

MIT Open Access Articles

Measurement of the ν_e and total 8B solar neutrino fluxes with the Sudbury Neutrino Observatory phase-III data set

The MIT Faculty has made this article openly available. **Please share** how this access benefits you. Your story matters.

Citation: Aharmim, B. et al. "Measurement of the ν_e and total 8B solar neutrino fluxes with the Sudbury Neutrino Observatory phase-III data set." *Physical Review C* 87.1 (2013). ©2013 American Physical Society

As Published: <http://dx.doi.org/10.1103/PhysRevC.87.015502>

Publisher: American Physical Society

Persistent URL: <http://hdl.handle.net/1721.1/78271>

Version: Final published version: final published article, as it appeared in a journal, conference proceedings, or other formally published context

Terms of Use: Article is made available in accordance with the publisher's policy and may be subject to US copyright law. Please refer to the publisher's site for terms of use.





Measurement of the ν_e and total ^8B solar neutrino fluxes with the Sudbury Neutrino Observatory phase-III data set

B. Aharmim,⁶ S. N. Ahmed,¹⁴ J. F. Amsbaugh,¹⁸ J. M. Anaya,⁹ A. E. Anthony,^{16,a} J. Banar,⁹ N. Barros,^{8,19} E. W. Beier,¹³ A. Bellerive,⁴ B. Beltran,¹ M. Bergevin,^{7,5} S. D. Biller,¹² K. Boudjemline,^{4,14} M. G. Boulay,¹⁴ T. J. Bowles,⁹ M. C. Browne,^{18,9} T. V. Bullard,¹⁸ T. H. Burritt,¹⁸ B. Cai,¹⁴ Y. D. Chan,⁷ D. Chauhan,⁶ M. Chen,¹⁴ B. T. Cleveland,¹² G. A. Cox,¹⁸ C. A. Currat,⁷ X. Dai,^{14,12,4} H. Deng,¹³ J. A. Detwiler,⁷ M. DiMarco,¹⁴ P. J. Doe,¹⁸ G. Doucas,¹² M. R. Dragowsky,^{9,7,b} P.-L. Drouin,⁴ C. A. Duba,¹⁸ F. A. Duncan,^{15,14} M. Dunford,^{13,c} E. D. Earle,¹⁴ S. R. Elliott,^{9,18} H. C. Evans,¹⁴ G. T. Ewan,¹⁴ J. Farine,^{6,4} H. Fergani,¹² F. Fleurot,⁶ R. J. Ford,^{15,14} J. A. Formaggio,^{11,18} M. M. Fowler,⁹ N. Gagnon,^{18,9,7,12} J. V. Germani,^{18,9} A. Goldschmidt,^{9,d} J. T. M. Goon,¹⁰ K. Graham,^{4,14} E. Guillian,¹⁴ S. Habib,¹ R. L. Hahn,³ A. L. Hallin,¹ E. D. Hallman,⁶ A. A. Hamian,¹⁸ G. C. Harper,¹⁸ P. J. Harvey,¹⁴ R. Hazama,^{18,e} K. M. Heeger,^{18,f} W. J. Heintzelman,¹³ J. Heise,^{2,9,14,g} R. L. Helmer,¹⁷ R. Henning,^{7,h} A. Hime,⁹ C. Howard,¹ M. A. Howe,¹⁸ M. Huang,^{16,6,i} P. Jagam,⁵ B. Jamieson,² N. A. Jelley,¹² K. J. Keeter,^{15,j} J. R. Klein,^{16,13} L. L. Kormos,¹⁴ M. Kos,¹⁴ A. Krüger,⁶ C. Kraus,^{14,6} C. B. Krauss,¹ T. Kutter,¹⁰ C. C. M. Kyba,^{13,k} R. Lange,³ J. Law,⁵ I. T. Lawson,^{15,5} K. T. Lesko,⁷ J. R. Leslie,¹⁴ J. C. Loach,^{12,7} R. MacLellan,¹⁴ S. Majerus,¹² H. B. Mak,¹⁴ J. Maneira,⁸ R. Martin,^{14,7} N. McCauley,^{13,12,l} A. B. McDonald,¹⁴ S. R. McGee,¹⁸ C. Mifflin,⁴ G. G. Miller,⁹ M. L. Miller,^{11,m} B. Monreal,^{11,n} J. Monroe,¹¹ B. Morissette,¹⁵ A. W. Myers,^{18,o} B. G. Nickel,⁵ A. J. Noble,^{14,4} H. M. O'Keefe,^{12,p} N. S. Oblath,^{18,11} R. W. Ollerhead,⁵ G. D. Orebi Gann,^{12,13} S. M. Oser,² R. A. Ott,¹¹ S. J. M. Peeters,^{12,q} A. W. P. Poon,⁷ G. Prior,^{7,c} S. D. Reitzner,⁵ K. Rielage,^{9,18} B. C. Robertson,¹⁴ R. G. H. Robertson,¹⁸ E. Rollin,⁴ M. H. Schwendener,⁶ J. A. Secrest,^{13,r} S. R. Seibert,^{16,9,13} O. Simard,⁴ J. J. Simpson,⁵ P. Skensved,¹⁴ M. W. E. Smith,^{18,9} T. J. Sonley,^{11,p} T. D. Steiger,¹⁸ L. S. Stonehill,^{9,18} G. Tešić,⁴ P. M. Thornewell,^{12,9} N. Tolich,¹⁸ T. Tsui,² C. D. Tunnell,^{16,s} T. Van Wechel,¹⁸ R. Van Berg,¹³ B. A. VanDevender,^{18,o} C. J. Virtue,⁶ B. L. Wall,¹⁸ D. Waller,⁴ H. Wan Chan Tseung,^{12,18} J. Wendland,² N. West,¹² J. B. Wilhelmy,⁹ J. F. Wilkerson,^{18,h} J. R. Wilson,^{12,t} J. M. Wouters,^{9,u} A. Wright,^{14,v} M. Yeh,³ F. Zhang,⁴ and K. Zuber^{12,w}

(SNO Collaboration)

¹Department of Physics, University of Alberta, Edmonton, Alberta, Canada T6G 2R3

²Department of Physics and Astronomy, University of British Columbia, Vancouver, British Columbia, Canada V6T 1Z1

³Chemistry Department, Brookhaven National Laboratory, Upton, New York 11973-5000, USA

⁴Ottawa-Carleton Institute for Physics, Department of Physics, Carleton University, Ottawa, Ontario, Canada K1S 5B6

⁵Physics Department, University of Guelph, Guelph, Ontario, Canada N1G 2W1

⁶Department of Physics and Astronomy, Laurentian University, Sudbury, Ontario, Canada P3E 2C6

⁷Institute for Nuclear and Particle Astrophysics and Nuclear Science Division, Lawrence Berkeley National Laboratory, Berkeley, California 94720, USA

⁸Laboratório de Instrumentação e Física Experimental de Partículas, Avenida Elias Garcia 14, 1°, Lisboa, Portugal

⁹Los Alamos National Laboratory, Los Alamos, New Mexico 87545, USA

¹⁰Department of Physics and Astronomy, Louisiana State University, Baton Rouge, Louisiana 70803, USA

¹¹Laboratory for Nuclear Science, Massachusetts Institute of Technology, Cambridge, Massachusetts 02139, USA

¹²Department of Physics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford OX1 3RH, United Kingdom

¹³Department of Physics and Astronomy, University of Pennsylvania, Philadelphia, Pennsylvania 19104-6396, USA

¹⁴Department of Physics, Queen's University, Kingston, Ontario, Canada K7L 3N6

¹⁵SNOLAB, Sudbury, Ontario, Canada P3Y 1M3

¹⁶Department of Physics, University of Texas at Austin, Austin, Texas 78712-0264, USA

¹⁷TRIUMF, 4004 Wesbrook Mall, Vancouver, British Columbia, Canada V6T 2A3

¹⁸Center for Experimental Nuclear Physics and Astrophysics, and Department of Physics, University of Washington, Seattle, Washington 98195, USA

¹⁹Institut für Kern- und Teilchenphysik, Technische Universität Dresden, Dresden, Germany

(Received 14 July 2011; revised manuscript received 16 April 2012; published 18 January 2013)

This paper details the solar neutrino analysis of the 385.17-day phase-III data set acquired by the Sudbury Neutrino Observatory (SNO). An array of ^3He proportional counters was installed in the heavy-water target to measure precisely the rate of neutrino-deuteron neutral-current interactions. This technique to determine the total active ^8B solar neutrino flux was largely independent of the methods employed in previous phases. The total flux of active neutrinos was measured to be $5.54_{-0.31}^{+0.33}(\text{stat.})_{-0.34}^{+0.36}(\text{syst.}) \times 10^6 \text{ cm}^{-2} \text{ s}^{-1}$, consistent with previous measurements and standard solar models. A global analysis of solar and reactor neutrino mixing parameters yielded the best-fit values of $\Delta m^2 = 7.59_{-0.21}^{+0.19} \times 10^{-5} \text{ eV}^2$ and $\theta = 34.4_{-1.2}^{+1.3}$ degrees.

DOI: [10.1103/PhysRevC.87.015502](https://doi.org/10.1103/PhysRevC.87.015502)

PACS number(s): 26.65.+t, 14.60.Pq, 13.15.+g, 95.85.Ry

I. INTRODUCTION

The Sudbury Neutrino Observatory (SNO) experiment [1] has firmly established that electron-type neutrinos (ν_e) produced in the solar core transform into other active flavors while in transit to Earth [2–7]. This direct observation of neutrino flavor transformation, through the simultaneous observation of the disappearance of ν_e and the appearance of other active neutrino types, confirmed the total solar neutrino flux predicted by solar models [8,9] and explained the deficit of solar neutrinos that was seen by other pioneering experiments [10–14]. The SNO results, when combined with other solar neutrino experiments and reactor antineutrino results from

the KamLAND experiment [15], demonstrated that neutrino oscillations [16–18] are the cause of this flavor change.

In the first two phases of the SNO experiment the determination of the total active ^8B solar neutrino flux and its ν_e component required a statistical separation of the Cherenkov signals observed by the detector’s photomultiplier tube (PMT) array. In the third phase of the experiment, an array of ^3He proportional counters [19] was deployed in the detector’s heavy-water target. The neutron signal in the inclusive total active neutrino flux measurement was detected predominantly by this “Neutral-Current Detection” (NCD) array and was separate from the Cherenkov light signals observed by the PMT array in the ν_e flux measurement. This separation of neutrino signals resulted in reduced correlations between the total active neutrino flux and ν_e flux measurements, allowing an improved determination of the solar neutrino mixing angle. This technique to measure the total active ^8B solar neutrino flux was largely independent of the methods employed by SNO in previous phases.

The results from the third phase of the SNO experiment were reported in a letter [7] and confirmed those from previous phases. We present in this article the details of this measurement and an analysis of the neutrino oscillation parameters. In Sec. II we present an overview of the SNO experiment and the solar neutrino measurement with the NCD array. The phase-III data set that was used in this measurement is described in Sec. III. A thorough understanding of the responses of the PMT array and the NCD array is important to the determination of the ν_e and the total active neutrino flux. Details of the optical response of the PMT array, the reconstruction of its data and the determination of the effective electron kinetic energy scale are provided in Sec. IV. The electronic and energy responses of the NCD array are discussed in Sec. V. The determination of the neutron detection efficiencies for both the PMT and the NCD arrays, which are crucial to the measurement of the total active solar neutrino flux, is presented in Sec. VI. The evaluation of backgrounds in the solar neutrino flux measurement, with the exception of α decays in the construction materials in the NCD array, is summarized in Sec. VII. The α decays were a non-negligible background for the detection of signal neutrons and were treated differently in the analysis. An extensive pulse-shape simulation was developed to understand the response of the NCD counters to these α decays, as presented in Sec. VIII. In Sec. IX, we discuss the analysis that determined the total active solar neutrino flux and the electron-type neutrino flux. These measured fluxes, along with results from previous SNO measurements and other solar and reactor neutrino experiments, were then used in the determination of the neutrino mixing parameters as described in Sec. X. A description of the cuts we used to remove instrumental backgrounds in the NCD array data can be found in Appendix A. A discussion of the parametrization of nuisance parameters in the neutrino flux analysis is provided in Appendix B.

II. OVERVIEW OF THE SNO EXPERIMENT

The SNO experiment was unique in its capability to detect all active neutrino flavors with equal sensitivity, which allowed

^aPresent address: Center for Astrophysics and Space Astronomy, University of Colorado, Boulder, Colorado, USA.

^bPresent address: Department of Physics, Case Western Reserve University, Cleveland, Ohio, USA.

^cPresent address: CERN, Geneva, Switzerland.

^dPresent address: Lawrence Berkeley National Laboratory, Berkeley, California, USA.

^ePresent address: Department of Physics, Hiroshima University, Hiroshima, Japan.

^fPresent address: Department of Physics, University of Wisconsin, Madison, Wisconsin, USA.

^gPresent address: Sanford Laboratory at Homestake, Lead, South Dakota, USA.

^hPresent address: Department of Physics, University of North Carolina, Chapel Hill, North Carolina, USA.

ⁱPresent address: Center of Cosmology and Particle Astrophysics, National Taiwan University, Taiwan.

^jPresent address: Black Hills State University, Spearfish, South Dakota, USA.

^kPresent address: Institute for Space Sciences, Freie Universität Berlin, Leibniz-Institute of Freshwater Ecology and Inland Fisheries, Germany.

^lPresent address: Department of Physics, University of Liverpool, Liverpool, United Kingdom.

^mPresent address: Center for Experimental Nuclear Physics and Astrophysics, and Department of Physics, University of Washington, Seattle, Washington 98195, USA.

ⁿPresent address: Department of Physics, University of California, Santa Barbara, California, USA.

^oPresent address: Pacific Northwest National Laboratory, Richland, Washington, USA.

^pPresent address: Department of Physics, Queen’s University, Kingston, Ontario, Canada.

^qPresent address: Department of Physics and Astronomy, University of Sussex, Brighton, United Kingdom.

^rPresent address: Department of Chemistry and Physics, Armstrong Atlantic State University, Savannah, Georgia, USA.

^sPresent address: Department of Physics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford OX1 3RH, United Kingdom.

^tPresent address: Department of Physics, Queen Mary University, London, United Kingdom.

^uDeceased.

^vPresent address: Department of Physics, Princeton University, Princeton, New Jersey, USA.

^wPresent address: Institut für Kern- und Teilchenphysik, Technische Universität Dresden, Dresden, Germany.

the experiment to demonstrate neutrino flavor transformation conclusively. In this section, a description of the SNO detector and its neutrino detection channels are provided. This is followed by a discussion of the physical characteristics of the NCD array, which was used to measure the total active solar neutrino flux in phase III.

A. The SNO detector

The SNO detector was located in Vale's Creighton Mine (46°28'30" N latitude, 81°12'04" W longitude) near Sudbury, Ontario, Canada. The center of this real-time heavy-water ($^2\text{H}_2\text{O}$, D_2O hereafter) Cherenkov detector was at a depth of 2092 m (5890 ± 94 m of water equivalent). At this depth, the rate of cosmic-ray muons entering the detector was approximately three per hour. The solar neutrino target was 1000 metric tons (tonnes) of 99.92% isotopically pure D_2O contained inside a 12-m-diameter acrylic vessel (AV). An array of 9456 20-cm Hamamatsu R1408 PMTs, installed on an 18-m-diameter stainless steel geodesic structure (PSUP), was used to detect Cherenkov radiation in the target. A nonimaging light concentrator [20] was mounted on each PMT to increase the effective photocathode coverage to nearly 55% of 4π . The AV and the PSUP were suspended in an underground cavity filled with approximately 7 kilotonnes of ultrapure light water (H_2O), which shielded the D_2O volume against radioactive backgrounds from the cavity rock. The inner 1.7 kilotonnes of H_2O between the AV and the PSUP also shielded the target against radioactive backgrounds from the geodesic structure and PMTs. On the outer surface of the PSUP, 91 outward-facing PMTs were installed to tag cosmic-ray events. An array of 23 PMTs was mounted in a rectangular frame that was suspended facing inwards in the outer H_2O region. These PMTs, along with the 8 PMTs installed in the neck region of the AV, were used to reject instrumental background light. A full description of the SNO detector can be found in Ref. [1]. In the third phase of the SNO experiment, an array of ^3He proportional counters was deployed in the D_2O volume. Details of this array are presented in Sec. II B and in Ref. [19].

The SNO detector detected solar neutrinos through the following processes:

$$\text{CC} : \nu_e + d \rightarrow p + p + e^- - 1.442 \text{ MeV},$$

$$\text{NC} : \nu_x + d \rightarrow p + n + \nu_x - 2.224 \text{ MeV},$$

$$\text{ES} : \nu_x + e^- \rightarrow \nu_x + e^-,$$

where ν_x refers to any active neutrino flavor ($x = e, \mu, \tau$). The charged-current (CC) reaction is sensitive exclusively to ν_e , whereas the neutral-current (NC) reaction is equally sensitive to all active neutrino flavors. Chen [21] realized that the NC measurement of the total active solar neutrino flux tests the solar model predictions independently of the neutrino-oscillation hypothesis, and a comparison of this flux to the CC measurement of ν_e flux tests neutrino flavor transformation independently of solar models. The neutrino-electron elastic scattering (ES) reaction is used to observe neutrinos of all active flavors in SNO and other real-time water Cherenkov and liquid scintillator detectors. Its cross section for ν_e is approximately six times larger than ν_μ and ν_τ for ^8B solar

neutrinos, but is smaller than the CC or NC cross sections in the energy region of interest.

In the first phase of the SNO experiment, which used an unadulterated D_2O target, NC interactions were observed by detecting the 6.25-MeV γ ray following the capture of the neutron by a deuteron. Under the assumption of an undistorted ^8B neutrino spectrum, the hypothesis of the observed CC, NC, and ES rates due solely to ν_e interactions was rejected at 5.3σ . Details of the solar neutrino analysis in this phase can be found in Ref. [5]. Approximately 2 tonnes of sodium chloride (NaCl) were added to the D_2O in the second phase of the SNO experiment. This addition enhanced the neutron detection efficiencies and allowed the statistical separation of CC and NC signals without making any assumption about the energy dependence of neutrino flavor change. As a result the accuracy of the ν_e and the total active neutrino flux measurements were significantly improved. A full description of the solar neutrino analysis in this phase can be found in Ref. [6]. Recently, the results of a solar neutrino analysis that combined the phase-I and phase-II data sets were reported in Ref. [22].

B. The Neutral-Current Detection array

The NCD array, consisting of 36 strings of ^3He and 4 strings of ^4He proportional counters, was deployed in the D_2O target, after the removal of NaCl, in the third phase of the experiment. Figure 1 shows a side view of the SNO detector with the NCD array in place. The NCD counter strings were arranged on a square grid with 1-m spacing as shown in Fig. 2. The acrylic anchors to which the NCD strings were attached were bonded to the AV during its construction, and their positions were surveyed precisely by laser theodolite and were taken as reference. Details of the deployment of the NCD string can be found in Ref. [19].

Each NCD string was 9 to 11 m in length and was made up of three or four individual 5-cm-diameter counters that were laser welded together. Ultralow-radioactivity nickel, produced by a chemical vapor deposition (CVD) process, was used in the construction of the counter bodies and end caps. This process suppressed all but trace amounts of impurities. The nominal thickness of the counter wall was $\sim 370 \mu\text{m}$. Each counter was strung with a 50- μm -diameter low-background copper wire that was pretensioned with a 30-g mass. The gas in the counters was a mixture of 85% ^3He (or ^4He) and 15% CF_4 (by pressure) at 2.5 atmospheres (1900 Torr).

The coordinates of the top of the NCD strings were not *a priori* known because they were pulled slightly out of the anchors' reference positions by the attached NCD cables (see Fig. 1). There were altogether three campaigns to measure the positions. Before and after data taking in phase III, laser range-finder (LRF) surveys were made of the counter tops using a custom-built LRF that could be introduced through the AV neck and immersed in the D_2O . During data taking, optical reconstruction of the average positions was obtained from the shadowing of calibration sources (see Sec. IV A). These three measurements were in good agreement, with the shadowing results being about two times more precise than the LRF results.

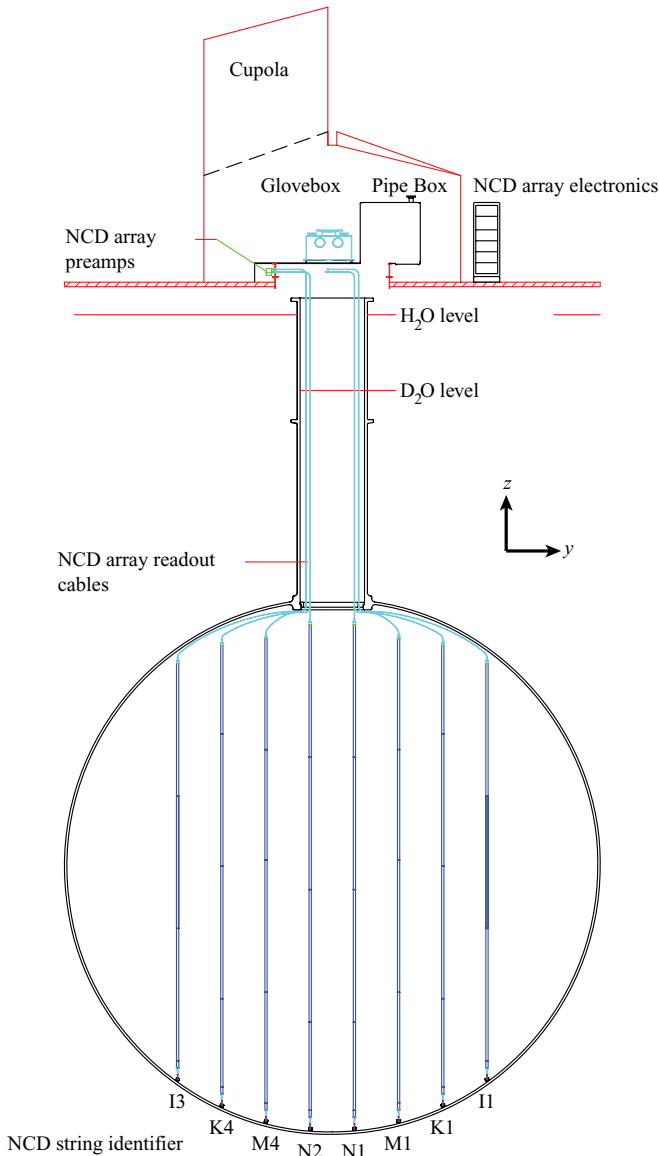
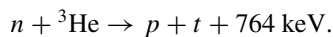


FIG. 1. (Color online) Side view of the SNO detector in phase III. The center of the AV is the origin of the Cartesian coordinate system used in this paper. The NCD counter strings were arranged on a square grid with 1-m spacing (shown in Fig. 2). Each string was made up of three or four individual 5-cm-diameter counters that were laser welded together. Weld joints are depicted along the length of the strings in this figure. Only the first row of NCD strings from the y - z plane are displayed in this figure.

