 31	AMPH 34	AMPH 35	AMPH 37	AMPH 40	AMPH 40A	AMPH 41
SiO2	37.60	36.80	37.58	36.97	39.45	37.66	37.70	37.91	39.98	39.53	42.18	36.54
TiO2	4.67	4.29	4.82	4.16	5.88	2.87	4.36	5.88	4.66	5.00	4.50	4.98
Al2O3	16.36	16.01	14.73	16.37	13.74	16.46	14.79	15.23	13.78	13.35	11.01	16.38
Cr2O3	0.01	0.01	0.03	0.00	0.05	0.02	0.06	0.05	0.07	0.00	0.09	0.00
FeO	13.29	15.10	13.03	14.69	10.87	14.62	15.12	10.38	13.96	12.29	10.20	14.79
MnO	0.23	0.31	0.27	0.24	0.15	0.28	0.20	0.18	0.18	0.21	0.15	0.27
MgO	10.12	9.00	10.74	8.70	12.19	9.20	9.17	12.33	10.43	11.85	10.44	8.80
CaO	11.70	11.38	11.82	11.20	12.04	11.59	12.18	11.39	11.58	11.54	18.38	11.46
Na2O	2.88	3.06	3.07	3.02	2.88	3.02	3.06	3.05	2.92	3.01	1.18	2.95
K2O	0.49	0.45	0.42	0.47	0.45	0.51	0.49	0.47	0.53	0.50	0.20	0.51
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	97.35	96.41	96.51	95.81	97.69	96.23	97.11	96.84	98.09	97.27	98.34	96.68

TAA B 26	SPINEL 22	SPINEL 24a	SPINEL 24 b	SPINEL 38	PLAG 40a	PLAG2 40b	PLAG 43	PLAG 49	PLAG 50	PLAG 55a	PLAG 55b
SiO2	0.00	0.00	0.00	0.00	49.20	46.56	45.66	47.22	46.84	46.42	47.46
TiO2	17.14	23.07	21.98	18.72	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Al2O3	10.14	9.15	9.95	9.65	27.83	30.31	32.71	34.66	33.06	33.24	33.55
Cr2O3	11.03	7.70	7.69	13.98	0.00	0.00	0.00	0.00	0.00	0.00	0.00
FeO	54.16	54.11	55.62	50.89	1.01	1.01	1.25	0.89	0.73	1.34	0.76
MnO	0.59	0.56	0.56	0.59	0.00	0.00	0.00	0.00	0.00	0.00	0.00
MgO	6.00	6.35	5.44	6.33	0.46	0.49	0.22	0.15	0.05	0.19	0.10
CaO	0.17	0.16	0.11	0.19	13.02	16.34	17.77	17.94	17.56	17.33	17.61
Na2O	0.00	0.00	0.00	0.00	2.28	2.03	1.17	1.20	1.43	1.49	1.58
K2O	0.00	0.00	0.00	0.00	2.15	0.20	0.18	0.09	0.17	0.11	0.15
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.43	0.40	0.42	0.48	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	99.66	101.50	101.77	100.84	95.94	96.95	98.95	102.15	99.83	100.13	101.22

**APPENDIX TABLE 4: Major element compositions of daughter mineral**

MAG B 47	SPINEL m 15	SPINEL 20	SPINEL 32	SPINEL 33A	SPINEL 33	SPINEL in OI 33	SPINEL 34	SPINEL in OI 134	SPINEL 41	SPINEL 43	SPINEL 45
SiO2	0.36	0.06	0.07	0.06	0.12	0.00	0.00	0.00	0.00	0.00	0.02
TiO2	19.19	18.56	3.28	3.74	3.64	2.56	3.11	2.07	2.47	3.32	7.55
Al2O3	8.27	12.67	16.22	17.93	16.22	15.70	16.48	18.94	15.35	16.94	15.10
Cr2O3	18.31	16.31	37.98	38.53	38.64	37.67	40.57	40.28	37.07	35.58	13.73
FeO	41.73	41.97	33.87	31.14	32.65	32.19	28.36	26.14	33.82	33.81	50.95
MnO	0.45	0.40	0.32	0.44	0.42	0.37	0.45	0.38	0.40	0.41	0.44
MgO	9.41	10.61	9.67	7.65	8.48	10.17	9.17	11.57	10.49	10.16	10.64
CaO	0.06	0.07	0.05	0.15	0.14	0.02	0.21	0.03	0.03	0.06	0.03
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.06	0.12	0.14	0.31	0.21	0.25	0.21	0.32	0.27	0.26	0.30
Total	97.83	100.78	101.58	99.94	100.52	98.93	98.55	99.72	99.90	100.55	98.76

MAG B 25	AMPH 42	AMPH 43	AMPH 44	AMPH 45	AMPH 47	AMPH 48	AMPH 49	AMPH 50	AMPH 51	AMPH 53	AMPH 54
SiO2	36.68	39.23	35.32	36.06	41.30	39.04	39.13	36.62	37.86	41.11	37.79
TiO2	4.05	4.88	4.88	4.51	2.39	6.42	5.79	3.20	5.17	3.98	4.99
Al2O3	16.93	13.78	16.82	17.00	10.78	11.92	13.50	18.55	16.53	13.24	14.38
Cr2O3	0.07	0.00	0.03	0.02	0.00	0.13	0.00	0.00	0.05	0.00	0.01
FeO	14.44	12.57	14.66	14.04	11.36	12.04	12.70	17.57	12.99	11.59	14.41
MnO	0.24	0.13	0.20	0.22	0.17	0.16	0.14	0.34	0.20	0.15	0.20
MgO	8.65	11.55	8.97	9.12	10.77	8.19	10.44	6.62	9.48	12.43	10.17
CaO	11.76	11.27	11.67	11.31	16.77	18.10	11.60	11.11	11.67	11.54	11.63
Na2O	3.01	3.03	3.05	2.97	2.15	1.52	2.87	2.86	2.69	2.89	3.11
K2O	0.46	0.56	0.50	0.44	0.26	0.23	0.63	0.48	0.50	0.60	0.45
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	96.27	96.99	96.10	95.69	95.94	97.74	96.80	97.34	97.13	97.53	97.13

**APPENDIX TABLE 4: Major element compositions of daughter mineral**

MAG B 47	SPINEL in Ol 46	SPINEL 51A c	SPINEL 51A m	SPINEL 52	SPINEL I53	SPINEL 54	SPINEL63	SPINEL 67	PLAG 45A	PLAG 51	BIOTITE 49
SiO2	0.07	0.00	0.11	0.06	0.00	0.05	0.00	0.04	66.73	65.66	34.18
TiO2	2.46	2.15	17.99	0.42	3.35	2.57	2.46	3.56	0.00	0.00	11.00
Al2O3	15.19	15.36	5.97	20.54	15.83	14.12	15.19	15.11	19.27	20.72	13.83
Cr2O3	36.13	41.65	17.72	45.38	41.81	42.54	40.45	37.57	0.00	0.00	0.00
FeO	34.96	30.85	54.75	22.75	30.10	27.89	30.43	35.28	0.73	0.72	9.80
MnO	0.25	0.34	0.75	0.30	0.55	0.31	0.98	0.34	0.00	0.00	0.07
MgO	10.43	10.83	2.95	11.44	8.45	11.15	9.98	8.03	0.00	0.00	13.77
CaO	0.00	0.02	0.04	0.02	0.29	0.14	0.13	0.04	0.11	0.23	0.40
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	4.85	4.71	0.47
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	8.69	7.67	7.81
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.10	0.27	0.19	0.08	0.26	0.15	0.59	0.13	0.00	0.00	0.00
Total	99.60	101.46	100.46	100.99	100.63	98.91	100.20	100.10	100.37	99.70	91.31

MAG B 25	Ti-MNT 1	Ti-MNT 4	Ti-MNT 5	Ti-MNT 6	Ti-MNT 8	Ti-MNT 9	Ti-MNT 10	Ti-MNT 13	Ti-MNT 15	Ti-MNT 18	Ti-MNT 24
SiO2	0.19	0.09	0.12	0.41	0.12	0.16	0.11	0.14	4.77	0.12	0.12
TiO2	15.23	16.33	17.15	20.49	15.51	17.73	14.91	16.42	14.06	16.88	17.13
Al2O3	8.81	8.52	7.97	6.08	8.30	7.91	8.45	6.75	5.66	8.14	8.02
Cr2O3	0.03	0.05	0.09	0.07	0.03	0.02	0.05	0.05	0.07	0.04	0.03
FeO	71.35	70.06	71.58	66.23	71.67	68.84	71.79	71.68	69.20	70.68	70.02
MnO	0.73	1.13	0.89	1.06	0.93	0.54	0.54	0.59	0.75	0.80	0.67
MgO	1.39	2.24	2.53	2.25	3.01	2.47	2.62	2.67	5.19	1.66	2.04
CaO	0.33	0.37	0.24	0.19	0.19	0.37	0.28	0.06	0.10	0.12	0.19
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.00	0.35	0.26	0.30	0.40	0.00	0.00	0.02	0.09	0.02	0.04
Total	98.06	96.26	99.37	97.06	96.98	97.63	97.21	95.77	95.79	95.90	97.00

**APPENDIX TABLE 4: Major element compositions of daughter mineral**

MAG B 47	BIOTITE 51A	BIOTITE 54	APATITE 52	APATITE 54	APATITE 61	SPHENE 16	SPHENE 31	SPHENE 53
SiO2	43.86	35.50	4.86	2.33	2.67	31.05	30.50	30.95
TiO2	4.11	1.02	0.00	0.00	0.00	31.75	36.95	34.55
Al2O3	13.11	18.37	2.45	0.66	0.51	5.06	1.89	3.83
Cr2O3	0.10	0.00	0.00	0.00	0.00	0.00	0.00	0.00
FeO	9.81	13.08	1.08	1.13	1.04	1.37	0.85	0.69
MnO	0.12	0.15	0.00	0.00	0.00	0.00	0.00	0.00
MgO	17.30	16.54	0.86	1.87	0.78	0.32	0.00	2.55
CaO	0.06	0.08	49.33	50.94	51.53	29.32	29.16	27.41
Na2O	0.80	0.29	0.00	0.00	0.00	0.05	0.02	0.00
K2O	8.46	7.46	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.11	0.00	33.59	0.00	36.57	0.00	0.00	0.00
NiO	0.00	0.10	0.00	0.00	0.00	0.00	0.00	0.00
Total	97.84	92.58	92.17	56.92	93.09	98.92	99.37	99.98

MAG B 25	Ti-MNT 25	Ti-MNT 27	Ti-MNT 28	Ti-MNT 35	Ti-MNT 35	Ti-MNT 38	Ti-MNT 40	Ti-MNT 42	Ti-MNT 44	Ti-MNT 48	Ti-MNT 53	Ti-MNT 54
SiO2	0.14	0.47	0.11	0.10	0.89	0.19	0.91	0.15	0.19	1.14	1.90	5.70
TiO2	15.89	20.20	17.23	15.09	13.98	12.92	17.90	16.93	16.62	20.46	16.53	26.15
Al2O3	7.21	7.70	7.71	8.11	9.04	10.34	5.47	8.86	8.33	4.22	6.99	4.63
Cr2O3	0.05	0.03	0.04	0.06	0.04	0.06	0.10	0.04	0.05	0.12	0.52	0.01
FeO	70.54	65.83	70.48	70.35	69.31	71.20	69.49	71.09	68.33	70.83	67.46	54.55
MnO	0.42	0.70	0.51	0.46	0.33	0.48	0.51	0.58	0.50	0.79	0.34	0.31
MgO	3.78	1.40	1.68	2.72	2.30	2.11	2.01	1.41	1.98	0.15	2.52	1.92
CaO	0.14	0.25	0.17	0.19	0.36	0.35	0.58	0.21	0.58	0.37	0.30	0.91
Na2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
K2O	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.04	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00	0.00
Total	96.92	97.35	96.41	96.51	95.81	97.69	96.95	97.11	96.84	98.09	97.27	94.17

MAG B 25	ILM 16	ILM 34	ILM 52	SPINEL 7	SPINEL 55	PLAG 5	PLAG 23
SiO2	0.21	11.24	5.58	0.00	0.07	59.57	58.33
TiO2	47.24	38.24	44.86	5.72	3.05	0.00	0.00
Al2O3	0.39	3.59	2.30	16.30	17.22	22.30	26.64
Cr2O3	0.03	0.00	0.34	20.13	30.56	0.00	0.00
FeO	49.48	40.32	43.36	46.69	38.94	0.97	0.81
MnO	1.06	0.50	0.72	0.88	0.27	0.00	0.00
MgO	0.46	2.28	1.76	9.19	9.53	0.04	0.07
CaO	0.35	4.56	2.03	0.19	0.08	2.67	7.81
Na2O	0.00	0.00	0.00	0.00	0.00	3.00	6.70
K2O	0.00	0.00	0.00	0.00	0.00	1.92	0.50
P2O5	0.00	0.00	0.00	0.00	0.00	0.00	0.00
NiO	0.02	0.00	0.00	0.78	0.14	0.00	0.00
Total	96.68	96.27	96.99	96.10	95.69	95.94	97.74

APPENDIX TABLE 5: Major element compositions of carbonate globules

MAGB25	4a	4b	5a	9a	9b	10a	10b	10c	11a	20a	25a	25b
SiO <sub>2</sub>	0.19	0.06	0.06	0.05	0.12	0.00	0.00	0.05	0.02	3.87	0.10	0.10
FeO	14.73	41.22	33.87	10.37	37.02	39.16	12.34	12.09	31.72	42.45	34.08	43.10
MnO	1.99	2.26	3.14	0.41	0.85	1.64	0.28	6.42	0.55	0.88	1.23	1.20
MgO	4.05	6.03	8.94	34.62	11.67	8.16	34.56	11.10	4.72	3.39	8.25	7.12
CaO	36.10	7.20	11.54	7.08	6.09	6.48	5.05	23.79	19.02	9.48	15.32	9.06
Na <sub>2</sub> O	0.08	0.10	0.14	0.22	0.18	0.12	0.21	0.28	0.17	0.19	0.30	0.12
SrO	0.11	0.05	0.02	0.00	0.00	0.00	0.00	0.00	0.07	0.06	0.00	0.00
BaO	0.00	0.07	0.02	0.00	0.01	0.00	0.00	0.00	0.01	0.01	0.11	0.08
CO <sub>2</sub>	43.40	39.09	41.71	50.20	41.03	39.09	49.57	42.45	40.03	43.52	43.07	42.28
Total	100.65	96.09	99.43	102.96	96.97	94.65	102.00	96.18	96.31	103.85	102.48	103.05

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MAGB25	25c	25d	25e	32a	35a	36a	44a
SiO <sub>2</sub>	0.05	0.71	0.10	0.03	3.00	0.09	0.11
FeO	17.38	26.86	30.27	28.97	46.39	13.94	0.87
MnO	0.26	0.46	0.66	0.56	0.81	3.08	0.78
MgO	30.49	11.09	13.91	2.39	4.12	8.01	1.01
CaO	4.25	18.20	12.85	24.67	4.61	29.42	51.91
Na <sub>2</sub> O	0.42	0.15	0.20	0.09	0.23	0.28	0.06
SrO	0.02	0.05	0.12	0.10	0.01	0.04	0.07
BaO	0.03	0.04	0.05	0.04	0.00	0.00	0.00
CO <sub>2</sub>	47.80	44.31	44.58	40.23	41.59	42.62	43.10
Total	100.68	101.86	102.74	97.09	100.75	97.47	97.91

**APPENDIX TABLE 6: Major element compositions of sulfides**

MAGB47	42a	68a	70a	70b
Fe	33.57	58.28	60.01	32.69
Co	0.68	0.18	0.13	0.93
Ni	31.59	2.16	0.32	32.11
Cu	0.27	0.22	0.16	0.28
Zn	0.21	0.12	0.18	0.16
Pb	0.00	0.00	0.00	0.00
S	33.03	37.95	39.39	33.65
Total	99.34	98.90	100.20	99.81

MAGB25	1a	1b	1c	5a	5b	12a	13a	16a	25a	25b	28a	28b	28c	32a	32b	33a	36a	42a
Fe	39.98	40.05	59.48	44.58	43.85	58.79	32.56	61.77	45.03	44.79	48.78	56.87	56.14	46.11	45.11	45.87	33.04	58.79
Co	0.13	0.07	0.09	0.28	0.23	0.05	0.05	0.09	0.09	0.13	0.09	0.37	0.16	0.11	0.67	0.08	0.16	0.03
Ni	0.33	0.55	0.06	0.80	0.83	0.13	0.51	0.19	0.64	0.86	0.34	0.63	0.34	0.09	0.48	0.04	0.67	0.10
Cu	22.07	21.73	0.06	0.22	0.32	0.03	30.22	0.10	0.11	0.30	9.08	1.93	0.65	0.01	0.00	0.04	23.64	0.04
Zn	0.15	0.03	0.03	0.00	0.00	0.04	0.22	0.00	0.00	0.02	0.04	0.04	0.01	0.01	0.00	0.00	0.06	0.07
Pb	0.18	0.11	0.08	0.00	0.00	0.11	0.06	0.00	0.00	0.00	0.00	0.00	0.00	0.04	0.00	0.00	0.00	0.00
S	35.58	35.70	37.52	50.08	51.10	36.43	33.64	32.36	52.87	54.22	39.06	38.97	39.19	51.45	52.74	51.42	36.92	36.49
Total	98.42	98.25	97.32	95.95	96.33	95.57	97.26	94.51	98.75	100.33	97.40	98.80	96.50	97.82	99.02	97.44	94.49	95.53

TAAB26	14a	15a	34a	34b	38a	52a
Fe	55.79	44.64	58.37	58.93	56.63	30.76
Co	0.14	0.50	0.13	0.14	0.21	0.08
Ni	0.59	12.83	0.88	0.41	1.60	0.52
Cu	0.36	0.19	0.11	0.07	0.22	30.00
Zn	0.09	0.11	0.08	0.02	0.12	0.21
Pb	0.00	0.00	0.10	0.06	0.00	0.00
S	37.55	35.84	39.57	38.64	38.14	34.58
Total	94.52	94.12	99.23	98.26	96.91	96.15

APPENDIX TABLE 7: Trace element compositions of melt inclusions

MAGB47	13a	13b	16a	18a	21a cpx	25a	38a	39a	44a	44b	45Aa
Rb	3.48	8.00	1.81	15.68	1.40	5.40	21.22	8.50	8.87	2.61	34.56
Ba	1.74	193.42	0.07	765.99	0.10	1.85	522.82	235.66	8500.68	3311.41	5295.27
Nb	4.13	72.65	0.03	207.27	0.88	0.33	92.76	31.56	184.60	22.73	27.04
La	0.55	70.16	0.02	44.84	2.00	0.45	60.15	7.51	61.18	58.32	41.25
Ce	0.65	137.53	0.02	104.30	6.54	0.44	121.49	18.11	125.96	114.80	62.27
Sr	10.05	3276.29	0.34	2096.13	50.23	8.64	473.09	148.45	1810.12	295.09	825.82
Nd	0.23	85.52	0.11	55.13	6.78	0.22	54.40	9.79	60.22	55.00	26.62
Sm	0.09	17.89	0.15	11.35	2.41	0.03	10.74	2.19	13.48	11.41	4.28
Zr	30.22	427.48	0.69	284.56	41.65	2.94	303.59	97.58	282.31	279.31	56.42
Hf	0.43	8.36	0.12	6.47	1.22	0.06	7.90	1.60	7.00	6.54	6.78
Eu	0.04	6.94	0.02	4.59	0.60	0.01	1.26	0.89			4.55
Ti	1833.47	22747.72	92.04	42314.17	5744.25	90.54	19245.94	6877.13	38226.08	18353.15	704.87
Gd	0.08	20.41	0.05	13.57	2.12	0.06	6.22	2.15	8.95	6.94	6.34
Dy	0.14	12.04	0.08	8.31	2.02	0.05	7.18	1.66	9.87	7.11	4.18
Y	2.37	58.73	0.16	38.66	9.94	1.05	31.97	9.02	44.67	30.99	15.22
Er	0.19	5.13	0.13	4.02	0.93	0.06	3.61	0.72	12.67	7.74	1.01
Yb	0.16	3.08	0.11	2.49	0.49	0.08	1.89	0.64	2.73	1.97	0.56
Li	1.41	5.66	1.74	14.27	11.96	0.48	11.85	3.64	0.90	2.78	10.21
Be	0.75	1.85	0.01	2.15	0.11	0.69	5.05	1.65	2.95	3.30	2.18
Sc	14.39	39.41	4.98	31.36	117.60	33.57	22.89	21.67	26.74	17.67	24.06
Cr	252.90	575.10	389.18	164.53	5036.82	553.06	168.72	254.74	140.11	85.14	20.34

MAGB25	1a	4a	5a	8a	9a	9b cpx	11Aa	12a	16a	25a	25b
Rb	32.39	2.21	6.18	35.19	3.65	1.03	44.63	7.38	26.52	26.57	2.76
Ba	587.39	293.24	131.38	733.09	177.55	0.13	722.34	16.98	585.83	523.93	144.59
Nb	48.94	5.75	30.67	87.67	22.45	0.48	53.51	10.89	60.23	74.90	7.94
La	94.26	0.27	12.81	48.34	1.25	2.44	95.52	26.41	76.56	131.09	6.31
Ce	161.21	0.19	20.27	77.52	1.39	8.93	159.46	77.89	120.67	231.70	10.95
Sr	1072.33	28536.95	25224.46	1195.59	26401.69	50.28	1211.03	210.03	1060.42	972.70	14454.72
Nd	53.05	16.13	18.55	18.76	11.02	8.93	51.41	60.59	34.07	78.60	4.64
Sm	7.67	8.03	5.78	2.54	4.29	2.71	7.64	15.16	5.10	14.93	0.89
Zr	253.95	187.75	254.51	314.77	148.96	44.57	162.47	368.86	169.63	271.93	33.78
Hf	4.35	2.09	3.33	4.72	1.66	1.46	3.16	8.45	2.92	4.90	0.38
Eu	2.98	0.28	0.44	2.57	0.11	0.79	3.72	4.50	0.44	4.47	0.18
Ti	910.95	71.28	533.74	1994.10	162.57	6664.85	1052.48	33175.39	1147.65	1528.91	96.50
Gd	10.75		1.37	3.81		2.09	10.65	19.05	3.20	11.00	0.61
Dy	3.94		0.11	1.58		2.25	4.09	10.10	2.91	7.14	0.45
Y	17.36	23.65	27.29	6.41	21.71	9.29	16.17	40.20	15.39	33.88	2.70
Er	1.86		0.22	0.50	0.03	0.83	1.74	4.12	1.98	2.41	0.21
Yb	0.75	0.81	0.59	0.57	0.54	0.68	0.71	2.53	0.91	1.13	0.14
Li	23.47	0.04	2.16	3.89	0.44	0.37	18.15	2.81	17.91	6.69	0.23
Be	3.53	0.12	0.46	3.39	0.31	0.56	3.46	0.75	7.03	5.38	2.31
Sc	26.09	28.60	23.97	24.63	21.81	74.11	26.29	103.37	5.36	12.90	5.83
Cr	2.58	4.51	12.42	3.38	5.55	2995.12	5.17	189.75	1.05	1.56	3.53

APPENDIX TABLE 7: Trace element compositions of melt inclusions

MAGB47	46a	46b	49a	50a	54a	54b	55a	55b	57a	57b
Rb	1.78	1.28	17.58	2.06	6.12	2216.25	4.81	3.82	5.76	11.07
Ba	1.06	19.02	720.84	27.87	30.46	1.39	63.60	143.51	232.99	321.83
Nb	8.22	17.61	37.46	24.00	30.08	86.13	22.45	78.65	52.90	15.79
La	4.59	7.28	44.88	7.05	5.77	0.22	8.46	28.88	96.80	179.75
Ce	8.82	15.85	89.50	14.38	12.34	3.74	21.25	64.77	178.98	264.30
Sr	25.72	78.62	449.34	117.10	94.60	2.76	79.79	362.80	594.94	147.32
Nd	4.73	8.91	45.28	6.81	4.08	12.21	12.33	30.68	84.76	106.54
Sm	0.96	1.95	10.03	1.70	0.55	12.08	2.66	6.34	15.57	15.87
Zr	48.66	69.53	146.69	104.06	50.66	12.79	133.37	213.96	326.25	150.60
Hf	0.74	1.52	5.18	1.09	0.47	2.44	2.13	3.83	7.14	4.46
Eu	0.27	0.53	5.69	0.48	0.34	3.59	0.58	1.42	4.64	5.77
Ti	3086.71	4396.12	28348.80	5800.80	1554.71	115.25	7750.01	15780.62	27918.77	12698.95
Gd	0.74	1.29	14.25	1.54	0.94	1.05	2.24	4.15	21.59	26.83
Dy	0.57	1.33	8.16	1.02	0.44	3.93	2.04	4.11	10.24	9.07
Y	3.42	5.28	39.23	6.36	2.45	68.14	10.01	22.16	59.90	66.14
Er	0.22	0.60	3.13	0.47	0.23	2.57	1.04	2.30	5.02	4.13
Yb	0.18	0.39	2.58	0.39	0.15	0.90	0.74	1.31	2.57	1.25
Li	0.43	0.42	7.61	3.08	4.09		3.16	2.75	5.01	7.86
Be	0.38	0.48	3.11	0.96	1.78	1.94	2.04	1.94	1.94	1.57
Sc	11.09	7.08	45.05	20.05	18.14	0.33	17.41	20.24	58.84	49.67
Cr	225.75	146.86	302.33	302.26	198.01	10805.22	190.87	195.49	333.19	312.73