Neutrons from the NC reaction were detected in the NCD array via the reaction



The ${}^4\text{He}$ strings were not sensitive to neutrons and were used to characterize non-neutron backgrounds in the array.

During normal operation the anode wires were maintained at 1950 V, resulting in a gas gain of ~ 220 . A primary ionization of the counter gas would trigger an avalanche of secondary ionizations, which led to a current pulse on the anode wire. This pulse traveled in both directions, up and down a NCD

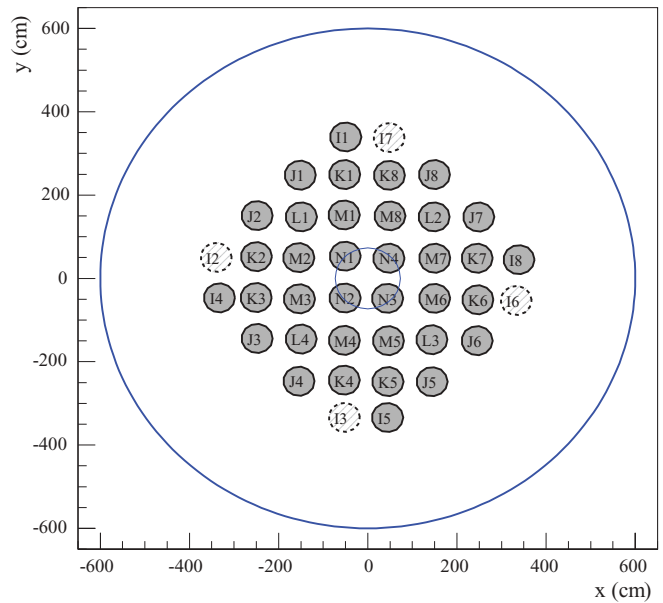


FIG. 2. (Color online) The positions of the strings in the NCD array projected onto the plane of the AV equator (x - y plane). The array was anchored on a square lattice with a 1-m grid constant. The strings labeled with the same letter denote strings of the same length and distance from the center of the AV. Strings I2, I3, I6, and I7 contained ${}^4\text{He}$ instead of ${}^3\text{He}$. The outer circle is the AV equator and the inner circle is the neck of the AV, through which the NCD strings were deployed. The NCD string markers are not drawn to scale.

string. The delay line at the bottom of each string added approximately 90 ns to the travel time of the portion of the pulse that traveled down the string. The termination at the end of the delay line was open, so the pulse was reflected without inversion. The direct and reflected portions of the pulse were separated by approximately 90–350 ns, depending on the origin of the pulse along the length of the NCD string. At the top of each NCD string was a 93- Ω impedance coaxial cable that led to a current preamplifier. This preamplifier linearly transformed the signal to a voltage amplitude with a gain of 27.5 mV/ μA .

The NCD array had two independently triggered readout systems. The “Shaper-ADC” system used a pulse-shaping and peak-detection network to integrate the signal pulse and measure its energy. This fast system was triggered by the pulse integral crossing a threshold and could handle the kilohertz event rates expected from a galactic supernova. The “Multiplexer-Scope” (MUX-scope) system digitized and recorded the entire 15- μs pulse. It consisted of four independent sets of electronics, or MUX boxes, each of which could accept signals from up to 12 strings. Each channel was triggered by the pulse amplitude crossing a threshold. Pulses were amplified by the logarithmic amplifier (“log-amp”) in a MUX box to increase the range of pulse sizes that could be digitized at a 1-GHz sampling rate by the 8-bit digitizer in one of the two digital oscilloscopes. The multiplexer controller triggered the oscilloscope that was not busy (or toggled between them when neither was busy), allowing for a maximum digitization rate of 1.8 Hz. If the oscilloscopes

were busy and the MUX system triggered, a “partial MUX” event was recorded without the digitized pulse. The MUX-scope system adopted in the 1990s is not a solution to be recommended today, but it was sufficient to handle typical solar-neutrino signal and background event rates. Signals from the PMT array and from these two readout systems of the NCD array were integrated in a global trigger system that combined the data streams with event timing information. Further details of the design and construction of these counters and their associated electronic systems can be found in Ref. [19].

III. DATA SET

A. Data selection

The measurements reported here are based on analysis of 385.17 ± 0.14 live days of data recorded between November 27, 2004, and November 28, 2006. Although the SNO detector was live nearly 90% of the time during this data-acquisition period, data from calibration periods, roughly 30% of the detector live time, were not used for solar neutrino analysis. Other losses of live time resulted from detector maintenance periods, and the loss of underground laboratory power, communication, or environmental systems. The selection of solar neutrino data runs for analysis was based on evaluation as outlined in previous papers [5,6]. In addition to the offline inspections of run data from the PMT array, data-quality checks of the NCD array data were implemented. These checks validated the running condition, such as the trigger thresholds, of the NCD array.

To accurately determine the total active solar neutrino flux using the NCD array it was essential to utilize only data from strings that were operating properly. Six ^3He strings were defective and their data were excluded in the analysis. One of the counters in the string K5 was slowly leaking ^3He into an intercounter space. Two strings, K2 and M8, had mechanical problems with the resistive coupling to the top of the counter string, as confirmed in postmortem examination at the end of the experiment, resulting in unstable responses. The string K7 showed similar behavior, but a physical examination of this string at the end of the experiment did not indicate a loose coupling. The strings J3 and N4 were observed to produce anomalous instrumental background events in the neutron signal window. A loose resistive coupling was found during a physical examination of J3, but not in N4. The removal of these problematic strings in the solar neutrino analysis did not have significant impact to the statistical uncertainties of the measurement, as several strings were in the outer region of the detector where the neutron capture efficiency was low.

A variety of other kinds of instrumental events were recorded in the shaper and digitized-data paths. Data reduction cuts were developed to remove these instrumental backgrounds. In Appendix A examples of these background events and a summary of the cuts to remove them are provided. Physics events in a counter would trigger both the shaper-ADC and the MUX-scope subsystems; thus, a large fraction of instrumental backgrounds was removed simply by accepting

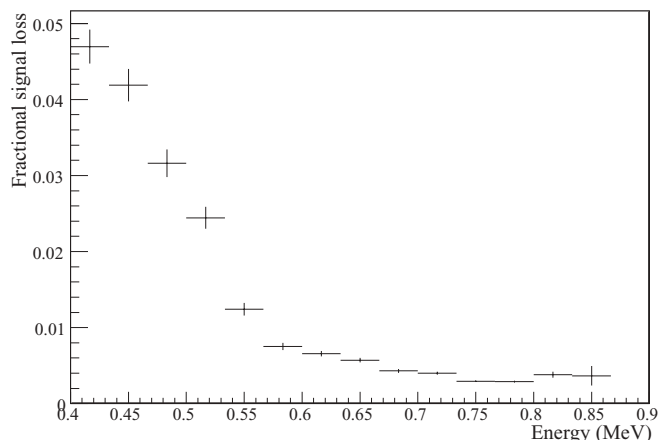


FIG. 3. Energy dependence of fractional signal loss of NCD array data owing to instrumental background cuts, as measured with ^{252}Cf and Am-Be neutron calibration data. The abscissa is the event energy recorded by the shaper-ADC system. The neutron signal, as shown in Fig. 22(d) peaks at the Q value of 764 keV. The fractional signal loss was sub-1% near the neutron peak.

only events with both triggers present. NCD array events that passed this selection criterion were subsequently analyzed by algorithms that were designed to identify nonionization pulses such as microdischarges and oscillatory noise. Two independent sets of cuts were developed. One of these sets examined the logarithmically amplified digitized wave forms in the time domain, while the other set utilized the frequency domain. The two sets of cuts were shown to overlap substantially, with 99.46% of cut events removed by both sets of cuts, 0.02% removed only by the time-domain cuts, and 0.52% removed only by the frequency-domain cuts. Both sets of cuts were used in reducing the data set. The number of raw triggers from the NCD array data stream was 1417811, and the data set was reduced to 91636 “NCD events” after application of data reduction cuts. Figure 3 shows the energy dependence of the fractional signal loss determined from ^{252}Cf and Am-Be neutron calibration sources. A suite of instrumental background cuts for the PMT array data was developed in previous phases of the experiment [5,6]. These cuts were reevaluated and recalibrated to ensure their robustness in the analysis of the third-phase data. The number of raw triggers from the PMT array data stream was 146431347, with 2381 “PMT events” passing data reduction and analysis selection requirements similar to those in Ref. [5]. These selected PMT events have reconstructed radial distance $R_{\text{fit}} \leq 550$ cm and reconstructed electron effective kinetic energies $T_{\text{eff}} \geq 6.0$ MeV.

B. Live time

The raw live time of selected runs was calculated from the time differences of the first and last triggered event using the main trigger system’s 10-MHz clock, which was synchronized to a global positioning system. Owing to the combination of the data streams from the PMT and NCD arrays, a run boundary cut was applied to ensure that both systems were taking data by defining the start of the run as 1.1 s after the first event

of either array, whichever came later. A reverse order cut was applied to define the end of each run. These calculated times were verified by comparing the results against those measured by a 50-MHz detector-system clock and a 10-MHz clock used by the NCD array's trigger system.

Several data selection cuts removed small periods of time from the data set during normal data taking conditions. The largest of these removed time intervals following high-energy cosmic-ray events and intervals containing time-correlated instrumental events. These cuts removed events from both the PMT and the NCD arrays. The final live time for the neutrino analysis was calculated by subtracting the total time removed by these cuts from the raw live time. This resulted in a reduction of 1.96% of the raw live time.

The final live time was checked with analyses of data from two detector diagnostic triggers: the pulsed global trigger (PGT) and the NCD system's random pulser (NRP). The PGT was a detectorwide trigger issued at a frequency of 5 Hz based on timing from the 50-MHz system clock. The NRP randomly pulsed a spare channel on the NCD system at an average rate of 7.75 mHz. Systematic uncertainties in live time were evaluated by comparing the PGT and NRP measurements to the 10-MHz clock measurement. The total live-time uncertainty was calculated to be 0.036%.

The "day" data set, which was acquired when the solar zenith angle $\cos\theta_z > 0$, has a live time of 176.59 days. The "night" data set, for which $\cos\theta_z \leq 0$, has a live time of 208.58 days. Detector maintenance work, calibrations, and radiochemical assays were generally performed during daylight hours. Data taken during these activities were not included in the solar neutrino analyses. In addition, seasonal variations in the lengths of day and night, when convolved with the SNO detector's exposure period, introduced an additional difference in the day and night live times.

IV. RESPONSE OF THE PMT ARRAY

In the solar neutrino measurement, the SNO PMT array observed Cherenkov radiation from high-energy electrons resulting from direct neutrino interactions, β decays of radioactive backgrounds, and Compton scattering of γ rays from nuclear deexcitations and radiative captures. A thorough understanding of the propagation and the detection of Cherenkov photons in the SNO detector was vital to reconstruct the observables, such as energy and vertex position, of each triggered event in the PMT array. These observables were used in the analysis of the solar neutrino flux in this paper. Details on the extensive optical and energy calibration of the PMT array in previous phases of SNO can be found in Refs. [5,6].

In the third phase of SNO, optical and energy calibration procedures were modified from those in previous phases to account for complexities that did not exist before. The nickel body of the NCD strings scattered and absorbed Cherenkov photons. The orientation of the NCD array and its signal cables also accentuated the vertical (z) asymmetry in the response of the PMT array. These effects had to be incorporated in the reconstruction of Cherenkov events to determine precisely the energy and spatial distributions of neutrino signals and

radioactive backgrounds in the heavy-water detector. In this section, we first present the response of the PMT array to optical photons in Sec. IV A, taking into account the optical effects introduced by the presence of the NCD array. This is followed by a study of event vertex reconstruction, which required the arrival time of the detected photons as input, in Sec. IV B, and a study of the energy response of the PMT array, which depended on reconstructed event vertex position, in Sec. IV C.

A. Optical response

The measurement of the SNO optical model parameters (see Sec. IV A of Ref. [6]) was done by a χ^2 fit that minimized the differences between the measured and predicted light intensities at each PMT, for a set of calibration runs ("scan") taken with a diffuse, near-isotropic light source (the "laserball") [23] at a number of positions inside the D_2O volume. These calibration methods had to be significantly modified with respect to the previous phases to take into account the optical effects of the NCD array. The construction of a high-isotropy laserball source and the optimization of the calibration plans enabled good sampling of the PMT array despite partial shadowing from the NCD array. The changes in the analysis of the calibration data, and its results, are described below.

1. Source and NCD string positions

The source position in individual laserball runs was determined by minimizing the differences between the calculated and measured PMT prompt peak time, similar to what was done previously, adding cuts to remove PMTs with very few counts. The removed PMTs were likely to have been hit by scattered light or were in the shadow of the NCD array.

In SNO phase III, the laserball source positions and the shadowing patterns observed in the PMT array were used to obtain the NCD string positions, which were, in turn, compared to the installation reference, or nominal, positions. These reference positions were the locations of the NCD string anchors that were surveyed by laser theodolite during the construction of the AV. The method selected an ensemble of about 30 source positions to triangulate and reconstruct the position of each NCD string [24]. For each NCD string, sets of PMTs were selected so that the light path between them and the laserball lay near to the string in x and y coordinates. The PMT occupancy for that run was filled in a two-dimensional map corresponding to the x - y line including the point of closest approach of the source-PMT path to the NCD string, and all the other points along that source-PMT direction (projected in the x - y plane). Because the count rate depended on the conditions of a run (such as laser pulse rate, source stability, PMT thresholds, source position, etc.), the occupancy of the selected PMTs in each run was normalized by their mean occupancy. A map of these relative occupancies was built and fitted to a two-dimensional Gaussian function. Figure 4 shows an example of a reconstructed string position. The extraction of all NCD string positions gave an average difference between

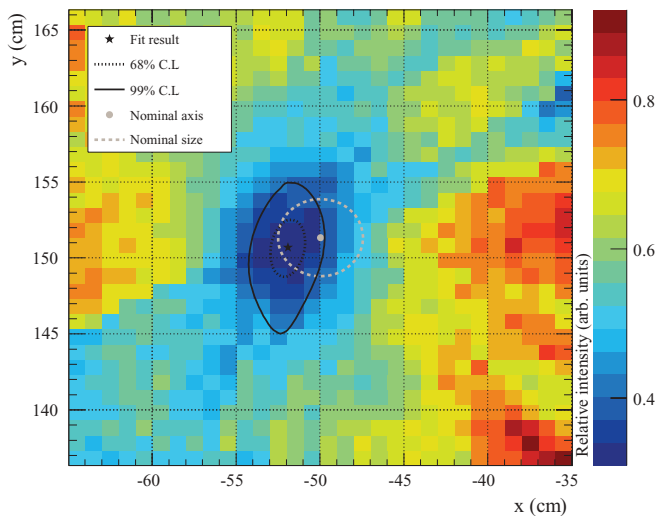


FIG. 4. (Color) An example of NCD string position reconstruction from the data of one of the laserball scans at 500 nm. The contours represent the 68% and 99% CL of the fitted position of NCD string M1. Also shown in this figure is the nominal position of the string. The difference between the fitted position and the nominal position is ~ 2 cm in this case.

the fitted and nominal coordinates, the average horizontal displacement, of about 2 cm in the x and y directions, which was consistent with the estimated average uncertainty of the string positions.

The NCD string position fit was repeated by dividing the various trajectories into three z bins, such that the (x, y) coordinates were obtained as a function of z . The slope of (x, y) vs z gave the deviation of the counters from their nominal vertical position. The best-fit angular deviation of all the counters was found to be less than 1° and was consistent with the measured average displacements. Therefore, the rest of the analysis assumed the counters were perfectly vertical but had an average horizontal displacement. This choice simplified the numerical simulation model.

The various systematic uncertainties of source position reconstruction, which summed up to approximately 2 cm, were propagated to get an estimated resolution of the (x, y) coordinates in the triangulation method. The method yielded an average uncertainty of 2.2 ± 0.3 cm on the individual NCD string positions, projected at $z = 0$. The spread of 0.3 cm arose partly because of the geometry of the NCD array with respect to the calibration planes that limited the laserball positioning to $(x, 0, z)$ and $(0, y, z)$ coordinates. Thus, strings in the outer rings were sometimes shadowed by other strings in inner rings, resulting in larger uncertainties for those strings. In addition, the uncertainty contained a small scan-to-scan variation which was taken into account in the spread. In the optical analysis, the 2.2-cm uncertainty and its spread were input parameters to the shadow-removal code that treated all laserball and string pairs in the same way, independently of their position, to remove the shadowed PMTs with an estimated efficiency of 99%. The robustness of the analysis was improved as the effects of scattered light, which could trigger the shadowed PMTs, was mitigated.

2. Optical effects of the NCD array

The attenuation lengths of various optical media and the PMT angular response were determined by analyzing the PMT data that were not affected by unwanted optical effects, including those caused by the NCD array. Given the fitted position of the laserball and NCD strings, the shadowed PMTs were removed from the analysis on a run-by-run basis.

In addition to the shadowing effects from the NCD strings and their signal cables, the anchors that held them down to the bottom of the AV were also taken into account because they were made of UV-absorbing acrylic. The implementation of the anchor cut showed improvements in the determination of the efficiency for PMTs located at the bottom of the detectors.

To handle PMT-to-PMT variations in efficiency in previous phases, the light intensity at a PMT for a given run was always normalized to the intensity of that PMT in a run where the laserball was deployed at the center of the detector. In the third phase, this technique would result in the systematic removal of the shadowed PMTs in that central run. Therefore, the optical model used the PMT relative efficiencies measured in the “preparatory” phase, a run period after the removal of salt but prior to the installation of the NCD array, taking into account changes in PMT gain and threshold. Even when the NCD string shadow and anchor cuts removed many PMTs in the analysis, the new fit method and optimally chosen laserball positions compensated for this loss, resulting in the overall statistics in a given scan at most 50% lower than in previous phases.

Light reflections off the surface of NCD strings were predominantly diffuse, and it was impossible to associate specific trajectories between the source and the PMTs. Therefore, the effect of such reflections on the PMT counts must be estimated and corrected for runs at various source positions on a run-by-run basis. An analytic correction was derived for each source-PMT trajectory by calculating the fraction of solid angle corresponding to the optical paths between the source and a given PMT that included a reflection in one NCD with respect to the direct paths. The correction for diffuse reflection off the NCD strings was found to be less than 5% of a PMT’s occupancy on average. The same correction could be inferred using various MC scenarios, employing the ratio of PMT calibrations from reflection-on and reflection-off simulations. The analytic and MC-based corrections agreed to within 10% and the difference between the two was applied as a systematic uncertainty on the optical parameters.

3. Determination of the optical parameters

In addition to an optical data set taken during the preparatory phase as a reference of the detector state, eight optical calibration scans were performed during the third phase, among which five were selected for use in the analysis. The five selected scans covered enough range in time to monitor any time dependence of the optical constants. This was an improvement over the calibration in previous phases, as there was only one full scan and a partial scan in phase I and three scans in phase II. For each scan, data were taken at six different wavelengths (337, 365, 386, 420, 500, and 620 nm) at multiple source positions in the SNO detector. The time span of the calibration sets allowed us to monitor the

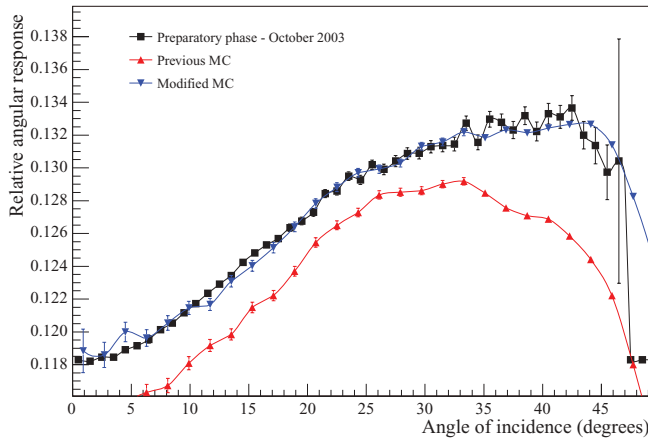


FIG. 5. (Color online) Angular response curves at 386 nm for the PMT-reflector assembly generated by the MC simulation both before and after the optimizations discussed in the text, in comparison to the optical scan in the preparatory phase in October 2003. Note that the y axis zero is suppressed. Details on the design specifications and angular response of the PMT light concentrator can be found in Ref. [20].

stability of the optical model parameters, which included the heavy- and light-water attenuation lengths, the PMT angular response, and the laserball’s light isotropy.

The modeling of the angular response of the PMTs was improved using optical scan data from the preparatory phase. An empirical collection efficiency function in the simulation modifies the response of the PMT as a function of the position at which a photon strikes the photocathode, thus altering the angular response. This function has five tunable parameters, which were previously optimized to reproduce laserball scans at 386 nm, the most probable wavelength for registering a hit in the detector. In this phase, a joint χ^2 fit was performed at all six wavelengths at which laserball calibration data were taken, fitting for both the shape of the response at each wavelength and the relative normalizations. The calibration data at each wavelength were first normalized to the amplitude at normal incidence, and then scaled by the quantum efficiency for a typical PMT at that wavelength. The χ^2 of the empirical collection efficiency function fit at each wavelength was weighted by the relative likelihood of a successful hit being caused by a photon at that wavelength in a Cherenkov light event to optimize the fit at the most probable wavelengths. This resulted in a greater weighting for the more probable wavelengths and a very small weighting for the data at 620 nm, for example, where the probability of a photon triggering a PMT was very low. In previous phases, the empirical collection efficiency reproduced the laserball scans at one wavelength (386 nm) only. In phase III, the angular response shape thus produced by the simulation showed a significant improvement in the agreement with calibration data at all the most probable wavelengths. Figure 5 illustrates this shape and scale modeling improvement at 386 nm.