MAGB25	25c carbonate	25d	25e carbonate	25f	25g cpx	35a	35b	35c	36a	38a	39a
Rb	0.84	8.94	4.20	1.32	2.37	14.13	1.54	4.85	0.98	41.78	4.89
Ba	33.27	15.03	100.08	20.19	14.54	100.35	133.09	101.44	125.79	2957.32	163.19
Nb	29.47	43.27	90.69	10.03	4.92	11.03	5.74	17.31	11.87	29.78	34.25
La	29.91	1.02	60.22	7.39	9.37	10.45	0.28	22.33	3.08	31.45	79.94
Ce	54.66	1.49	136.77	10.04	28.70	15.91	0.22	38.96	3.05	47.00	154.80
Sr	312.80	402.42	10281.54	921.03	103.20	20990.34	18889.52	16023.07	18309.69	1424.06	15218.46
Nd	19.41	0.54	57.07	2.43	24.48	16.98	1.45	12.15	0.55	16.43	73.42
Sm	2.76	0.23	10.87	0.52	7.33	4.23	0.24	1.49	0.28	2.75	14.48
Zr	112.01	151.14	553.76	41.75	170.90	70.22	13.42	87.19	78.30	159.54	409.52
Hf	1.17	2.58	5.31	0.45	4.87	1.48	0.14	1.29	0.98	4.82	8.92
Eu	0.68	0.03	3.16	0.08	1.99	0.47		0.36		7.74	3.84
Ti	805.13	716.73	16049.68	244.29	13156.52	516.94	86.50	496.09	78.55	480.65	6947.67
Gd	1.35	0.10	7.25	0.13	5.63	2.43		0.70	0.05	3.87	16.08
Dy	1.24	0.08	6.75	0.31	5.74	1.69	0.08	0.67	0.10	2.00	9.31
Y	7.67	1.29	50.28	2.64	26.67	33.16	12.09	13.77	9.98	8.07	46.78
Er	0.56	0.14	5.29	0.28	2.60	1.72	0.15	0.41	0.24		4.57
Yb	0.59	0.20	8.93	0.51	1.58	2.53	0.17	0.34	0.16	0.29	2.93
Li	1.97	17.17	603.97	32.91	1.25	8.95	0.38	4.30	0.17	1.67	1.35
Be	1.64	3.46	20.53	3.71	1.63	6.50	1.73	1.89	2.95	1.61	1.05
Sc	0.25	4.97	15.64	0.83	47.16	27.15	5.50	3.24	1.58	40.56	20.19
Cr	3.67	25.89	274.82	19.36	73.45	83.84	2.55	7.70	1.11	9.20	31.00



APPENDIX TABLE 7: Trace element compositions of melt inclusions

TAAB26	1a	3a	4a	4b	5a	7a	8a	8a average	9a	16a	18a
Ba	1077.34	459.73	934.49	822.41	892.44	580.37	948.39	994.62	419.53	754.26	941.13
Nb	42.57	57.99	61.71	49.29	95.98	140.12	96.08	73.59	59.70	29.08	91.35
La	51.04	45.51	49.29	37.20	23.74	108.26	57.71	54.69	47.23	16.74	52.53
Ce	76.44	86.07	99.08	81.89	42.36	246.36	122.40	119.53	90.09	41.83	122.31
Sr	1022.24	708.63	1086.49	1022.43	801.57	1521.38	857.63	912.78	1253.56	286.04	1465.04
Nd	20.52	37.42	50.99	38.44	30.41	126.35	64.03	65.67	31.34	36.51	70.78
Sm	2.61	6.59	9.49	8.34	14.41	22.21	13.05	13.51	6.75	9.50	17.90
Zr	326.65	441.09	372.23	358.96	401.21	874.88	655.24	854.95	348.69	306.11	544.63
Ti	5113.88	7384.32	22821.60	9390.45	29915.19	33529.36	17901.41	15497.09	9092.98	22280.12	33493.87
Dy	1.11	5.37	5.94	4.62	7.95	14.53	9.57	9.14	2.34	8.34	10.21
Y	10.47	18.57	25.47	21.41	12.95	60.07	37.10	31.46	13.30	36.17	45.42
Yb	0.41	1.45	1.36	1.14	3.53	5.45	2.96	2.69	2.08	2.45	3.56

MAGB25	44a	52a	54a
Rb	10.13	19.36	28.32
Ba	138.08	545.71	341.87
Nb	29.36	59.24	119.55
La	39.33	57.12	31.37
Ce	69.11	104.97	84.75
Sr	564.40	1061.80	1010.66
Nd	20.61	41.29	57.17
Sm	2.42	8.79	13.19
Zr	145.90	163.57	362.69
Hf	2.50	2.41	7.31
Eu	0.97	1.02	4.75
Ti	371.06	2839.44	26968.74
Gd	0.97	0.84	19.28
Dy	1.13	5.82	9.73
Y	7.14	27.31	45.31
Er	0.70	3.10	4.34
Yb	0.60	1.22	2.81
Li	3.31	14.17	16.86
Be	3.76	4.75	2.07
Sc	2.80	6.00	65.08
Cr	22.73	10.63	273.30

**APPENDIX TABLE 7: Trace element compositions of melt inclusions**

TAAB26	20a	21a	22a	29a	29a average	30a	30a average	32a average	59a	60a
Ba	1086.94	923.58	1369.79	312.35	285.86	673.33	636.33	1001.29	1293.78	1078.10
Nb	52.36	102.23	130.93	115.98	104.80	50.16	47.40	61.07	95.91	64.23
La	60.86	55.22	75.52	161.75	177.41	39.26	34.75	52.42	55.74	41.41
Ce	103.82	120.73	137.69	287.76	324.67	84.98	72.11	102.80	116.49	72.88
Sr	1215.50	1318.70	1733.73	941.71	985.07	637.82	602.46	1404.35	2031.49	1397.17
Nd	36.01	61.47	48.91	134.26	152.39	44.85	39.23	50.64	63.53	28.45
Sm	5.27	12.62	7.49	26.64	28.81	9.51	9.06	9.53	11.47	4.97
Zr	432.24	741.84	456.58	1065.89	1394.17	485.71	472.11	496.47	541.68	457.36
Ti	11524.91	11110.07	21430.24	24561.40	24980.75	13944.40	12159.60	16056.62	22750.82	9460.36
Dy	4.11	7.58	4.14	13.69	15.83	6.57	6.45	5.86	8.90	2.97
Y	20.33	31.41	21.60	69.38	79.78	30.45	27.22	26.86	46.94	17.31
Yb	1.35	2.26	1.99	5.51	7.84	3.21	2.15	2.40	3.47	1.28

## 4. Chapter Four

### U series isotopic variability in Galapagos lavas, evidence of a warm-spot

#### 4.1 Abstract

Evaluation of U series disequilibrium, trace element composition and Sr, Nd and Pb isotopes of Galapagos lavas indicates that magma mixing between plume and asthenospheric melts has been the main process responsible for the geochemical variation observed in the archipelago. Correlation between He isotopes and Ti/Ti\*, K/Rb and Nb/La ratios suggest that the mantle plume has positive anomalies of Nb and Ti and negative anomalies of K.  $^{230}\text{Th}$  excesses measured in the lavas indicate that the basalts from Galapagos originated total or partially in the garnet stability field. Mantle upwelling velocity for the Galapagos plume (Fernandina) ranges from  $\approx 1$  to 2 cm/y with a maximum porosity of 0.3% indicating that Galapagos is a mildly buoyant plume. Very slow mantle upwelling rates and very low porosity for Pinta (0.5 to 1 cm/y and 0.1%) and Floreana (0.1 cm/year and  $<0.1\%$ ) support the hypothesis that the movement of the plume across the  $91^{\circ}50'$  transform fault produced slow upwelling and small extents of melting.

#### 4.2 Introduction

U decay series isotopes provide information on magmatic processes and mantle source compositions that are inaccessible by any other isotopic systems. Thus,  $^{238}\text{U}$ - $^{230}\text{Th}$  disequilibrium provides quantitative evaluation of the present-day depletion of the mantle source, the extent of melting represented by erupted lavas, and helps to constrain estimation on the mantle porosity and upwelling velocities during melting (Elliot, 1997 and references therein). Two end-member models have been used to explain the extent of disequilibrium in volcanic rocks:

1) as an index of elemental fractionation during melting (Sims et al., 1995),  
and

2) as an indication of  $^{230}\text{Th}$  ingrowth during time-dependent melting and melt transport (McKenzie, 1985; Williams and Gill, 1989; Spiegelman and Elliott, 1993).

U series isotopes have also been useful in dating young lavas, estimating magma chamber residence times (Allègre and Condomines, 1976; Goldstein et al., 1994, Sigmarsson, 1996), and in probing the slab-component input in subduction zone volcanism (Gill and Williams, 1990; Hawkesworth, et al., 1997, Elliot et al., 1997). Moreover  $^{238}\text{U}$ - $^{230}\text{Th}$  disequilibrium has been used as a geodynamic indicator of the strength of the plumes traveling through the convecting mantle (Chabaux and Allègre, 1994) and as a sign of ridge-hotspot interaction (Bourdon et al., 1996). Thus, U decay series isotopes have become one of the most powerful geochemical tools to answer fundamental questions in igneous petrogenesis (Condomines et al., 1988, Goldstein et al., 1992; Rubín and Macdougall, 1992; Cohen and O’Nion, 1993; Hémond et al., 1994; Chabaux and Allègre, 1994; Lundstrom et al., 1995; Bourdon et al., 1996; among others).

The Galapagos archipelago has been the object of many geologic studies since Darwin’s (1860) first report on the islands. Although the Galapagos islands have been interpreted as the surface manifestation of a mantle plume, geological, geophysical and geochemical studies have found it challenging to fit the Galapagos plume to a “standard” hotspot model. In contrast to Hawaii (the “archetype” of the plume model), the Galapagos islands do not form a linear chain. The volcanism is not strictly time-transgressive in the direction of the plate motion, nor is there a consistent petrological or geochemical evolution of the volcanism as seems to be the case in Hawaii. In fact, the differences between Galapagos and the standard hotspot model make this archipelago unique. For example, the historic to Holocene volcanism occurs throughout a region of 30,000 km<sup>2</sup>, and the geographical distribution of various petrological and geochemical lava types within the archipelago is quite diverse (McBirney and Williams, 1969; Geist, 1992; White et al., 1993 and references therein). These differences have been explained by entrainment of asthenospheric mantle into the plume subject to velocity shear, and by interaction of the plume with the Galapagos Spreading Center (Geist et al., 1988; Richards and Griffiths, 1989; White et al., 1993; Ito et al., 1997). Thus, the Galapagos islands provide a unique natural laboratory to test petrologic, tectonic and fluid dynamic models of ridge-hotspot interaction. Although the Galapagos plume is one of the two best known examples for ridge-hotspot interaction, data for U series isotopes in Galapagos lavas do not exist.

We present here the first data for  $^{238}\text{U}$ - $^{230}\text{Th}$ - $^{232}\text{Th}$  isotopes on a suite of 29 basalt samples from the Galapagos islands. The samples are geochemically well-characterized and of known age, and were chosen to represent the different petrographic and geochemical groups observed in the Galapagos archipelago (McBirney and Williams, 1969; Geist, 1992; White et al., 1993, Kurz and Geist, 1999). The main objective of this work is to constrain the mantle composition, extent of melting, mantle upwelling velocity and porosity beneath the Galapagos islands. We compare our U decay series results with those obtained from three dimensional numerical model of the interaction between mantle plumes and migrating mid-ocean ridges (Ito et al., 1997). The diffuse nature of the Galapagos plume suggests it is significantly less intense thermally than the Hawaiian hotspot .

### ***4.3 Geodynamic Background***

The Galapagos archipelago is located in the eastern equatorial Pacific, 1000 km west of South America, and approximately 150 km south of the active Galapagos Spreading Center (GSC) (Figure 1). The islands and numerous seamounts are constructed mainly upon a shallow submarine volcanic platform overlying a very young lithosphere (< 10 Ma). The GSC is a moderately fast spreading ridge ( $\approx 6$  cm/year); judging from the volcanic traces of the hotspot (the Cocos and the Carnegie Ridges), the ridge has migrated northward away from the Galapagos hotspot in the last 5 to 6 Ma (Hey, 1977). The transform fault at  $90^{\circ}50'$  offsets the GSC by 150 km and is the boundary between symmetric spreading to the west and asymmetric spreading to the east. This fracture zone divides the Galapagos platform into two main sectors (the western and eastern sectors) with an estimated 5 Ma age discontinuity across the lithospheric sectors (Hey, 1977). Analysis of bathymetric and gravity data indicates that the eastern sector of the archipelago is underlain by a weak, younger lithosphere with an elastic thickness of 6 km or less and close to Airy compensation. In contrast, the western sector is flexurally supported by a lithosphere with an elastic thickness of 12 km (Feighner and Richards, 1994). It is important to remark that the transition from the eastern to the western sector is sharp and can be modelled as a fault zone (Figure 1). The crustal thickness increases from approximately 10 km on the margin of the Galapagos platform to as much as 18 km beneath southeastern Isabela.

Galapagos volcanoes are aligned in two rough trends N30W and N60E (Darwin, 1860). Prominent block faulting can also be seen in the central and southern islands

(McBirney and Williams, 1969). Altogether, the evidence suggests that faulting either controls or greatly modifies the major structural features of the Galapagos hotspot region (Feighner and Richards, 1994). There is not yet a good explanation for these alignments, as they bear no simple relationship to plate motion, local transform directions or likely direction of plume flow (Geist and Reynolds 1998).

The first geologic report on the Galapagos was made by Darwin (1860); since then, many reports and studies have followed. The first detailed geologic and petrologic study of the archipelago as a whole was done by McBirney and Williams, (1969); this was followed by the work of Geist et al, (1988); Geist, (1992); and White et al, (1993). Many studies have contributed in considerable detail to our understanding of individual volcanoes and islands (see Geist, 1992; White et al, 1993; and reference therein). We present here a brief synthesis of these works. To the west of the 90°50' transform fault, the eruptive edifices are large central shield volcanoes of unusual shape ('inverted soup-bowl') with well-developed calderas, and with lavas erupting from radial and ring fissures. These volcanoes have the most frequent historic volcanic and seismic activity due to their location close to the present position to the plume at Fernandina, (Kurz and Geist, 1999). In contrast, the volcanoes to the east of the fracture zone are seismically quiet and they present more varied morphologies and eruptive histories. This difference in size of the volcanic edifices probably reflects the effect of lithospheric thickness, with older and thicker lithosphere being able to support larger volcanic structures.

From the 21 emergent volcanoes in the Galapagos, 9 have been active historically and four others have been active in the Holocene. There is no clear age progression in the geographic distribution of the volcanoes. However, the ages of the oldest lava dated in each volcano form a progression from young in the west to old in the east (White et al., 1993; Kurz and Geist, 1999). The islands, as they drift away from the hotspot, have had intermittent eruptions for at least 2.4 Ma, as in the case of San Cristobal island (White et al., 1993; Kurz and Geist, 1999). The continuous volcanism in this most easterly island indicates either an extended zone of melting downstream from the plume or the existence of postshield extensional volcanism across the Galapagos platform.

Geochemically, Galapagos lavas have lower SiO<sub>2</sub> and higher FeO, Na<sub>2</sub>O, CaO, K<sub>2</sub>O and TiO<sub>2</sub> than MORB (at the same MgO content). Lavas erupted from the large western volcanoes are in general moderately fractionated tholeiites, while lavas erupted

from small cones and vents from the eastern islands are compositionally diverse, ranging from picrites to basanitoids. Most of the isotopic variation observed in the Galapagos archipelago occurs between islands or even between volcanoes in the same island, as in the case of Isabela island. However, lavas erupted in some of the eastern islands (San Cristobal, Santa Cruz Santiago) show significant intra-island variation in isotopes, largely related to geographic or temporal sequences. Major and trace element contents and ratios have been used to suggest that melting appears to be less extensive and shallower to the east and south, a possible indication of a cooler mantle away from the center of the plume (White et al., 1993; Geist, 1992; Geist et al., 1988). However, it has been extremely difficult to unequivocally separate effects due to the extent and depth of melting from those of source composition. For example, Geist (1992) proposed that the depth of melt segregation can be estimated in Galapagos lavas from an inverse correlation between  $Si_8$  and  $Fe_8$ . Although  $Fe_8$  and  $Na_8$  are frequently used as indicators of pressure and extents of melting respectively (Geist et al., 1988; White et al., 1993), in Galapagos these indices correlate with He isotopes, which is an indicator of source composition (Kurz and Geist, 1999). White et al. (1993) estimated the degree and depth of melting by considering ratios of LREE to MREE and MREE to HREE, for samples of similar Nd isotopes composition. Thus, White et al. concluded that lavas from the eastern islands, especially Floreana and San Cristobal, were probably produced by a low extent of melting in the spinel stability field (shallow depths). However, these same authors also show that in Galapagos lavas  $(La/Sm)_{PM,8}$  ratios ( $La/Sm$  ratios normalized to primitive mantle and at  $MgO$  content  $\approx 8$ , averaged by volcano) correlate with averaged isotopic compositions, suggesting a source control on the REE.

Isotopes and incompatible element ratios from Galapagos lavas define a “horse shoe”-like geographic pattern, with the most depleted signatures (MORB) appearing in the center of the archipelago and the most enriched signatures emerging on the eastern, northern and southern peripheries (Geist et al. 1988; White et al., 1993). This horse shoe pattern has been explained by entrainment of asthenospheric mantle into the plume, controlled by velocity shear produced by the moving lithosphere (Richards and Griffiths, 1989).

The GSC basalts at  $92^\circ W$  have anomalously enriched Sr isotopic compositions, clear evidence of the interaction between the ridge and the Galapagos plume (Schilling et al., 1982; Verma and Schilling, 1982). Surprisingly, the geochemical anomaly does not

occur at the closest position of the GSC to the Galapagos archipelago, but further to the northwest, at the intersection of the Darwin-Wolf lineament and the GSC. Morgan (1978) suggested that the plume flow was channeled at shallow level beneath the Darwin-Wolf lineament. However, basalts dredged from this lineament become isotopically more plume-like toward the GSC (Harpp et al. , 1990). This observation requires plume flow in the opposite direction to that in which the shallow asthenosphere is being dragged by the moving plate. Therefore, the flow from the plume toward the ridge must be deep rather than shallow, perhaps related to deep asthenospheric counterflow toward the ridge (White et al., 1993).

#### ***4.4 Sampling and Analytical techniques***

We sampled historic to Holocene lavas from numerous volcanoes and islands from the Galapagos archipelago. Our strategy was to obtain representative samples from the different lava types that define the geographical “horse shoe” pattern in isotopes and incompatible element ratios observed in the Galapagos (Geist et al., 1988; Geist, 1992; White et al., 1993). The sampling was simplified by the fact that the observed isotopic variation in the Galapagos archipelago occurs between islands or even between volcanoes in the same island; the intra-volcanic variations are generally small. We selected samples from the large central shield volcanoes from the western area and from volcanoes representing more varied morphologies and eruptive histories in the eastern sector of the archipelago. At the same time, this strategy took into account basalts erupted through lithosphere of different ages and elastic thickness by sampling east and west of the fracture zone defined by Feighner and Richards (1994). We selected 29 samples for U-series analysis:

a) samples from Fernandina Vn. (8), Cerro Azul Vn. (4), Wolf Vn. (4), Sierra Negra (3), Darwin Vn. (2) and Ecuador (1) represent the large shield volcanoes, closest to the center of the plume, erupted on the older and thicker lithosphere to the west of the 90°50' fracture zone.

b) basalts from San Cristobal (3), Santiago (1), Floreana (1), Pinta (1) and Marchena (1) represent small volcanoes erupted away from the center of the plume on a thin and young lithosphere, east of the fracture zone.



When possible several samples with similar Sr, Nd, Pb and He (i.e. constant source composition) isotopic composition were taken from a single volcano to evaluate the melting process. However, we faced the problem that lavas from isotopically homogeneous volcanoes (western volcanoes) generally are less primitive (e.g. low MgO contents), and are typically homogeneous in trace element ratios (i.e., Fernandina lavas). Therefore, the limited variation in trace element ratios does not allow us to make a precise evaluation of melting processes. In the eastern islands, where the trace element variations are larger, either the lavas are variable in isotopic composition, or they are too old for U decay series analysis (i.e., Floreana lavas). Moreover, lavas from the northern islands, like Marchena, are very poor in olivine phenocrysts (the mineral used for surface exposure dating techniques), making it impossible to obtain ages for those basalts. This is the main reason that the number of analyses in the eastern and northern islands are fewer than in the western volcanoes.

The selection of samples for U decay series analysis was constrained mainly by the age data available for the lavas. Many of the samples are from well recorded historic eruptions, such as in Fernandina and Marchena. All our samples have been dated by surface exposure dating, with the exception of those samples obtained from well recorded historic eruptions. This technique, using cosmogenic  $^3\text{He}$  (Kurz, 1986; Kurz et al., 1990, Kurz and Geist, 1999), has been ideal for the Galapagos archipelago where other dating techniques have failed. The  $^{14}\text{C}$  dating technique is difficult to apply because the arid climate in the Galapagos makes charcoal virtually nonexistent. Also,  $^{40}\text{K}$ - $^{40}\text{Ar}$  and  $^{39}\text{Ar}$ - $^{40}\text{Ar}$  techniques have been difficult to apply to lavas younger than 100,000 years, mainly due to the low potassium content of most Galapagos lavas.

Most of the samples for surface exposure dating were collected from the top 5 to 10 cm of lava flows to minimize the correction for depth in the lava flows. Wherever possible, we sampled pahoehoe lavas or large blocks from aa flows where original orientation was preserved (Kurz and Geist, 1999). The preservation of the lava flow surface is extremely good in the Galapagos due to the slow rate of weathering in the arid climate. A serious limitation to the surface exposure dating was the scarcity of olivine phenocrysts, the ideal mineral phase for surface exposure dating (Kurz and Geist, 1999). For example, the lack of olivines was a serious limitation in lavas from the western and northern Galapagos volcanoes (Sierra Negra, Alcedo, Wolf, Fernandina, Marchena, among the most important). The ages were calculated using sea level and high latitude production rate scaled

to equatorial areas (Lal, 1991). The uncertainties reported only consider the analytical uncertainty and not the errors in the absolute production rate (Kurz and Geist, 1999). For more detail on the analytical techniques please refer to Kurz (1986) and Kurz et al. (1996).

The analytical technique and uncertainties for Sr, Nd, Pb and He isotopes are reported in Kurz et al., (1995; 1996) and Kurz and Geist, (1999). Major and trace element data reported here were done by XRF and ICP-MS at Washington State University. The precision for XRF data (major and trace elements) for seven repeat analyses of BCR-P was < 1% ( $2\sigma$ ) for all major elements, except MnO, K<sub>2</sub>O, Na<sub>2</sub>O and P<sub>2</sub>O<sub>5</sub> with <3% ( $2\sigma$ ). The precision for XRF analyses for trace elements was <5% ( $2\sigma$ ) for most elements, <10% ( $2\sigma$ ) for Rb, Ga, Cu and Zn, <20% ( $2\sigma$ ) for Sc, and <30% ( $2\sigma$ ) for Ba. The precision for ICP-MS data on a single sample BCR-P run over a four month period (n=24) was <5% ( $2\sigma$ ) for most trace elements, <10% ( $2\sigma$ ) for Cs, Rb, Pb and <20% ( $2\sigma$ ) for Th and U.

To determine U-Th concentrations and Th isotope ratios the fresh interiors of the samples were crushed ( $\approx$ 1 kg) in a steel jaw crusher. The coarse fraction (between 1 to 2 cm), was cleaned in an ultrasonic bath with ultrapure water. The chips were dried at low temperature on a hot plate, and inspected under the binocular microscope to avoid any altered material. To avoid problems of sample inhomogeneity, we selected approximately 30 to 50 g of material to be powdered. The chips were powdered in an agate shatter box for no longer than a couple of minutes to minimize contamination. 1 to 2 grams of sample were dissolved and aliquots from the solution were used to chemically separate U and Th by standard chromatographic techniques (similar to the procedure reported in Goldstein et al., 1989a and b). For Th isotope determinations, a grade A spectrographically pure carbon planchette, pre-cleaned with isopropanol in an ultrasonic bath, is heated on a hot plate between 80 and 100°C. A low temperature polyethylene tube 900 $\mu$  in inside diameter is positioned perpendicular to the planchette such that when the tube touches the hot planchette a ring of melted plastic 750  $\mu$  in diameter is formed on the carbon. The unspiked Th concentrated from the chromatographic column is dissolved in 1 $\mu$ l of 0.3N HNO<sub>3</sub> and loaded onto the planchette at 80°C, inside the plastic ring. The ring is used as a dam to avoid spreading of the Th solution as it dries down into the planchette (spreading of the Th load produces a decrease in Th concentration per unit area, resulting in low intensity measurements). The solution is dried down into the carbon generating a conductive matrix that allows the Th to be analyzed by ion probe. The plastic ring was monitored to be absolutely sure it did not contribute any contaminant or artifact to the measurements.

The Th isotopic ratios were determined on the IMS Cameca 1270 ion microprobe at the Woods Hole Oceanographic Institution (Layne and Sims, 1999). We used a nominal primary accelerating voltage of 12.5 keV with a  $^{16}\text{O}^-$  beam, a secondary accelerating voltage of 10 keV, and currents ranging from 90 to 225 nAmps in a  $250\mu$  square raster. The dynamic transfer optics were utilized to recover 100% of the rastered beam. To resolve isobaric interferences, secondary ions were analyzed at a mass resolution of 2000, with no energy filtering. Each analysis comprised 150 cycles of alternating measurements of  $^{232}\text{Th}$  in the Faraday cup (2 s counting),  $^{230}\text{Th}$  in the electron multiplier (6 s counting), background in the electron multiplier (2 s counting) and background in the Faraday cup (2 s counting). Calibration for the gains of the detectors were done by measuring  $^{232}\text{Th}$  in a pure metal in both the Faraday cup and the electron multiplier. We used a gravimetric standard solution from the University of California Santa Cruz (generously supplied by J. Gill and C. Lundstrom), previously analyzed by TIMS as a standard during the period of our Th isotopes measurement. Layne and Sims (1999) measured the Th metal and the standard one after the other in sequence to correct for gain calibration and to evaluate reproducibility. The ion probe analyses for the standard  $^{232}\text{Th}/^{230}\text{Th} = 1.71 \times 10^5 \pm 0.7\%$  ( $n=20$ ), reproduced within errors the TIMS value of  $1.706 \times 10^5 \pm 0.8\%$  ( $2\sigma$ ,  $n=30$ ) (Figure 2 $\sigma$ ). The standard was measured two to three times a day, interspersed with our runs. The total Th content of the loaded standard was similar to that of the samples. During the course of our runs only once we determine the gain calibration by measuring  $^{232}\text{Th}$  in a pure metal with the Faraday cup and electron multiplier. To correct for mass fractionation and detector gain variation in our samples, we normalized the  $^{232}\text{Th}/^{230}\text{Th}$  values obtained for the standards measured each day to a standard value for  $^{232}\text{Th}/^{230}\text{Th} = 1.71 \times 10^5$  ( $n=20$ ). The rationale behind directly normalizing the measured values for the standard to a  $^{232}\text{Th}/^{230}\text{Th} = 1.71 \times 10^5$  is the constancy of the obtained value for gain calibration using Th metal during several days (Layne and Sims, 1999; G. Layne pers. com., 1999). Typical  $^{230}\text{Th}$  and  $^{232}\text{Th}$  count rates for our samples were approximately  $> 200$  cps for  $^{230}\text{Th}$  and  $> 10^7$  cps  $^{232}\text{Th}$ . The typical in-run precision for the Th isotopic measurements in the Galapagos samples ranges from  $\approx 0.5\%$  ( $2\sigma$ ) to  $1\%$  ( $2\sigma$ ) similar to or slightly better than the external reproducibility. The agreement for a duplicate analysis of sample NSK97-224 is  $\approx 1.0\%$ .

Th and U concentrations were determined by isotopic dilution using a Finnigan MAT Element ICP-MS at the Woods Hole Oceanographic Institution. A small aliquot (3 -

5%) from the initial solution is spiked with  $^{229}\text{Th}$  and  $^{233}\text{U}$  spikes such that the  $^{232}\text{Th}/^{229}\text{Th}$  and  $^{238}\text{U}/^{233}\text{U}$  ratios are approximately 30 and 10 respectively. Special care was taken to insure sample/spike equilibration, by drying down the solution several times with  $\text{HClO}_4$ . The Th and U are separated by standard chromatographic methods and the Th-U concentrate is diluted in 5 ml of 1-2 N  $\text{HNO}_3$  solution. The ICP-MS is tuned and calibrated twice a day with a 0.5 ppm solution of 16 different elements (from Li to U). Each sample analysis comprised 2000 cycles of alternating measurements of  $^{232}\text{Th}$ ,  $^{229}\text{Th}$ ,  $^{238}\text{U}$  and  $^{233}\text{U}$ . The typical in-run precision for the  $^{232}\text{Th}/^{229}\text{Th}$  and  $^{238}\text{U}/^{233}\text{U}$  measurements are  $<0.3\%$  ( $2\sigma$ ). The concentrations of  $^{229}\text{Th}$  and  $^{233}\text{U}$  in our spikes are known to approximately  $0.5\%$  ( $2\sigma$ ). To avoid memory effect, washout time of 15 to 20 minutes were used between samples. To monitor the memory effect, we ran blanks of 1-2 N  $\text{HNO}_3$  solution between every four samples.

We used the U standard NBS 960 to correct for U mass fractionation, and assumed the same fractionation per amu for Th. The observed mass fractionation in the standard was  $\approx 0.2\%$  per amu. Each sample was bracketed between runs of the standard, and corrected individually based on the mean mass fractionation value of the two standard. To evaluate external reproducibility of NBS 960 U standard, we corrected the standard for fractionation in the same way as we did for the samples. Namely, Standard #1 and #3 were used to correct standard #2 for fractionation; standard #2 and #4 were used for standard #3, etc. Results are shown in Figure 2b. The external reproducibility of the NBS960 U standard was  $\approx 0.2\%$  ( $2\sigma$ ,  $n = 23$ ). Two total procedure blanks gave values of 50 and 160 pg for Th and U respectively. All samples have been corrected for blank contribution.

## **4.5 Results**

### **4.5.1 Major elements, trace elements and Sr, Nd, Pb and He isotopes**

Most of the major, trace element and isotope data used in this work have been already published. The dataset used here is mainly from Kurz and Geist, (1999), complemented with data from Geist and coworkers (Geist, 1992 and reference therein), and from White et al. (1993); we have included some unpublished data from Kurz and Geist (personal communication). Table 1 lists representative major, trace element, and isotope analyses from the published and unpublished data set of Kurz for those samples analyzed for U series isotopes.

All of the lavas are basalts, mostly tholeiitic with only a few alkali basalts (Figure 3a). Mg# ranges from 0.56 to 0.26 indicating that all the samples have undergone fractional crystallization, mainly at low pressure (Geist, 1992; White et al., 1993) (Figure 3b). In some cases, like Cerro Azul, Nauman and Geist, (1998) have suggested that the range in major element contents in contemporaneous lavas (from tholeiites to alkali basalt), is the result of high pressure fractional crystallization.

Figure 4 shows primitive-mantle-normalized trace element diagrams for representative samples analyzed for U decay series. There are important variations in trace element patterns between lavas erupted in the western volcanoes (e.g., Fernandina and Darwin) and those from the eastern islands (e.g., San Cristobal and Floreana). Lavas from the western volcanoes have higher MREE/HREE ratios and less intra-volcano variability than lavas from the eastern islands; features already observed by White et al., (1993). These authors suggested that the variation in MREE/HREE ratios between lavas from the western volcanoes and those from the eastern islands was produced by difference in the depth of melting, from the garnet to the spinel stability field respectively. However, the lower MREE/HREE ratios in basalts from the eastern area are mainly produced by the variations in MREE rather than the variations in the HREE (Figure 5a). For example, Floreana lavas have similar LREE/HREE ratios and higher LILE, but lower MREE contents than those measured in basalts from the western volcanoes (Figure 5 a and b). This is an unexpected result assuming the variation in trace element content is produced by difference in the extent of melting or depth of melt segregation. Moreover, lavas from Darwin and Wolf volcanoes have relatively low LREE/HREE but high MREE/HREE ratios compared to other lavas from the western volcanoes (Figure 5b).

Lavas from the western volcanoes are characterized by positive anomalies of Nb, Ta and Ti and negative anomalies of K and Y. These anomalies tend to decrease or disappear in lavas with flat trace element pattern (low Hf/Lu) from the eastern islands, with the exception of a Ti anomaly in the Floreana basalts (Figure 4a and 6). Positive Sr anomalies seems to be ubiquitous in the Galapagos archipelago. The Sr anomaly tends to be small in the western volcanoes, probably due to plagioclase fractionation, and large in basalt from the eastern islands (Figure 4a). Lavas from Floreana and San Cristobal islands have low Th and U contents relative to elements of similar compatibility, features that are not observed in basalts from the western volcanoes (Figure 4a).

The Marchena and Pinta samples have not been analyzed for trace or major elements. White et al, (1993) reported REE data for samples from both islands. The REE content in Marchena samples defines a flat slightly concave downward pattern. In contrast, Pinta samples have enriched LREE similar to the lavas from the southern Isabela volcanoes.

Sr, Nd, Pb and He data have been discussed extensively in Geist et al. (1988); White et al. (1993) and Kurz and Geist (1999). The data set utilized here confirms the same remarks previously made in those papers. Namely, the western volcanoes tend to be homogeneous isotopically and with higher Sr, Pb and He and lower Nd isotopic composition than MORB. The exceptions are lavas from Darwin and Wolf volcanoes, which have MORB-like isotopic signature. In contrast, the eastern islands show a relatively large isotopic heterogeneity frequently with MORB isotopic signatures. Most of the isotopic variation in the archipelago occurs between volcanoes, and this generates a geographic distribution of the isotopic variations, which define a horse shoe pattern. A good correlation exists between Sr, Nd and Pb isotopes, but the correlation of these with He isotopes is not significant with the exception of Nd isotopes (Kurz and Geist, 1999). Also, these correlations hold when average isotopic compositions of volcano are compared (White et al, 1993). Four main components are needed to bracket the isotopic composition of the Archipelago: a MORB component defined by the Galapagos Spreading Center and Genovesa island, two components more enriched isotopically than MORB defined by Floreana and Pinta and the high He component characterized by Fernandina (White et al., 1993, Graham et al., 1993; Kurz and Geist, 1999). As is the case in all OIBs with high  $^3\text{He}/^4\text{He}$  ratios (Hawaii, Iceland, Samoa), the component with the highest  $^3\text{He}/^4\text{He}$  ratios (Fernandina), is associated with intermediate Sr, Nd and Pb isotopic ratios (Kurz and Geist, 1999).

The relationship between incompatible trace element ratios and isotopes is interesting. White et al, 1993, and Geist, 1992 showed that  $(\text{La}/\text{Sm})_{\text{PM},8}$  (normalized to primitive mantle and to  $\text{MgO} = 8$ ), averaged by volcano, correlates with averaged Sr, Nd and Pb isotopic composition. Kurz and Geist (1999) found that Nb/La ratios correlate positively with  $^3\text{He}/^4\text{He}$  isotope ratios and negatively with Nd isotope ratios. Fitton et al. (1998) suggested that the enrichment of Nb in Galapagos lavas correlates with Sr, Nd, Pb and Hf isotope ratios, indicating mixing between plume and MORB components. Complementing these previous observations, we also found that He isotopes ratios

correlate with Ti/Ti\* (Ti anomaly), Ti/Zr and K/Rb ratios (Figure 7). Moreover, Nb/Zr, Ba/La, K/Rb and Ba/Ce ratios correlate with Sr, Nd and Pb isotope ratios (Figure 8), and Ti/Ti\* defines a rough correlation with Nd isotope ratios. The exceptions to these correlations are the alkali basalts from Floreana, which seem to be enriched in LILE contents at similar LREE contents relative to other Galapagos samples. Moreover, the Floreana lavas are depleted in Th and U, and enriched in Ba compared to elements with similar incompatibility (Figure 4a).

#### 4.5.2 U decay series

U, Th concentrations and ( $^{230}\text{Th}/^{232}\text{Th}$ ), ( $^{238}\text{U}/^{232}\text{Th}$ ) and ( $^{230}\text{Th}/^{238}\text{U}$ ) ratios for Galapagos samples are reported in Table 2. The ( $^{230}\text{Th}/^{232}\text{Th}$ ) and ( $^{230}\text{Th}/^{238}\text{U}$ ) activity ratios have been corrected by eruption ages determined by surface exposure dating techniques (Kurz, 1986; Kurz et al., 1990, Kurz and Geist, 1999). Most of the samples have eruption ages well below 10,000 years. The exceptions are three samples from San Cristobal, with ages of 9200, 13,000 and 15,000, ( $\pm 1000$ ) years, and one sample from Floreana with an age of  $26,000 \pm 7000$  years. The ( $^{230}\text{Th}/^{238}\text{U}$ ) disequilibrium of 80% in the case of the Floreana sample is extreme. To evaluate the possibility of alteration effects, we used trace element ratios such as Rb/Cs and Ba/Rb which are easily modified by secondary process. The Floreana sample has Ba/Rb  $\approx 14$  and Rb/Cs  $\approx 89$ , indistinguishable from the canonical mantle values of  $13 \pm 1$  and  $90 \pm 10$  respectively (Hart and Reid, 1991; Hofmann and White, 1983). This indicates that alteration is not the explanation for the large ( $^{230}\text{Th}/^{238}\text{U}$ ) disequilibrium in the Floreana sample.

Extended magma chamber residence times will produce a decrease in the extent of disequilibrium, leading to a smaller inferred U/Th fractionation during the melting process. This is a concern for the western shield volcanoes that, in general, have erupted evolved and homogeneous lavas, indicating some processing in magma chambers. However, lavas from the western volcanoes with variable Mg# and similar Sr, Nd and Pb isotopic compositions have similar ( $^{230}\text{Th}/^{232}\text{Th}$ ) and ( $^{230}\text{Th}/^{238}\text{U}$ ) ratios (Figure 9). For example, the Cerro Azul lavas range from Mg# 0.52 to 0.28, but have almost constant ( $^{230}\text{Th}/^{232}\text{Th}$ )  $\approx 1.04$ . Moreover, Fernandina, Cerro Azul and Ecuador, who share very similar Sr, Nd and Pb isotopic ratios (White et al., 1993; Kurz and Geist, 1999) show a large variation in Mg# but are similar in Th isotopes, suggesting no significant effects due to magma chamber residence time. Therefore, although this process could have produced slight

variation in the Th isotopes, for example in some of the lavas from Fernandina (Figure 9), the effect of magma chamber residence time on Th isotopes appears to be insignificant.

$(^{230}\text{Th}/^{232}\text{Th})$  versus  $(^{238}\text{U}/^{232}\text{Th})$  activity ratios for Galapagos lavas are displayed in Figure 10a. The two most important observations are:

a) The measured  $(^{230}\text{Th}/^{238}\text{U})$  ratios in the samples is much larger than 1, indicating significant disequilibrium with  $^{230}\text{Th}$  enriched relative to  $^{238}\text{U}$ . The extent of  $(^{230}\text{Th}/^{238}\text{U})$  disequilibrium observed in the whole suite of Galapagos samples is surprising: from  $\approx 20$  up to  $\approx 80\%$ . Compared to a compilation of other OIBs (Lundstrom, personal communication), Galapagos samples extend to extreme disequilibrium, especially when compared to Hawaii (Figure 10b).

b)  $(^{230}\text{Th}/^{232}\text{Th})$  ratios correlate positively with  $(^{238}\text{U}/^{232}\text{Th})$ . Most of the age corrected samples, except those from Floreana and San Cristobal, plot on a quasi linear trend extending toward higher  $(^{230}\text{Th}/^{232}\text{Th})$  and higher  $(^{238}\text{U}/^{232}\text{Th})$ , and overlapping with the field for analyzed MORBs (Figure 10). Moreover, samples analyzed from the GSC (Williams and Gill, 1989) and the EPR at 9°N (Goldstein et al., 1993; Volpe and Goldstein, 1993) appear to plot on an extension of the Galapagos trend. Floreana and San Cristobal lavas plot slightly above the linear array, showing a less well defined trend toward the GSC data. Sierra Negra and Cerro Azul samples show a variation of up to 12% in  $(^{238}\text{U}/^{232}\text{Th})$  at almost constant  $(^{230}\text{Th}/^{232}\text{Th})$  for each suite. This variation in  $(^{238}\text{U}/^{232}\text{Th})$  produces a weak cross-deviation from the linear trend defined by the rest of the Galapagos samples.

$(^{230}\text{Th}/^{232}\text{Th})$  correlates with Nd isotope ratios (Figure 11a) and Sr isotope ratios (not shown). Although few analyses exist for both Pb and Th isotopes in the same samples, the isotope ratios average by volcano also show a good correlation between Th and Sr, Nd and Pb. In contrast,  $^3\text{He}/^4\text{He}$  ratios do not correlate with  $(^{230}\text{Th}/^{232}\text{Th})$ .

Th isotope ratios correlate with trace element ratios such as Nb/Zr, K/Nb, K/Rb, and Sr/Sr\* (for lavas that did not fractionate plagioclase, Sr/Sr\*>1) (Figure 11b). In general  $(^{230}\text{Th}/^{238}\text{U})$  ratios do not correlate with either trace element contents, ratios, or any indicator of extent of melting, with the exception of Ba/Nb (Figure 11c), Ba/La, Ba/Ce ratios. This correlation is mainly defined by samples from Floreana and San Cristobal, while most of the western volcanoes plot very close to each other.



## 4.6 Discussion

### 4.6.1 Extent of melting and pressure of melt segregation.

Previous work on Galapagos lavas has suggested that the variation in trace and major elements can be used to infer the extent of melting and the pressure of melt segregation (Geist, 1992, White et al, 1993). These observations are in direct contradiction with several results found in this work. Namely, the correlations between Nb/Zr ratios and Sr, Nd and Pb isotopes are difficult to explain solely by different extents of melting (Figure 8). Kurz and Geist, (1999) found a correlation between Nb/La and Nd isotopes, and stated that the variation in Nb/La ratios was a result of source heterogeneity. In this case, it can be argued that the extent of melting was sufficiently high and that Nb and La bulk partition coefficient ( $D_{Nb}$ ,  $D_{La}$ ) are low and close enough for the lavas to have Nb/La ratios similar to that of their source (assuming simple batch melting process). In contrast, the  $D_{Nb}$  and  $D_{Zr}$  are not that similar, especially during melting in the garnet stability field, making it difficult to use the previous reasoning regarding the correlation between isotopes and Nb/Zr ratios. The difference between Nb/La and Nb/Zr can be seen in Figure 12. The samples define two trends: one is steep and contains most of the lavas, the other is more shallow and defined mainly by the Floreana samples. These two trends are compared with two simple batch melting models for garnet and spinel peridotite. However, the correlation between isotopes, Nb/La and Nb/Zr, except for Floreana, suggests that most of the trace element variations are controlled by either mixing of melts or mixing of mantle sources. Also, this hypothesis is consistent with the observed correlation between  $La/Sm_{PM,8}$  and Sr, Nd and Pb isotopic ratios averaged by volcano (Geist, 1992; White et al., 1993). The hypothesis of mixing of mantle sources requires an extra step: lavas with progressively more enriched isotopes (high Sr, Pb, and low Nd isotopic ratios) have higher Nb/La and Nb/Zr, indicating a lower extent of melting for a more enriched source. Therefore, the correlation between Nb/Zr ratios and isotopes (Figure 8) requires the extent of melting to be inversely proportional to the fertility of the source.

The model for melting in the garnet stability field seems to reproduce the Floreana trend. The similarity between the Floreana trend and the melting model brings us to a second of the contradiction with previous papers. Specifically, the flat MREE to HREE pattern in the Floreana lavas (Figure 4b) has been traditionally used as an indication of melting in the spinel stability field (White et al., 1993). As we mentioned before, the low MREE/HREE ratios appear to be related more to the depletion of MREE than to changes in

the HREE (Figure 5). Floreana basalts have similar LREE/HREE ratios to lavas from Fernandina (Figure 4b and 5); this result is unexpected, assuming that the two suites of lavas were segregated from different depths (spinel and garnet stability field, respectively), and have undergone different extents of melting. The concave upward REE pattern of Floreana lavas, with enriched LREE and depleted MREE and HREE at relatively low total REE contents, is not easy to reproduce by simple melting models. However, a reasonable way to reproduce the pattern is by mixing between magmas produced by low extent of melting from a fertile source, segregated from the garnet stability field, and magmas produced by moderate extents of melting from a depleted source, segregated from the spinel stability field (Figure 13a). The high trace element content of the magma produced by a low extent of melting controls the budget of the most incompatible elements, and therefore dominates the isotopic composition of the mixture. Thus, melting processes in Floreana basalts seem to have displaced some of the incompatible trace elements ratios (e.g., Nb/Zr) away from the overall mixing trend (Figure 8 and 12).

The flat REE patterns observed in lavas from San Cristobal and Santa Cruz can be explained in the same manner as the Floreana basalts (Figure 13b). These simple models illustrate that although the HREE may not reveal any indication of garnet, the mixture encodes the garnet signature in the most incompatible elements such as U/Th ratios. In contrast to the lavas from the eastern islands, all the western volcanoes show steep REE pattern suggesting melting in the garnet stability field (White et al., 1993). Basalts from Wolf and Darwin volcanoes have high MREE/HREE ratios but low LREE/HREE ratios (Figure 4a). This tendency can be explained either by low extent of melting from a previously depleted mantle source or by magma mixing. In any case, the steep REE pattern is an indication of melting in the garnet stability field. Either hypothesis is consistent with the depleted isotopic composition observed in these lavas.

The high ( $^{230}\text{Th}/^{238}\text{U}$ ) disequilibrium found in Galapagos lavas (greater than 1 for all the samples analyzed) indicates fractionation of Th from U during magmatic processes (Figure 10a). Based on experimental data, the observed Th-U fractionation in Galapagos melts can be produced by melting in the presence of garnet (Beattie, 1993; LaTourrette et al., 1993, Hauri et al., 1994; Salters and Longhi, 1999). This is in agreement with our previous interpretation of the REE patterns and trace element ratios. Especially in the case of Floreana and San Cristobal lavas, which have flat REE patterns, the ( $^{230}\text{Th}/^{238}\text{U}$ ) ratios

are much greater than 1 ( $\approx 1.8$  and  $\approx 1.4$  respectively) reinforcing the hypothesis of mixing between magmas derived from both the garnet and spinel stability fields.

In the Th-U isochron diagram, Galapagos lavas, with the exception of Florena and San Cristobal, define a quasi linear trend between MORB (GSC-EPR 9°N) at one end, and Pinta at the other extreme (Figure 10 a and b). The positive correlation between ( $^{230}\text{Th}/^{232}\text{Th}$ ) and ( $^{238}\text{U}/^{232}\text{Th}$ ) can be accounted for in the same way as the Nb/La-Nb/Zr trend and the REE; namely, by mixing of magmas generated from a fertile and a depleted source or by mixing of these mantle sources. As before, the second case would require an inverse correlation between the extent of melting and the fertility of the source. Moreover, it would also require that the effect of melt transport on ( $^{230}\text{Th}/^{232}\text{Th}$ ) was either insignificant, or equivalent, for all the islands. Thus, the hypothesis of mixing mantle sources requires more restricted conditions than magma mixing. The correlation between ( $^{230}\text{Th}/^{232}\text{Th}$ ) and Sr and Nd isotope ratios (Figure 11a) or Sr, Nd and Pb averaged by volcano (not shown) obtained from White et al., (1993) suggests that the Th isotope composition represents either the Th isotope composition of the lavas produced by mixing, or the isotopic composition of their mantle sources.

( $^{230}\text{Th}/^{238}\text{U}$ ) does not correlate with any geochemical indicator of the extent of melting. The lack of correlation suggests that the variation in the extent of disequilibrium is controlled by a combination of melting, melt transport and mixing processes. Although Ba/La, Ba/Ce and Ba/Nb ratios correlate with ( $^{230}\text{Th}/^{238}\text{U}$ ), these ratios also define positive trends with Sr and Pb isotopes suggesting that the variation in the ratios is produced mainly by mixing.

To estimate the extent of melting for Galapagos lavas using U decay series isotopes is not an easy task. This is mainly because ( $^{230}\text{Th}/^{232}\text{Th}$ ) could indicate the isotopic ratio in the mantle source, the conditions of melt transport (Elliot, 1997), or mixing processes (Lundstrom et al., 1998). Samples from Cerro Azul have a variation in  $\text{La}/\text{Yb}_{\text{PM}}$  (from 2.66 to 5.98) and in ( $^{230}\text{Th}/^{238}\text{U}$ ) (from 1.240 to 1.374) at similar ( $^{230}\text{Th}/^{232}\text{Th}$ ) and similar Sr, Nd and Pb isotope ratios. The simplest interpretation of the data is that in Cerro Azul, the ( $^{230}\text{Th}/^{238}\text{U}$ ) of the lavas is mainly controlled by the extent of melting of the mantle source. More data will be necessary to evaluate this hypothesis. Although it would be premature to estimate the extent of melting, given the complexities of the data, it is clear from the measured ( $^{230}\text{Th}/^{238}\text{U}$ ) disequilibrium that the mantle sources with higher Sr and

Pb and lower Nd isotope ratios represented by Floreana and Pinta lavas have melted to a much lower extent than the MORB-like sources supplying the Wolf and Darwin volcanoes. Regardless of the assumptions made or melting model used, disequilibrium of 50 to 80% indicate that the extent of melting for Pinta and Floreana has been small (0.5 to 2%) compare to that for Wolf and Darwin volcanoes (5 to 7%). Moreover, all analyzed Galapagos lavas seem to have melted, or have a component (in the mixture) that melted, in the garnet stability field.

#### 4.6.2 Seeing through melting and the mixing, to source characteristics.

Giving the complexity of Galapagos lavas it is difficult to make firm statements about the mantle source. However, trace element ratios seems to show certain features that are not explained either by variable extents of melting or by mixing. For example, the correlation between Nb/La ratios and He and Nd isotopes suggest that the plume mantle source have a positive anomaly in Nb and Ta or that the mantle melts in the presence of mineral phases that would fractionate Nb from elements with similar compatibility (Kurz and Geist, 1999, Fitton, et al., 1998). In addition to Nb/La, also Ti/Ti\*, K/Rb correlate with isotopes, even with  $^3\text{He}/^4\text{He}$  and ( $^{230}\text{Th}/^{232}\text{Th}$ ) (Figure 7 and 11b). If high  $^3\text{He}/^4\text{He}$  ratios are a good indicator of the plume source, it is then unavoidable to suggest that the plume mantle has positive anomalies of Nb and Ti and negative anomalies of K. Note that we used Nb/La and K/Rb as approximations to the K and Nb anomalies, because both elements are situated next to each other in the normalized plots the measured Nb/Nb\* and K/K\* are influenced by the K and Nb contents respectively. We also detected negative Y anomalies, but they do not correlate as well with the isotopes as do the other anomalies. These various trace element anomalies appear to decrease toward the eastern islands where the trace element patterns are flatter (low MREE/HREE, low Hf/Lu) and the isotopic ratios become more depleted, suggesting an increase in the proportion of MORB (asthenospheric) component in the mixture (Figure 6). Finally, all the lavas, except those that have fractionated plagioclase, have Sr/Sr\* higher than 1. In contrast to the K, Nb, Y and Ti anomalies, the Sr/Sr\* anomalies are higher in lavas from the eastern islands. The Sr anomaly does not correlate with isotopes, most probably because the lavas in the western volcanoes have fractionated plagioclase, which decreases this anomaly. For Sr/Sr\*>1, the anomaly does not correlate with Eu/Eu\* or CaO/Al<sub>2</sub>O<sub>3</sub> ratios suggesting that the anomaly is not produced by plagioclase accumulation, but is probably related to the mantle source composition.