Figure 6 shows the relative PMT angular response for the five scans in the third phase. The measurement from the preparatory phase is also shown for reference. The average change in the response from the preparatory phase to the third

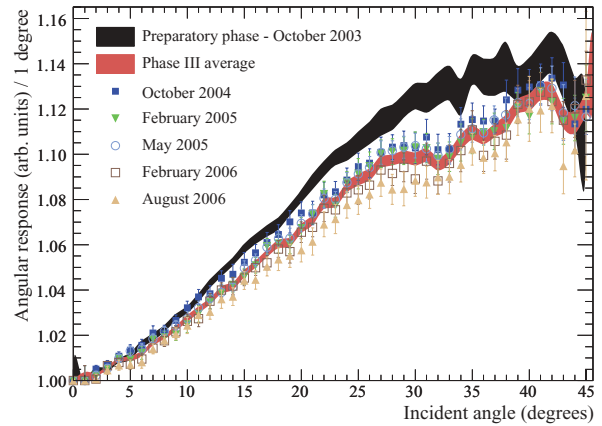


FIG. 6. (Color) PMT relative angular response at 421 nm.

phase, after analyzing the data at all six wavelengths, was about 4% at higher incidence angles. The change from the first to the last scan in phase III was around 2%. The decrease in the response was consistent with observations in previous phases, where it was attributed primarily to aging of the light concentrators.

Figure 7 shows the heavy- and light-water attenuation coefficients obtained at 421 nm. The time dependence of the attenuation lengths was very small, showing better stability than in phase II. Therefore, an average value of the optical

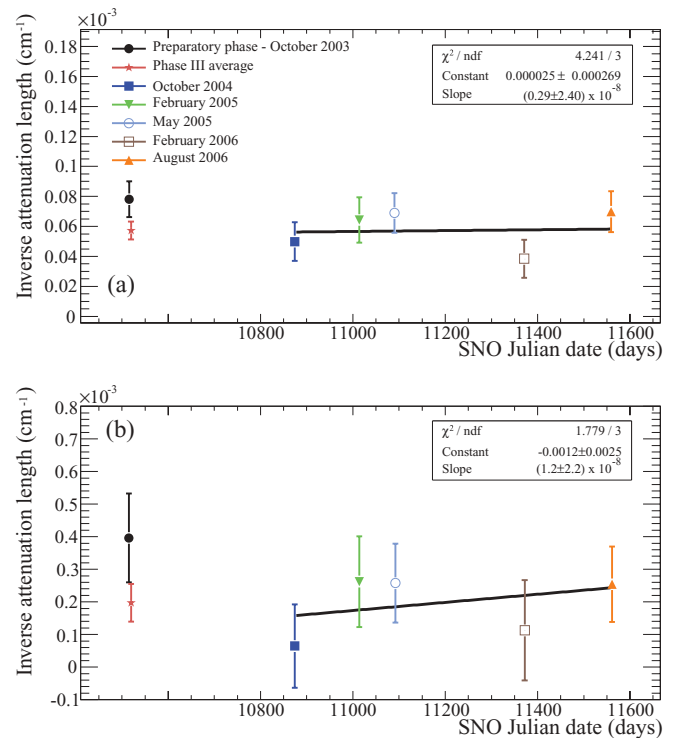


FIG. 7. (Color online) (a) D₂O and (b) H₂O inverse attenuation lengths as a function of time at 421 nm. The lines show the linear fits covering the commissioning and data-taking periods of phase III. In each optical scan, the inverse attenuation length for the two media were fitted simultaneously with other optical parameters, as discussed in the text.

parameters was used in all Monte Carlo simulations in the third phase.

The new systematic uncertainties evaluated for the optical parameters included the precision of the PMT efficiency estimations, the NCD string shadow cut efficiency, the NCD string reflection corrections, and the up-down asymmetry in the PMT array response. The new systematic uncertainties contributed to a 10%–25% increase of the uncertainties in the optical parameters compared to previous phases, and because the total uncertainties were dominated by the systematic component, the loss in statistics induced a negligible increase in the total uncertainties of the optical parameters. After averaging the results from all the analyzed scans, the uncertainties on the D_2O and H_2O attenuation coefficients and PMT response were generally well below 10%, 15%, and 1.5%, respectively. The effect of these uncertainties on event vertex position reconstruction accuracy and energy estimation was estimated to be less than 0.1% and 0.25%, respectively.

B. Vertex reconstruction of Cherenkov events

Algorithms that maximize the likelihood of event vertex position and direction, given the distribution of PMT trigger times and positions, were used to reconstruct Cherenkov events in SNO. The event position and direction were used to separate statistically different detection channels in the solar neutrino analysis. In addition, this information was used to determine the energy of the observed event, which was also used as one of the observables in the flux analysis. The following sections describe event reconstruction in phase III and the methodology used to determine the associated systematic uncertainties. The primary calibration source for evaluating the systematic uncertainties associated with event vertex reconstruction was the ^{16}N source [25], whose signals in the SNO detector have light distribution and timing characteristics for Cherenkov events. The ^8Li β source [26], which had a substantially lower source rate than the ^{16}N source, was also deployed to check the reconstruction performance for single electrons.

1. Event vertex

Vertex reconstruction in the third phase was performed by maximizing the likelihood function

$$\mathcal{L} = \prod_{i=1}^{N_{\text{hits}}} f(t_{\text{res}}|\text{hit}; \vec{r}_{\text{PMT}}, \vec{r}_{\text{fit}}), \quad (1)$$

where t_{res} is the time-of-flight corrected PMT trigger time

$$t_{\text{res}} = t_{\text{PMT}} - t_{\text{fit}} - \frac{|\vec{r}_{\text{fit}} - \vec{r}_{\text{PMT}}|}{c_{\text{avg}}}, \quad (2)$$

where \vec{r}_{fit} and t_{fit} are the reconstructed event position and time, respectively, N_{hits} is the number of selected PMT hits, and c_{avg} ($=21.87$ cm/ns) is the group velocity of the mean detected photon wavelength at 380 nm. The function $f(t_{\text{res}}|\text{hit}; \vec{r}_{\text{PMT}}, \vec{r}_{\text{fit}})$ is the probability density function (PDF) that a particular PMT fires at time t_{res} , given its position \vec{r}_{PMT} and the reconstructed event position. In SNO's third phase, the t_{res} PDF was dependent on \vec{r}_{PMT} and \vec{r}_{fit} owing to partial

or complete shadowing by the NCD array. This shadowing effect was incorporated by generating t_{res} distributions for nonshadowed and completely shadowed PMTs using Monte Carlo (MC) simulations and by interpolating between the two for partially shadowed PMTs using an algorithm that computed shadowing analytically. To reduce the effects of reflected photons, the PDF was approximated as a constant for t_{res} values greater than 15 ns and a time cut of ± 50 ns around the median PMT hit time was imposed. This determination of the PDF with shadowing effects allowed an overall improvement of 5% in spatial resolution.

To evaluate the differences between the true and the reconstructed event vertex positions, or the “vertex shift,” the average reconstructed event position of ^{16}N calibration data relative to the source position was compared to that computed from simulated data. Figure 8 shows the difference between the data and MC vertex shift as a function of the source position for scans along the main axes of the detector. It shows a spread of 4 cm in the three directions, and this value was taken as the vertex shift uncertainty. This uncertainty was found to be correlated to the PMT timing calibration of the detector. An overall offset of 5 cm was also observed in the z direction.

The systematic uncertainty on “vertex scaling,” a position-dependent inward or outward shift of reconstructed position, can have a direct effect on the fiducial volume for events detected by the PMT array. It was measured by determining the range of the slope of a first-order polynomial that allowed the inclusion of 68% of the data points in Fig. 8. This uncertainty was believed to be caused by physical factors such as a mismatch of the speed of light in different media and was expected to be the same in all directions. It was estimated to be 0.9% of the Cartesian coordinates.

Vertex resolution was another systematic uncertainty that could affect the fiducial volume. It was assessed by taking the difference between the data and the MC fitted position resolution for all ^{16}N calibration data taken inside the AV and propagated by smearing the coordinates of simulated events with a Gaussian random variable such that the width increased by the measured discrepancy. Figure 9 shows that this uncertainty varies significantly as a function of the z position of the source. Some smaller fluctuations were also observed for ^{16}N calibration scans in the x - y plane and were incorporated in the analysis. The apparent similarity of the results in x and y directions and the cylindrical symmetry of the detector in the third phase suggested the use of the same parametrization of the systematic effect for these two directions. The uncertainty on the vertex resolution was expressed as a second-order polynomial for x and y directions and as a first-order polynomial for the z direction. Tables I and II present the fitted values for the parameters of these polynomials, along with their associated correlation matrices.

2. Event direction

The direction of events in the detector was determined at a later stage, decoupled from position reconstruction. It relied on a likelihood function composed of a Cherenkov light angular

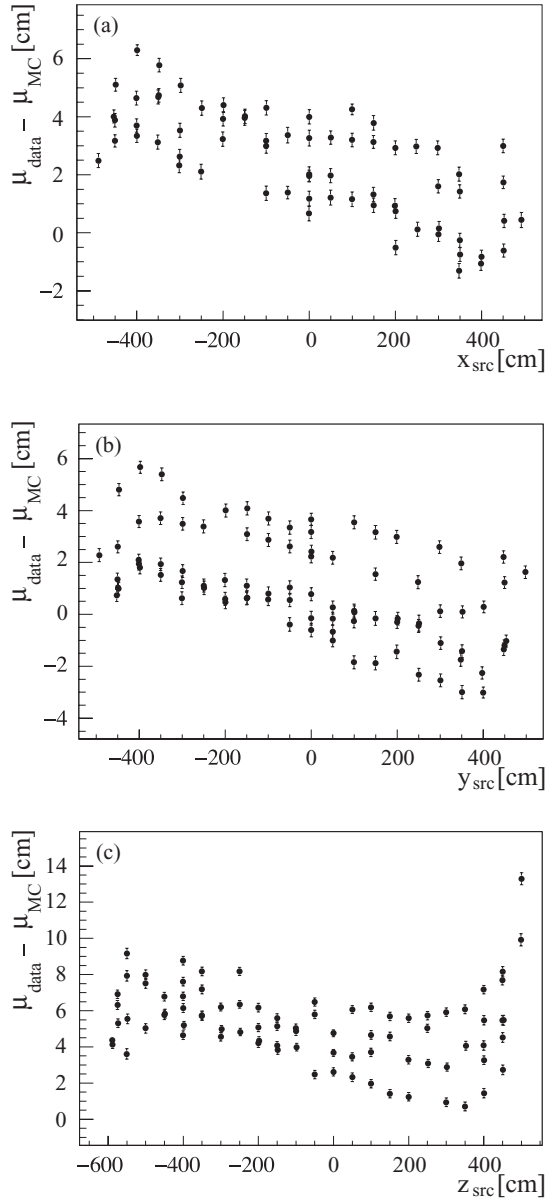


FIG. 8. Difference between data and MC vertex shifts in the three axis directions as a function of source position for scans along the main axes of the detector. Panels (a)–(c) represent the shifts in x , y , and z coordinates, respectively.

distribution function and a PMT solid angle correction. The direction of the event was crucial in separating Cherenkov events from the ES reaction from those from the CC and NC reactions. The direction of the scattered electron in the ES event was strongly correlated with the neutrino source: the Sun. To avoid biases caused by scattered light, only PMT hits occurring within a time window of ± 10 ns of the prompt light peak were selected.

Because most signals and backgrounds were not correlated with the direction of the Sun's position, the angular resolution uncertainty did not generally have a significant effect. The direction of electron-scattering (ES) events was, however, strongly correlated with the incoming neutrino direction and

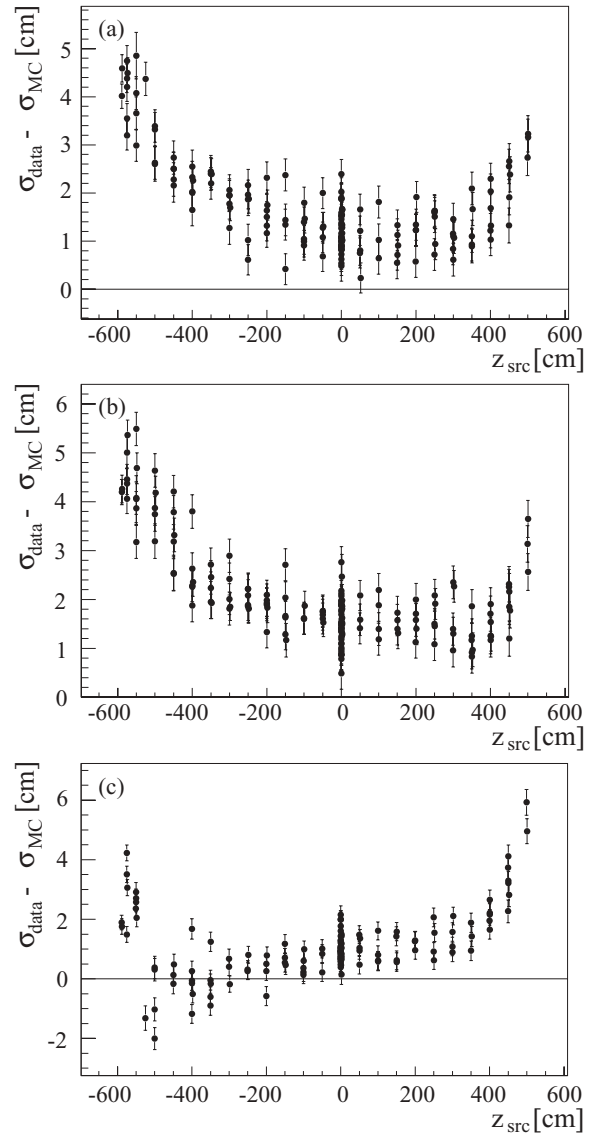


FIG. 9. Difference between data and MC fitted vertex width in (a) x , (b) y , and (c) z coordinates as a function of the position of the source (z_{src}) for ^{16}N runs taken along the z axis of the detector.

was well modeled by the function

$$P(\cos\theta_{\odot}) = \alpha_M \frac{\beta_M e^{\beta_M(\cos\theta_{\odot}-1)}}{1 - e^{-2\beta_M}} + (1 - \alpha_M) \frac{\beta_S e^{\beta_S(\cos\theta_{\odot}-1)}}{1 - e^{-2\beta_S}}, \quad (3)$$

where β_S is the parameter for the exponential component associated with the main peak and β_M is associated with the multiple scattering component. To determine the systematic uncertainty on these parameters, the function in Eq. (3) was fitted using ^{16}N calibration data and simulations. In this analysis, $\cos\theta_{\odot}$ was replaced by the cosine of the angle between the reconstructed and the true electron direction, the latter being approximated by the fitted vertex position relative to the source, for events reconstructed 120 cm or more away from the source position. Owing to the correlation between

TABLE I. Fitted values for the parameters in the polynomial $a_0 + a_1z + a_2z^2$ used to evaluate the systematic uncertainty on vertex resolution.

Direction	a_0 (cm)	a_1 ($\times 10^{-2}$)	a_2 ($\times 10^{-5} \text{ cm}^{-1}$)
x, y	1.19 ± 0.52	-0.10 ± 0.11	0.71 ± 0.36
z	1.29 ± 0.51	0.21 ± 0.15	NA

these resolution parameters and the complexity associated with the smearing of angular resolution for single events using Eq. (3), the uncertainty was propagated using the expression

$$\cos \theta'_\odot = 1 + (\cos \theta_\odot - 1)(1 \pm \delta), \quad (4)$$

where $\delta = 0.12$ is the relative uncertainty on β_M and β_S parameters. This parametrization was shown to be a good approximation for ES events.

C. Energy calibration of Cherenkov events

The fundamental measure of event energy in SNO was the number of Cherenkov photons produced by fast electrons, and the most basic energy observable for Cherenkov events was the number of triggered PMTs (N_{hit}). The energy reconstruction algorithm discussed in Sec. IV C1 below attempts to determine the number of photons (N_γ) that would have been produced by an electron, given the reconstructed position (\vec{r}_{fit}) and direction (\hat{u}_{fit}) of the event in the detector, to yield the number of triggered PMTs observed. The PMT response to optical photons, which was an important component of the algorithm, is described in Sec. IV C2. An estimate of event kinetic energy (T_{eff}) can then be derived from the one-to-one relationship (\mathcal{F}_T) between electron kinetic energy (T_e) and the mean of the corresponding distribution of N_γ [27]. The derivation of \mathcal{F}_T was performed via MC simulation. The only free parameter in the simulation, the average PMT collection efficiency (ϵ_{PCE}) for the PMT array, was determined by comparing the energy scale of ${}^{16}\text{N}$ calibration data and its simulation (see Sec. IV C3). A study of the energy systematic uncertainties, of which the energy scale and the energy resolution are the most significant to the solar neutrino flux measurement, is provided in Sec. IV C4.

TABLE II. Correlation matrices for the fitted parameters of the polynomials used to evaluate the systematic uncertainty on vertex resolution.

ρ	x, y directions		
	a_0	a_1	a_2
a_0	1.00	-0.13	-0.74
a_1	-0.13	1.00	0.31
a_2	-0.74	0.31	1.00
ρ	z direction		
	a_0	a_1	
a_0	1.00	0.15	
a_1	0.15	1.00	

1. The energy reconstruction algorithm

For an initial estimate of an event's equivalent electron kinetic energy T_{eff} , a corresponding estimate of N_γ could be calculated via \mathcal{F}_T^{-1} . The number of photons expected to trigger the i th PMT, N_i , was calculated as

$$N_i = N_\gamma \frac{\sum_\lambda R_i(\lambda, \vec{r}, \hat{u}) \frac{1}{\lambda^2}}{\sum_\lambda \frac{1}{\lambda^2}}, \quad (5)$$

where R_i , discussed in detail in Sec. IV C2, is the response of the i th PMT to a photon of wavelength λ originating from the position \vec{r} and the direction \hat{u} . The sum over λ was done in 10-nm steps from 220 to 710 nm, the wavelength range over which the detector was sensitive. The total number of direct PMT hits (i.e., not from reflected light) predicted by the energy reconstruction ($N_{\text{predicted}}$) is then

$$N_{\text{predicted}} = \sum_i^{N_{\text{PMTs}}} N_i M(N_i). \quad (6)$$

The correction function M accounts for the possibility of multiple photons counting in the hit PMT and has the form of [28]

$$M = \frac{1 - e^{-\mu}}{\mu}, \quad (7)$$

where μ is the number of photons expected to hit the PMT. PMTs triggered by multiple photoelectrons were indistinguishable from those triggered by a single photoelectron, as the charge resolution of the PMTs was insufficient to resolve the difference.

The initial estimate of N_γ was modified by

$$N_\gamma \rightarrow \frac{N_{\text{eff}}}{N_{\text{predicted}}} N_\gamma, \quad (8)$$

where N_{eff} is the effective number of PMT hits, the number of PMT hits within the prompt light window after corrections for dark noise were made. This process was iterated until agreement was reached between $N_{\text{predicted}}$ and the noise-corrected number of prompt PMT hits.

2. PMT optical response (R_i)

The optical response of the i th PMT, R_i , in Eq. (5) was calculated as

$$R_i = \epsilon_i(\lambda, \vec{r}_{\text{fit}}, \vec{p}_i, \hat{n}) \Omega_i(\vec{r}_{\text{fit}}, \vec{p}_i, \hat{n}) D(T_e | \vec{r}_{\text{fit}}, \vec{p}_i, \hat{u}_{\text{fit}}) \times F(\vec{r}_{\text{fit}}, \vec{p}_i) \exp\left(-\sum_{m=1}^3 d_m(\vec{r}_{\text{fit}}, \vec{p}_i) \alpha_m(\lambda)\right), \quad (9)$$

where λ is photon wavelength, \vec{r}_{fit} is the reconstructed position, \hat{u}_{fit} is the reconstructed direction, \vec{p}_i is the vector between the PMT position and \vec{r}_{fit} , \hat{n} describes the orientation of the PMT, and T_e is the true electron kinetic energy. The terms in the sum over the media ($m = \text{D}_2\text{O}$, acrylic, and H_2O) are the product of the average optical path length (d_m) and the inverse attenuation length (α_m) measured in Sec. IV A3 above. The probability of photon transmission through the AV (F) was calculated from the Fresnel coefficients, D is the Cherenkov angular distribution, Ω_i is the solid angle of the PMT as seen from \vec{r}_{fit} , and ϵ_i is the efficiency of photons to trigger a PMT

upon their entering the light-collecting region of the PMT and reflector assembly. Each factor in Eq. (9) required the average optical path of a photon to be calculated from \vec{r}_{fit} , \vec{p}_i , and the detector geometry. The calculation was done to the center of the PMT photocathode surface. Ω_i was determined from the optical paths to multiple points around the PMT light concentrator assembly.

The PMT efficiency (ϵ_i) was broken down into the factors

$$\epsilon_i = \epsilon_{\text{PCE}} E(\lambda, \cos \theta_n) E_i^{\text{optical}} E_i^{\text{electronic}} \epsilon^{\text{NCD}}, \quad (10)$$

where ϵ_{PCE} is the aforementioned average PMT collection efficiency, E is the relative response as a function of incidence angle (θ_n) for a typical PMT, and E_i^{optical} and $E_i^{\text{electronic}}$ are the relative optical and electronic channel efficiencies. E was determined at multiple wavelengths by a combination of optical calibrations and MC simulations (see Sec. IV A3). It was normalized to the quantum efficiency for a typical PMT, which was also a function of wavelength, at normal incidence. In the central region of the detector, the obstruction of photons by the NCD array reduced the average number of direct PMT hits by up to 20%. The efficiency factor (ϵ^{NCD}), calculated based on a MC simulation of Cherenkov light from single electrons in the D₂O, accounted for this effect.

3. Energy calibration

Once the detector optical parameters have been determined, the PMT collection efficiency (ϵ_{PCE}) and the energy calibration function (\mathcal{F}_T) were still to be set. High-rate (~ 200 Hz) central ^{16}N source calibration runs were compared to MC simulations using an initial estimate of $\epsilon_{\text{PCE}} = 0.645$ (as was determined for phase II). This value was adjusted to match the means of the N_γ distributions obtained from the source data and from simulations. The result was a value of $\epsilon_{\text{PCE}} = 0.653$ for the third phase.