As before, Floreana seems to be an exception, sharing characteristics from both the western volcanoes, (high  $^3\text{He}/^4\text{He}$  ratios, positive anomaly of Ti), and the eastern islands (flat REE, small to nonexistent anomalies of Y and K, and a large Sr anomaly). Moreover, Floreana lavas are enriched in Ba and depleted in Th and U relative to elements of similar incompatibility (Figure 4a). We believe the dichotomy observed in Floreana is produced by mixing magmas representing very low degree melts from an enriched source, and moderate degree melts from a depleted source. Thus, in Floreana many of the characteristics ascribed to either asthenospheric or plume mantle sources are clouded by melting processes, which seem to have caused the ratios of incompatible trace element (Nb/Zr, Th/U) to diverge from the mixing trends.

The trace element anomalies observed in Galapagos lavas are also common in Hawaii and Iceland (Fitton, et al 1998, Hofmann, 1998 and Sobolev et al, 1998), all known areas with high  $^3\text{He}/^4\text{He}$  ratios. The origin of the Nb, Ti, Sr, K and Y anomalies in the Galapagos mantle plume is controversial. However, the enrichment of HFSE and Sr, the depletion in K, and the correlation with isotopes, especially  $^{206}\text{Pb}/^{204}\text{Pb}$  ratios (up to 19.5 in volcanoes from Isabela) are consistent with the model that invokes tapping of recycled oceanic crust (Hofmann, 1982; Hofmann and White, 1980). Even the characteristics of Floreana (Ti, Nb and Sr enrichments, Th-U depletions, and  $^{206}\text{Pb}/^{204}\text{Pb}$  ratios up to 20) are consistent with the hypothesis of recycled oceanic crust gabbros proposed by Sobolev et al. (1998) and Hofmann, (1998) for Hawaii and Iceland. If this is the case, the correlation between the trace element anomalies (index of ocean crust recycling) and the  $^3\text{He}/^4\text{He}$  ratios (indicator of undegassed mantle) is puzzling. One simple and circumstantial explanation could be the thermal entrainment of undegassed mantle in a plume formed by recycled material (Hart et al., 1992). On the other hand, minerals such as Ca and Mg-Perovskites could fractionate HFSE and K from other incompatible elements (Kato et al., 1988, Shimizu pers. comm., 1999). Thus, a second hypothesis is that the correlation between the trace element anomalies (index of Ca-Perovskite fractionation) and  $^3\text{He}/^4\text{He}$  ratios (undegassed mantle) could be a simple indication of the early Earth differentiation. However, the near primitive mantle Ba/U ratios for lavas from Fernandina islands, which have the highest  $^3\text{He}/^4\text{He}$  ratios, are inconsistent with both the above hypotheses. The relative uniformity and near primitive mantle values of ratios such as Ce/U and to lesser extent Ba/U, Ba/Ce in OIBs compared to MORB is surprising and challenges both the recycling and the early Earth differentiation hypothesis (Halliday et al,

1995). Finally, more work is required to distinguish the effects of recycling from those of local assimilation of oceanic crust. Local assimilation could be an important source of geochemical variation in Galapagos lavas, where the oceanic crust platform ranges up to 18 km in thickness (Feighner and Richards, 1994). Nevertheless, the high  $^{206}\text{Pb}/^{204}\text{Pb}$  (up to 20, White et al., 1993) and the high  $^3\text{He}/^4\text{He}$  isotopes, (up to 30, Kurz and Geist, 1999) are not easily explained by interaction of Galapagos magmas with a young lithosphere (<10 Ma).

#### 4.6.3 Mantle porosity and upwelling velocities

In the last 10 years, a series of elaborate models were developed to explain the effects of melting and melt transport on U decay series isotopes. There are two end-member models: dynamic melting and equilibrium porous flow. Both models consider that melting processes in the presence of garnet is responsible for the Th-U fractionation producing lavas with higher Th/U ratios than those of their mantle source (Beattie, 1993; LaTourrette et al., 1993, Hauri et al., 1994; Salters and Longhi, 1999).

Dynamic melting was applied to U series isotopes by McKenzie (1985) and modified by Williams and Gill (1989). It considers that the melting process can be described by accumulated fractional melting with a threshold porosity (Shaw, 1970). In this model, the melt is extracted continuously and accumulated in chemical isolation from the solid residue, with the result that Th and U are extracted from the solid very early during the melting.

The second model, equilibrium porous flow (EPF), was proposed by Spiegelman and Elliott (1993). It postulates that as the melt migrates, it reacts with the solid residue, reaching equilibrium at each step. This process allows the U to remain in the solid matrix for a longer time, producing greater disequilibrium than in the dynamic melting model.  $^{230}\text{Th}$  cannot remain in the melting columns longer than about 6 half-lives. Therefore, a slow effective velocity in a long melting column will erase the extent of disequilibrium due to the fast decay of  $^{230}\text{Th}$  relative  $^{238}\text{U}$ . The beauty of EPF is that the  $^{230}\text{Th}$  excesses are controlled both by melting and by differential (melt/matrix) transport. At the bottom of the column, for degrees of melting comparable to the bulk partition coefficient (D), the initial excesses are large, principally controlled by the chemical fractionation during melting. As the extent of melting gets larger, this effect disappears and the transport processes dominate the ingrowth. As long as the  $^{238}\text{U}$  parent isotope continues to spend more time in the

column than its daughter, the excess will persist even to high extent of melting (Spiegelman and Elliott, 1993). Thus, in contrast to dynamic melting, in EPF the length of the melting column influences the resulting  $^{230}\text{Th}$  excesses. At slow upwelling rates and faster effective velocity for Th than U, long melting columns will produce larger  $^{230}\text{Th}$  excesses than shorter columns. At present, there is no clear evidence that would allow us to choose between dynamic or equilibrium porous flow models. The only reason we consider dynamic melting more realistic is that there is no evidence in major and trace element data for basalt and peridotite that support the hypothesis that the melt equilibrate with the surrounding mantle matrix as proposed by (Spiegelman and Elliot, 1993).

We have applied both the dynamic melting and equilibrium porous flow models to the Galapagos plume. The main objective is to derive information on mantle porosity and upwelling rates. Both models should be considered approximations, due to the uncertainties in the mineral modes, partition coefficients, mantle compositions, melt productivity, etc. We used two sets of partition coefficient (Beattie, 1993; LaTourrette et al., 1993). The bulk partition coefficients (D) define the maximum obtainable disequilibrium for any model, which is the ratio between the parent and daughter bulk partition coefficient. Moreover, the Ds control the effective velocity of the element and the rate at which the element is transferred from the solid to the melt.

In general, past U series work on magmatic processes has been done using garnet/clinopyroxene ratios for the mantle source that appear to be high (0.12/0.08, Beattie, 1993) compared to ratios obtained from melting experiments in the garnet stability field (0.06/0.25, Walter, 1998). This observation constrains either the source composition (peridotite versus eclogite-pyroxenite) or the partition coefficients to be used. For example, to obtain 80% disequilibrium as observed in the Floreana lavas, assuming Walter's mineral modes, we are forced to use Beattie, (1993) partition coefficients. Any other partition coefficient set will not reproduce the observed disequilibrium [ $^{230}\text{Th}/^{238}\text{U}$ ] = 1.8] because the ratio between the parent/daughter bulk partition coefficient set is smaller than 1.8. On the other hand, if we use LaTourrette's partition coefficients, the observed disequilibrium would require the existence of eclogite (garnet/clinopyroxene  $\approx$  0.5/0.5) in the melting column.

For these exercises, we have assumed the generally used garnet/clinopyroxene = 0.12/0.08 ratios (Beattie, 1993). Both models applied here utilize modal, one-stage

melting, that is, melting occurs only in the garnet stability field (Figure 14, 15). Moreover, in the EPF model  $^{230}\text{Th}$  excesses are dependent on the length of the column. Therefore, we show the results using melting column lengths varying from 0.5 to 10% melting for Floreana, Pinta, and Fernandina respectively (Figure 15). In contrast, if we utilize non-modal melting models, the source mineral modes of Walter (1998), and consider two stage melting models, in the garnet and spinel field, to obtain the same disequilibrium the output of the model will result in a smaller mantle porosity and slower upwelling rates. Moreover, the EPF model will require longer melting columns, starting at greater depth, and will involve larger extent of melting. Thus, the results obtained here are an upper bound for the mantle porosity and upwelling velocities and a minimum limit for the length of the melting column in the Galapagos plume.

Figure 14 and 15 show contours of constant  $^{230}\text{Th}$  excess as a function of the mantle porosities and upwelling velocities for both Dynamic melting and EPF. Each contour represent the maximum  $^{230}\text{Th}$  excess found in each volcano or island. The first important observation from Figure 14 is that lavas from Darwin and Wolf volcanoes, or even from Santiago island, give faster apparent mantle upwelling velocities and larger porosities than Fernandina which is the assumed center of the hot-spot. The lower  $^{230}\text{Th}$  excess in those lavas compared with Fernandina basalts is mainly produced by the larger MORB component involved in the mixing for Wolf and Darwin volcanoes. Therefore, from these results we can argue that the smaller  $^{230}\text{Th}$  excesses found in lavas from Darwin and Wolf indicate:

- 1) the MORB-like mantle below Darwin and Wolf volcanoes upwell faster than the center of the plume (Fernandina). This hypothesis is unlikely, because the center of the plume is considered to be the hottest and therefore should upwell with the fastest velocities.

- 2) The mantle below Darwin and Wolf has higher porosity than that of Fernandina. Again this hypothesis seems unlikely; because the center of the plume (Fernandina) should be the hottest, and lavas from Fernandina indicate a more fertile mantle source (higher Sr, Pb and lower Nd isotope ratios) than those from Darwin and Wolf, we would expect higher extent of melting, and therefore higher porosity in the mantle source of Fernandina than in those of Darwin and Wolf.



3) The lower  $^{230}\text{Th}$  excess in lavas from Darwin and Wolf volcanoes than those of Fernandina are produced by mixing between lavas originated from the plume mantle with lavas from the MORB mantle. Because we do not know the end-member that represent the MORB component we could argue that the MORB end member was in secular equilibrium. Therefore, the lower  $^{230}\text{Th}$  excess could be the result of simple mixing with no indication of mantle porosity or upwelling velocities. It cannot be unambiguously argued that the  $^{230}\text{Th}$  excess in Darwin and Wolf lavas indicates faster upwelling velocities or larger porosities than that for Fernandina.

Thus, we will consider here only Fernandina as the best example to establish the melting conditions in the plume. We choose Fernandina as representative of the plume because its lavas have the highest  $^3\text{He}/^4\text{He}$ , indicative of its proximity to the present position of the plume (Kurz and Geist, 1999). We will also consider the melting conditions for Pinta and Floreana samples, because they exemplify lavas erupted on the thin lithosphere, which represent the lowest extents of melting, and the smallest contamination with MORB in the archipelago (Figure 10).

To reproduce the measured ( $^{230}\text{Th}/^{238}\text{U}$ ) in Fernandina, Pinta and Floreana dynamic melting will require mantle upwelling velocities of approximately  $8.5 \times 10^{-3}$ ,  $4.5 \times 10^{-3}$ ,  $2 \times 10^{-3}$  m/y and porosities of  $3 \times 10^{-3}$ ,  $2 \times 10^{-3}$ ,  $1 \times 10^{-3}$ , respectively. In contrast, the EPF allows faster upwelling mantle velocities than dynamic melting (approximately  $3 \times 10^{-2}$ ,  $1 \times 10^{-2}$ ,  $5 \times 10^{-3}$  m/y and similar porosities of  $3 \times 10^{-3}$ ,  $1.5 \times 10^{-3}$ ,  $1 \times 10^{-3}$  respectively). Two important points can be obtained from this exercise:

1) The mantle upwelling velocity for the Galapagos plume (Fernandina) is extremely slow ( $\leq 1$  to  $3$  cm/y), compared to the tholeiites from Hawaii (40 to 100 cm/y, Sims et al., 1999, using same bulk Ds and modal proportions).

2) Very slow mantle upwelling for Pinta (0.5 to 1 cm/y) and Floreana (0.2 to 0.5 cm/year) indicates that the plume is dragged approximately toward the east almost horizontally after impinging in the lithosphere, thus limiting the decompression and producing low extent of melting.

Ito et al. (1997), presented a three-dimensional numerical model of hot-spot ridge interaction for the Galapagos and GSC, with variable viscosities. The model match the residual bathymetry and mantle Bouguer anomaly along isochrons. The amplitude of the

gravity anomalies can be explained by a plume source 100 to 200 km in radius and a temperature anomaly of 100 to 50 °C respectively. Thus, the plume is only mildly buoyant (buoyancy flux of  $2 \times 10^3 \text{ kg s}^{-1}$ ). Considering the case of a plume source 200 km in radius and a temperature anomaly of 50 °C, the plume volume flux and the mantle upwelling velocities modeled by Ito et al. are  $\approx 9.1 \times 10^6 \text{ Km}^3/\text{Ma}$ . and  $\approx 7 \text{ cm/y}$  respectively. Thus, the mantle upwelling velocities obtained from modeling the residual bathymetry and mantle bouger anomaly are slightly higher than the upwelling mantle velocities of  $\approx 3 \text{ cm/y}$  estimated from ( $^{230}\text{Th}/^{238}\text{U}$ ) disequilibrium for the center of the hotspot (Fernandina). Both studies (the numerical model and the  $^{230}\text{Th}$  excesses) indicate that the Galapagos plume is quite weak, with a plume temperature anomaly of the order of 100 to 50°C. The agreement between the numerical model for Galapagos by Ito et al.,(1997) and our work here confirms the correlation observed by Chabaux and Allègre, (1994) between  $^{230}\text{Th}$  excesses and buoyancy flux.

Many of the characteristics observed in the Galapagos islands can be explained by different aspects of the interaction between the plume, the GSC and the Feighner and Richards' fault (90°50' transform fault?). Bathymetric and gravimetric data indicate that this fault divides the archipelago into two sectors: the west sector with an old, thick and cold lithosphere ( $\approx 10 \text{ Ma}$ , 12 km elastic thickens) and an eastern sector with a young, thin and warm lithosphere ( $\approx 5 \text{ Ma}$ , 6 km elastic thickness) (Feighner and Richards, 1994). The plume impinges under the western sector, where the sloping lithosphere forms a rheological boundary layer (RBL) that thins toward the ridge. The plume is dragged by the moving plate toward the SE, into the eastern sector, where the lithosphere is thin and relatively flat. Fluid dynamics laboratory experiments found a correlation between the magnitude of the RBL slope, and the proportion of material that the plume feeds to the nearby ridge (Kincaid et al, 1995). The three-dimensional variable-viscosity experiment indicates that the greater the RBL slope, the more effectively it diverts the buoyant low viscosity plume material toward the ridge. In the western sector, the plume is warmer and less viscous, and the lithosphere has a more pronounced slope than in the east. To the east of the fault, the lower thermal difference between the flatter lithosphere and the colder plume facilitates the dragging of the plume away from the GSC. Although the effect of the lithosphere on the plume may be small because the Galapagos is a mildly buoyant plume (Ito et al., 1997), the ability of the lithosphere to divert the plume toward the ridge will be larger in the western sector than in the eastern sector. Thus, this different behavior on

opposite sides of the fault can explain why the geochemical anomaly does not occur at its closest position to the Galapagos archipelago, but to the northwest, at the intersection of the Darwin-Wolf lineament and the GSC. Moreover, the lower thermal difference between the flatter lithosphere and the colder plume in the eastern sector facilitates the dragging of the plume away from the GSC and allows the upwelling MORB material to fill the central area of the Archipelago producing the “horse shoe pattern”.

After impinging under the western sector, the plume cools and is dragged toward the east, moving across the fault. The west to east difference in thickness of the lithosphere ( $\approx 6$  km) allows the relatively cold plume, that was moving almost horizontally, to move upward, producing decompression and a low extent of melting (approximately 2%, assuming melting rate of 0.3%/km and a 6 km distance). As the dragged plume moves to the SE, the central part of the archipelago will be underlain largely by thermally entrained asthenospheric mantle, as postulated by Richard and Griffiths (1989). As with the plume, when the asthenosphere crosses the fault, decompression and melting occurs. Thus, the effect of the transform fault and the abrupt thinning of the lithosphere may account for the broad distribution of the volcanoes and the long periods of eruption in the eastern sector. Moreover, this process can explain the negative Bouguer anomaly observed in the center of the archipelago (Case et al., 1973, Feighner and Richards, 1994). The effect of transform faults on melting was proposed by Reynolds and Langmuir (1997) to explain the trace element composition of MORB erupting close to the intersection of the MARK area with the Kane transform fault. We believe that consideration of variations in lithospheric thickness and transform fault effects are of utmost importance in understanding Galapagos magmatism.

#### ***4.7 Conclusions***

The correlation between isotopes and trace element ratios (such as Nb/La, Nb/Zr, Ba/La, K/Rb and Ba/Ce) suggests that most of the trace element variations in the Galapagos are controlled by mixing of magmas from at least two different mantle sources. The extent of melting represented by the extreme components of the mixing seems to be inversely proportional to the fertility of the sources. This simple model illustrates that such a mixture carries a garnet signature in the most incompatible elements (such as U, Th, Nb, Zr) even though the variation of the MREE to HREE does not reveal any indication of melting in the garnet stability field.

The high ( $^{230}\text{Th}/^{238}\text{U}$ ) disequilibrium found in Galapagos (ranging from 1.16 to 1.81), indicates that melting occurs in the garnet stability field in agreement with the trace element results. The positive correlation between ( $^{230}\text{Th}/^{232}\text{Th}$ ) and ( $^{238}\text{U}/^{232}\text{Th}$ ), with  $^{230}\text{Th}$  excesses decreasing toward higher ( $^{230}\text{Th}/^{232}\text{Th}$ ) ratios, can be accounted for by mixing of magmas generated by low extents of melting from a fertile source with magmas produced by relatively large extents of melting from a depleted source. The correlations between ( $^{230}\text{Th}/^{232}\text{Th}$ ) and Sr, Nd and Pb isotopes confirm that the Th isotopic ratios are mainly source-controlled and are the result of magma mixing between the plume and asthenosphere melts.

Correlation between He isotopes and Ti/Ti\*, K/Rb and Nb/La ratios suggests that the mantle plume has positive anomalies of Nb and Ti and negative anomalies of K. The decrease of these anomalies toward the eastern islands where the trace element pattern are flatter and the isotopic ratios become more depleted, suggests that the increase of the proportion of a MORB component in the mixture tend to erase these anomalies. The origin of the Nb, Ti, Sr, K and Y anomalies in the mantle plume of Galapagos is still a matter of debate.

Dynamic melting and equilibrium porous flow models constrain mantle upwelling velocities for the Galapagos plume (Fernandina) to  $\leq 3$  cm/y and maximum porosities of 0.3%, indicating that Galapagos is a mildly buoyant plume; this is consistent with a 3D numerical model of plume-ridge interaction. Evidence for very slow mantle upwelling and low porosity for Pinta (0.5 to 1 cm/y and 0.15%) and Floreana (0.2 to 0.5 cm/y and 0.1%) supports the hypothesis that the lateral transport of the plume across the 90° 50' transform fault produced additional slow upwelling and further small extent of melting.

Dynamic melting and equilibrium porous flow models may give erroneous estimations of mantle porosity and upwelling rates when the suite of lavas has undergone mixing processes.

## 4.8 References

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#### **4.9 Table Caption**

**Table 1:** Representative majors and trace elements and isotope analyses from the published and unpublished data set of Kurz. Pb isotopes are reported relative to value for NBS 981  $^{206}\text{Pb}/^{204}\text{Pb} = 16.9356$ ;  $^{207}\text{Pb}/^{204}\text{Pb} = 15.4891$  and  $^{208}\text{Pb}/^{204}\text{Pb} = 36.7006$ . Nd isotopes were corrected for fractionation to  $^{146}\text{Nd}/^{144}\text{Nd} = 0.7219$  and normalized to La Jolla value of 0.51184. Sr values are normalized to NBS 987= 0.710240. Typical uncertainties for Sr and Nd are approximately 18 ppm. Uncertainties for Pb are 1, 1.4 and 1.6 per mil for  $^{206}\text{Pb}/^{204}\text{Pb}$ ,  $^{207}\text{Pb}/^{204}\text{Pb}$ ,  $^{208}\text{Pb}/^{204}\text{Pb}$  respectively.

**Table 2:** Th and U concentration determined by isotopic dilution and Th isotopes for Galapagos lavas. Ages determined by surface exposure dating techniques by Kurz. The Th isotopes have been corrected by the reported ages, assuming the ages represent the time of eruption. Parenthesis denote activity ratios;  $\lambda_{230} = 9.195 \times 10^{-6} \text{ yr}^{-1}$ ,  $\lambda_{232} = 4.948 \times 10^{-11} \text{ yr}^{-1}$ ,  $\lambda_{238} = 1.551 \times 10^{-10} \text{ yr}^{-1}$ .

#### **4.10 Figure Caption**

**Figure 1:** Map of the Galapagos Archipelago. GSC is the Galapagos Spreading Center. The dash line represents the inferred fault from Feighner and Richards (1994). The fault separates thick (12 km elastic thickness), old ( $\approx 10$  Ma) and strong lithosphere to the west from a thin (6 km elastic thickness), young ( $\approx 5$ Ma) and weak lithosphere to the east. The inset shows a regional view of the GSC.

**Figure 2:** a) Th isotope reproducibility for the UCSC standard measured by IMS 1270 ion probe. The TIMS value is reported for comparison. b) ICP-MS reproducibility for U isotope ratios for the NBS 960 standard measured, after correction for mass fractionation to the certified value of 137.88.

**Figure 3 :** a)  $\text{SiO}_2$  versus  $\text{Na}_2\text{O}+\text{K}_2\text{O}$  according to the classification of Le Bas et al., (1986). All of the samples plot in the field for basalts. Light dash-line separates the fields for alkaline and tholeiitic basalts (Macdonald and Katsura, 1964). Several samples (two from Floreana and two from Cerro Azul) are alkali basalts. b)  $\text{TiO}_2$  versus Mg# for Galapagos lavas. Most of the lavas have undergone significant crystal fractionation. However, this differentiation did not reach saturation with ilmenite-magnetite in the analyzed samples.

**Figure 4 :** a) Primitive mantle-normalized diagram for representative samples from the western (Fernandina and Darwin volcanoes) and eastern (Floreana and San Cristobal island) sector of the Galapagos Archipelago. Nb, Ti, K and Y anomalies are larger in lavas from the western volcanoes, while the Sr anomaly is larger in the lavas from the eastern islands. b) Primitive mantle-normalized REE. Note the flattening of the MREE to HREE in lavas from the eastern sector. Normalizing values are from Sun and McDonough,(1989), except for Pb and Cs 0.15 and 0.021 respectively.

**Figure 5 :** a) Sm versus Yb contents (ppm). The lavas from the eastern islands have lower Sm/Yb ratios than those from the western volcanoes. Note that the variation in Sm/Yb ratios is mainly controlled by the variation in Sm content. The lines defines constant values of Sm/Yb ratios b)  $(La/Yb)_{PM}$  versus  $(Sm/Yb)_{PM}$  ratios. Most of the lavas from the eastern islands have lower  $(La/Yb)_{PM}$  and  $(Sm/Yb)_{PM}$  ratios than those of the western volcanoes, with the exception of the Floreana lavas, which have similar  $(La/Yb)_{PM}$  ratios to the lavas from the western volcanoes. Moreover, note that the Darwin and Wolf volcanoes have lavas with high  $(Sm/Yb)_{PM}$  but low  $(La/Yb)_{PM}$  ratios compared to the other western volcanoes. See text for explanation. The lines define values of constant  $(La/Sm)_{PM}$  ratios.

**Figure 6:**  $Ti/Ti^* = [2Ti_{PM}/(Eu_{PM}+Gd_{PM})]$  and  $Y/Y^* = [2Y_{PM}/(Dy_{PM}+Ho_{PM})]$  versus Hf/Lu ratios. With the exception of Floreana, the anomalies in Ti and Y tend to disappear in lavas from the eastern islands.

**Figure 7:**  $^3He/^4He$  versus K/Rb and  $Ti/Ti^*$ . The correlation between the anomalies of Ti and K (we used K/Rb as an approximation of  $K/K^*$ ; see text) indicates that the mantle source responsible for the high  $^3He/^4He$  is also responsible for the positive Ti and negative K anomalies.

**Figure 8:**  $^{143}Nd/^{144}Nd$ ,  $^{206}Pb/^{204}Pb$  and  $^{87}Sr/^{86}Sr$  versus Nb/Zr. Nb/Zr correlates with isotopes, with the exception of the Floreana alkaline basalts for Sr and Nd isotope ratios, and one sample from San Cristobal that has anomalously high Sr isotopes. Although Nb/Zr ratios do not correlate with Sr and Nd isotopes, they correlate with Pb isotopes suggesting that even in Floreana lavas, the trace element are in some way affected by mixing. The correlation between Nb/Zr and isotopes indicates that the trace elements in Galapagos lavas are controlled mainly by source variations and mixing processes. Reference as in Figure 7

**Figure 9:** ( $^{230}\text{Th}/^{232}\text{Th}$ ) and ( $^{238}\text{U}/^{232}\text{Th}$ ) versus Mg#. Samples having a range in Mg# with similar Sr, Nd and Pb isotopic composition have similar ( $^{230}\text{Th}/^{232}\text{Th}$ ) and ( $^{238}\text{U}/^{232}\text{Th}$ ) ratios; the best example of this is the suite of lavas from Cerro Azul volcano.