Figure 10 shows the relative photon-collection efficiency of the detector using the same central ^{16}N calibration runs as above. The slight time variation, not accounted for by energy

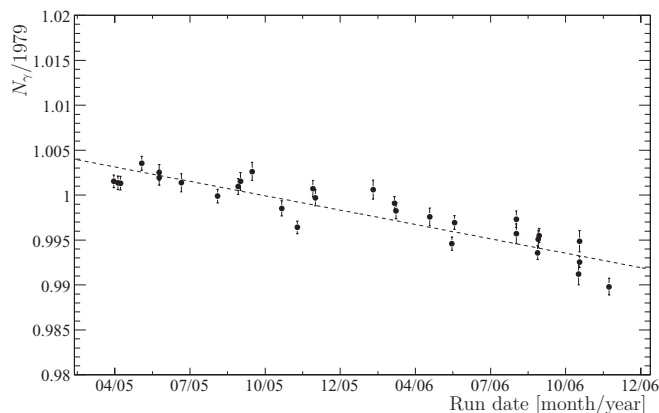


FIG. 10. Relative photon-collection efficiency of selected ^{16}N calibration runs as a function of time. The dashed line represents δ_{drift} that was scaled to unity on September 27, 2005. The error bars show only the statistical uncertainties. The impact owing to systematic effects is explained in the text, which takes the drift and the difference between the calibrated MC and the ^{16}N calibration data into account.

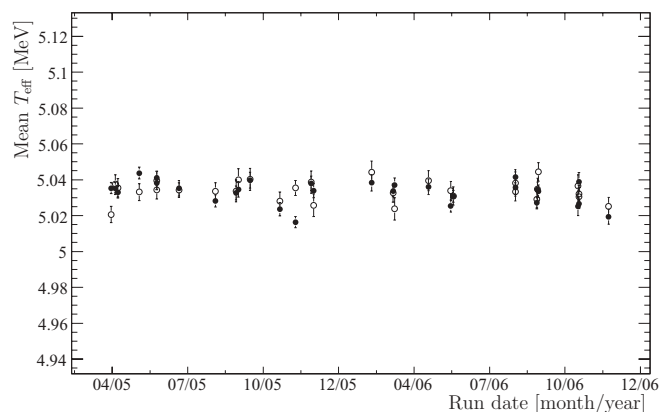


FIG. 11. The mean effective electron kinetic energy of selected central ^{16}N calibrations (solid points) and MC simulation (open points).

reconstruction, was described by the function

$$\delta_{\text{drift}} = 1.197 - 1.751 \times 10^{-5}t, \quad (11)$$

where t is the number of days since the reference date December 31, 1974 (UTC). The function was arbitrarily normalized to unity on September 27, 2005. As the PMT collection efficiency in the MC simulation was tuned to track the time variation observed in Fig. 10, a correction $T_{\text{eff}} \rightarrow T_{\text{eff}}/\delta_{\text{drift}}$ was required to be performed on both the reconstructed energy of data and simulated events.

To derive the energy calibration function \mathcal{F}_T , a series of monoenergetic electron MC simulations, from 2 to 130 MeV, were performed. The mean of the N_γ distribution for each set of simulated electron events, corrected by δ_{drift} , was then matched to the known electron kinetic energy T_e . \mathcal{F}_T consisted of the interpolation between these values. For those readers who need an approximate form of the average electron energy response function of the SNO detector in phase III, it can be characterized by a Gaussian function with resolution $\sigma_T = -0.2955 + 0.5031\sqrt{T_e} + 0.0228T_e$, where T_e is in MeV.

Figure 11 shows the resulting mean T_{eff} of selected ^{16}N calibration runs (solid points) after application of the drift correction. The mean T_{eff} of the full MC simulation with $\epsilon_{\text{PCE}} = 0.653 \delta_{\text{drift}}$ are also shown (open points).

4. Energy systematic uncertainties

Two of the most important systematic uncertainties on the measured neutrino fluxes were the energy scale and energy resolution of the PMT array. These uncertainties were determined by comparing data and MC simulations using the ^{16}N source. This source was deployed on nearly a monthly basis and was used to probe not only the center of the detector but also to scan along the x , y , and z axes. A total of 1053 ^{16}N runs were used in the energy systematics analysis.

To correctly determine the energy of an event, the energy estimator must use the number of working PMTs. The number of nonworking PMTs considered as working by the energy estimator could be approximated by counting the number of PMTs that fell outside a region of 5σ from the average PMT occupancy. The ratio of this number by the total number of

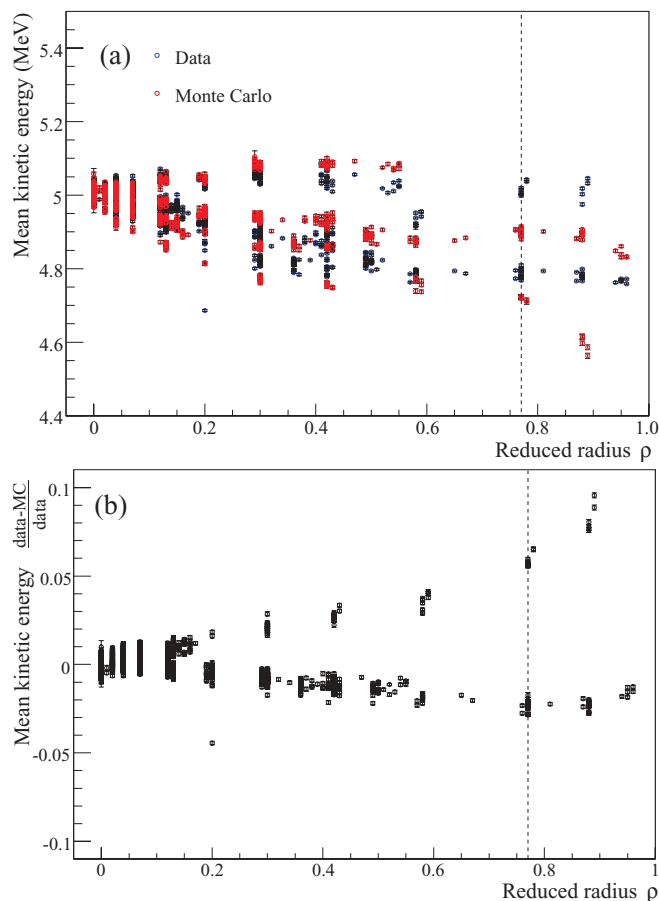


FIG. 12. (Color) Panel (a) shows the mean effective kinetic energy T_{eff} versus volume-weighted radius ρ for the ^{16}N source deployed throughout the detector. Panel (b) shows the fractional difference between data and MC [the ratio $(\text{data-MC})/\text{data}$] for the ^{16}N source deployed throughout the detector. The dashed line at $\rho = 0.77$ represents the edge of the fiducial volume at 550 cm. Points at the same value of ρ could have different energy response in data or MC because of local point-to-point variation in detector response. The AV neck and the NCD array accentuated the variations at large ρ .

working PMTs for a run was taken as the potential influence of the uncertainty of the detector state on the estimated energy. The uncertainty owing to the detector state was determined to be 0.03%.

The temporal stability of the energy response of the data and MC was evaluated using the ^{16}N runs taken at the center of the detector. A comparison of the mean kinetic energy and resolution distributions in data and simulation was used to determine the temporal stability systematic uncertainty. The energy drift/stability uncertainty on the energy scale was found to be 0.40% and that on the resolution was determined to be 1.19%.

One of the largest contributions to the energy scale and to the energy resolution uncertainties was attributable to spatial variations in the detector. Figure 12 shows a comparison of T_{eff} between data and MC simulations as a function of the volume-weighted position $\rho = R_{\text{fit}}/R_{\text{AV}}$, where $R_{\text{AV}} = 600$ cm is the radius of the AV. The point-to-point variations and radial biases of the detector response were determined by

dividing the detector into radial and polar angle bins assuming an azimuthal symmetry. The average differences between data and MC simulations and the variance of the mean kinetic energy and resolution were determined in each of these bins. The volume-weighted average of these differences among the bins was then taken as the spatial variation on the energy scale and resolution. The uncertainty on the spatial variation of the energy scale was determined to be 0.64% and the spatial variation of the energy resolution was determined to be 1.04%.

The ^{16}N source was typically run at a rate on the order of several hundred hertz, whereas neutrino data were taken with the detector operating at an event rate of an order of magnitude lower. To evaluate the potential rate dependence, the ^{16}N source was periodically run at “low rate” (several hertz). Comparing the mean kinetic energy of low-rate runs taken close in time to high-rate runs, the uncertainty on the energy scale related to the rate dependence was determined to be 0.20%.

A series of dedicated PMT high-voltage scans were performed to quantify the dependence of the detector response on the PMT gain. Inspecting the value of the half-maximum height at the upper end of the single photoelectron charge distribution allowed for an estimation of the gain effects on the energy scale uncertainty. The ratio of the value of the upper half-maximum height of the single photoelectron charge distribution at the nominal ^{16}N energy compared to this slope allowed for a conservative estimation of the energy scale systematic owing to gain changes. This uncertainty was determined to be 0.13%.

A series of dedicated threshold scans were performed to quantify the dependence of the detector response on the PMT channel threshold. Comparing the ratio of the energy scale to the value of the half-maximum height at the lower end of the single photoelectron charge distribution allowed for an estimation of uncertainty on the energy scale owing to threshold changes. This uncertainty was conservatively estimated at 0.11%.

The energy scale could be affected by the change in the timing position of the prompt light peak. This is because the energy estimator utilized only the prompt light within a limited time window of 20 ns. To determine an uncertainty owing to changes in the time residual position compared to the energy estimator’s time window, the mean and width of the prompt timing peak were determined by a Gaussian fit on all the central ^{16}N runs. The extracted timing peak means and widths were quite stable and similar to what were seen during previous phases. Because there were no indications that the timing had changed from the last phase, the uncertainty from the second phase, 0.10%, was used and was considered conservative.

Incomplete modeling of the ^{16}N source contributed to the energy scale uncertainty as MC simulations of the source were used to compare to real data. Effects such as approximations to the source geometry, uncertainties in the ^{16}N decay branching ratios, and the finite step size in EGS4 [29] simulation were studied. Their combined contribution to the energy scale uncertainty was determined to be 0.65%.

TABLE III. Summary of PMT array's energy scale and resolution systematic uncertainties.

Scale uncertainty	
Source	Uncertainty
Detector state	0.03%
Drift/stability (data MC)	0.40%
Spatial variation	0.64%
Gain	0.13%
Threshold	0.11%
¹⁶ N source modeling	0.65%
Rate dependence	0.20%
Time calibration	0.10%
Total	1.04%
Resolution uncertainty	
Source	Uncertainty
Spatial variation	1.04%
Resolution shift	
Source	Shift
Drift/stability (data MC)	1.19%

Table III summarizes the contributions to the systematic uncertainty on the energy scale and resolution. The energy scale uncertainty contributions were added in quadrature to yield an overall uncertainty of 1.04%. The resolution uncertainty contributions were determined to be a shift of 1.19% along with an uncertainty of 1.04%.

V. ELECTRONIC CALIBRATION OF NCD ARRAY

In phase III of SNO, the neutrons from NC interactions were predominantly detected by the NCD array. Thus, understanding the response of the NCD array to neutrons was crucial to an accurate determination of the total active solar neutrino flux. In this section, the calibration of the ionization signal in the NCD detector strings is presented. The determination of neutron detection efficiency in the NCD and the PMT arrays are presented in Sec. VI.

As described in Sec. II B the data stream of the NCD array consisted of events from the shaper-ADC and the MUX-scope subsystems. The shapers could provide the total charge in an event and the MUX-scope subsystem could digitize the log-amplified wave form of the signal. The primary goal of the electronic calibration was to measure the parameters of the electronic model, so that the transformations of the counter signals as they propagated through the front-end electronic and data-acquisition systems were quantified. The energy of individual events in the NCD arrays could then be determined and used as an observable in the solar neutrino flux analysis. A calibration system was implemented to pulse the preamplifiers. The output of a programmable wave-form generator was attenuated by 30 dB and injected into a pulse distribution system (PDS) board, whose amplified outputs could be sent to selected preamplifiers through computer control. The electronic calibration and the performance of the

shaper-ADC and the MUX-scope subsystems are presented in the following.

A. Linearity

The gain and linearity of the shaper-ADC and MUX-scope channels were calibrated by sending rectangular pulses of known amplitudes to the preamplifiers, one preamplifier at a time. These calibrations were performed once a week at five different pulse amplitudes. Extended electronic calibrations with 20 different pulse amplitudes over an expanded range were performed monthly. Rectangular pulses were used because their start and stop times were easily determined, which facilitated the integration of digitized wave forms in the analysis. The measured charge of the signal as a function of the calculated input charge was fit to a linear function, which measured the gain and offset of each channel and tested the channels' linear response. Because the digitized wave forms were logarithmically amplified and recorded by the digital oscilloscopes, they were first delogged [inverting Eq. (12) below] for the MUX-scope channel linearity analysis. This tested the linearity of the MUX-scope system as well as the measured log-amp parameters. The shaper-ADC channels were found to be linear to within 0.5% across all channels. The transfer function of the logarithmic amplifier was responsible for an observed nonlinearity of up to $\sim 5\%$, and a model was developed to account for this behavior. Figure 13 shows the temporal variation of the relative gain in the shaper-ADC and MUX channels, measured from the monthly extended calibration runs in phase III, for string N1.

The shaper-ADC channel corresponding to the mean of neutron signal peak in Am-Be calibrations was used to determine the conversion gain for the measured shaper-ADC charges to event energies in the ³He counters. For the ⁴He counters, which were insensitive to neutrons, the energy peak from ²¹⁰Po α decays was used.

B. Threshold

The threshold calibration involved injecting offset, single-cycle sine waves with a constant width and varying amplitudes to all the preamplifiers simultaneously. The range of these pulser output amplitudes extended above and below each channel's threshold level. Sine waves were chosen for this measurement because the calibration pulse amplitude and calculation of the total injected charge were more stable with sine waves than with signals that are not smooth, such as rectangular or triangular pulses. The high-frequency components of nonsmooth waveforms could produce transient currents, which could be difficult to calculate accurately for each channel.

The threshold levels of the shaper-ADC and MUX channels were determined by finding the pulser amplitude at which half of the expected events were observed. The algorithm searched over the range of pulser amplitudes, estimating the charge and current thresholds of the shaper-ADC and MUX channels. The thresholds were stable over the course of the experiment except for when they were intentionally changed to account for sporadic electromagnetic pickup or a malfunctioning string. In

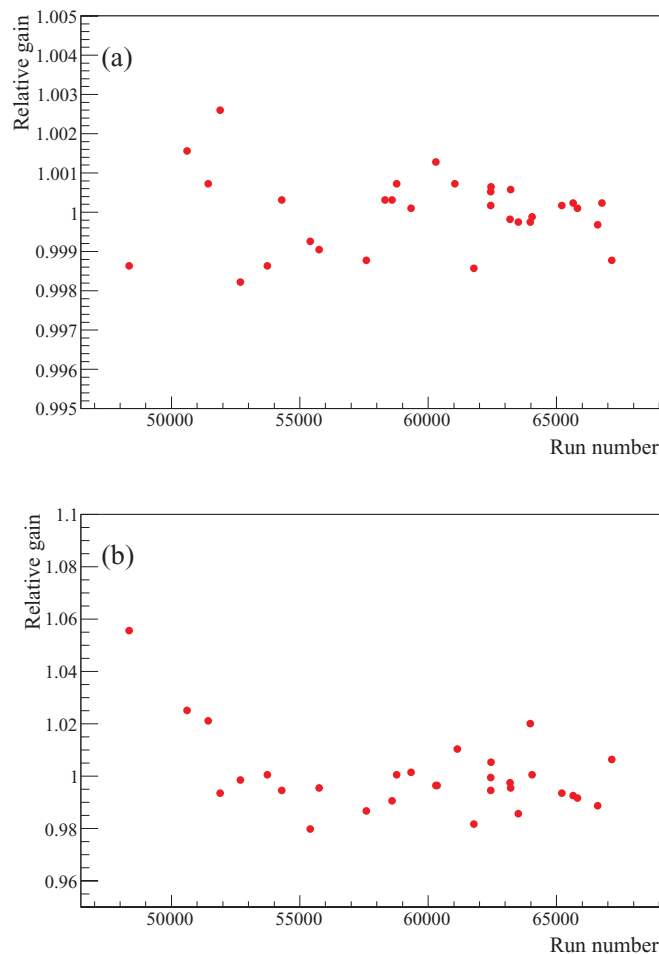


FIG. 13. (Color online) Temporal variation of the relative gain in (a) the shaper-ADC (top) and (b) MUX channels for string N1. In these plots the gain measured from each monthly extended electronic calibration was normalized to the mean gain. The data shown here extended from the commissioning (prior to run 50 000) to the completion of phase III data taking (\sim run 67 000).

the latter case, the thresholds for the channels associated with the malfunctioning string were set to their maximum values to ensure that they were offline. Figure 14 shows the temporal variation of the shaper-ADC channel threshold for string N1 during phase III.

C. Log-amp

The log-amp calibration pulse was an offset, single-cycle $1\text{-}\mu\text{s}$ -wide sine wave preceded $6\text{ }\mu\text{s}$ by a narrow rectangular trigger pulse. These pulses were injected to each preamplifier channel at 3 Hz for a duration of 15 s. A sine wave was selected because its smoothly varying shape and frequency characteristics were similar to the expected counter signals. If the sine wave were used to trigger the channel, it could be possible that the beginning of the pulse would not be recorded because the time for the sine wave to go from zero to the MUX threshold level might be longer than the electronic delay time. The width of the trigger pulse was set to a small value to reduce

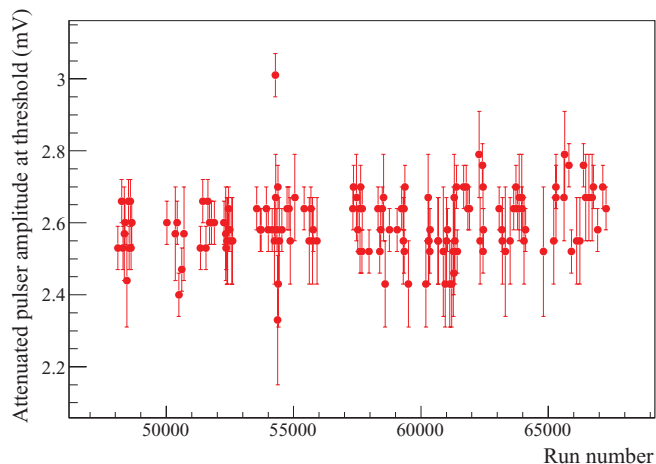


FIG. 14. (Color online) Temporal variation of the shaper-ADC channel threshold for string N1. In the threshold calibration, pulses from the pulser were attenuated by 30 dB. The channel threshold was determined by finding the attenuated pulser amplitude at which half of the expected events were observed. The data shown here extended from the commissioning (prior to run 50 000) to the completion of phase-III data taking (\sim run 67 000).

the amount of baseline offset produced by its integrated charge before the sine wave arrived at the input.

The logarithmic amplification in the MUX electronic chain was modeled as

$$V_{\log}(t) = A \cdot \log_{10} \left(1 + \frac{V_{\text{lin}}(t - \Delta t)}{B} \right) + C_{\text{chan}} + V_{\text{PreTrig}}, \quad (12)$$

where V_{\log} and V_{lin} are the logarithmic and linear voltages, Δt represents the time delay for each channel in the MUX, and A , B , C_{chan} , and V_{PreTrig} are constants determined by calibrations. Details of this parametrization of the MUX-scope electronic chain can be found in Ref. [30].

The log-amp calibration analysis involved a χ^2 minimization that estimated the set of five log-amp parameters that best fit a simulated signal to each measured calibration pulse. A weighted average of the parameter values extracted for each event, along with the uncertainty of the weighted average, was calculated for each NCD electronic channel. Because some of these parameters were dependent on which of the two oscilloscopes recorded the event, there were two sets of log-amp parameters to be measured for each string. The electronic calibrations measured these parameters and the current threshold level of digitization separately for each string.

VI. NEUTRON DETECTION EFFICIENCY CALIBRATION

The NC interaction produced a uniform distribution of neutrons in the D_2O volume. To determine the total active solar neutrino flux via the NC interaction rate accurately, the neutron detection efficiency for both the NCD and the PMT arrays must be well understood. The primary method for determining the neutron capture efficiency of the arrays was to deploy an evenly distributed ^{24}Na source in this volume [31]. Two such calibrations were performed in October 2005 and in November

2006. In this section, a detailed study of the neutron detection efficiency for the NCD and the PMT arrays is presented.

In these “spike” calibrations, about 1 L of neutron-activated brine containing ^{24}Na was injected into the D_2O volume and mixed. The ^{24}Na isotope, with a 14.96-h half-life, decays to ^{24}Mg , almost always (99.9%) producing a β decay electron with an end point of 1.39 MeV and two γ rays with energies of 1.37 and 2.75 MeV. The 2.75-MeV γ ray is capable of photodisintegrating a deuteron, the binding energy of which is 2.225 MeV. A perfectly mixed spike would produce a uniform distribution of neutrons up to within about 30 cm (one Compton scattering length) of the AV, where the neutron intensity would drop off because the 2.75-MeV γ ray has a significant chance of escaping the D_2O volume. A correction factor, based on simulations of neutrons from NC interactions and from photodisintegration owing to an evenly distributed ^{24}Na brine, was applied to compensate for the differences in the neutron energy and neutron radial distribution between the solar neutrino and the spike calibration data. After this correction, the neutron capture efficiency from the ^{24}Na calibration was equal to the efficiency for neutrons produced by the NC interaction. However, in calculating the efficiency by this technique and its uncertainty, possible deviations from perfect mixing that might occur near boundaries, such as the walls of the AV and the NCD strings, were also considered, as described below.

The temporal behavior of the NCD array response was monitored by deploying ^{252}Cf and AmBe sources in different parts of the D_2O volume using the source manipulator system [1]. The data from these point calibrations were also used to calibrate the MC code. The technique that was used to tune the MC is discussed in Sec. VI E.

In our previous paper [7], the reported neutron capture efficiency of the NCD array was measured with the ^{24}Na spikes, and the neutron detection efficiency of the PMT array was calculated from the calibrated MC code. A subsequent analysis of the PMT array’s neutron detection efficiency using the ^{24}Na data is reported in Sec. VI D. The results from direct ^{24}Na calibration and calculations from the tuned MC are consistent.

A. Inputs for determining the neutron capture efficiency with a uniformly distributed ^{24}Na source

The input elements needed to determine the neutron capture efficiency for the NCD array and the neutron detection efficiency for the PMT array are made explicit in the following formula:

$$\epsilon_{\text{sol}} = f_{\text{nonunif}} \cdot f_{\text{edge}} \cdot \epsilon_{\text{spike}}, \quad (13)$$

where ϵ_{sol} is the capture or detection efficiency for neutrons produced by solar neutrinos and ϵ_{spike} is the efficiency determined from the ^{24}Na spike calibration. The two factors multiplying ϵ_{spike} correct for differences in the distribution of neutrons in solar neutrino and ^{24}Na calibration data. The factor f_{edge} accounts for the differences in the neutron energy and neutron radial distribution, while f_{nonunif} is a factor that accounts for the effect of possible nonuniformity of the activated brine in the D_2O .

The neutron capture efficiency of the NCD array and the neutron detection efficiency of the PMT array measured from ^{24}Na calibration are given by the ratio of the observed signal rate R_{spike} at a reference time and the ^{24}Na source strength $A_{^{24}\text{Na}}$ at that time:

$$\epsilon_{\text{spike}} = \frac{R_{\text{spike}}}{A_{^{24}\text{Na}}}. \quad (14)$$

1. $A_{^{24}\text{Na}}$: The ^{24}Na source strength measurement

The strength of the ^{24}Na source was determined using three different detectors: a germanium detector (*ex situ* measurement), the SNO PMT array (*in situ*), and the NCD array (*in situ*). In the following discussion, the quantity to be calibrated is the rate of neutrons produced by an encapsulated sample of the ^{24}Na brine placed at the center of the SNO detector. For the calibration with the germanium detector, the total rate of 2.75-MeV γ rays produced by the ^{24}Na brine sample was measured and a MC simulation program was used to calculate the expected rate of neutrons produced in the heavy water if the ^{24}Na were positioned at the center of the SNO detector [32]. For the calibration measurement using the SNO PMT and NCD arrays, a comparison was made between the neutron rates observed in these detector arrays when the ^{24}Na brine sample was placed at the center of SNO and the rate observed when a well-calibrated ^{252}Cf neutron source was placed at the same location, with a small further correction determined from simulations to deal with the different neutron spatial and energy distributions of the two sources.

In the germanium detector measurement, a small sample of the activated brine, with a mass measured to better than 0.5%, was placed on the detector (the liquid was contained in a Marinelli beaker [31]) and the rate of the 1.37- and 2.75-MeV γ rays was measured. The measured rate was converted to the γ radiation rate of the sample by correcting for the detector dead time, γ -ray acceptance, and the effect of the detector’s dead layer. Given this measurement and the photodisintegration cross section of the deuteron at 2.75 MeV (with a 2% uncertainty [32]), the effective source strength if it were placed at the center of the SNO detector at a reference time could be determined in units of neutrons per second.

Unlike the above measurement, the SNO PMT and NCD arrays were used to measure directly the neutron production rate owing to a small activated brine sample in the D_2O . A 10-g sample (measured to better than 0.5%) was placed in a sealed container and deployed to the center of the D_2O volume using the calibration source manipulator. The 2.75-MeV γ rays from the activated brine mostly escaped the container without any interactions; about 1 in 385 of these γ rays photodisintegrated a deuteron to produce a free neutron. γ rays that interacted with the container lost energy by Compton scattering; when the effect of this energy loss was included, the average photodisintegration probability decreased to about 1/390 or 1/395, depending on the type of source container used (different ones were used in 2005 and 2006). The decrease in probability was estimated with a 0.7% uncertainty using MC simulation.

In the *in situ* SNO PMT array measurement, the PMT array detected the Cherenkov radiation produced by the 6.25-MeV γ

ray following the capture of these neutrons on deuterons. It was necessary to apply event selection cuts on the reconstructed radius and the energy to isolate these events from those owing to background noise from the β and γ rays produced directly by the ^{24}Na decay. T_{eff} was required to be between 5 and 9.5 MeV. This event selection criterion discriminated against radioactive background events, which had an average energy of about 3.0 to 3.5 MeV. The reconstructed event vertex was required to be 200 to 450 cm away from the source. This selection effectively removed background γ -ray events whose range was limited by the Compton scattering length.

The calibration of the neutron rate from the brine sample was obtained by comparing this rate to the rate from the ^{252}Cf neutron source, whose strength was known with an uncertainty of 0.7% [6]. The rate measurement from the ^{252}Cf source was obtained within a few days of the ^{24}Na measurement so that the detector condition would be as similar as possible. Neutrons produced by the ^{252}Cf source had a similar, but not identical, radial capture profile in the D_2O compared to that from the brine. Whereas the neutrons from the ^{252}Cf source were produced inside the source container, those from the ^{24}Na source were produced by photodisintegration in a sphere of radius about 30 cm. After production, the neutrons typically diffused by about 100 cm, so these initial differences were, to a great extent, mitigated. Yet the capture profiles were different enough to introduce significant uncertainty in the ^{24}Na source strength measurement. Detector simulation was used to determine the effect of this difference; the capture efficiency of neutrons passing the radial selection cut was found to be about 2% smaller for the ^{24}Na source than for the ^{252}Cf source. The combined statistical and systematic uncertainties were about as large as this correction.

The *in situ* measurement with the NCD array was performed almost identically as that with the PMT array and data from the same runs were analyzed. The main difference was the detection of the neutrons with the NCD array instead of with neutron captures by deuterons. As with the PMT array measurement, the neutron detection rate in the NCD array was measured with the activated brine in the detector, and this rate was divided by the measured rate of neutrons from the ^{252}Cf source. The ratio multiplied by the known ^{252}Cf source strength gave a good estimate of the brine source strength. Again, as with the PMT array data, the difference in the radial neutron capture profile required a correction which was obtained using detector simulation. The neutron capture efficiency for the ^{24}Na source was about 2% greater than that for the ^{252}Cf source. The uncertainty on this figure was much smaller (about 0.4%) than for the corresponding one for the PMT measurement because neutrons were captured much closer to the production region so that there was less reliance on the accuracy of modeling neutron diffusion to large radii.

The results for the total neutron production rate $A_{^{24}\text{Na}}$ from the 2005 and 2006 measurements are shown in Fig. 15 and Table IV. These numbers were derived from the actual measured values of the source strength in the following manner:

$$A_{^{24}\text{Na}} = f_{\text{P}} \cdot \frac{m_{\text{main}}}{m_{\text{samp}}} \cdot e^{-\Delta t/\tau_{^{24}\text{Na}}} \cdot A_{\text{samp}}, \quad (15)$$

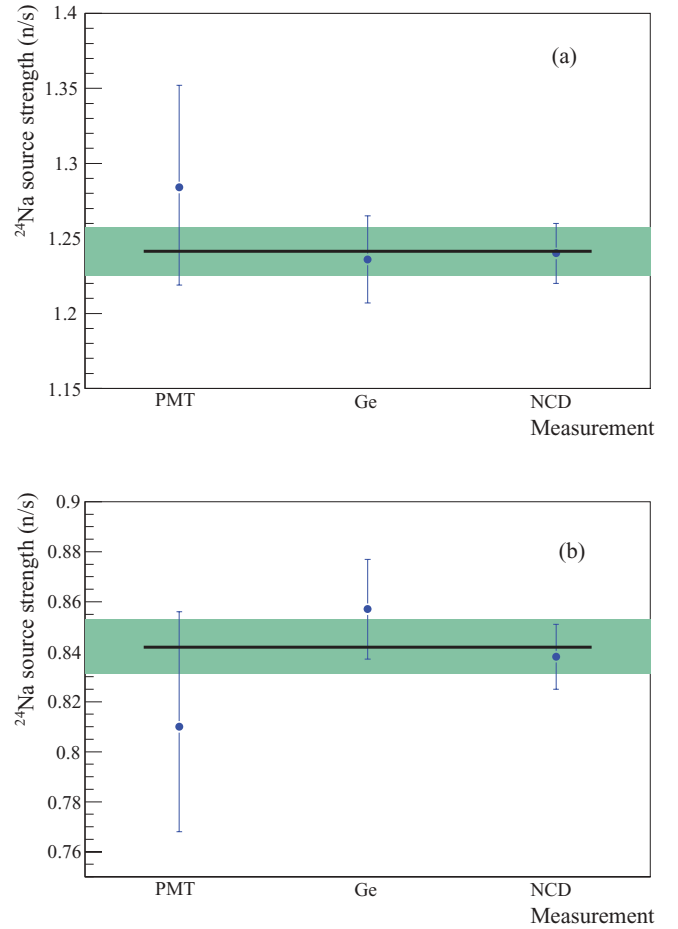


FIG. 15. (Color online) The results of the brine source strength measurements in (a) 2005 and (b) 2006. The vertical axis in each frame shows the source strength $A_{^{24}\text{Na}}$ (in neutrons per second) in the SNO detector, deduced from the three different sample source strength A_{samp} measurements (horizontal axis), at a reference time [see text and Eq. (15) for details]. This reference time ($t = 0$) was defined as the beginning of the first run when the brine was judged to be well mixed. The error bars show combined statistical and systematic uncertainties. The line through the data points is the result of the best fit, while the band shows the uncertainty of the fit.

where A_{samp} is the actual measured source strength of the sample of mass m_{samp} (10–30 g), while $A_{^{24}\text{Na}}$ is the derived source strength of the main body of brine ($m_{\text{main}} = 968.7$ g in the 2005 spike or 549.8 g in the 2006 spike) that was injected into the SNO detector several days later. The exponential correction factor $e^{-\Delta t/\tau_{^{24}\text{Na}}}$ corrects for the exponential decay of the source strength between the time of measurement and the reference time. The time offset Δt for the run in 2005 was about 4 days, while in 2006 it was 6 days. The initial

TABLE IV. The result of fitting a constant value to the data points in Fig. 15.

Data Set	Fit $A_{^{24}\text{Na}}$ (n/s)	Percent uncertainty	Fit $\chi^2/\text{d.o.f.}$
2005	1.241 ± 0.016	1.3	0.22
2006	0.842 ± 0.011	1.3	0.59

source strength was much larger in 2006, so more time was required to allow the source strength to decay to levels that the SNO data-acquisition system could handle (see Sec. II B). The quantity $\tau_{^{24}\text{Na}}$ is the ^{24}Na lifetime, 21.58 h. The quantity f_p applies only to the *in situ* measurements and corrects for the effects (discussed several paragraphs above) owing to the radial distribution of neutron production and to a small fraction of γ rays that scattered off the source container: $f_p = 1.0122 \pm 0.0044(\text{stat.}) \pm 0.0053(\text{syst.})$ for the 2005 run, and $f_p = 1.0288 \pm 0.0050(\text{stat.}) \pm 0.0053(\text{syst.})$ for the 2006 run, where the statistical uncertainty is attributable to MC statistics. The value of f_p differed between 2005 and 2006 because different source containers were used.

The brine source strength measurements from the three techniques were combined by taking a weighted average. The weighting of each data point was inversely proportional to the quadratic sum of the statistical and systematic uncertainties. Thus, the PMT array measurement, which had the largest uncertainty, made only a minor contribution to the final result. The germanium detector and NCD array measurements, having similar uncertainty magnitudes, contributed about equally. The large uncertainty in the PMT array measurement was, in large part, attributable to the difficulty in determining the neutron diffusion profile beyond 200 cm from the source, which was necessary to obtain a pure sample of neutron events.

B. Neutron capture efficiency of the NCD array

1. R_{spike} : Neutron capture rate measurement by the NCD array

The neutron capture rate in the NCD array as a function of time for the ^{24}Na spike in 2005 is shown in Fig. 16; the plot for the spike in 2006 is similar. The horizontal axis can be divided into three regions.

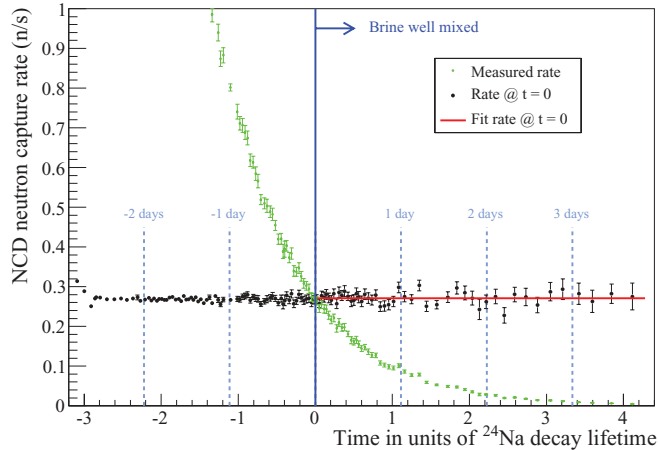


FIG. 16. (Color) Neutron capture rate in the NCD array as a function of time from the ^{24}Na spike in 2005. The vertical axis is in units of neutron events per second, while the horizontal axis is time in units of ^{24}Na mean lifetimes (one lifetime = 21.58 h). The exponentially decaying points are the actual measured rate $R(t_i)$, while the data with the flat trend [$R_i(0)$] is the rate after correcting for the exponential decay factor. The line through the points at time >0 is the best fit to the data. Time $t = 0$ was 4.53 ^{24}Na mean lifetimes after the spike was added.

TABLE V. The neutron capture rate by the NCD array at the reference time ($t = 0$), obtained by fitting a constant value to the set of measurements $\{R_i(0)\}$. The statistical uncertainty is from the fit, while the systematic uncertainty is from the instability of the rate (see text for discussion).

Year	Best-fit rate (n/s)	$\chi^2/\text{d.o.f.}$
2005	$0.2708 \pm 0.0020(\text{stat.}) \pm 0.0027(\text{syst.})$	45.6/48
2006	$0.1811 \pm 0.0016(\text{stat.}) \pm 0.0018(\text{syst.})$	35.6/38

- (i) During the first few hours, the spike was highly nonuniform, and the rate varied in an erratic manner.
- (ii) For the two ^{24}Na lifetimes preceding the reference time $t = 0$, the spike was not quite uniform, as assessed using the distribution of Cherenkov light produced directly by the β and γ rays from the decay of ^{24}Na . However, the NCD array appeared to be quite insensitive to this moderate nonuniformity of the spike, as the measured rate was indistinguishable from the equilibrium distribution. During this time period, the rate appeared to decay exponentially, as can be seen in the figure. Time $t = 0$ was 4.53 ^{24}Na mean lifetimes after the spike was added.
- (iii) After $t = 0$, the spike was determined to be well mixed according to the Cherenkov light distribution. The rate continued to decay exponentially.

The neutron rate at $t = 0$ could be derived from the measured rate $R(t_i)$ at the average time t_i of the i th run as follows:

$$R_i(0) = R(t_i) e^{t_i/\tau_{^{24}\text{Na}}}. \quad (16)$$

The set of measurements $\{R_i(0)\}$ was found to be equal to each other within the statistical uncertainty; the result of a fit of a constant value to this set of measurements is shown by the horizontal line through the data in Fig. 16. The quantity R_{spike} was obtained from this fit. The results from the runs in 2005 and 2006 are summarized in Table V.

In addition to the statistical uncertainty, several sources of systematic uncertainty existed for the neutron capture rate in the NCD array. In general, the rate can be written as follows:

$$R = \frac{N \cdot f_{\text{inst}}}{t_{\text{data}} \cdot L \cdot \epsilon_{\text{comb}}}, \quad (17)$$

where N is the number of detected events, t_{data} is the length of time data were taken, L is the fraction of the time the detector was live (referred to as the “live fraction”), ϵ_{comb} is a product of cut and threshold efficiencies, and f_{inst} is a factor used to remove the estimated contribution of instrumental noise. Studies showed that uncertainties from these input terms were negligible in comparison to a 1% long-term fluctuation in the neutron detection rate, as assessed using standard point ^{252}Cf and Am-Be sources. Because the spike calibration runs were essentially two snapshots of the detector performance, the long-term fluctuation implied that the rate could have been different by $\pm 1\%$ if the calibration were performed at any other time. For this reason, we assigned a systematic uncertainty of 1% to R_{spike} .

TABLE VI. The NCD array's capture efficiency for neutrons produced by ^{24}Na brine injected and well mixed in the SNO detector. The combined statistical and systematic uncertainty is shown here.

Year	ϵ_{spike}	Percent uncertainty
2005	0.2182 ± 0.0046	2.1
2006	0.2151 ± 0.0043	2.0

2. ϵ_{spike} : *Combining the rate and source strength measurements*

The capture efficiency by the NCD array of neutrons produced by the activated brine that was injected and mixed in the SNO detector is defined as the ratio of the number of neutrons captured in the NCD string live volume to the total number of neutrons produced. This ratio can be obtained experimentally from the two measured quantities $A^{24\text{Na}}$ (Sec. VIA1) and R_{spike} (Sec. VIB1): $\epsilon_{\text{spike}} = R_{\text{spike}}/A^{24\text{Na}}$. The central value of ϵ_{spike} can be obtained by simply dividing the numbers in Table V by those in Table IV. The error propagation was performed with care because the numerator and denominator depend at least partly on measurements in the NCD array. The systematic uncertainty of the rate in the NCD array was dominated by the long-term fluctuation of 1%. If the time between the source strength measurement and the spike rate measurement could be considered short, then the instability systematic should cancel out and the 1% systematic uncertainty in the numerator and denominator could be ignored. However, if this time period was not sufficiently short, then these uncertainties should be combined in quadrature. Although there was strong evidence for stability within a day or two of running, no data exist to demonstrate stability over a 4- to 6-day period, as was the case for the present analysis. Thus, we decided to combine the uncertainties in quadrature. The result for ϵ_{spike} is shown in Table VI.

3. f_{edge} : *Correction factor for neutron density near the acrylic vessel*

Although the activated brine produced a neutron distribution in D_2O that was similar to that produced by solar neutrinos, the neutron density near the AV was different. For ^{24}Na decays occurring within about 30 cm (one Compton scattering length) of the AV, the probability of the 2.75-MeV γ ray to escape from the D_2O volume was significant, so the neutron density dropped quickly as a function of the distance to the AV wall. The neutrons produced by the 2.75-MeV γ rays in ^{24}Na decays started with an energy of 260 keV and then were moderated by scattering, whereas those from the NC interaction started with a range of energies. This difference in the initial neutron energy could affect the neutron density at the edge of the vessel. Neutrons produced near the vessel wall were significantly less likely to be detected by the NCD array than those produced elsewhere. Thus, the volume-averaged capture efficiency of the NCD array for neutrons from the spike was somewhat larger than that for neutrons from solar neutrinos.

This difference near the AV could be determined accurately by simulations and was accounted for with the correction factor

f_{edge} :

$$f_{\text{edge}} = 0.9702 \pm 0.0078. \quad (18)$$

4. f_{nonunit} : *Correction factor for source nonuniformity*

Another potential source of difference in the neutron distribution between that from the activated brine and from solar neutrinos was imperfect mixing of the brine. Although it was not possible to directly measure the salinity as a function of position in the D_2O , there were a number of indications that the brine was well mixed and that any residual nonuniformity would not have a large impact on the measurements of the neutron capture efficiency. As can be seen in Fig. 16, the NCD array's neutron detection rate stabilized well after about a day of mixing, whereas the signal from the γ rays emitted by the ^{24}Na did not stabilize completely for about 3 more days [31]. This indicated that the neutron capture rate of the NCD array was not sensitive to the remaining inhomogeneities for several days prior to the time $t = 0$ used for the start of the analysis of the ^{24}Na data and therefore could certainly be considered stable after $t = 0$. Whereas the stability in the rate only indicated that the brine distribution reached some equilibrium configuration, the only reasonable regions where it might conceivably be nonuniform were near the inside of the AV wall or near the NCD strings where boundary layers might have different salinity. Therefore, these regions were considered carefully and used to establish systematic uncertainties on the uniformity of the brine.

A possible way of measuring the brine distribution was by detecting the Cherenkov radiation from the Compton scattering of γ rays and β 's from the decay of ^{24}Na with the SNO PMT array and comparing with the expectations for a uniform distribution. This method, however, lacked precision because the calculations of the detector response were not very accurate for the low energies deposited by these decay products. When the Cherenkov light data were compared with MC simulation of perfectly uniform brine, variations at the level of about $\pm 10\%$ were seen, which were consistent with our ability to model the detector at such low energies. Similar variations were seen in the studies of solar neutrino data at such energies.