**Figure 10:** ( $^{230}\text{Th}/^{232}\text{Th}$ ) versus ( $^{238}\text{U}/^{232}\text{Th}$ ). a) Isochrone diagram for Galapagos lavas. Dashed lines are given for ( $^{230}\text{Th}/^{238}\text{U}$ ) ratios of 1.2, 1.5 and 2.0; the solid line is the equiline. The measured ( $^{230}\text{Th}/^{238}\text{U}$ ) ratios in all of the samples is much higher than 1, indicating melting in the garnet stability field. ( $^{230}\text{Th}/^{232}\text{Th}$ ) correlates roughly with ( $^{238}\text{U}/^{232}\text{Th}$ ), approaching the equiline as the ( $^{230}\text{Th}/^{232}\text{Th}$ ) ratio increases. The samples define two linear trends: one is well defined and contains most of the samples analyzed; the second trend is less well defined with samples from Floreana and San Cristobal. b) Comparison between the Galapagos lavas and a compilation of OIB and MORB data (C. Lundstrom, personal communication). Note that one extreme of the linear trend defined by the Galapagos samples (dashed line) overlaps with the field defined by MORB, especially with samples from the GSC and EPR 9°N, which are geographically the closest MORB samples to the Archipelago. At the other end of the trend, the Galapagos samples show extreme disequilibrium, when compared to other OIB lavas, especially when compared with Hawaii. Analytical uncertainties for Galapagos samples are smaller than the symbols

**Figure 11:** a and b) ( $^{230}\text{Th}/^{232}\text{Th}$ ) versus  $^{143}\text{Nd}/^{144}\text{Nd}$  and K/Rb ratios. The ( $^{230}\text{Th}/^{232}\text{Th}$ ) correlates with Sr Nd and Pb isotopes and with some trace element ratios, suggesting that source variations and mixing processes have also controlled the Th isotopic composition of Galapagos samples. c) ( $^{230}\text{Th}/^{238}\text{U}$ ) versus Ba/Nb. This correlation is mainly defined by samples from Floreana and San Cristobal. Although Ba/La, Ba/Ce and Ba/Nb ratios correlate with ( $^{230}\text{Th}/^{238}\text{U}$ ), these ratios also define positive trends with Sr and Pb isotope ratios, suggesting that the variation in the ratios is produced mainly by source variation and mixing. References as in Figure 10

**Figure 12:** Nb/La versus Nb/Zr ratios. With the exception of two alkali basalts from Floreana, the Galapagos lavas define a steep trend. Dashed-lines represent results of non-modal batch melting models in the garnet and the spinel stability field, the number next to the (+) symbols represents the extent of melting in %. We assumed an enriched source that melts in the garnet stability field, from an initial Nb/La = 1.3, Nb/Zr = 0.07. Note that the Nb/La and Nb/Zr ratios are slightly more enriched than the PM  $\approx$  1.02 and 0.063 respectively, because we assumed a source with a positive anomaly in Nb. We assumed a

depleted MORB source with  $Nb/La = 0.6$  and  $Nb/Zr = 0.03$ , with melting in the spinel stability field. The mineral mode for the enriched garnet source is taken from Beattie (1993), and the melting reaction is from Walter (1998), for 3 GPa. For the depleted source we used the mineral mode and melting reaction from Kinzler and Grove (1992). Most of the Galapagos lavas plot in the area defined by mixing between lavas representing a low extent of melting from an enriched source (in the presence of garnet) and a moderate extent of melting from a depleted source (in the spinel field). The partition coefficients are from a compilation made by Takazawa (1996) for La and Zr, and from Kelemen et al. (1993), for Nb. Note that the alkaline basalts from Floreana seem to reproduce melting along a trend comparable to the garnet trend. This would suggest that melting processes in Floreana basalts seem to have displaced some of the incompatible trace elements ratios (such as Nb/La and Nb/Zr) away from the overall mixing trend.

**Figure 13:** Primitive mantle-normalized REE for Floreana and San Cristobal lavas. In both cases, the flat MREE to HREE patterns can be explained by mixing between magmas produced by a low extent of melting from an enriched garnet source and melts produced by moderate extent of melting from a depleted spinel source. In both cases the pattern for the mixture is shown by thick gray lines. a) Floreana lavas are reproduced by a 7.3:1 mixture of 5% melt from a 5% previously depleted primitive spinel mantle and a 0.1% melt from primitive garnet mantle. b) San Cristobal lavas are reproduced by a mixture 1:32 of 0.5% melt from primitive garnet mantle and 7% melt from a 5% previously depleted spinel primitive mantle. We used the source mineralogy and melting reaction of Kinzler and Grove (1992) for the spinel mantle, and the source mineralogy from Beattie (1993) and the 3 GPa melting reaction of Walter (1998) for the garnet source. The initial composition is the primitive mantle of Sun and McDonough (1989). Partition coefficients are from a compilation made by Takazawa (1996). These models are not unique solutions and many other combinations can explain the data. The main goal of the modeling is to demonstrate that flat MREE to HREE patterns are not necessarily an indication of melting in the spinel stability field. Rather, lavas with flat REE patterns can carry a garnet signature in the highly incompatible elements, as in the case of Th/U fractionation (see Figure 10).

**Figure 14:** Mantle upwelling velocities versus mantle porosities. These figures show contours of constant ( $^{230}\text{Th}/^{238}\text{U}$ ). We choose the maximum  $^{230}\text{Th}$  excess observed in each volcano/island analyzed from Galapagos, such that each contour represent one volcano. The figure shows the results of dynamic melting processes (McKenzie, 1985) applying two

different sets of partition coefficients (Beattie, 1993; LaTourrette et al., 1993). In dynamic melting, Th and U are readily stripped from the mantle solid, therefore the extent of melting does not play an important role in the extent of disequilibrium. Using LaTourrette's  $D_U$  and  $D_{Th}$  will require smaller mantle porosities and upwelling rates than those obtained using Beattie's  $D$ .

**Figure 15:** Mantle upwelling velocities versus mantle porosities for equilibrium porous flow. As before, we used the same two sets of partition coefficients. We also considered four different melt column lengths: 1.7 km ( $\approx 0.5\%$  melting), 3 km ( $\approx 1\%$  melting), 17 km ( $\approx 5\%$  melting) and 33 km ( $\approx 10\%$  melting), assuming a constant melt productivity of  $\approx 0.3\%/km$ . In these models we have used the analytical solution of Sims et al., (1999), which considers a constant porosity as an approximation to the numerical solution. As shown by Sims et al. (1999) for Th-U disequilibrium, the analytical solution gives comparable results to the numerical solution. The model assumes modal melting, the source mineral composition taken from Beattie, (1993) and melting in the garnet field. As in dynamic melting, LaTourrette's  $D$ 's require slower mantle upwelling rates and porosities than Beattie's  $D$ 's, especially for the case of very high  $^{230}Th$  excesses. Note that the differences between the models generated by using both sets of partition coefficients is larger at a low extent of melting ( $F \approx D$ ), where melting processes play a more important role in the extent of Th-U disequilibrium. As  $F$  increases ( $F \gg D$ ), the transport processes control the extent of disequilibrium. Thus, porosity and length of the columns begin to be major factors in affecting the  $^{230}Th$  excesses, decreasing the effect of using different  $D$ s.

TABLE 1

Sample ID island/volcano	NSK97-202 Fernandina	NSK97-206 Fernandina	NSK97-214 Fernandina	NSK97-215 Fernandina	NSK97-216 Fernandina	NSK97-224 Fernandina	NSK97-228 Fernandina	SG 93-16 Fernandina
SiO <sub>2</sub>	49.39	49.61	49.44	48.5	48.88	49.18	48.9	49.38
Al <sub>2</sub> O <sub>3</sub>	15.87	15.52	14.77	13.98	15.05	15.71	16.08	15.59
TiO <sub>2</sub>	2.5	2.58	3.12	2.72	3.42	3.09	3.13	2.72
FeO	9.66	9.83	11.3	10.42	11.46	10.22	10.72	10.45
MnO	0.17	0.17	0.19	0.17	0.19	0.18	0.18	0.18
MgO	7.63	7.5	6.41	10.4	5.84	6.21	5.81	6.52
CaO	11.88	11.89	11.6	10.68	11.34	11.75	11.45	12.09
K <sub>2</sub> O	0.31	0.32	0.39	0.36	0.54	0.47	0.48	0.38
Na <sub>2</sub> O	2.36	2.34	2.46	2.49	2.91	2.86	2.92	2.5
P <sub>2</sub> O <sub>5</sub>	0.25	0.25	0.31	0.27	0.38	0.34	0.34	0.27
Total	100	100	100	100	100	100	100	100
<b>XRF</b>								
Ni	97	87	49	225	37	51	37	58
Cr	323	288	155	641	143	186	185	193
Sc	35	34	33	35	35	35	30	38
V	309	316	366	331	383	357	351	330
Ba	35	37	37	34	41	58	29	40
Rb	6	6	6	8	9	10	5	6
Sr	311	304	321	303	350	344	354	330
Zr	126	128	155	135	196	171	172	146
Y	23	23	29	23	31	29	29	25
Nb	17.4	17.3	21.0	19.1	26.0	23.8	22.3	19.3
Ga	22	22	26	22	23	27	25	25
Cu	80	75	83	63	87	78	72	94
Zn	101	96	110	104	110	106	107	97
<b>ICPMS</b>								
Cs	0.06	0.06	0.07	0.06	0.08	0.09	0.08	0.07
Rb	5.4	5.3	7.6	6.0	9.7	8.2	8.6	6.2
Ba	68	67	89	73	114	99	101	80
Th	0.97	0.94	1.31	1.05	1.64	1.35	1.36	1.14
U	0.28	0.27	0.38	0.3	0.49	0.39	0.4	0.34
Nb	15.2	15.1	20.0	16.9	24.3	21.5	21.6	17.1
Ta	1.08	1.08	1.5	1.24	1.75	1.51	1.59	1.23
La	10.5	10.4	14.1	11.8	17.9	15.7	15.5	12.9
Ce	24.3	24.0	32.3	27.0	40.2	35.4	34.8	28.1
Pb	1.42	1.16	0.93	1.15	1.1	3.21	0.96	1.18
Pr	3.22	3.24	4.23	3.66	5.27	4.72	4.65	3.79
Sr	305	301	351	289	371	360	383	323
Nd	15.6	15.2	19.5	17.1	23.9	21.6	21.5	17.7
Sm	4.50	4.48	5.77	4.88	6.92	6.26	6.13	5.02
Hf	3.06	3.12	4.15	3.35	4.96	4.34	4.5	3.59
Eu	1.60	1.61	2.04	1.78	2.37	2.10	2.18	1.84
Gd	4.74	4.69	6.01	5.09	7.04	6.32	6.50	5.32
Tb	0.80	0.79	1.03	0.87	1.17	1.07	1.09	0.92
Dy	4.69	4.76	6.08	5.13	6.94	6.30	6.40	5.47
Y	23.9	24.1	30.1	24.3	34.5	31.6	32.1	26.8
Ho	0.93	0.93	1.21	0.99	1.37	1.21	1.25	1.04
Er	2.27	2.35	3.04	2.46	3.42	3.10	3.21	2.59
Tm	0.31	0.33	0.41	0.34	0.47	0.43	0.44	0.36
Yb	1.82	1.85	2.46	2.00	2.80	2.48	2.57	2.07
Lu	0.27	0.28	0.36	0.29	0.41	0.37	0.38	0.31
Sc	35.6	35.5	44.3	33.6	42.2	40.0	39.1	38.7
<sup>3</sup> He/ <sup>4</sup> He	25	23	30	22	22	23	25	27
<sup>87</sup> Sr/ <sup>86</sup> Sr	0.703272	0.703304	0.70329	0.703264	0.703251	0.703249	0.703207	0.703234
<sup>143</sup> Nd/ <sup>144</sup> Nd	0.512921	0.512929	0.512937	0.512922	0.512958	0.51294	0.512947	0.512965
<sup>206</sup> Pb/ <sup>204</sup> Pb								
<sup>207</sup> Pb/ <sup>204</sup> Pb								
<sup>208</sup> Pb/ <sup>204</sup> Pb								

TABLE 1

Sample ID island/volcano	NSK97-238 Santiago	NSK97-252 San Cristobal	NSK97-254 San Cristobal	SG93-22 Floreana	SG 93-5 Cerro Azul	SG 93-7 Cerro Azul	SG 93-8 Cerro Azul	SG 93-12 Cerro Azul
SiO2	48.4	47.3	47.44	46.4	48.09	50.93	48.89	49.08
Al2O3	14.85	16.77	16.4	14.39	14.79	13.6	15.49	14.34
TiO2	2.13	1.01	0.98	1.67	2.04	1.45	2.66	3.4
FeO	10.67	8.83	9.04	9.78	9.96	9.47	10.22	12.37
MnO	0.18	0.16	0.17	0.17	0.18	0.16	0.19	0.22
MgO	10.22	11.17	11.8	13.3	10.2	11.84	7.03	5.72
CaO	10.25	12.37	11.55	10.6	11.75	9.9	11.28	10.24
K2O	0.31	0.12	0.19	0.83	0.34	0.18	0.66	0.78
Na2O	2.74	2.16	2.31	2.66	2.45	2.34	3.23	3.44
P2O5	0.24	0.1	0.11	0.22	0.21	0.14	0.35	0.42
Total	100	100	100	100	100	100	100	100
<b>XFF</b>								
Ni	241	211	259	324	201	356	97	46
Cr	495	481	479	698	508	583	227	90
Sc	34	40	40	33	37	30	35	34
V	246	228	210	247	288	189	326	412
Ba	27	28	44	257	77	38	107	138
Rb	5	3	3	20	6	5	13	15
Sr	234	173	173	390	284	182	358	356
Zr	139	66	67	91	113	77	178	211
Y	29	21	21	21	22	17	29	35
Nb	11.1	5.2	6.5	23.5	17.8	9.1	28.8	36.4
Ga	21	15	16	13	19	17	21	24
Cu	80	281	330	76	68	77	74	71
Zn	92	58	60	68	85	80	97	118
<b>ICPMS</b>								
Cs	0.05	0.02	0.04	0.23	0.06	0.04	0.13	0.18
Rb	5.1	1.2	2.3	19.4	6.1	3.6	12.5	14.9
Ba	49	24	38	279	88	44	148	188
Th	0.76	0.24	0.37	1.3	1.08	0.57	1.99	2.35
U	0.21	0.06	0.1	0.22	0.29	0.16	0.54	0.64
Nb	9.6	3.6	4.9	22.2	15.5	7.7	27.6	33.9
Ta	0.65	0.25	0.31	1.22	1.01	0.52	1.79	2.19
La	9.0	3.4	4.3	16.0	12.0	5.8	20.8	25.0
Ce	21.0	8.5	9.5	26.6	26.1	12.9	43.2	51.6
Pb	1.07	0.9	0.75	2.99	1.1	0.99	1.89	2.27
Pr	2.99	1.31	1.36	3.04	3.4	1.82	5.61	6.55
Sr	230	174	167	390	283	187	378	362
Nd	14.9	6.8	6.6	12.5	15.5	8.9	24.9	28.9
Sm	4.90	2.40	2.18	3.24	4.42	3.05	6.53	7.74
Hf	3.53	1.65	1.54	1.8	2.97	2.1	4.6	5.45
Eu	1.79	0.96	0.89	1.18	1.56	1.13	2.22	2.59
Gd	5.61	3.16	2.84	3.50	4.73	3.60	6.68	7.71
Tb	0.94	0.61	0.55	0.64	0.80	0.64	1.08	1.27
Dy	5.86	4.08	3.87	4.03	4.88	3.77	6.56	7.77
Y	29.2	23.0	22.1	21.9	23.3	19.7	32.4	37.1
Ho	1.12	0.89	0.83	0.82	0.95	0.75	1.26	1.48
Er	2.93	2.56	2.37	2.34	2.40	1.93	3.20	3.74
Tm	0.39	0.37	0.34	0.32	0.34	0.26	0.44	0.51
Yb	2.37	2.36	2.32	1.96	1.96	1.55	2.63	3.00
Lu	0.35	0.37	0.36	0.32	0.30	0.23	0.39	0.45
Sc	32.4	45.4	41.6		36.77	31.74	36.73	34.86
3He/4He	9	8	10	14	14	14	14	12
87Sr/86Sr	0.703036	0.702991	0.703485	0.703415	0.703323	0.703366	0.703288	0.703344
143Nd/144Nd	0.512998	0.513041	0.513042	0.512946		0.512958	0.512929	0.512917
206Pb/204Pb				19.783				
207Pb/204Pb				15.619				
208Pb/204Pb				39.515				



**TABLE 1**

Sample ID	GI92-5	GI92-8	GI92-17	GI92-18	GI92-19
island/volcano	Sierra Negra	Sierra Negra	Darwin	Darwin	Ecuador
SiO <sub>2</sub>	49.33	49.06	48.98	48.8	47.93
Al <sub>2</sub> O <sub>3</sub>	13.97	14.18	16.61	16.74	14.32
TiO <sub>2</sub>	3.01	3.36	2.79	2.58	2.29
FeO	12.2	12.71	10.37	10.05	11
MnO	0.19	0.21	0.17	0.16	0.17
MgO	6.5	5.99	5.65	6.07	10.9
CaO	11.19	10.62	11.88	12.39	9.9
K <sub>2</sub> O	0.44	0.49	0.39	0.32	0.52
Na <sub>2</sub> O	2.86	3.03	2.89	2.64	2.69
P <sub>2</sub> O <sub>5</sub>	0.29	0.34	0.27	0.24	0.28
Total	100	100	99.99	100	100
<b>XRF</b>					
Ni	53	40	53	56	265
Cr	158	104	99	152	499
Sc	40	36	34	35	27
V	367	362	331	309	226
Ba	63	92	22	24	95
Rb	8	9	6	3	10
Sr	296	302	374	382	344
Zr	160	183	166	155	150
Y	32	35	27	26	25
Nb	25.3	28.5	14.5	12.8	22.1
Ga	24	24	26	23	19
Cu	105	92	120	125	68
Zn	110	117	89	82	94
<b>ICPMS</b>					
Cs	0.08	0.1	0.05	0.06	0.11
Rb	7.5	8.7	4.7	3.6	9.2
Ba	97	115	51	44	120
Th	1.13	1.39	0.78	0.66	1.27
U	0.25	0.36	0.2	0.18	0.28
Nb	23.3	26.8	13.4	11.9	20.5
Ta	1.5	1.75	0.91	0.8	1.2
La	15.6	17.9	13.0	11.5	15.7
Ce	34.3	39.4	30.1	26.8	32.3
Pb	0.98	1.38	2.49	2.27	1.68
Pr	4.68	5.31	4.1	3.68	4.09
Sr	296	302	374	382	344
Nd	21.4	24.5	19.0	17.4	18.8
Sm	6.19	6.99	5.78	5.20	5.23
Hf	3.94	4.53	3.61	3.36	3.36
Eu	2.21	2.46	2.06	1.89	1.81
Gd	6.12	6.95	5.75	5.36	5.22
Tb	1.10	1.24	0.98	0.89	0.89
Dy	6.56	7.25	5.81	5.29	5.12
Y	31.7	35.7	27.8	25.4	26.0
Ho	1.28	1.40	1.14	1.01	0.99
Er	3.35	3.75	2.95	2.70	2.67
Tm	0.43	0.50	0.39	0.35	0.35
Yb	2.59	2.91	2.25	2.01	2.04
Lu	0.39	0.44	0.33	0.29	0.30
Sc					
<sup>3</sup> He/ <sup>4</sup> He	16	17	12	12	10
<sup>87</sup> Sr/ <sup>86</sup> Sr	0.703365	0.703504	0.702898	0.702856	0.7032
<sup>143</sup> Nd/ <sup>144</sup> Nd	0.512904	0.512941	0.513011	0.51302	0.512951
<sup>206</sup> Pb/ <sup>204</sup> Pb	19.342		18.797	18.78	19.294
<sup>207</sup> Pb/ <sup>204</sup> Pb	15.569		15.518	15.512	15.575
<sup>208</sup> Pb/ <sup>204</sup> Pb	38.989		38.296	38.26	38.968

TABLE 2

Sample ID	island/volcano	Th ppm (ID)	U ppm (ID)	(238U/232Th)	(230Th/232Th)	2St err %	(230Th/238U)	age (Ka)	(230Th/232Th)	age	(230Th/238U)	age
NSK97-202	Fernandina	0.95	0.25	0.812	1.063	1.03	1.310	1.87 ±0.4	1.067			1.315
NSK97-206	Fernandina	0.97	0.26	0.816	1.055	0.90	1.292	3±0.3	1.061			1.300
NSK97-214	Fernandina	1.16	0.31	0.821	1.035	0.53	1.261	>0.9	1.037			1.263
NSK97-215	Fernandina	1.20	0.32	0.801	1.012	0.50	1.264	<0.1	1.012			1.265
NSK97-216	Fernandina	1.54	0.42	0.822	1.072	0.80	1.303	historical (1982)	1.072			1.303
NSK97-224	Fernandina	1.39	0.38	0.824	1.034	0.45	1.256	historical (1995)	1.034			1.256
NSK97-224R	Fernandina	1.39	0.38	0.824	1.045	0.74	1.269	historical (1995)	1.045			1.268
NSK97-228	Fernandina	1.42	0.38	0.819	1.047	0.77	1.278	<0.6	1.048			1.279
SG 93-16	Fernandina	1.13	0.31	0.830	1.067	0.66	1.287	1±0.6	1.070			1.289
NSK97-238	Santiago	0.75	0.20	0.811	1.040	0.64	1.284	<0.9	1.042			1.286
NSK97-126	San Cristobal	0.64	0.16	0.765	1.100	0.90	1.438	9.2 ±0.7	1.130			1.477
NSK97-252	San Cristobal	0.24	0.07	0.833	1.132	0.96	1.359	15 ±1	1.177			1.412
NSK97-254	San Cristobal	0.37	0.10	0.821	1.087	1.05	1.325	13±0.5	1.121			1.366
SG93-22	Floreana	1.52	0.29	0.574	0.938	0.60	1.635	26 ±7	1.036			1.806
SG 93-5	Cerro Azul	1.06	0.27	0.759	1.040	0.58	1.371	<1	1.043			1.374
SG 93-7	Cerro Azul	0.50	0.14	0.843	1.042	1.12	1.236	9±4	1.045			1.240
SG 93-8	Cerro Azul	1.99	0.51	0.772	1.043	0.57	1.351	<0.8	1.045			1.353
SG 93-12	Cerro Azul	2.32	0.58	0.757	1.027	0.58	1.357	1.7 ±0.3	1.032			1.363
GI92-5	Sierra Negra	1.31	0.34	0.786	1.010	0.82	1.285	0.6 ±0.4	1.011			1.287
GI92-8	Sierra Negra	1.61	0.47	0.884	1.027	0.69	1.163	0.9 ±0.4	1.028			1.164
SN 91-60	Sierra Negra	1.63	0.43	0.798	1.030	0.46	1.291	1.6 ±1	1.034			1.295
GI92-17	Darwin	0.91	0.27	0.905	1.120	0.97	1.238	0.8 ±0.3	1.122			1.240
GI92-18	Darwin	0.70	0.21	0.907	1.140	0.78	1.257	1.7 ±0.4	1.144			1.261
GI92-19	Ecuador	1.44	0.36	0.761	1.053	0.74	1.384	<0.8	1.055			1.387
W95-25	Wolf	0.81	0.24	0.912	1.160	1.03	1.272	<0.2	1.160			1.272
W95-54	Wolf	1.22	0.38	0.932	1.186	0.83	1.272		1.186			1.272
W95-74	Wolf	0.74	0.24	1.001	1.186	0.77	1.185	<1	1.188			1.187
W95-84	Wolf	1.05	0.33	0.940	1.164	0.69	1.238	<1.4	1.167			1.241
SKV98-108	Pinta	1.25	0.28	0.677	1.014	0.77	1.499	<1.5	1.019			1.506
SKV98-101	Marchena	0.73	0.20	0.834	1.092	0.65	1.309	historical (1991)	1.092			1.309