Another piece of evidence suggesting that the brine was well-mixed in the central regions away from the AV or the NCD array came from the comparison of the distribution of Cherenkov events from ^{24}Na γ rays observed in the 2005 and 2006 data. The injection and mixing methods for the 2 years were very different [31]. In 2005, the brine was injected at several positions along the central vertical axis and then the water circulation was turned on. In this configuration D_2O was pulled out from the bottom of the detector and returned at the top. In 2006, a more sophisticated method involving flow reversal and temperature inversion was employed. Using this method, an extensive eddy current was set up by causing the entire bulk of the D_2O to rotate in one direction, then reversing the direction of the flow. In both years, the data showed that the brine distribution reached a stable equilibrium. If one examines the Cherenkov light data from each year alone, one cannot say with much confidence whether or not the equilibrium configuration was uniform because of the limitations

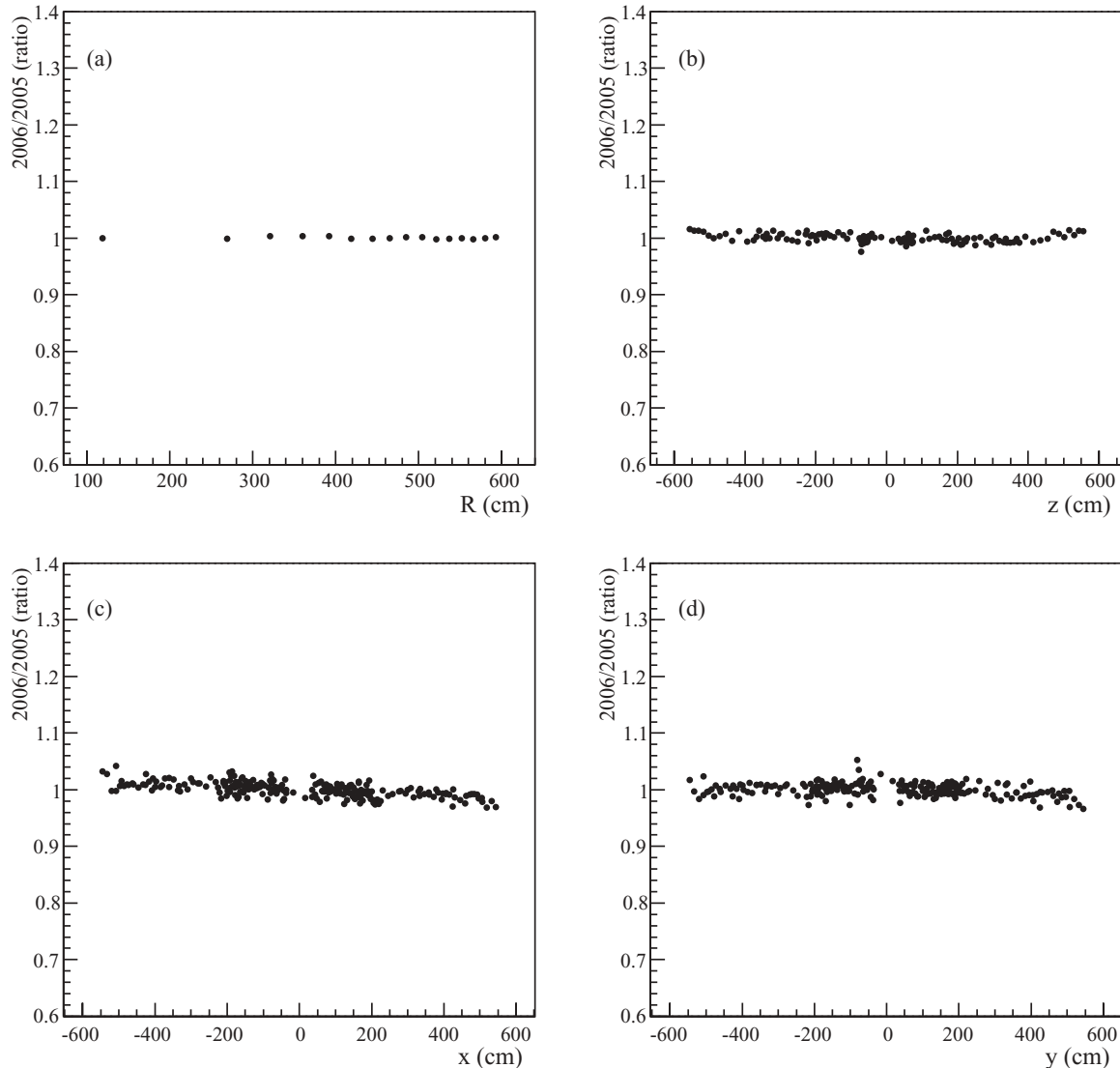


FIG. 17. Ratio of the brine concentration as observed by reconstructed Cherenkov events in 2006 to that in 2005, as a function of detector coordinates. (a) R = reconstructed radial distance from the center of SNO; (b) z ; (c) x ; and (d) y .

associated with modeling the low-energy response. However, because the brine was mixed so differently, it is implausible that a nonuniform equilibrium configuration in 2005 could be the same as that in 2006. Figure 17 shows, moreover, that the spatial distributions of the Cherenkov events in the brine were very similar. This comparison does not depend on the MC simulations but is simply a study of the Cherenkov light data observed in the two cases. A plausible explanation for the similarity is that the brine was well mixed in the central regions away from the NCD array and AV for both calibration sessions.

Based on the above argument, we assumed that the spike was well-mixed, so the correction factor f_{nonunif} has a central value of 1.0. The uncertainty on this value was obtained by considering the areas where stable inhomogeneities could possibly be established in the detector, namely near the NCD strings and near the wall of the AV.

In the vicinity of the NCD strings, calculations based on the laminar flow rates measured from the velocities observed from the Cherenkov light events indicated that the boundary layer could be on the order of 5 cm. A 0.5% effect of such a boundary layer on neutron capture efficiency was estimated from redistributing the salinity from within the layer uniformly in the D_2O target.

Careful studies of the fall-off of reconstructed γ -ray signals from the ^{24}Na brine in the vicinity of the AV, coupled with uncertainties in the knowledge of the optical properties of the detector in this region, led to an upper limit for a boundary layer thickness of 3.6 cm. Estimates based on flow rates in this vicinity gave smaller values for this boundary-layer thickness. The effect of such a boundary layer on neutron capture efficiency was considered by redistributing the salinity from within the layer uniformly in the D_2O target. This 1.5% effect was then combined in quadrature with the uncertainty from

the NCD string region to give a full uncertainty 1.6% on the efficiency. From this conservative analysis of the systematic uncertainties, we obtained

$$f_{\text{nonunif}} = 1.000_{-0.016}^{+0.016}. \quad (19)$$

5. ϵ_{sol} : Capture efficiency for neutrons produced by solar neutrinos

The NCD array's capture efficiency for neutrons produced by solar neutrinos was obtained from the product of ϵ_{spike} , f_{edge} , and f_{nonunif} [Eq. (13)]. Before performing this multiplication, however, the value and uncertainties for ϵ_{spike} from the runs in 2005 and 2006 were combined, taking into account uncertainty components that were correlated between the 2 years. The final result is

$$\epsilon_{\text{sol}} = 0.211 \pm 0.005. \quad (20)$$

This is in good agreement with the result from detector simulation, 0.210 ± 0.003 (see Sec. VIE), and with our previously published value of 0.211 ± 0.007 [7]. In this paper, we have improved on the determination of the neutron capture efficiency by the NCD array and have further examined the assumptions that were made in the first phase-III results reported in Ref. [7]. This resulted in a small reduction in the systematic uncertainty in the present analysis. This small improvement would have negligible effect on the measured NC flux and was not incorporated in the solar neutrino flux analysis reported in this paper.

C. Neutron detection efficiency of the NCD array

Several corrections must be applied to the capture efficiency, described in the previous section, to determine the NCD array's detection efficiency of NC neutrons in the solar neutrino analysis. These corrections, averaged over the duration of data taking in phase III, included the mean live fraction of the MUX (l_{MUX}) and the digitizing oscilloscope (l_{scope}), the average MUX threshold efficiency ($\langle\epsilon_{\text{MUX}}\rangle$), signal acceptance in the shaper energy window ($\langle\epsilon_{\text{shaper}}\rangle$), and the acceptance of data reduction cuts ($\langle\epsilon_{\text{cut}}\rangle$). The overall correction C is the product of these individual factors:

$$C = l_{\text{MUX}} \cdot l_{\text{scope}} \cdot \langle\epsilon_{\text{MUX}}\rangle \cdot \langle\epsilon_{\text{shaper}}\rangle \cdot \langle\epsilon_{\text{cut}}\rangle. \quad (21)$$

Table VII provides a summary of these factors.

The MUX live fraction in a data run was determined by comparing the readings on two live-time scalars: one that determined the total run time and the other that was stopped when the MUX system was unable to take in events. A pulser was installed to inject pulses at random times to provide additional validation. The mean MUX live fraction l_{MUX} was the run-time-weighted average of the solar neutrino run measurements. The scope live fraction as a function of event rate was determined using the number of observed partial MUX events. The mean scope live fraction, l_{scope} , was then established by numerically integrating the neutrino run time, weighted by the rate-dependent scope live fraction. This calculation of l_{scope} was verified using the random pulser system. The methodology for measuring the MUX threshold efficiency has been described in Sec. VB, and the results from

TABLE VII. Summary of correction factors in the determination of the neutron detection efficiency. The NCD array's detection efficiency of NC neutrons, with a shaper energy threshold $E_{\text{NCD}} > 400$ keV, is the product of the neutron capture efficiency (Sec. VIB) and the combined factor C . The parameters in this table are directly related to the efficiency and live fraction of the MUX-scope system and to the event selection cuts, which could be calibrated regularly and accurately. The factors shown in this table are the average values for solar neutrino data presented in this paper.

Correction factor	Value
MUX live fraction (l_{MUX})	0.9980 ± 0.0001
Scope live fraction (l_{scope})	0.957 ± 0.004
Average MUX threshold efficiency ($\langle\epsilon_{\text{MUX}}\rangle$)	0.99491 ± 0.00031
Average shaper energy window acceptance ($\langle\epsilon_{\text{shaper}}\rangle$)	0.91170 ± 0.00014
Data reduction cut acceptance ($\langle\epsilon_{\text{cut}}\rangle$)	0.99521 ± 0.00011
Combined (C)	0.862 ± 0.004

regular calibration runs were averaged to provide an estimate of $\langle\epsilon_{\text{MUX}}\rangle$. NCD array events with shaper energy $E_{\text{NCD}} > 0.4$ MeV were selected for the solar neutrino flux measurement described in Sec. IX. The fraction of shaper events above this energy threshold was evaluated for each NCD counter using AmBe calibration data. The mean acceptance ($\langle\epsilon_{\text{shaper}}\rangle$) for NC neutron events was then calculated by averaging these individual counter estimates, weighted by the expected fraction of NC neutrons that each counter would capture. The last correction factor is the neutron signal acceptance of the data reduction cuts, which are described in Sec. III A and Appendix A. The overall correction factor is

$$C = 0.862 \pm 0.004. \quad (22)$$

D. Neutron detection efficiency of the PMT array

Neutrons can also be captured by deuterons with the emission of a 6.25-MeV γ ray that could be detected by the PMT array. The efficiency for this detection channel was much smaller than that for the NCD array because of the large difference in the thermal neutron capture cross section between the deuteron and ^3He . In this section, we present the analysis of the PMT array's neutron detection efficiency. The basic analysis approach, that is, the evaluation of individual terms in Eq. (13), was nearly identical to that for the NCD array, discussed in Sec. VIB above.

The source nonuniformity factor f_{nonunif} was evaluated in the same manner as the case for the NCD array. We assumed that the ^{24}Na spike was well mixed, but allowed for possible inhomogeneity near the AV and the NCD strings. Based on simulation studies, these effects could lead to a combined uncertainty of 0.8% on the neutron detection efficiency for the PMT array, that is,

$$f_{\text{nonunif}} = 1.000 \pm 0.0080. \quad (23)$$

The difference in the detection efficiency for NC neutrons and the ^{24}Na photodisintegration neutrons, characterized by f_{edge} , was also different for the PMT array and the NCD array. This is because the radial dependence of the neutron capture

efficiency was different for the two detection mechanisms. This difference was determined to be 0.92% by simulation, or a f_{edge} value of

$$f_{\text{edge}} = 0.9789 \pm 0.0090. \quad (24)$$

For the measurement of R_{spike} in Eq. (14), a maximum likelihood analysis was performed to statistically separate the neutron signal from the ^{24}Na β - and γ -ray backgrounds. Cherenkov light events were selected using the energy and fiducial volume cuts for the solar neutrino analysis, that is, $T_{\text{eff}} \geq 6$ MeV and $R_{\text{fit}} \leq 550$ cm. A fit of the T_{eff} spectrum from the combined data of the ^{24}Na runs in 2005 and 2006 to a combination of neutron signal and ^{24}Na β - γ background was performed. The number of neutrons was found to be $n_n = 8205.3 \pm 121.9$ and the number of background events was found to be $n_\gamma = 1261.7 \pm 88.8$ with a correlation of -0.61 . Energy and reconstruction-related systematic uncertainties in the solar neutrino analysis were propagated in this analysis, and their combined effect was found to be ${}_{-1.38}^{+1.71}\%$. After symmetrizing this systematic uncertainty and combining it with the statistical uncertainty in the T_{eff} spectral fit in quadrature, n_n was determined to be 8205.3 ± 185.4 . With the neutron rates given in Table IV and the time span of the calibration runs, ϵ_{spike} could be evaluated, and the result for ϵ_{sol} is

$$\epsilon_{\text{sol}} = 0.0502 \pm 0.0014. \quad (25)$$

This is in agreement within uncertainties with the MC calculations ($\epsilon_{\text{sol}} = 0.0485 \pm 0.0006$) in Sec. VIE, which was used in our previously published results [7]. This small difference has a negligible effect on the solar neutrino flux results.

E. Neutron calibration with discrete sources

Calibration data were used to tune the MC simulation code, which was then used to predict the neutron detection efficiencies for both the PMT and NCD arrays. The uncertainties were calculated by propagating those on the tuning parameter and on other MC input parameters.

The calibration data used to tune the MC were acquired using AmBe sources, which were periodically deployed around the D_2O target during the data taking period. In each of the calibrations data were taken with the source positioned at a series of well-defined, repeatable locations. The tuning parameter was the hydrogen concentration in the heavy water; its measured value, from Fourier transform infrared spectroscopy, was $(9.8 \pm 0.5) \times 10^{-4}$ atom of hydrogen per atom of deuterium.

The method for tuning the hydrogen concentration was to take averaged source data for a given source location and calculate the relative detection efficiency of rings of counters around the position. The relative capture efficiency of a ring of counters close to the source to one further away was sensitive to the neutron diffusion length and therefore to the concentration of hydrogen, which has a large neutron capture cross section. These ratios were independent of source strength and, because they were of rings of counters, they were relatively insensitive to the exact source position. Simulations were run with a range of hydrogen concentrations and source locations, and a maximum likelihood fit was used to extract

the most probable hydrogen concentration. The measured hydrogen concentration was used as a constraint in the fit.

The central values for the neutron detection efficiencies in the PMT and NCD arrays were calculated using a simulation run at the best-fit hydrogen concentration of $(9.45_{-1.05}^{+0.50}) \times 10^{-4}$ atom of hydrogen per atom of deuterium, and the uncertainty propagated by rerunning the simulation with the hydrogen concentration set to its upper and lower bounds. Additional sources of uncertainty in the MC, such as parameters relating to the modeling of the counter and AV geometry, were studied separately. The only significant uncertainty came from the modeling of the shape of the NCD counter live regions, which was estimated to impart a 1.0% uncertainty on the NCD neutron detection efficiency.

The final prediction for the neutron capture efficiency of the NCD array was 0.210 ± 0.003 . The prediction for the neutron detection efficiency of the PMT array was 0.0485 ± 0.0006 .

VII. BACKGROUNDS

Several sources of radioactive backgrounds were present in the PMT array and NCD array data. The majority were associated with naturally occurring ^{238}U and ^{232}Th , cosmogenic activity and atmospheric neutrino interactions in the detector. The impact of these backgrounds on the neutrino analysis was reduced by optimization of analysis cuts. Those that remained were included in the fits for the background and the solar neutrino signals. A summary of background contributions is given in Table VIII.

In this section, we discuss the identification and measurement of the neutron and Cherenkov light backgrounds in the solar neutrino measurement. α decays from the construction materials were the largest source of backgrounds in the neutron signal region in the NCD array. This background source was difficult to calibrate owing to variation in the spatial distribution and in the composition of trace radioactivity in different counters. The treatment of α decay backgrounds are presented in the next section (Sec. VIII).

A. Photodisintegration backgrounds

γ rays with energy greater than 2.225 MeV can break apart a deuterium nucleus releasing a free neutron, which was indistinguishable from one produced by a NC interaction. Such γ rays are emitted by β - γ decays of ^{208}Tl and ^{214}Bi from the ^{232}Th and ^{238}U chains, respectively. An accurate measurement of these radioisotopes was crucial for the determination of the total ^8B neutrino flux. Concentrations of 3.8×10^{-15} gTh/gD₂O and 30×10^{-15} gU/gD₂O are each equivalent to the production of one neutron per day via photodisintegration. Two independent approaches were developed to measure these backgrounds. These are broadly classified as *ex situ* and *in situ* techniques.

1. Ex situ determination of radioactivity in D₂O

Three *ex situ* methods were developed to assay parent isotopes of ^{208}Tl and ^{214}Bi in the D_2O and H_2O regions of the detector. Common to all three techniques was extraction

TABLE VIII. Summary of backgrounds in the PMT and NCD arrays for the phase III 385-day solar neutrino data set.

Source	PMT events	NCD events
Neutrons generated inside D ₂ O		
² H photodisintegration (U,Th in D ₂ O)	7.6 ± 1.2	28.7 ± 4.7
² H photodisintegration (U,Th in NCD bulk)	4.3 ^{+1.6} _{-2.1}	25.8 ^{+9.6} _{-12.3}
² H photodisintegration (U,Th in Hot spots)	17.7 ± 1.8	64.4 ± 6.4
² H photodisintegration (U,Th in NCD cables)	1.1 ± 1.0	8.0 ± 5.2
<i>n</i> from spontaneous fission (U)	0.2 ± 0.1	0.3 ± 0.1
² H(α , αn) ¹ H (Th, ²²² Rn)	0.2 ± 0.1	0.4 ± 0.1
^{17,18} O(α , <i>n</i>) ^{20,21} Ne (Th)	0.3 ± 0.1	1.8 ± 0.4
Atmospheric ν	24.1 ± 4.6	13.6 ± 2.7
Cosmogenic muons	0.009 ± 0.002	0.04 ± 0.004
Reactor and terrestrial neutrinos	0.3 ± 0.1	1.4 ± 0.2
CNO solar ν	0.05 ± 0.05	0.2 ± 0.2
Total internal neutrons	55.8 ^{+5.6} _{-5.4}	144.7 ^{+13.4} _{-15.5}
Neutrons generated from AV and H ₂ O radioactivity		
² H photodisintegration (U,Th in H ₂ O)	2.2 ^{+0.8} _{-0.7}	7.1 ^{+5.4} _{-5.2}
(α , <i>n</i>) in AV	18.3 ^{+10.2} _{-7.3}	33.8 ^{+19.9} _{-17.1}
Total external-source neutrons	20.6 ^{+10.2} _{-7.3}	40.9 ^{+20.6} _{-17.9}
Cherenkov events from radioactivity inside the D ₂ O		
β - γ decays (U,Th)	0.70 ^{+0.37} _{-0.38}	N/A
Decays of spallation products in D ₂ O: ¹⁶ N following muons	0.61 ± 0.61	N/A
Cherenkov backgrounds produced outside D ₂ O		
β - γ decays (U,Th) in AV, H ₂ O, PMTs	5.1 ^{+9.7} _{-2.9}	N/A
Isotropic AV events	<0.3 (68% CL)	N/A

and filtering of a known amount of water from the detector and external counting of the resultant sample. Two methods extracted ²²⁴Ra and ²²⁶Ra, one using beads coated with manganese oxide (MnO_x) [33] and the other using filters loaded with hydrous titanium oxide (HTiO) [34,35]. For MnO_x and HTiO assays, up to 500 tonnes of water passed through the loaded columns over a 4- to 5-day period. In the MnO_x technique, Ra isotopes were identified by α spectroscopy of Rn daughters. In the HTiO method, Ra isotopes were stripped from the filters, concentrated, and identified using β - α coincidence counting of the daughter nuclides.

The equilibrium between ²³⁸U and ²¹⁴Bi was broken by the ingress of ²²²Rn (half-life = 3.82 d), primarily from the laboratory air and emanation from construction materials. The amount of ²²²Rn in the water was measured by degassing, cryogenically concentrating the dissolved gases and counting the sample using a ZnS(Ag) scintillator [36]. Each Rn assay processed approximately 5 tonnes of water in a 5-h period.

During the third phase of SNO, 20 MnO_x and 16 HTiO assays were conducted at regular intervals in the heavy-water region. The results from each independent assay method were in good agreement. The activity measured by each assay was a combination of activity from the D₂O and water systems piping. The variation of Th (Ra) activity in the water and piping was modeled as a function of time, taking into account other sources of Th (Ra) in the flow path. The resultant concentration was the live-time weighted combined HTiO and MnO_x activity, which was $0.58 \pm 0.35 \times 10^{-15}$ gTh/gD₂O.

The quoted uncertainty was combined from the systematic and statistical uncertainties.

A total of 66 Rn assays were performed in the D₂O region at regular intervals throughout the third phase. To calculate the mean ²²²Rn concentration, the individual assay results were time and volume weighted. The equivalent mean ²³⁸U concentration was $5.10 \pm 1.80 \times 10^{-15}$ gU/gD₂O, where the total uncertainty was combined from the systematic and statistical uncertainties.

Figure 18 is a summary of the D₂O assay results since the beginning of the SNO experiment. During SNO's phase-III operation, an aggressive program of system purification combined with minimum recirculation of the heavy water led to a factor of five reduction in thorium. The concentration of ²²⁴Ra was routinely measured at 0.1 atom/tonne.

2. *In situ* determination of radioactivity in D₂O and NCD housing

The *in situ* technique measured the ²³²Th and ²³⁸U content of the water and NCD array housings directly from the Cherenkov light data [37]. In the energy window $4.0 < T_{\text{eff}} < 4.5$ MeV, the selected events were dominated by ²¹⁴Bi and ²⁰⁸Tl, from the ²³⁸U and ²³²Th chains, respectively. The observed Cherenkov light was dominated by the direct β decay of ²¹⁴Bi to the ground state of ²¹⁴Po with an end point of 3.27 MeV. ²⁰⁸Tl decays almost always emitted a 2.614-MeV γ ray, accompanied by one or more lower energy γ ray and a β with an end point of up to 1.8 MeV. ²⁰⁸Tl events produced a more isotropic Cherenkov light distribution when compared

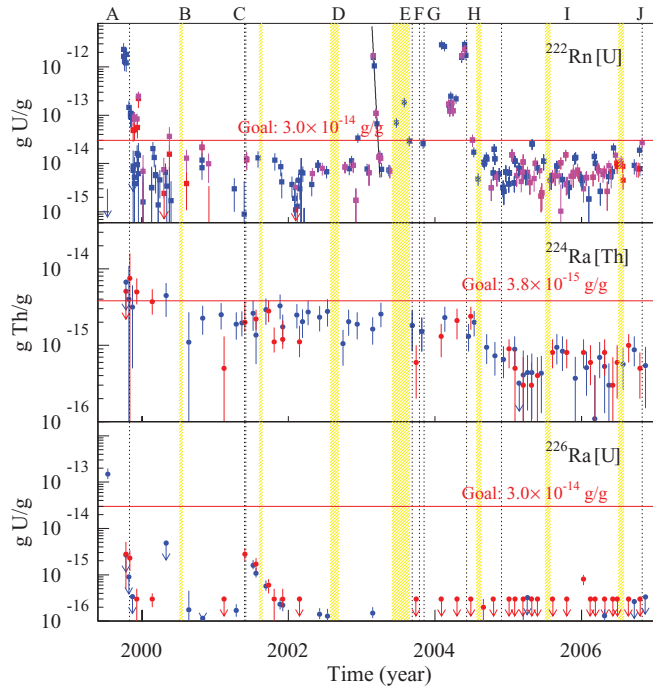


FIG. 18. (Color) D_2O radioactivity measurements by *ex situ* assays are shown for ^{222}Rn (top), ^{224}Ra (middle), and ^{226}Ra (bottom). The dashed lines are the bounds for different SNO detector configurations (from periods A to J, as indicated): A, phase-I (unadulterated D_2O) commissioning; B, phase-I operation; C, salination; D, phase-II (salt) operation; E, desalination; F, preparation for NCD array installation; G, NCD array deployment; H, NCD array commissioning; I, phase-III (NCD array) operation; and J, removal of the NCD array. Periods highlighted in yellow were times with reduced access to the underground laboratory. For the radon data, color represents different sampling points: red is at the top (near the chimney-sphere interface), purple one-third of the way down and blue at the bottom of the AV. The radon level was well below target for essentially the whole duration of the experiment, except for the high level at the beginning of the experiment, a large calibration spike in phase II, and during the deployment of the NCD array (when the cover gas protection was temporarily turned off). For the radium data, color indicates the technique used: red is for $HTiO$ and blue for MnO_x . Radium assays sampled the heavy water either from the top or the bottom of the AV. Again the initial higher concentrations quickly went below target. It should be noted that the data shown in this plot were not the only input to determining the radiopurity of the D_2O target.

with ^{214}Bi events and it was this difference in light isotropy that was used, in addition to differences in the radial distributions, to separate background components.

The light isotropy parameter was $\beta_{14} \equiv \beta_1 + 4\beta_4$, where

$$\beta_l = \frac{2}{N(N-1)} \sum_{i=1}^{N-1} \sum_{j=i+1}^N P_l(\cos \theta_{ij}). \quad (26)$$

In this expression P_l is the Legendre polynomial of order l , θ_{ij} is the angle between triggered PMTs i and j relative to the reconstructed event vertex, and N is the total number of triggered PMTs in the event. Details of β_{14} can be found in Ref. [6].

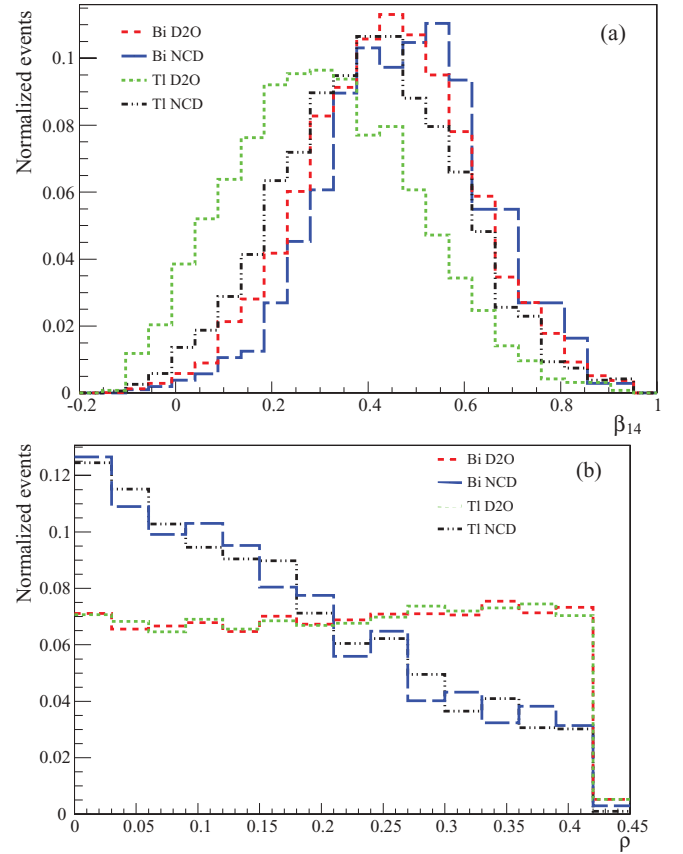


FIG. 19. (Color online) Probability distribution functions of (a) β_{14} and (b) the volume-weighted position $\rho = R_{fit}/R_{AV}$ for different background sources in the *in situ* measurement of D_2O and NCD array backgrounds.

Radioactive decays originating from the NCD array had an exponential radial profile while those from the D_2O had an approximately flat radial profile. The radial profiles were statistically indistinguishable for different radioisotopes that originated from the same location. Therefore, to distinguish between ^{208}Tl and ^{214}Bi , differences in event isotropy were used. By analyzing events that reconstructed with $R_{fit} < 450$ cm they can be classified as ^{208}Tl and ^{214}Bi in the D_2O or NCD strings. The *in situ* method provided continuous monitoring of backgrounds in the neutrino data set and a direct measurement of ^{214}Bi and ^{208}Tl , both of which could cause photodisintegration, without making any assumptions about equilibrium in the decay chain. The β_{14} and the volume-weighted radial position, ρ , distributions for the different background sources in this *in situ* measurement are shown in Fig. 19.

There were four main radioactive signals in the D_2O region: uranium- and thorium-chain activities in the D_2O and in the NCD strings. Assuming these were the dominant contributions in the *in situ* analysis signal region, a two-dimensional (radial and β_{14}) maximum likelihood fit was made to the data. The normalized PDF associated with each background was constructed from MC simulations of ^{214}Bi and ^{208}Tl from the D_2O and NCD strings. The simulated events were selected using the same cuts used on the data. Cherenkov light produced

by ^{208}Tl decays in the NCD string bodies was dominated by the 2.614-MeV γ rays. There was very little contribution from β 's as nearly all of them were stopped in the nickel of the NCD counter housing. Thus, ^{208}Tl decay events in the NCD string bodies were less isotropic and had a higher average value of β_{14} , when compared with events from ^{208}Tl decays in the D_2O . The mean value of β_{14} for these events was very similar to that of ^{214}Bi decays in D_2O . Bismuth decays in the NCD array produced even less light than ^{208}Tl decays. In addition to β 's being stopped in the nickel bodies, γ rays from ^{214}Bi decays have lower energies than those in ^{208}Tl decays. These ^{214}Bi decays were very similar in isotropy to ^{214}Bi D_2O events. Therefore, ^{214}Bi D_2O , ^{214}Bi , and ^{208}Tl ‘‘NCD bulk’’ isotropy distributions were very similar, but were significantly different from that of ^{208}Tl D_2O .

Combining radial and isotropy information, the ^{208}Tl D_2O , ^{214}Bi D_2O , and NCD bulk events could be separated. It was not possible to distinguish between ^{208}Tl and ^{214}Bi events originating from the NCD bulk. The ratio of U to Th in the NCD bulk obtained by coincidence studies of NCD array's α data [38] was used to separate the NCD events obtained by the 2D maximum likelihood fit into ^{208}Tl and ^{214}Bi . In the ^{232}Th chain, ^{220}Rn α decays to ^{216}Po , which decays by α emission. The signature for this coincidence was a 6.288-MeV α followed by a 6.778-MeV α with a half-life of 0.15 s. In the ^{238}U chain the coincidence was between ^{222}Rn and ^{218}Po , the signature being a 5.49-MeV α followed 3.10 min later by a 6.02-MeV α . The results from the α coincidence measurements were $2.8_{-0.8}^{+0.6} \times 10^{-12}$ gU/gNi and $5.7_{-0.9}^{+1.0} \times 10^{-12}$ gTh/gNi.

For comparison with *ex situ* assay and α coincidence measurements, the number of ^{214}Bi and ^{208}Tl events were converted into equivalent amounts of ^{238}U and ^{232}Th by assuming secular equilibrium and using MC simulations. The equivalent concentrations, integrated over the solar neutrino data set, were found to be $6.63_{-1.22}^{+1.05} \times 10^{-15}$ gU/g D_2O , $0.88_{-0.27}^{+0.27} \times 10^{-15}$ gTh/g D_2O , $1.81_{-1.12}^{+0.80} \times 10^{-12}$ gU/gNi, and $3.43_{-2.11}^{+1.49} \times 10^{-12}$ gTh/gNi. The *in situ* results of the NCD bulk were in good agreement with those obtained from the α coincidence analysis.

Results from the *in situ* and *ex situ* analyses of the D_2O were found to be consistent. As the two methods and their systematic uncertainties were independent, the best measurement of the equivalent concentrations of ^{232}Th in the D_2O was obtained by taking the weighted mean of the *in situ* and combined HTiO and MnO_x results. The weighted mean of the *in situ* and *ex situ* Rn results were used to obtain the best measurement of the ^{238}U content in the D_2O . The weighted mean concentrations were $(6.14 \pm 1.01) \times 10^{-15}$ gU/g D_2O and $(0.77 \pm 0.21) \times 10^{-15}$ gTh/g D_2O . Only the *in situ* results for the NCD array bulk concentrations were used because the α coincidence method could not provide data for all NCD counters and it could not sample the whole array.

3. Radioactive hot spots on NCD strings

Two areas of increased activity (hot spots) were identified in strings K5 and K2. The *in situ* method identified an excess of events close to each of these strings, but it could not prove conclusively that these events were caused by radioactivity.

TABLE IX. Equivalent masses of uranium and thorium for lower-chain activities in the K5 and K2 hot spots. More detailed and updated results may be found in Ref. [39].

	String	^{232}Th (μg)	^{238}U (μg)
Total <i>ex situ</i>	K5	1.28 ± 0.14	$0.10_{-0.05}^{+0.05}$
Total <i>in situ</i>	K5	$1.48_{-0.27}^{+0.24}$	$0.77_{-0.23}^{+0.19}$
Total <i>ex situ</i>	K2	1.43 ± 0.17	<0.40
Total <i>in situ</i>	K2	<0.93	$\equiv 0$

The isotropy distribution of the events associated with K5 was more isotropic than that of the NCD bulk, implying either radioactivity or scintillant on the surface of the NCD strings. If the hot spot was radioactivity, an excess of neutrons should have been captured by the contaminated string. However, K5 had a gain drift problem (Sec. III) and the number of neutrons captured by this string could not be quantified. No excess α were observed in the data from K2, suggesting that the contamination was embedded in the dead region of the counter, a conclusion that was supported by the *in situ* analysis. An extensive experimental program was developed to measure the radioactive content of these hot spots, and more details can be found in Ref. [39].

The lower-chain hot-spot activities expressed in terms of equivalent masses of ^{232}Th and ^{238}U are summarized in Table IX. For the analysis presented here, the neutron rates have been calculated using the weighted average of the *in situ* and *ex situ* data. The updated analysis presented in Ref. [39] yields changes that are negligible relative to the uncertainties in the final result.

B. Other neutron backgrounds

1. Internal-source neutrons

In addition to photodisintegration backgrounds, there were other neutron backgrounds that were generated in the D_2O . These included contributions from (α, n) reactions on nuclei, spontaneous fission from ^{238}U , cosmic-ray spallation, and anti-neutrinos from nuclear reactor and atmospheric neutrinos.

Neutrons can be produced by α reactions on ^2H , ^{17}O , and ^{18}O . The most significant contribution to this background arose from the 5.3-MeV α produced by ^{210}Po decay. In the third phase, Po isotopes on the external surface of the NCD array were of particular concern. Taking the average α activity of 18 samples, a total surface area of 2.40 m^2 that was counted using a multiwire proportional counter, yielded a neutron production rate of $1.32 \pm 0.28 \times 10^{-2}$ neutron/day generated from the entire array, 64.72 m^2 , given that 6.4×10^{-8} neutron were produced per α . This rate resulted in 5.1 ± 1.1 neutrons produced during phase III.

The contribution from spontaneous fission of ^{238}U was determined from the results of *ex situ* HTiO assays which placed limits on the concentration of ^{238}U in the detector.

The NUANCE [40] neutrino MC simulation package was used in the calculation of neutron backgrounds produced by atmospheric neutrino interactions. These atmospheric neutrino interactions were often associated with a burst of events in the

detector. After applying time-correlation cuts that removed event bursts and other data cuts to these simulated events, the expected number of observed neutrons from atmospheric neutrino interactions was determined to be 13.6 ± 2.7 for the NCD array and 24.1 ± 4.6 for the PMT array. The dominant systematic uncertainties associated with these estimates were those in the neutrino interaction cross section and atmospheric neutrino flux.

The muon flux incident on the SNO detector was measured to be $3.31 \pm 0.09 \times 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$ [41,42]. The possibilities that the muon tag could be missed and that neutrons could be produced in the surrounding rock, led to an estimate of their total rate of 0.18 ± 0.02 neutron per year, a negligible background.

Antineutrinos produced by nuclear reactors afar could also create neutrons in the D_2O . The magnitude of this background was calculated assuming an average reactor antineutrino spectrum and an average power output of all commercial reactors within 500 km of the SNO detector. Oscillations were taken into account in the calculation, and the estimate was 1.4 ± 0.2 neutrons per year. An estimate was also made of the number of neutrons produced by antineutrinos from radioactive decays within Earth and this was found to be 0.02 neutron per year.

The SNO detector was also sensitive to CNO neutrinos from the Sun. It was estimated that 1 neutron per year would be produced by this process. A signal of <0.2 event was expected in the NCD array.

2. External-source neutrons

Radioactive backgrounds in the AV and the surrounding H_2O could bring forth photodisintegration neutrons in the D_2O target. Neutrons could also be produced via (α, n) reactions in the AV. The total neutron backgrounds owing to these external sources were found to be $20.6^{+10.2}_{-7.3}$ events for the PMT array data and $40.9^{+20.6}_{-17.9}$ events for the NCD array data.

The *in situ* and *ex situ* techniques were applied to the H_2O region between the AV and PSUP. In total, 29 MnO_x and 25 $\text{HTiO H}_2\text{O}$ assays were conducted during the third phase. The results from the HTiO and MnO_x assays were in good agreement. The results were corrected for the neutrino live time and the weighted average was calculated to produce a single *ex situ* ^{232}Th -chain measurement for the H_2O region. The activity was found to be $26.9 \pm 12.3 \times 10^{-15} \text{ gTh/gH}_2\text{O}$. Owing to a weld leak in the H_2O sampling line, which was discovered towards the end of the phase, results from the Rn assays were not used in this analysis.

The *in situ* analysis window for the H_2O region was $4.0 < T_{\text{eff}} < 4.5 \text{ MeV}$ and $650 < R_{\text{fit}} < 680 \text{ cm}$. The equivalent ^{238}U and ^{232}Th concentrations were determined using an isotropy fit to the data. The background levels determined by the *in situ* analysis were $30.0^{+9.2}_{-19.4} \times 10^{-15} \text{ gTh/gH}_2\text{O}$ and $35.0^{+9.9}_{-5.4} \times 10^{-14} \text{ gU/gH}_2\text{O}$. The photodisintegration neutron background from Th activity in the H_2O was determined from a weighted mean of the *in situ* and *ex situ* assay results, while that from Rn was determined exclusively from the *in situ* results. The photodisintegration neutron backgrounds owing to Th and U

in the H_2O region were found to be $2.2^{+0.8}_{-0.7}$ events for the PMT array data and $7.1^{+5.4}_{-5.2}$ events for the NCD array data.

During its construction the AV was exposed to Rn in the underground laboratory air. The subsequent Rn daughters became embedded in the acrylic and could initiate (α, n) reactions on ^{13}C , ^{17}O , and ^{18}O . The activity on the surface of the AV was directly counted using silicon counters. Results from measurements performed at the end of the third phase were in agreement with those performed at the end of the second phase. Thus, the rate of these external-source neutrons from the vessel was taken to be the same as for phase II [6]. Adding the photodisintegration neutron backgrounds owing to intrinsic Th and U in the acrylic, which were determined from *ex situ* assays [5], the total external-source neutron backgrounds from the AV were found to be $18.3^{+10.2}_{-7.3}$ counts for the PMT array data and $33.8^{+19.9}_{-17.1}$ counts for the NCD array data.

C. Other Cherenkov light backgrounds

A 20-s veto removed the majority of radioactivity that followed a cosmic muon event. The duration of this veto window was about three half-lives of the cosmogenically produced ^{16}N , whose contribution to the total background was estimated at 0.61 ± 0.61 event in the PMT array data. Other Cherenkov light background events inside and outside the fiducial volume were estimated using calibration source data, measured activities, MC calculations, and controlled injections of Rn [31] into the detector. These backgrounds were found to be small above the analysis energy threshold and within the fiducial volume, and were included as an additional uncertainty on the flux measurements. Isotropic acrylic vessel background (IAVB) events were identified in previous phases [6]. It was estimated that <0.3 IAVB event (68% CL) remained in the PMT array data after data reduction cuts.

VIII. SIMULATION OF PULSES IN THE NCD ARRAY

The largest source of backgrounds in the neutron signal region in the NCD counters was α decays from the construction materials of the array. This was a very difficult background to calibrate, as any α particles from external calibration sources would not have sufficient energy to penetrate the counter wall. The spatial distribution and the composition of trace radioactivity in the counters also varied from counter to counter; therefore, background samples from the ^4He counters were not sufficient to fully characterize this all-important background to the neutron signal. An extensive MC, discussed in this section, was developed to simulate ionization pulses in the NCD counters and was used in defining the α background spectral shape for the solar neutrino analysis. Further details can be found in Refs. [43–45].

A. Physics model

The NCD counter simulation created ionization tracks for protons, tritons, α 's, and β 's in the NCD counter gas. α energy loss in the nickel wall was also calculated if necessary. Track formation for β 's was handled by EGS4 [29]. Proton, triton,

and α tracks were all calculated using the same procedure as follows. The track was divided into N (typically 5000–20 000) $1\text{-}\mu\text{m}$ -long segments such that each segment could be approximated as a point charge. The total current resulting from the whole track at time t was the sum of the individual currents from each track segment, i . The current induced on the anode wire from each segment was mainly a result of a positive charge, $q_i = en_i$, drifting towards the cathode [46]:

$$I_{\text{track}}(t) = \sum_{i=1}^N G_i n_i \frac{q_i}{2 \ln(b/a)} \frac{1}{t - t_0 + \tau}, \quad (27)$$

where G_i is the gas gain, n_i is the number of electron-ion pairs created in segment i , $a = 25 \mu\text{m}$ is the anode radius, $b = 2.54 \text{ cm}$ is the NCD-counter inner radius, t_0 is the “start time” for the current from the i th segment, and τ is the ion-drift time constant. The number of ion pairs depended on the stopping power, $\frac{dE}{dx}$; the mean energy required to produce an electron-ion pair in the gas, W ; and the segment length l , such that $n_i = \frac{dE}{dx} \frac{l}{W}$.

The description of an ionization track involved knowing where each segment was located and how much energy had been deposited there. Multiple scattering of the ionizing particle was simulated with the Ziegler-Biersack-Littmarck method [47]. The results of that simulation were in excellent agreement with the full TRIM MC calculation [48]. Values of $\frac{dE}{dx}$ were determined with stopping-power tables from TRIM for protons, tritons, and α 's.

The average energy required to produce an electron-ion pair in the NCD counter gas was measured using neutron sources and undeployed NCD counters. Integrating over many current pulses with average energy \bar{E} , the ratio G/W is proportional to the total current in a proportional counter, I [49]:

$$\frac{G}{W} = \frac{I}{\eta e \bar{E}}, \quad (28)$$

where η is the rate of neutron captures and e is the electron charge. \bar{E} was determined with the NCD counter MC to be $(701 \pm 7) \text{ keV}$. W is a characteristic of the NCD counter gas. It is approximately energy-independent, and is approximately equal for protons, tritons, and α 's [50]. We measured W by operating the counter in the “ion saturation” mode (200–800 V) and G/W by operating the counter at the standard voltage (1950 V). In the first case, we found $W = 34.1 \pm 12.4 \text{ eV}$, and in the latter case, $W = 34 \pm 5 \text{ eV}$, which was used in the MC.

A low-energy electron transport simulation was developed to evaluate the mean drift times, t_d , of electrons in the NCD counter gas mixture as a function of radial distance from the anode wire r . The results were in good agreement with GARFIELD [51] predictions, measurements by Kopp *et al.* [52] and further verifications made by inspecting specific types of α pulses. The $t_d(r)$ curve used in the simulation was

$$t_d = 121.3r + 493.9r^2 - 36.71r^3 + 3.898r^4, \quad (29)$$

with t_d in ns and r in cm.

Electron diffusion resulted in a radially dependent smearing effect on all pulses, and dominated the time resolution. A smearing factor σ_D was tabulated as a function of r and applied

in pulse calculations. σ_D and t_d were linearly related:

$$\sigma_D(t_d) = 0.0124 t_d + 0.559. \quad (30)$$

The mean NCD counter gas gain \bar{G} , as a function of voltage, was well described by the Diethorn formula [49]. However, ion shielding from the charge multiplication can significantly change the gas gain. A two-parameter model was developed to account quantitatively for this space-charge effect. The change in gas gain, δG , resulting from a change in wire charge density, $\delta\lambda(\bar{r})$, owing to the ions formed near the anode at voltage V is

$$\delta G \propto \bar{G} \ln(\bar{G}) \frac{\ln(b/a)}{2\pi\epsilon_0 V} \left[1 + \frac{1}{\ln(r_{\text{av}}/a)} \right] \delta\lambda(\bar{r}), \quad (31)$$

where ϵ_0 is the permittivity of free space and $r_{\text{av}} = 58 \pm 10 \mu\text{m}$ is the mean avalanche radius. $\delta\lambda$ can be obtained by dividing the induced charge by a characteristic shower width in the spatial dimension parallel to the anode wire, \mathcal{W} . The other parameter that needs to be optimized is the constant of proportionality in this equation. Electrons originating from some segment of a track are affected by the density changes $\delta\lambda_j$ owing to ions formed in previous electron cascades. Each of these ion clusters moves slowly towards the cathode while the primary electrons are being collected. In the presence of many ion clusters, the total change in the anode charge density at time t , experienced by electrons from the i th track segment is therefore

$$\delta\lambda_i = \frac{e}{\mathcal{W}} \sum_{j=1}^{i-1} \frac{\ln(b/\bar{r}_j(t))}{\ln(b/a)} G_j n_j + \frac{e}{\mathcal{W}} \frac{\ln(b/\bar{r})}{\ln(b/a)} n_i, \quad (32)$$

where n_j is the number of ion pairs formed in the j th segment. j loops over all previous ion clusters, which have moved to different radii $\bar{r}_j(t)$ at time t . $\bar{r}_j(t)$ is solved by integrating the relation $\frac{dr_j}{dt} = \mu_i \mathcal{E}$:

$$\bar{r}_j(t)^2 = \frac{2\mu_i V t}{\ln(b/a)} + r_{\text{av}}^2, \quad (33)$$

where μ_i is the ion mobility and \mathcal{E} is the cylindrical electric field.