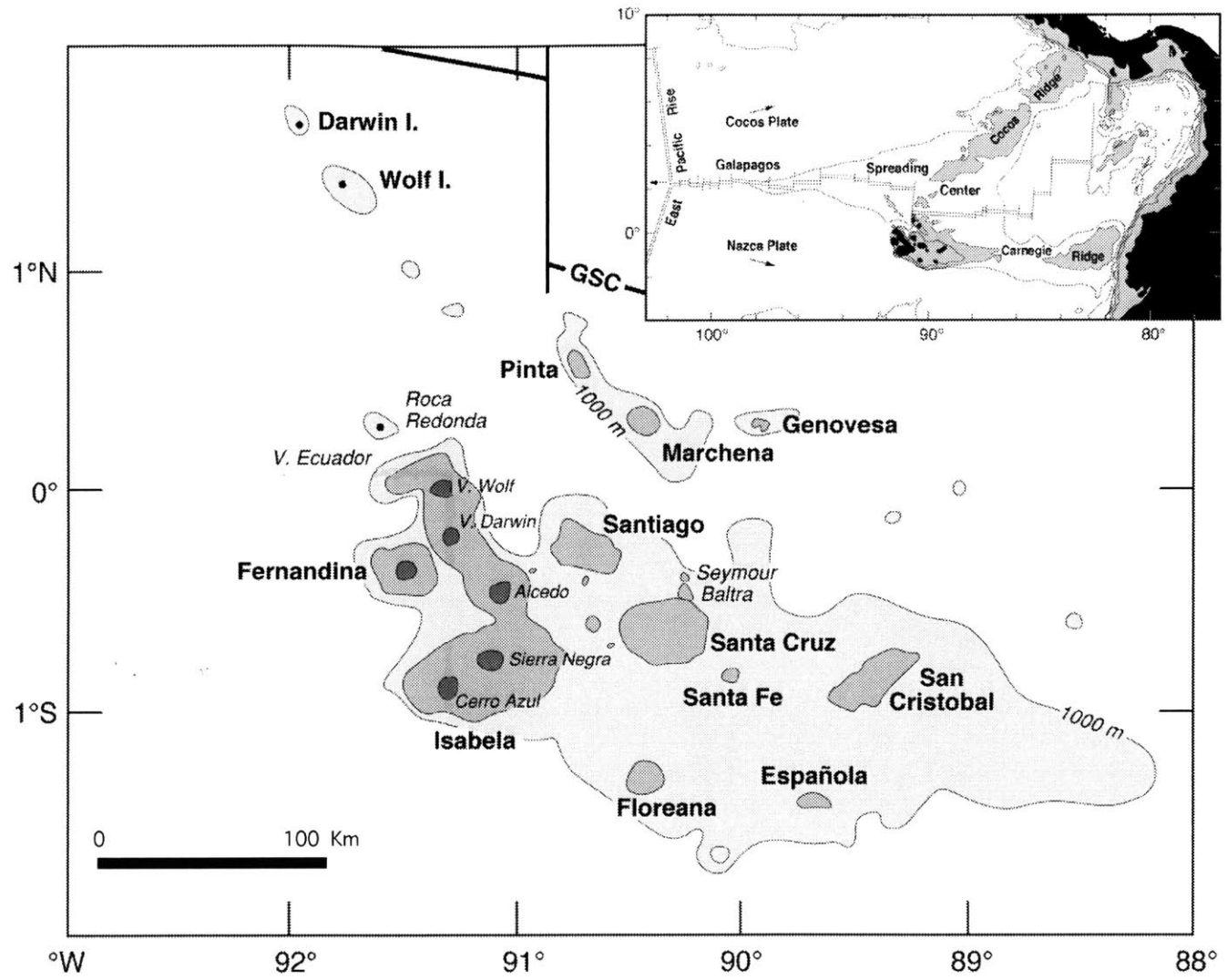


Figure 1

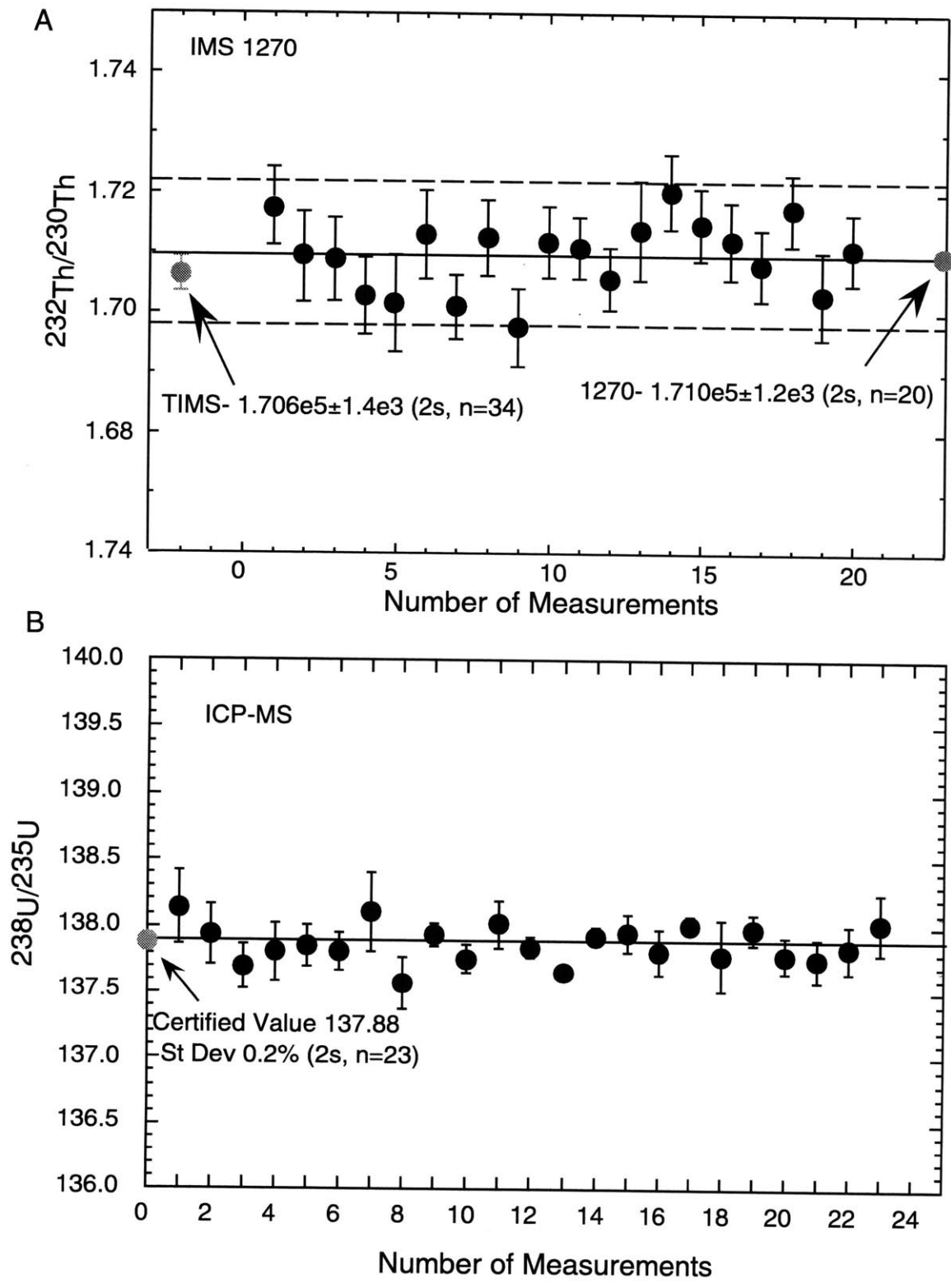


Figure 2

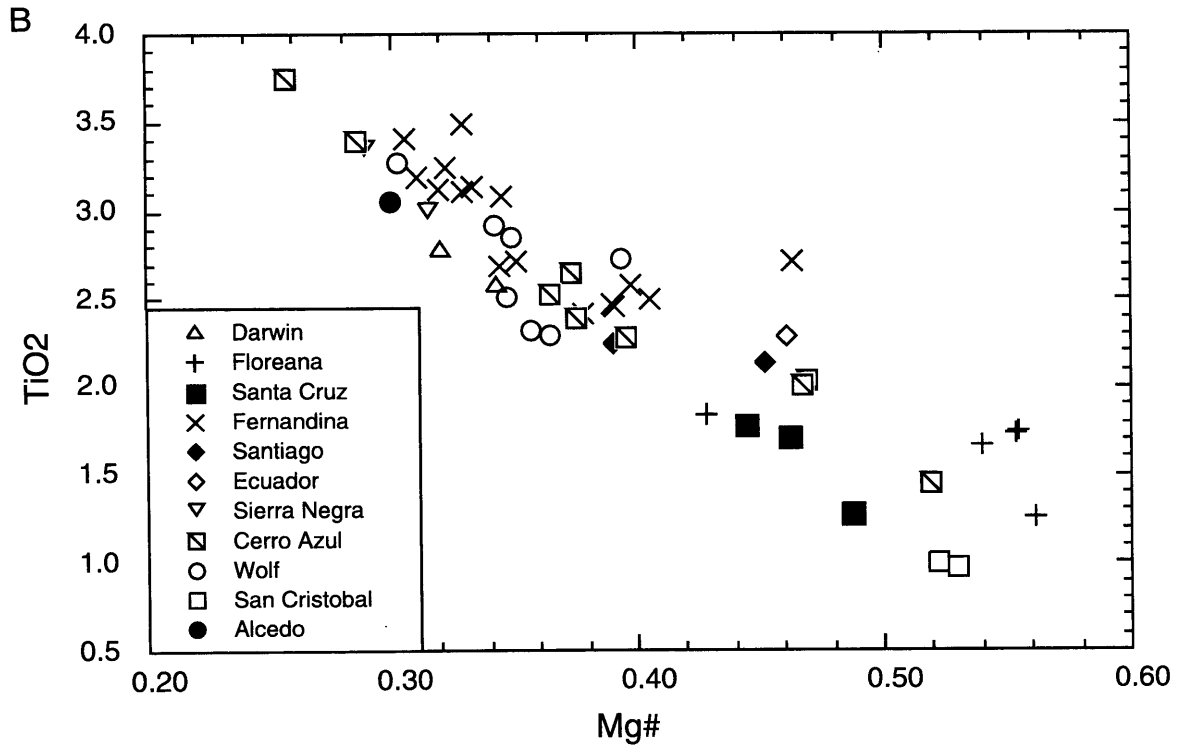
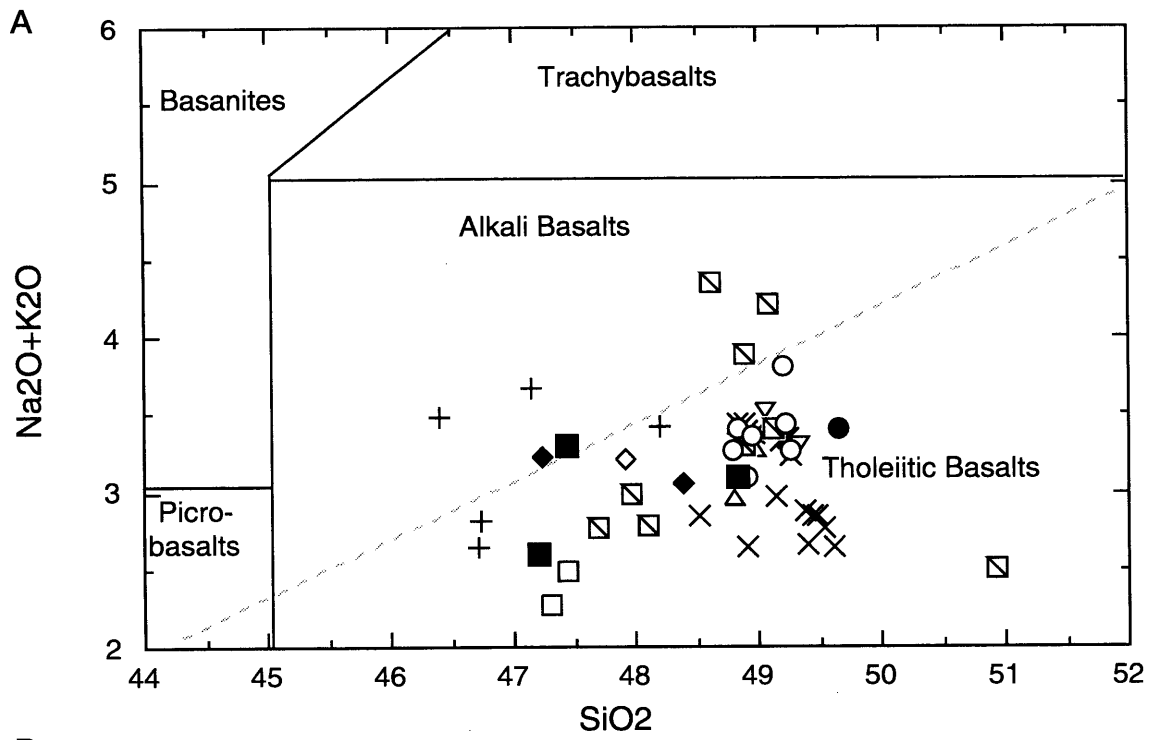


Figure 3

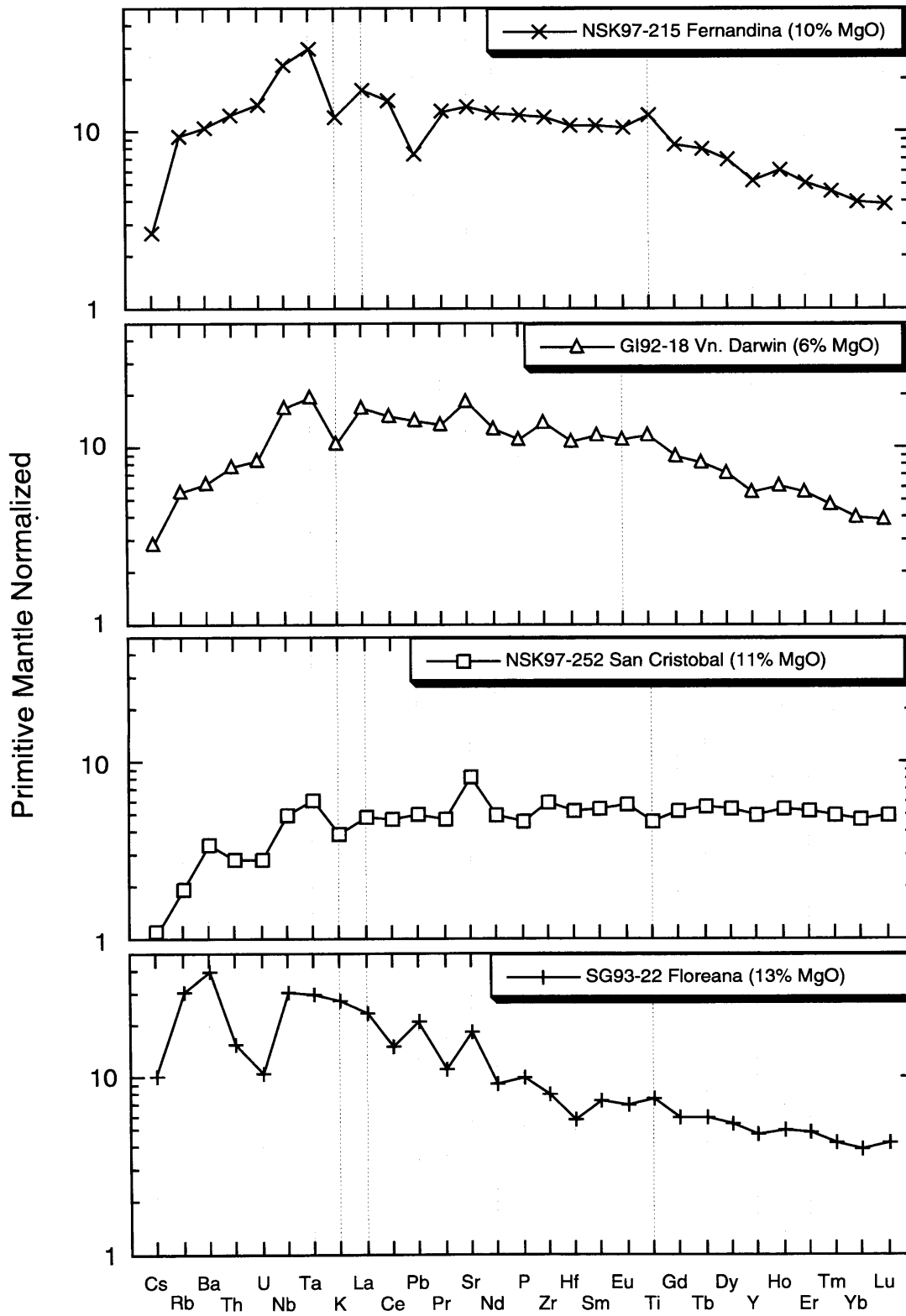


Figure 4 a

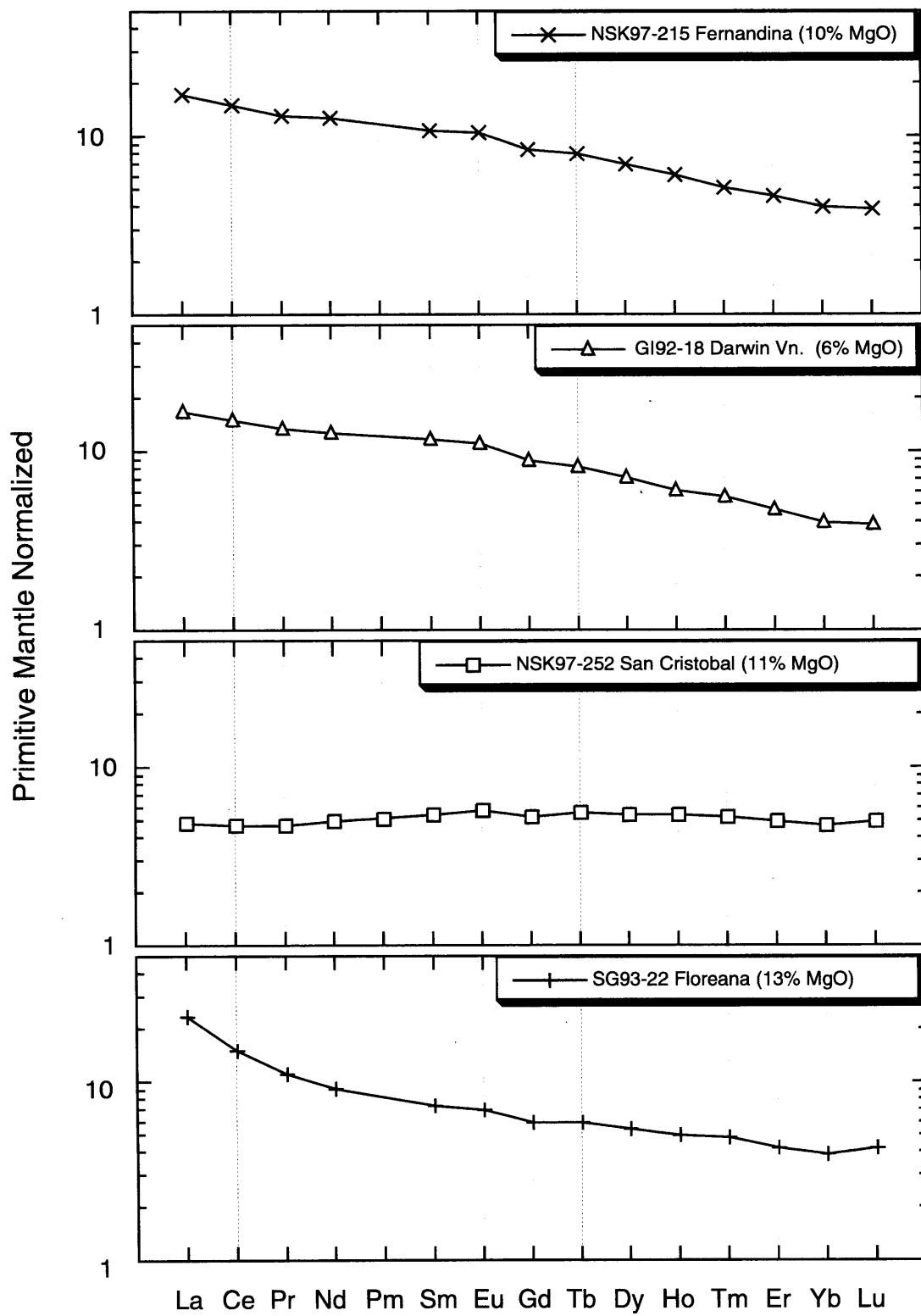


Figure 4b

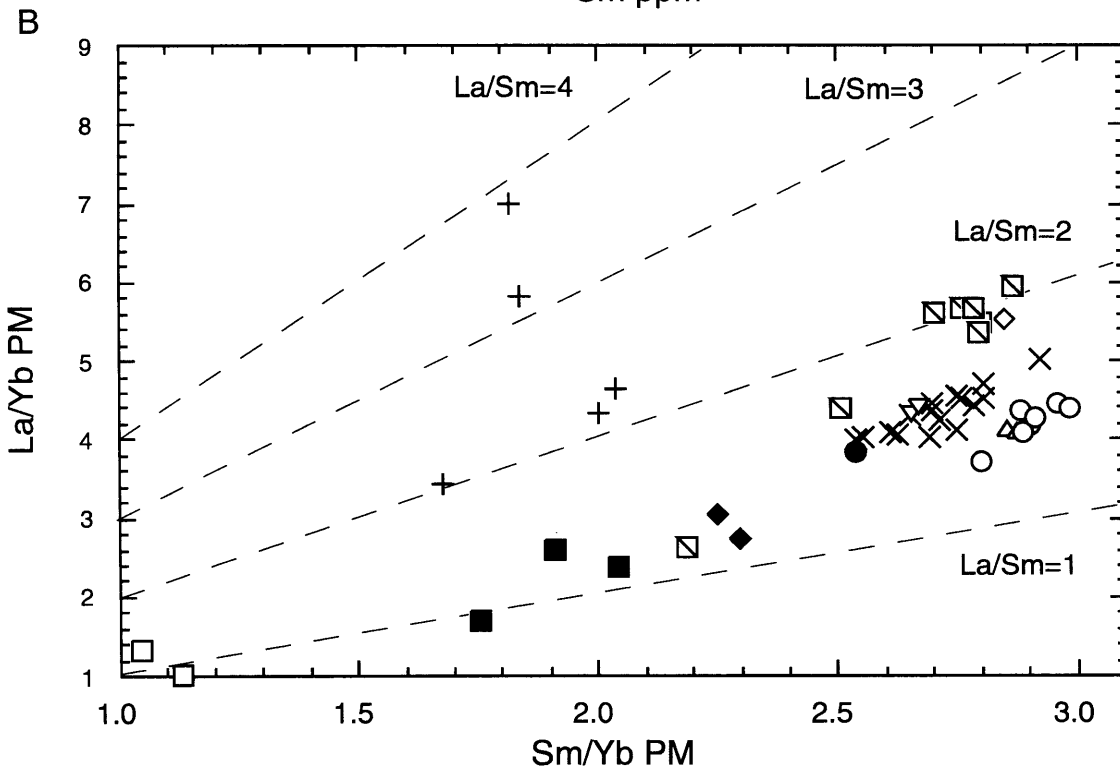
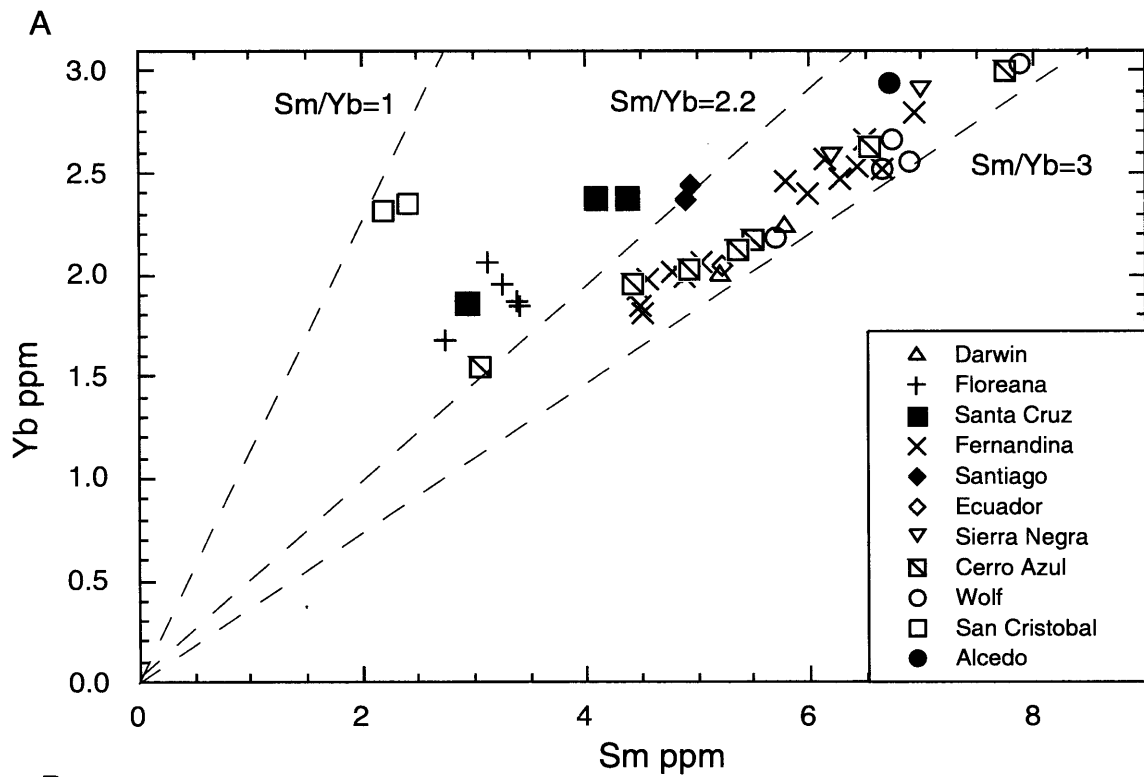


Figure 5



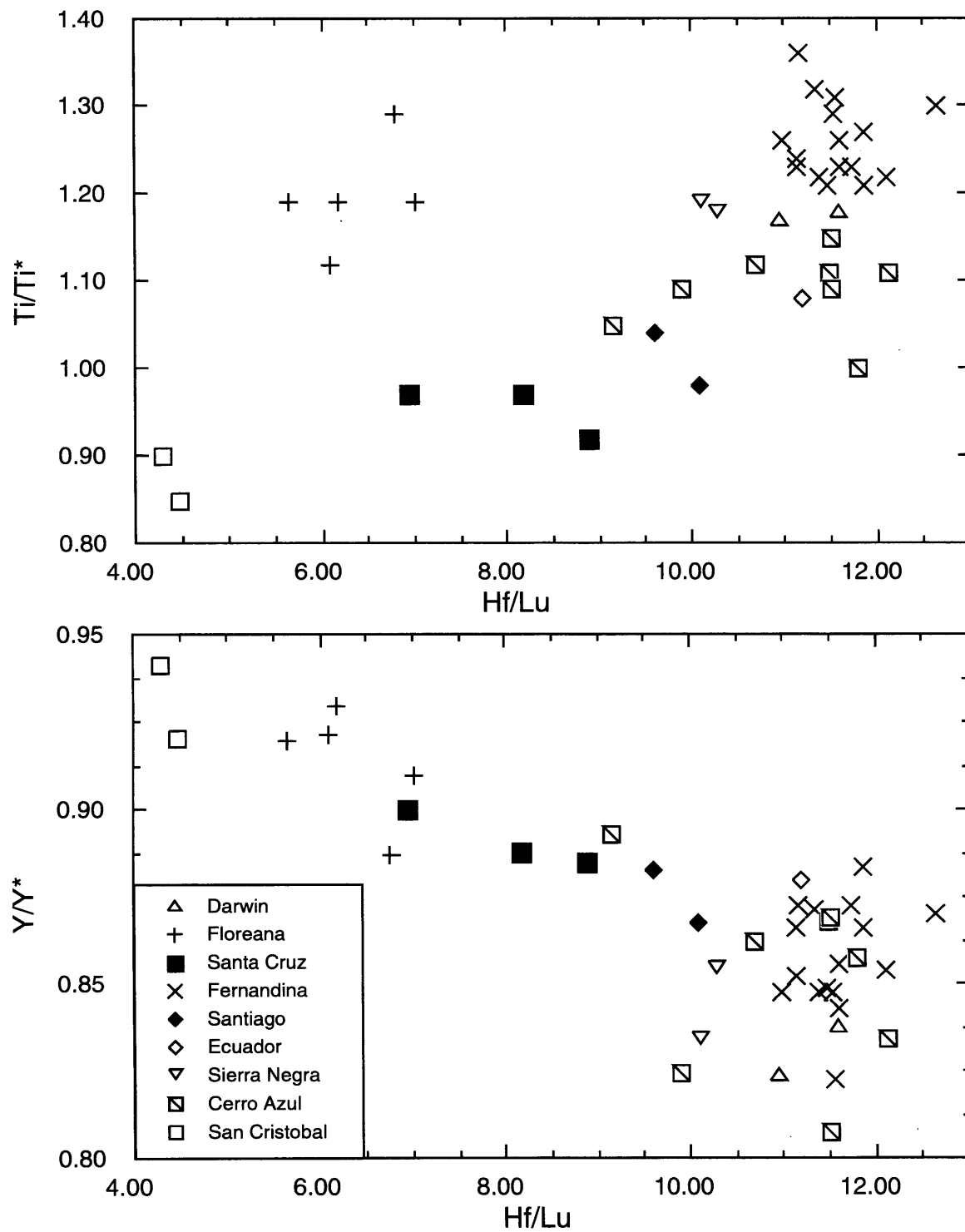


Figure 6

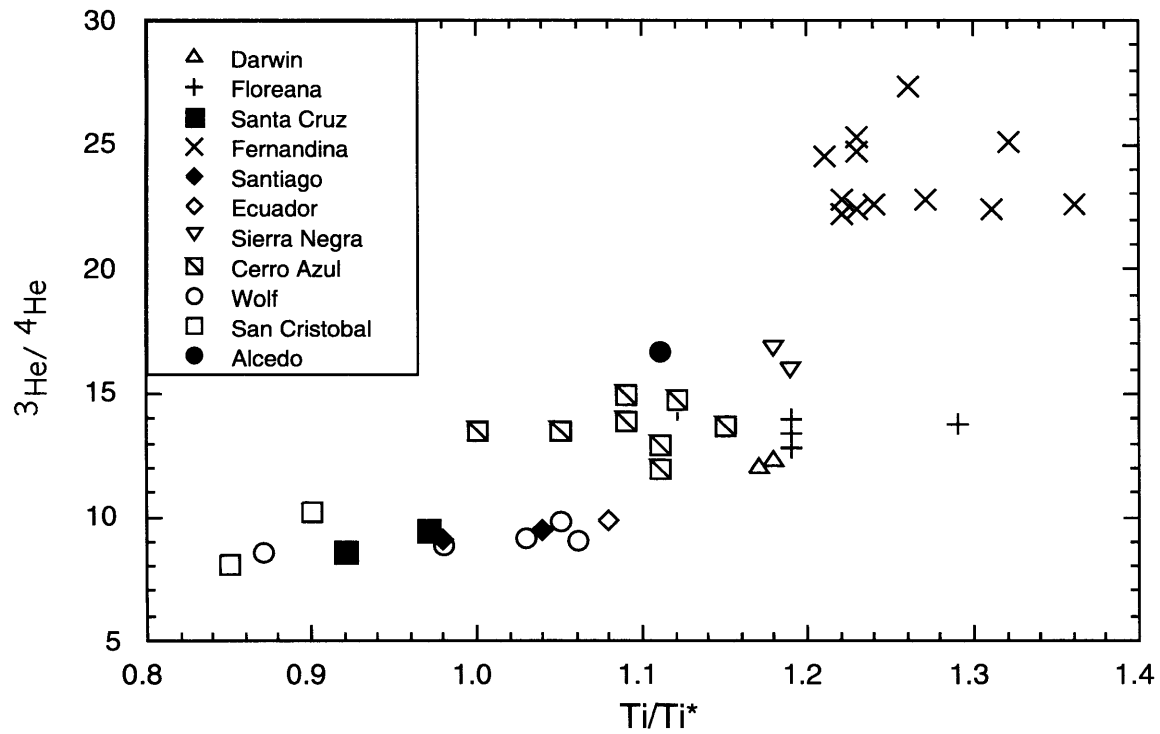
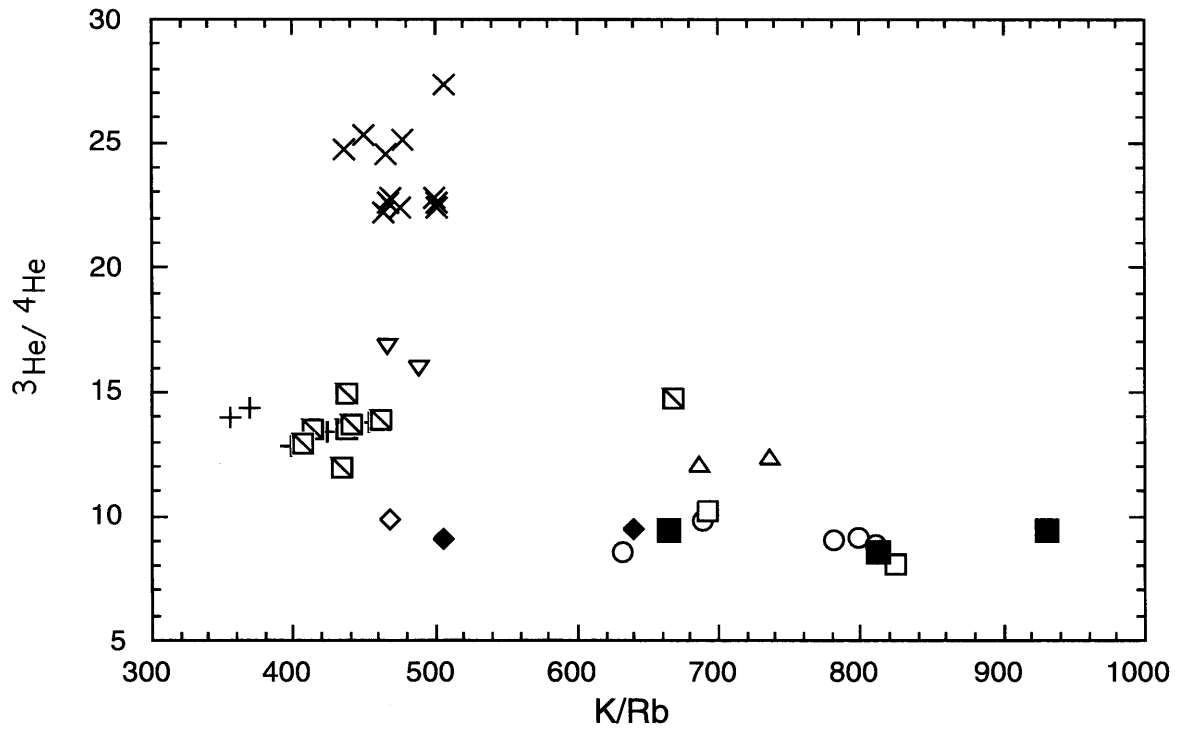


Figure 7

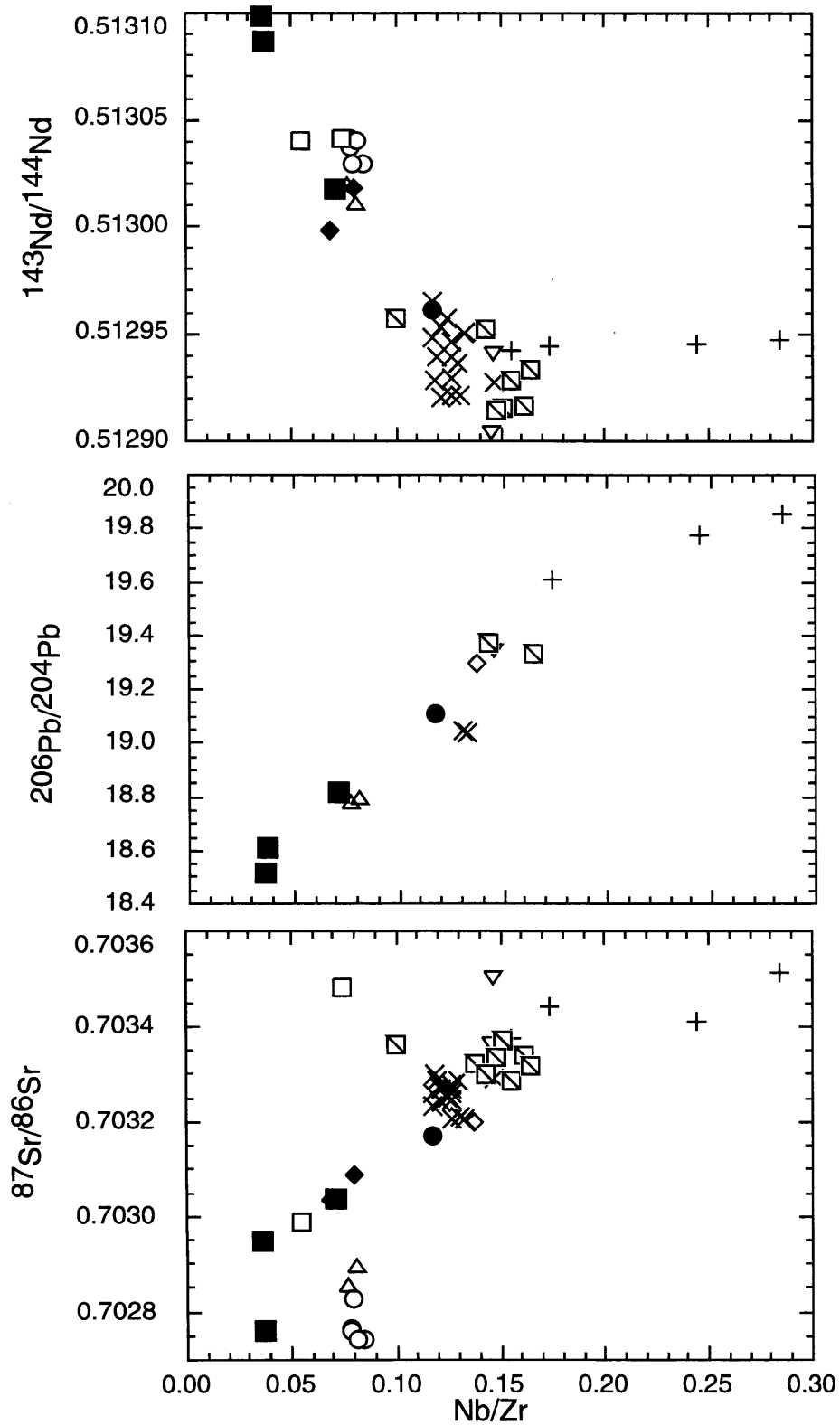


Figure 8

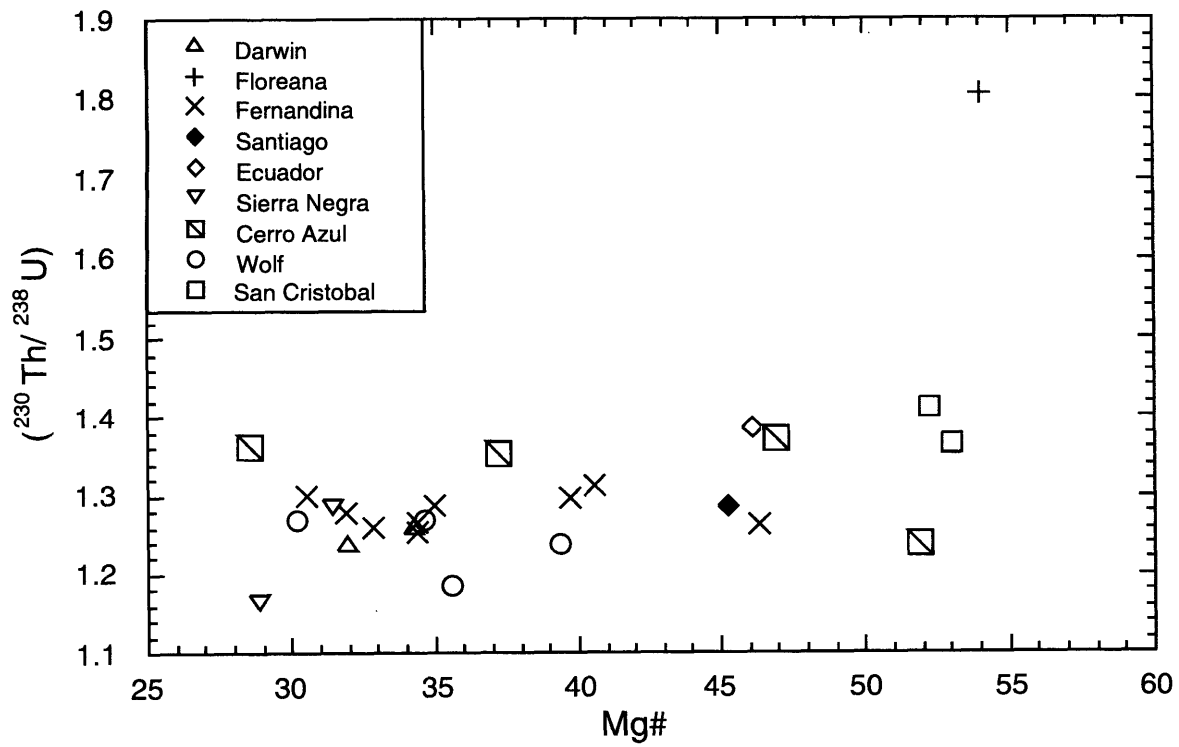
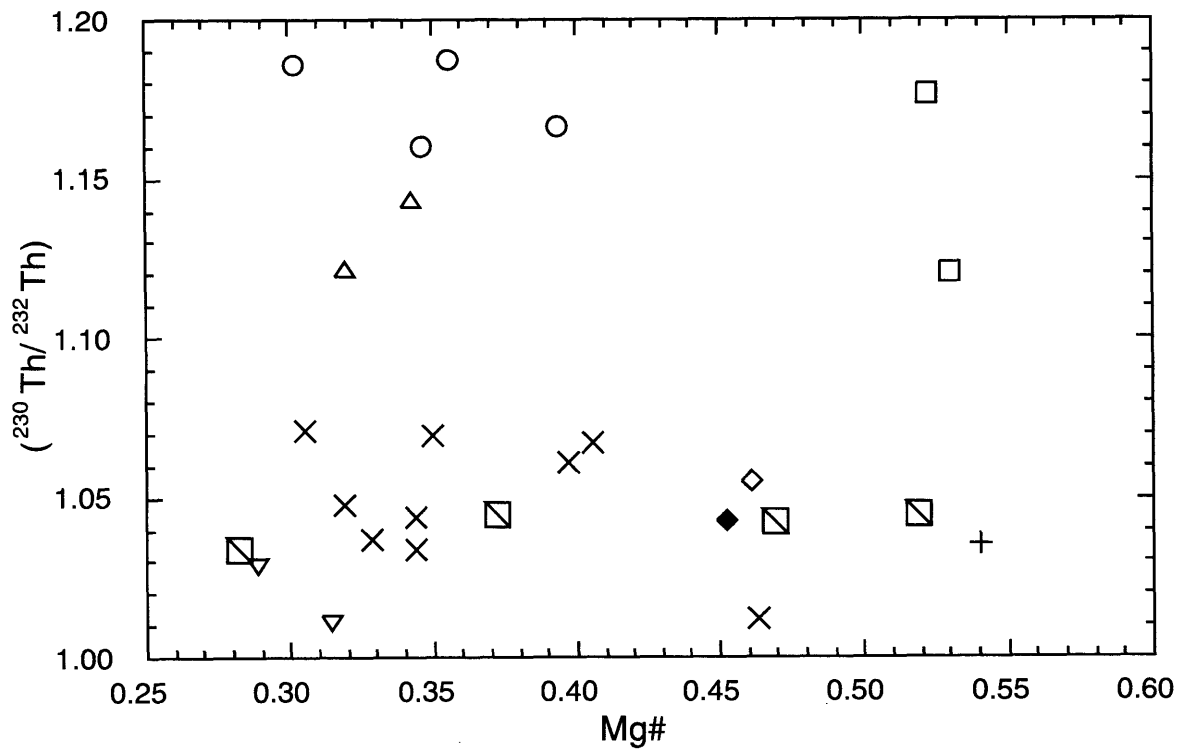


Figure 9

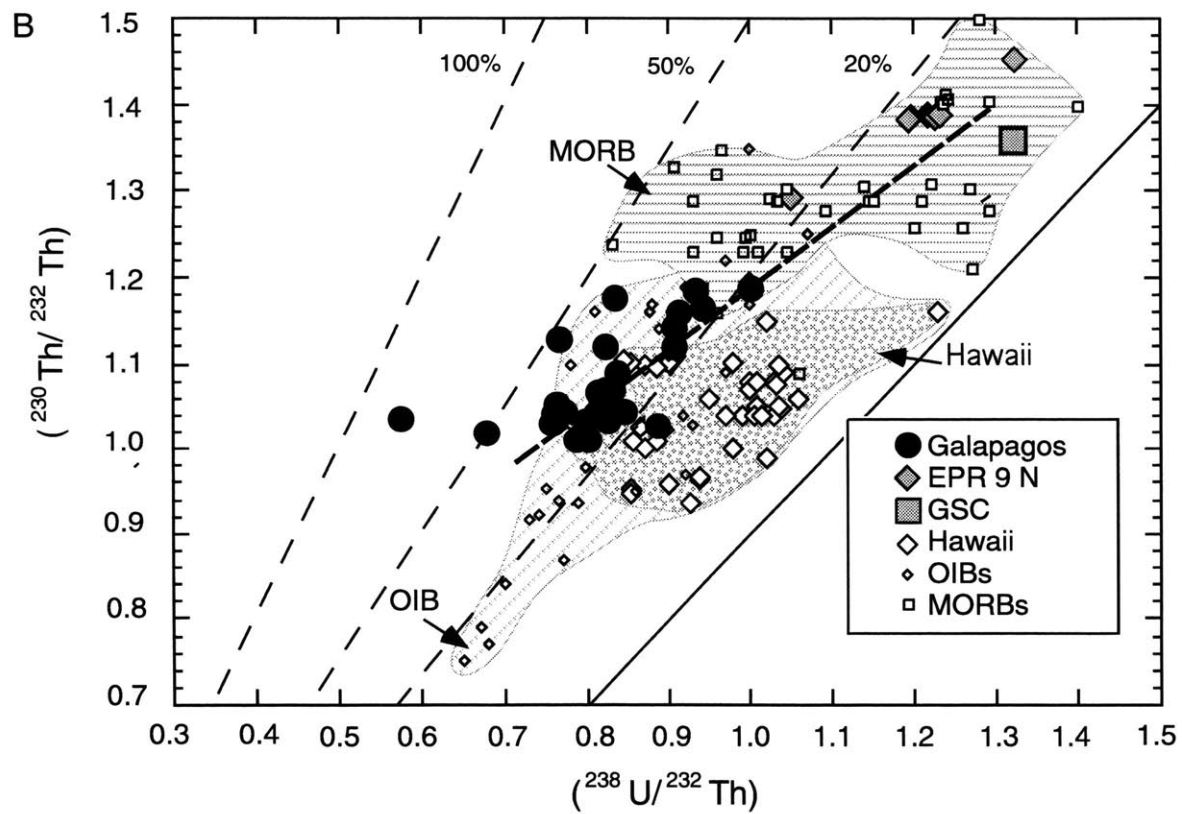
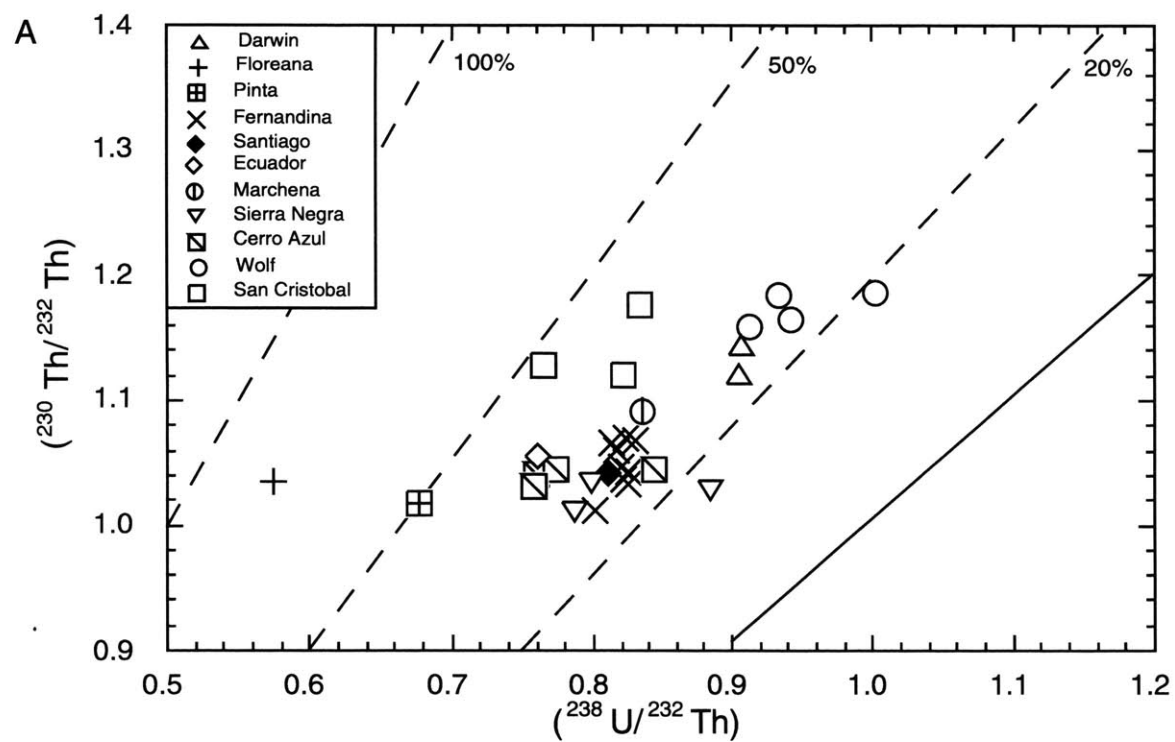


Figure 10

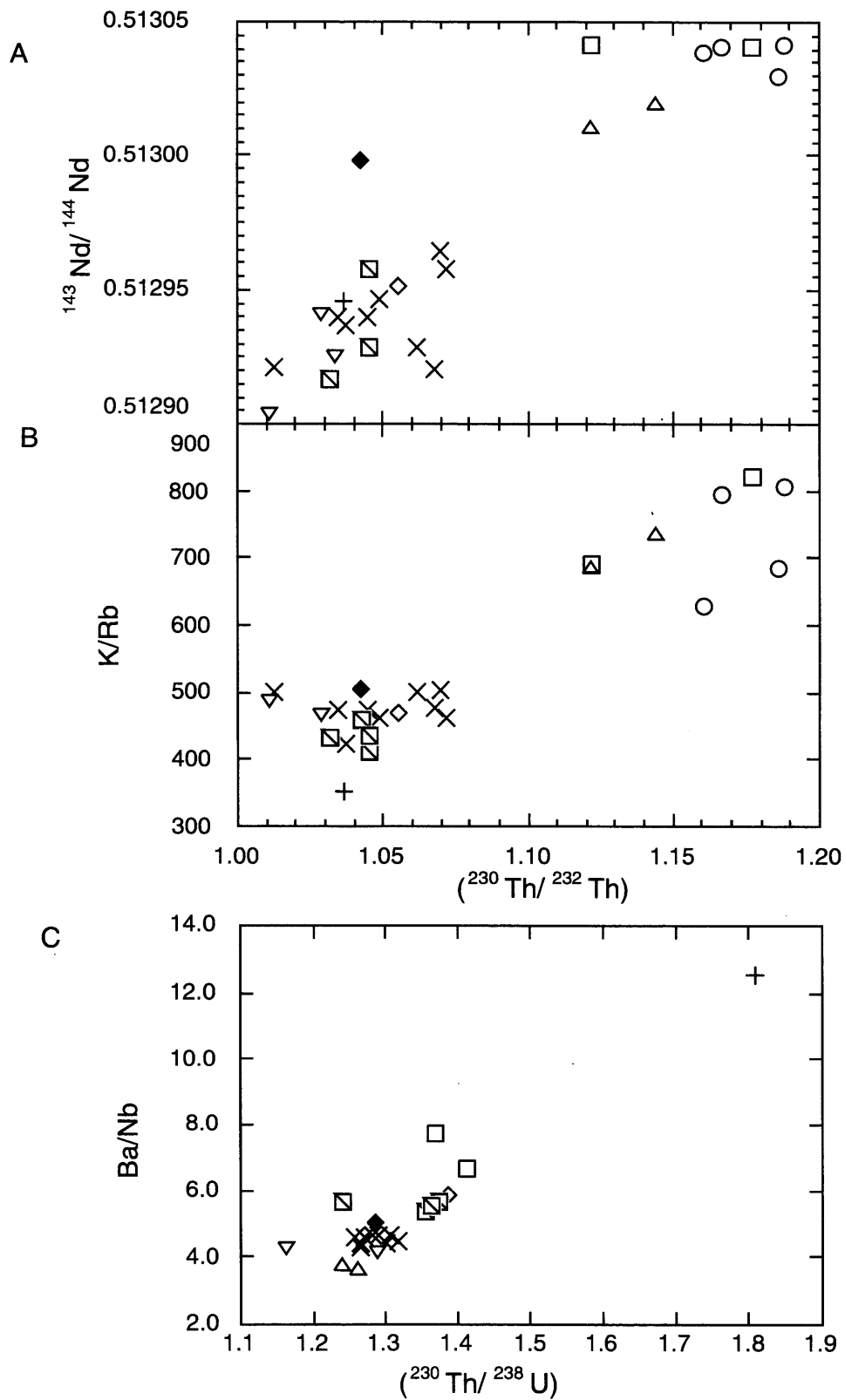


Figure 11

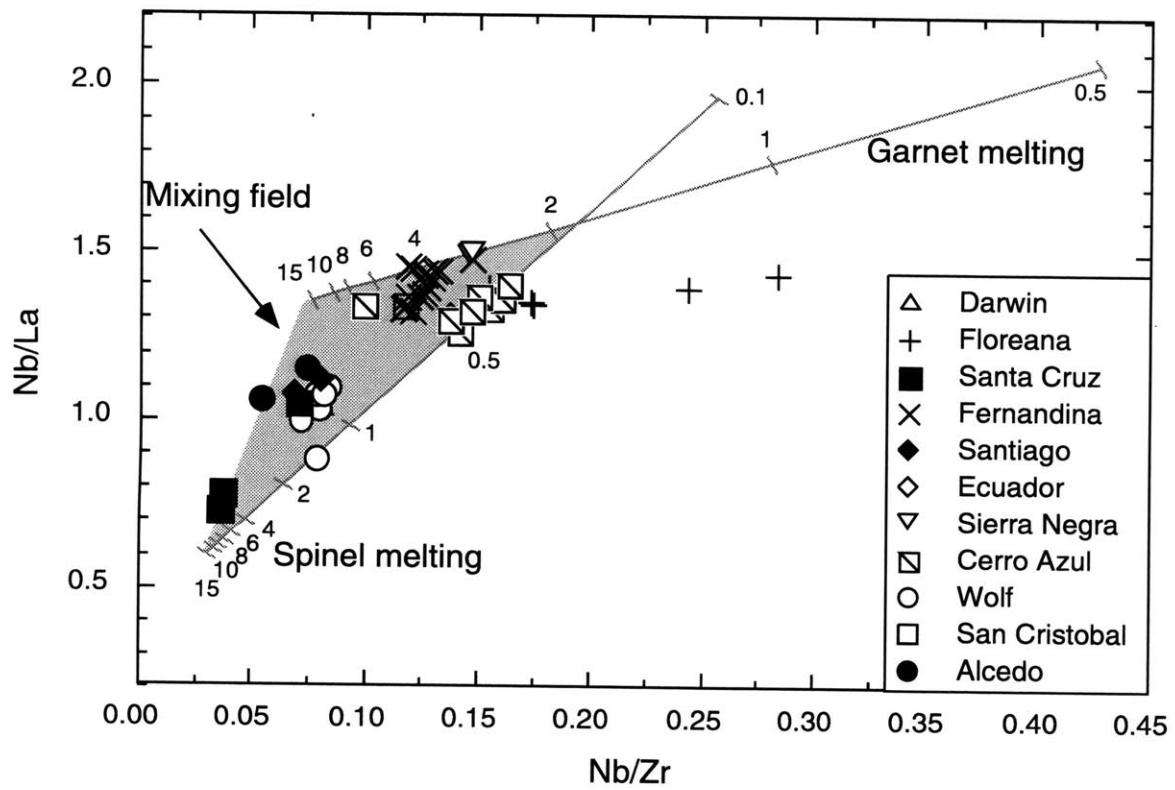


Figure 12

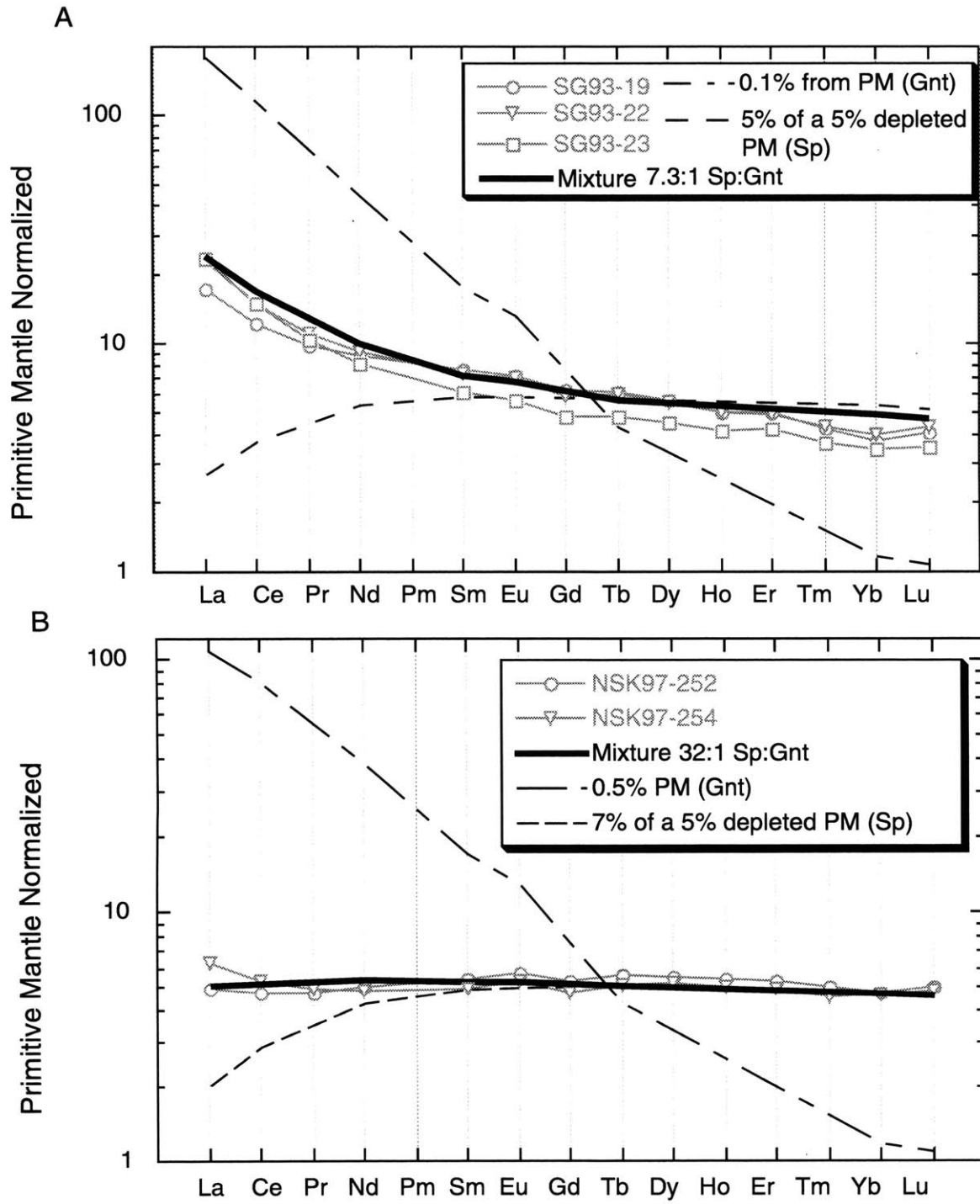


Figure 13



# Dynamic Melting

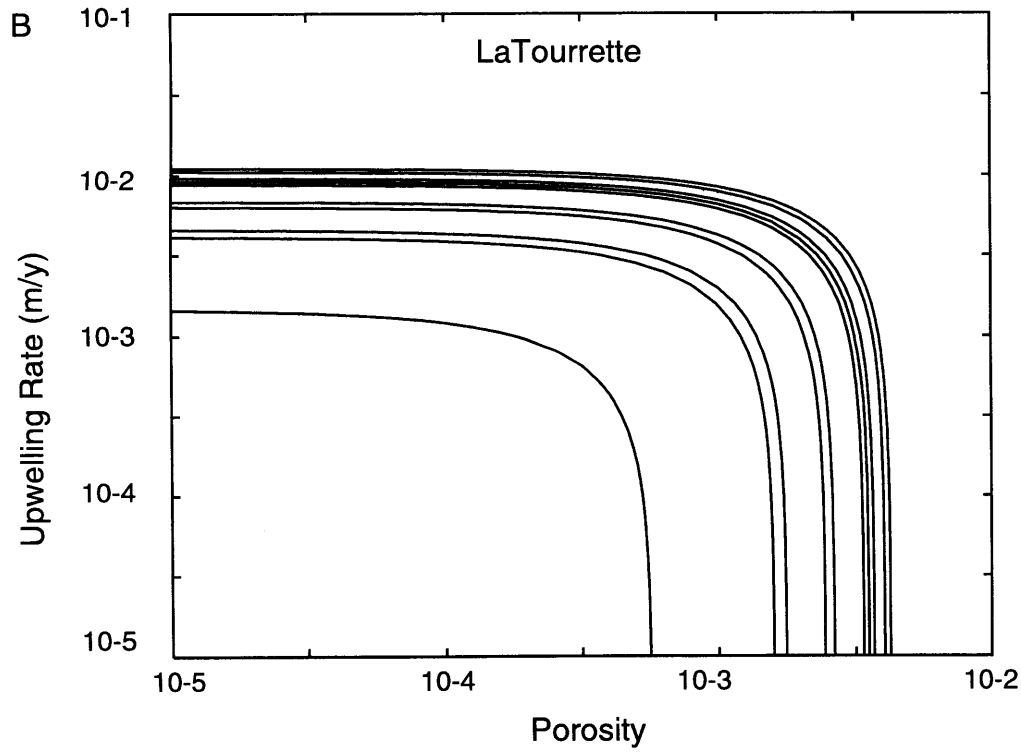
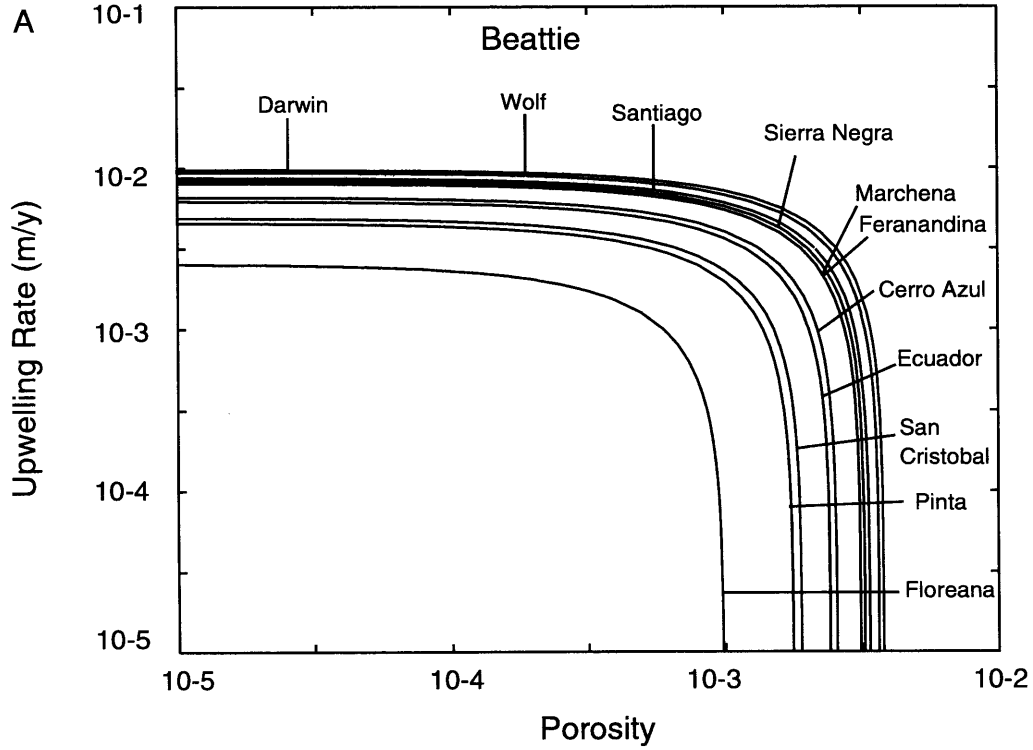
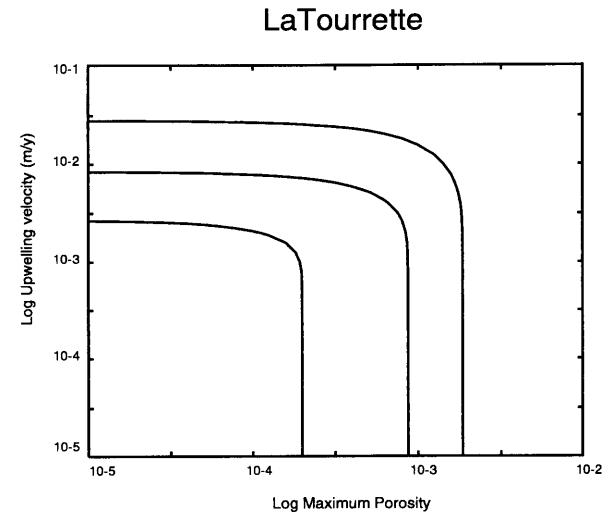
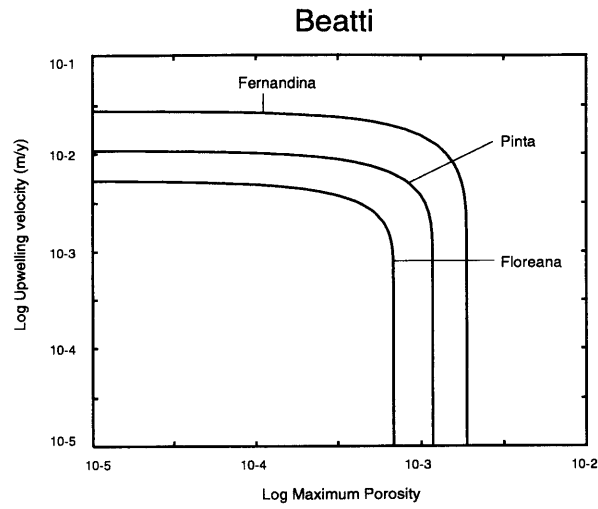


Figure 14

# Equilibrium Porous Flow

Extent of Melting

0.5%  
(1.7km)



1%  
(3 km)

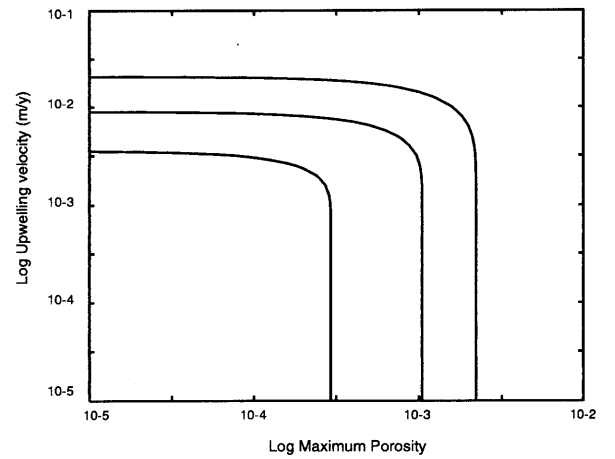
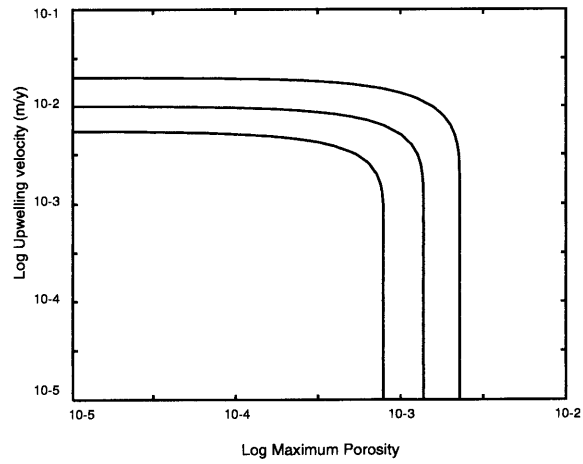


Figure 15

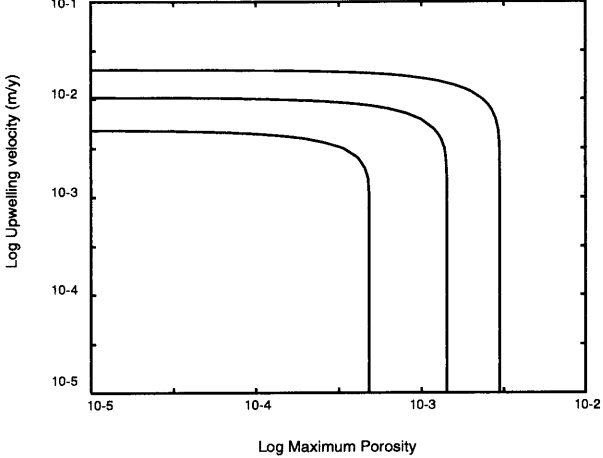
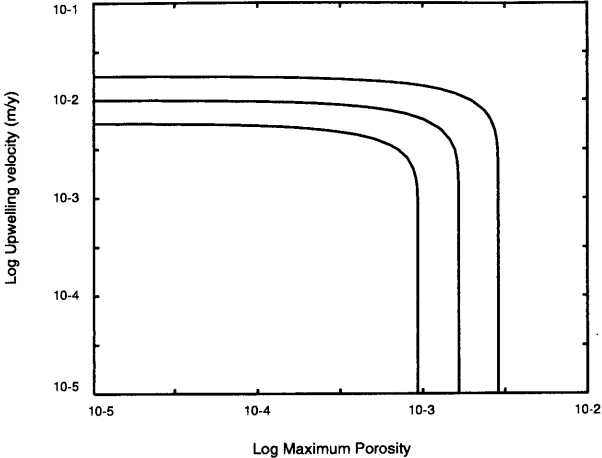
# Equilibrium Porous Flow

Extent of Melting

Beatti

LaTourrette

5%  
(17 km)



10%  
(33 km)

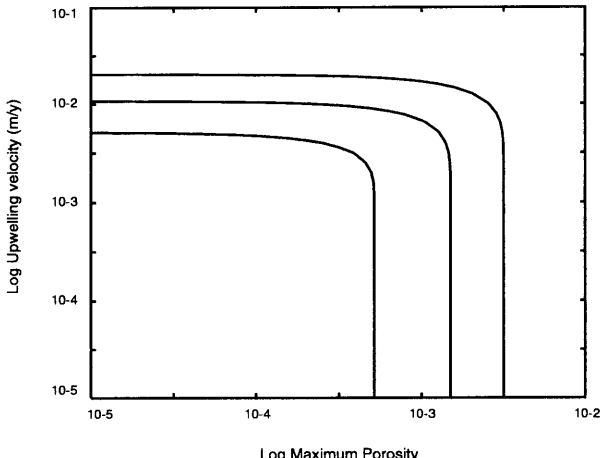
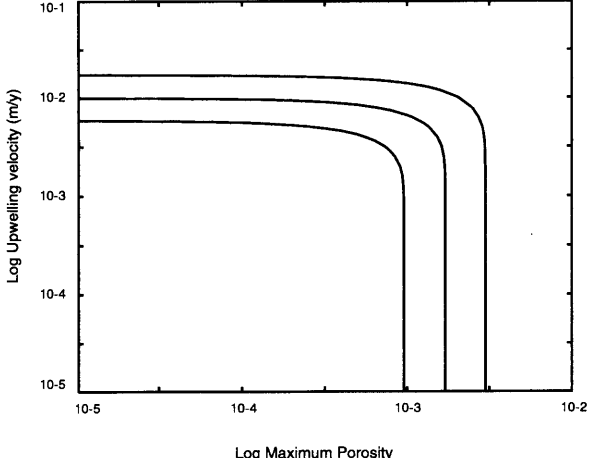


Figure 15

## Conclusions

The main goal of this study has been to use geochemical tracers to constrain the geologic processes responsible for the composition, formation and age of geochemically distinct Earth reservoirs. The strategy has been to work in a broad range of topics which are relevant to this end. This dissertation focus on four independent geochemical problems: defining the composition, formation and age of the lower continental crust and mantle lithosphere using Re-Os systematics, evaluating the scale of the mantle heterogeneities through a melt inclusion study from oceanic island basalts, and constraining the geodynamic parameters during upwelling and melting of mantle plume using U-series isotopes.

The Re-Os systematics in the deeper parts of the continental crust indicate that the lower crust has, on average, higher Os ( $\geq 2$  times) and lower Re concentrations than current estimates of the average upper crust, and may thus be comparatively less radiogenic. Materials with relatively high Os concentrations and radiogenic Os isotopic compositions exist in the lower crust, and may act as contaminants to basaltic magmas that differentiate there. The Re-Os system is extremely sensitive to AFC processes at relatively small degrees of differentiation (i.e.,  $\ll 20\%$  AFC), and AFC processes are more effective than bulk mixing in changing Os isotopic compositions of melts, due to Os depletion during fractional crystallization. Like other isotopic systems, the Re-Os system is not a good indicator of the crystallization age of rocks that formed through AFC processes.

Re-Os isotope systematics in Horoman peridotites indicate that  $^{187}\text{Os}/^{188}\text{Os}$  ratios are mainly controlled by variations in Re content, while Re contents are controlled by the variation of a “basaltic component” in the ultramafic rocks. Geochemical, geological and petrographic information are consistent with the hypothesis that the N-Type plagioclase lherzolites and the layered structure of the Horoman massif are produced by refertilization processes. That is, the addition of a garnet-rich mafic cumulate to a depleted peridotite. Re-Os systematics suggest that depletion model ages of  $\approx 1.8$  Ga represent the time of the melting event. The colinearity between mafic and ultramafic rocks in the Re-Os isochron diagram defines an apparent age of  $\approx 1$  Ga.. The similar “ages” determined by Re-Os and Sm-Nd isotopes and the high Re/Os ratios in the most fertile peridotites plotting to the right of the geochron, indicate that the mafic layers and the ultramafic rocks are genetically related by a refertilization process which took place  $\approx 1$  Ga ago.

The melt inclusion Pb isotope data for Mangaia, Tahaa and Pitcairn forms quasi-linear arrays that can be described by simple binary mixing between the HIMU, EMI and EMII end-member compositions and an unradiogenic component common to all three isotopic arrays. Therefore, the geochemical variation in trace and major elements from melt inclusions does not always represent variable melting processes of a single mantle source, but clearly reflect geochemical source heterogeneities as well. Moreover, the observed HIMU Pb isotopic composition of carbonate globules, their ocelli texture, and mineral associations, and the mineral and glass compositions in melt inclusions from two Mangaia basalts provide the first direct evidence for a carbonated mantle source for HIMU basalts. The negative correlation between the Cr content in daughter clinopyroxene and the Pb isotope composition of the melt inclusions from Mangaia and Tahaa suggests that the melt corresponding to the less radiogenic component may have suffered differentiation before its entrapment, or represents a source material in the mantle that has already been differentiated (eclogite?).

The high ( $^{230}\text{Th}/^{238}\text{U}$ ) disequilibrium found in the Galapagos lavas (ranging from 1.16 to 1.81), indicates that melting occurs in the garnet stability field in agreement with the trace element results. The positive correlation between ( $^{230}\text{Th}/^{232}\text{Th}$ ) and ( $^{238}\text{U}/^{232}\text{Th}$ ), with  $^{230}\text{Th}$  excesses decreasing toward higher ( $^{230}\text{Th}/^{232}\text{Th}$ ) ratios, can be accounted for by mixing of magmas generated by low extents of melting from a fertile source with magmas produced by relatively large extents of melting from a depleted source. The correlations between ( $^{230}\text{Th}/^{232}\text{Th}$ ) and Sr, Nd and Pb isotopes confirm that the Th isotopic ratios are mainly source-controlled and are the result of magma mixing between the plume melts and MORB. Dynamic melting and equilibrium porous flow models constrain mantle upwelling velocities for the Galapagos plume (Fernandina) to  $\leq 3$  cm/y and maximum porosities of 0.3%, indicating that Galapagos is a mildly buoyant plume. Evidence for very slow mantle upwelling and low porosity for Pinta (0.5 to 1 cm/y and 0.15%) and Floreana (0.2 to 0.5 cm/y and 0.1%) supports the hypothesis that the lateral transport of the plume across the 90° 50' transform fault produced additional slow upwelling and further small extent of melting.