The mean gas gain of the i th track segment is $\bar{G}_i = \bar{G} - \delta G_i$. The actual G_i applied to the i th segment is sampled from an exponential distribution with mean \bar{G}_i .

The smaller ion mobility, relative to that of the electrons, results in the long tail that is characteristic of pulses from ionization in the NCD counters. The evolution of a current pulse in a cylindrical proportional counter is described by Eq. (27). The ion time constant, τ , is inversely proportional to the ion mobility, μ :

$$\tau = \frac{a^2 p \ln(b/a)}{2\mu V}, \quad (34)$$

where p is the gas pressure.

We measured τ using neutron calibration data. Ionization tracks that are parallel to the anode wire have a relatively simple underlying structure; the primary ionization electrons all reach the anode at approximately the same time, with some spread owing to straggling. The ion-tail time constant was extracted by selecting the narrowest neutron pulses from calibration data sets and fitting each pulse with a Gaussian

convolved with the ion tail, a reflection and the electronic model. This model fitted the peaks well enough to allow for a characterization of the ion tail. We found $\tau = 5.50 \pm 0.14$ ns. This time constant corresponds to an ion mobility of $\mu = (1.082 \pm 0.027) \times 10^{-8}$ cm² ns⁻¹ V⁻¹. For the NCDs, the ions present consisted primarily of CF₃⁺, F⁻, and CF₃⁻. For comparison, in the approximate range of electric field strengths found in the NCDs, the mobilities of CO₃⁻, CO₂⁺, and O⁻ in He gas are $\approx (10-12) \times 10^{-8}$ cm² ns⁻¹ V⁻¹, $(10-15) \times 10^{-8}$ cm² ns⁻¹ V⁻¹, and $(13-20) \times 10^{-8}$ cm² ns⁻¹ V⁻¹, respectively [53].

B. Simulation of NCD array electronics

Propagation of the pulse along the NCD string was simulated with a lossy transmission line model. Half of the pulse was propagated down the string, through the delay line, and back to the point of origin of the pulse. The delay-line attenuation was also simulated as a lossy transmission line. Both halves of the pulse (reflected and direct) were then transmitted up to the top of the NCD string. The model parameters were based on SPICE simulations [54] and *ex situ* measurements.

Propagation in the NCD counter cable was simulated with a low-pass filter (RC ≈ 3 ns). There was a small reflection (reflection coefficient = 15%) at the preamp input owing to the slight impedance mismatch between the preamp input and the cable. This portion of the pulse traveled to the bottom of the NCD string and reflected back upwards.

The preamplifier was simulated with a gain (27 500 V/A), a low-pass filter (RC ≈ 22 ns), and a high-pass filter (RC = 58 000 ns). All RC constants in the electronic model were measured by fitting the model to *ex situ* injected pulses.

The frequency response of the multiplexer system before the logarithmic amplifier was simulated with a low-pass filter (RC ≈ 13.5 ns). The constants used to parametrize the logarithmic amplification were the same constants used to “de-log” real data pulses, which were determined by regular *in situ* calibrations during data taking. The circuit elements after the logarithmic amplification were simulated with the final low-pass filter (RC ≈ 16.7 ns). The pulse array values were rounded off to the nearest integer to replicate the digitization.

Noise was added to the pulses as the final stage in the simulation. It was added to the multiplexer and shaper branches of the electronics independently. For the multiplexer branch, the frequency spectrum of the noise on the current pulses was measured for each channel using the baseline portions of injected calibration data.

The shaper-ADC branch of the electronics was simulated by a sliding-window integral of the preamplified pulse. This number was then converted to units of ADC counts by doing an inverse linearity calibration. The calibration constants used in this “uncalibration” were the same constants that were used to calibrate the data. Noise was added to the shaper value with a Gaussian-distributed random number. The mean and standard deviation of the noise for each channel were determined uniquely for each of the 40 NCD strings; the typical RMS noise (in units of ADC values) was 2.0, with a variance of 0.7 across the array.

The multiplexer and shaper systems included independent triggers that used the true threshold values. Event triggers were determined by checking each pulse amplitude or shaper value against the appropriate threshold. The dead times of the two systems were then taken into account within each MC event. The NCD string signals were integrated with the PMT trigger simulation by inserting each signal into the time-ordered array of PMT signals. As the simulation scanned over the combined PMT and NCD array signals, any individual NCD string signal was sufficient to cause a global trigger of the detector.

Certain parts of the overall simulation were relatively slow owing to the loops over the large ($N = 17\,000$) arrays containing the simulated pulses. As a result we implemented a fast alternative to the full simulation. The ionization track was simulated to determine the timing of the event and the energy deposited in the gas. That energy was converted directly to an approximate shaper-ADC measurement and was smeared with a Gaussian to roughly account for the missing physics and electronic noise. This option allowed simulations for preliminary comparisons with the data because they did not require pulse-shape calculations.

C. Verification and systematic uncertainties

The ratio of the number of bulk uranium and thorium α events to that from cathode-surface polonium in the NCD strings could not be fully represented in the ⁴He string data owing to string-to-string variation in these backgrounds. Simulations were used to calculate the α energy spectrum PDFs instead. Therefore, it was important to accurately simulate these PDFs and to assess their systematic uncertainties for use in the region of interest for the solar neutrino analysis.

We optimized and validated the NCD array signal simulation by comparing with several types of data, including neutron source calibrations, high-energy α events, and ⁴He-string α data. This procedure was designed to ensure that the simulations accurately reproduced the data, without in any way tuning on a data set that contained the neutron signal, which was determined by signal extraction (Sec. IX).

The comparison of the simulation with neutron calibration data tested nearly all aspects of the simulation physics model. We compared ²⁴Na neutron calibration data with simulations for a number of pulse characteristics, such as the mean, width, skewness, kurtosis, amplitude, and integral, and timing variables, including the rise time, integral rise time, and full width at half maximum. These comparisons were used to estimate parameter values and uncertainties for electron and ion motion in the NCD counter gas, as well as the space charge model. Figure 20 shows a comparison of some of these pulse-shape variables between real and simulated ²⁴Na neutron data.

We estimated the fraction of surface polonium and bulk α events on each string by fitting the energy distribution above the neutron energy region, shown in the bottom plot in Fig. 21. After the fit, we calculated an event weight, which was a function of string number and α type (polonium, uranium, or thorium) describing the best-fit fraction of α 's on each string owing to each source. In general, polonium comprised $\sim 60\%$ of the α signal; however, there were $\pm 20\%$ variations between strings. The best-fit α fractions and the MC energy

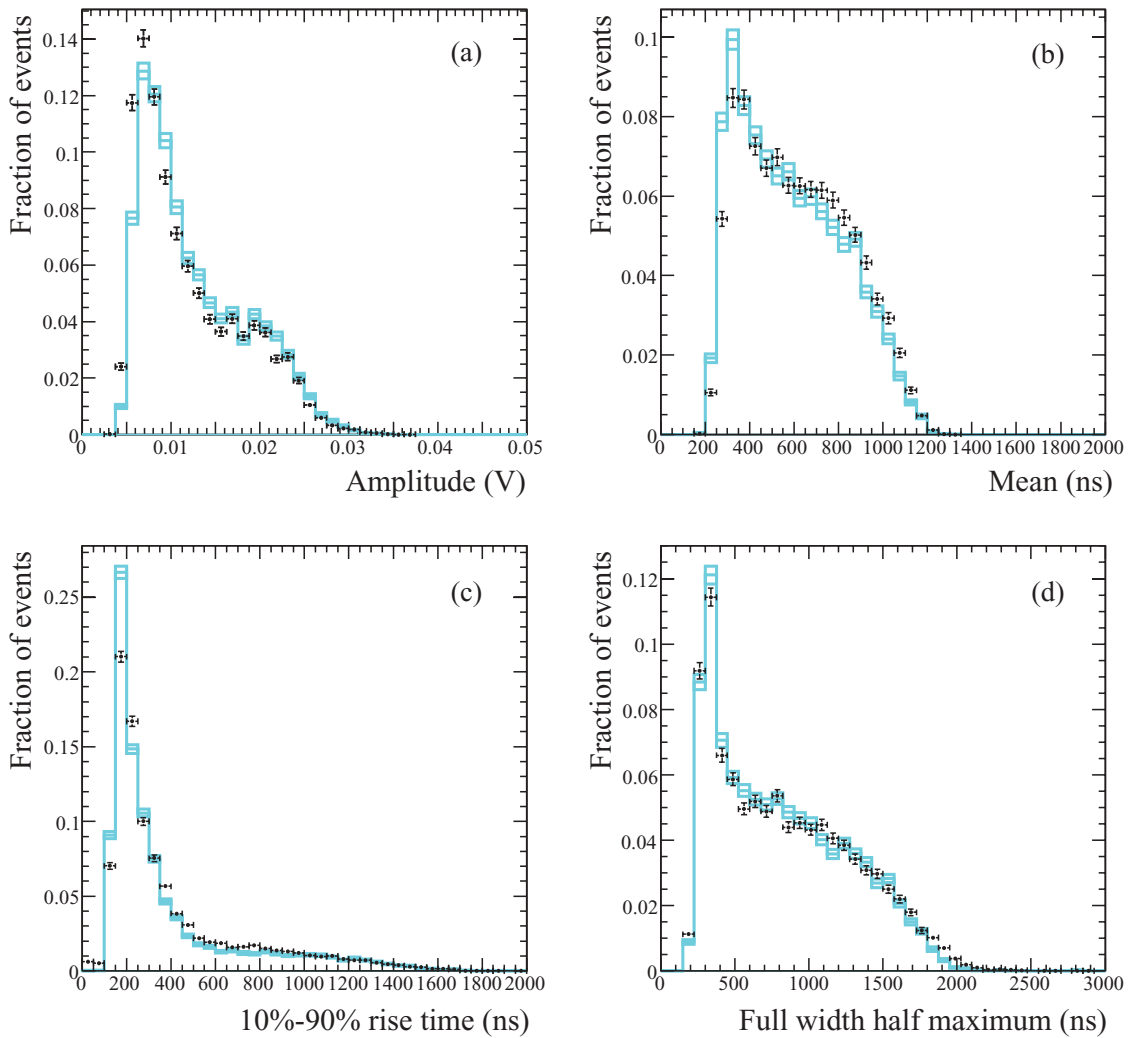


FIG. 20. (Color online) Comparison of pulse shape variables in ^{24}Na neutron calibration data (data points) and the NCD array MC (histogram) in the neutron energy window, 0.4 to 1.2 MeV, with statistical uncertainties only. The purpose of these figures is to compare different pulse-shape variables and moments of their distributions between simulation and data. The background rate during this calibration was negligible. (a) Fraction of events as a function of pulse amplitude, (b) time-axis mean of the pulse, (c) 10%–90% rise time, and (d) full width at half maximum. All distributions are normalized to unit area.

scale correction were applied on an event-by-event basis to produce the α energy spectrum PDF.

The systematic uncertainties included the depth of α -emitting contaminations (“ α depth”) within the NCD counter walls, the efficiency of data reduction cuts, the space-charge model parameters, the electron-drift curve, the ion mobility, and the surface polonium to NCD bulk activity fraction on each string. These effects reflected uncertainties in the parameters of the NCD array simulation physics model. Systematic uncertainties were assessed by generating a large set of variation MC samples, each with one input parameter varied by 1σ with respect to its default value. The size of the variation of the NCD array MC systematic uncertainties was estimated from *ex situ* data, off-line measurements of the NCD array signal processing electronic response, and *in situ* constraints from the NCD array data. The most significant sources of systematic uncertainty were attributable to variations of the

simulated α depth within the NCD counter walls and the data reduction cuts.

We calculated fractional first derivatives to describe the changes in the MC α energy spectrum allowed by the systematic uncertainties [43]. For each of the systematic uncertainties, the first derivatives for each bin of the energy distribution were calculated by taking the difference between histograms containing the standard MC prediction for the shape of the energy distribution and the variational MC energy-distribution shapes. Then, the total systematic uncertainty in each energy bin was assessed by summing the eight contributions in quadrature. The functional dependence of those derivatives on energy is shown in the top plot of Fig. 21.

The final simulated α energy spectrum in the solar neutrino analysis window for the NCD array is shown in the center plot of Fig. 21, with systematic uncertainties. This α spectrum was used as the α background PDF, together with the neutron signal

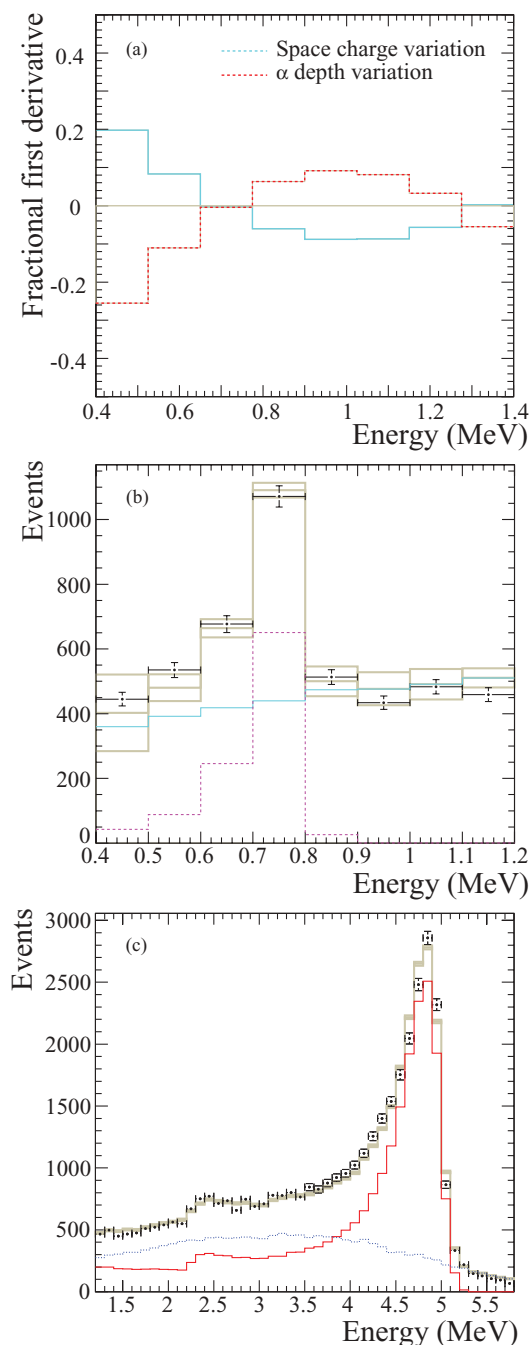


FIG. 21. (Color) (a) Fractional first derivatives vs event energy (MeV) for two systematic variations (space charge and α depth) of the simulation. (b) Number of events versus energy in the NCD array data analysis energy window, for data (black, with statistical uncertainties), neutron template from calibration data (purple, dashed), α cocktail template from simulation (cyan), and total predicted number of events (gray, with systematic uncertainties) in the “open” data set (see Sec. IX). (c) Energy spectrum above the NCD array data analysis energy window, for data (black, with statistical uncertainties), alpha “cocktail” (gray, with systematic uncertainties), polonium (red), and bulk (blue dashed) simulation. The “cocktail” is a collection of simulated α pulses with the appropriate mixture of NCD string cathode-surface polonium and bulk uranium and thorium α events.

template from ^{24}Na neutron calibration data, to determine the total number of neutron events.

IX. NEUTRINO SIGNAL DECOMPOSITION

This section describes the techniques used in the SNO phase-III neutrino flux measurement (referred to as signal extraction). Three different signal extraction methods were developed to determine from the phase-III solar neutrino data set the total active solar neutrino flux and the ν_e flux. These extended log-likelihood techniques were designed to perform a joint analysis of the data from the PMT and the NCD arrays. In the three methods, the energy spectrum of NCD array events was fit with the MC α background distribution described earlier, the neutron spectrum determined from ^{24}Na calibration, expected neutron backgrounds, and two instrumental background event distributions. Events from the PMT array were fitted in the reconstructed effective kinetic energy T_{eff} , the cosine of the event direction relative to the vector from the Sun $\cos\theta_{\odot}$, and the volume-weighted radius ρ . The nuisance parameters (systematic uncertainties), weighted by external constraints determined from calibrations and simulations, were allowed to vary in the fit of the neutrino signals. This “floating” technique enabled the determination of the correlations between the observed signals and the nuisance parameters.

There were, in general, two classes of systematic uncertainties: The first had a direct impact on the shapes of the PDF, such as energy scale or angular resolution for the PMT array data; the second are uncertainties on parameters that did not affect the PDF shapes, such as detection efficiencies. The three signal extraction techniques differed in the implementation of how the nuisance parameters were floated.

The first method was an extension to the signal extraction techniques [32] used in previous analyses [5,6]. Systematic uncertainties of the second kind were easily floated in this method. However, floating the first kind was much more challenging as it would require PDFs to be rebuilt between evaluation of the likelihood function in the minimization process. As a result, only a few significant systematic uncertainties were allowed to float in this method owing to computational limitations. Those systematic uncertainties that were not floated were estimated by repeating the fit with the parameters varied by their positive and negative 1σ deviations.

In the second method the nuisance parameters were not varied during the fit. Instead the signal extraction fit to the data was done many times with the necessary PDFs built from a set of nuisance parameters, whose values were drawn randomly according to their distributions. After this ensemble of fits, the flux result from each trial was reweighted in proportion to its likelihood. This sampling method circumvented the need to rebuild PDFs during the numerically intensive likelihood maximization process.

Markov-chain Monte Carlo (MCMC) was the third method used in signal extraction. The flux results from this method were published in the original letter on this phase of SNO [7], and the details in its implementation are presented in the following. These discussions include a description of

how the MCMC parameter estimation works, the extended log-likelihood function used to obtain parameter estimates, and finally the results of fits to the full third-phase data set.

It is noted that a blind analysis procedure was used. The “open” data set used during the development of the signal extraction procedures excluded a hidden fraction of the final data set and included an admixture of neutron events from cosmic-ray muon interactions. The blindness constraints were removed only after the robustness of the three analyses had been established. The finalization of the three signal extraction procedures before removing these constraints required an agreement of their results of ensemble tests with common simulated data sets and the flux results for the “open” data set to within the expected statistical spread. After the blindness constraints were lifted, the full data set was analyzed using the three independent methods and their results were compared to each other. This comparison of the results from the three analysis methods revealed two issues, which were understood prior to the publication of the letter [7]. These issues are discussed later in this section.

A. Markov-chain Monte Carlo parameter estimation

In SNO’s previous phases, the negative log-likelihood (NLL) function was simply minimized with respect to all parameters to get the best-fit value. Minimizing the NLL was very challenging with so many fit parameters, and the likelihood space near the minimum could be uneven. This means that common function minimization packages, such as MINUIT [55], could run into numerical convergence problems. The use of MCMC circumvented this convergence problem by interpreting NLL as the negative logarithm of a joint probability distribution P for all of the free parameters, that is, $P = \exp[\ln(L)]$. The nuisance parameters were integrated to determine the posterior distributions for the fluxes. The origins of this procedure go back to Bayesian probability theory, and in fact our approach could be considered to be a Bayesian analysis with uniform priors assumed for the fluxes. The MCMC method’s insensitivity to the unevenness in the NLL space means that it is better suited to handling large numbers of nuisance parameters.

The basic idea of MCMC is to take a random walk through the parameter space, where each step is taken with a probability given by the likelihood of the new step (L_{prop}) compared to the previous step (L_{curr}). The probability of accepting the proposed new parameters is $\min(1, L_{\text{prop}}/L_{\text{curr}})$. If the step is accepted the parameter values are updated, or else the MCMC repeats the current point in the chain and generates a new proposal for the parameters. By the Metropolis-Hastings theorem [56,57], the resulting distribution of parameters from the chain will have a frequency distribution given by L . The choice of an appropriate proposal kernel is critical to this method. If the width of the proposal distribution were too wide, steps are rarely taken, and if the width of the proposal distribution were too narrow, then the MCMC would not sample enough of the parameter space to find the best-fit region and instead could fall into a local minimum. In the MCMC signal extraction method, the step proposal distribution was a Gaussian of mean zero and width that was $\sim 1/3$ of the expected statistical uncertainty or

the constraint uncertainty. This choice was checked by the convergence and distributions of different MCMC chains that were started at different random starting points.

B. Observables, signals, and fit ranges

1. PMT array data

For the PMT array data, 33 signals and backgrounds were included in the fit. An energy-unconstrained fit was done, meaning that CC and the ES fluxes were fitted for each T_{eff} bin. The energy binning used for this unconstrained fit was 0.5-MeV bins between T_{eff} of 6 and 12 MeV, and a single bin from 12 to 20 MeV, totaling 13 bins per signal. The other signal and backgrounds in the fit included the NC signal; photodisintegration neutron backgrounds owing to radioactivity in the D_2O target, the bulk nickel in the NCD strings, and the hot spots on the strings K2 and K5; photodisintegration and (α, n) neutron backgrounds in the AV (which also included the backgrounds from the cables of the NCD array); and neutrons from atmospheric neutrino interactions.

The observables used in the fit to the PMT data were each event’s reconstructed T_{eff} , ρ , and $\cos\theta_{\odot}$. Cherenkov light candidate events were selected with the criteria $\rho < (\frac{550}{600})^2 = 0.77025$, $-1 \leq \cos\theta_{\odot} \leq 1$, and $6.0 \text{ MeV} < T_{\text{eff}} < 20.0 \text{ MeV}$. The signal and background PDFs for the PMT array data were three dimensional in ρ , $\cos\theta_{\odot}$, and T_{eff} . The exceptions were the photodisintegration backgrounds owing to K2, K5, and the bulk nickel, which were factorized [i.e., $P(T_{\text{eff}})P(\cos\theta_{\odot})P(\rho)$] to avoid any problems associated with low statistics in their simulations.

2. NCD array data

The NCD array signals and backgrounds in the fits were the NC signal and various sources of neutron backgrounds. These backgrounds include photodisintegration owing to radioactivity in the D_2O target, nickel housing of the NCD array, and hot spots on strings K2 and K5; external source neutrons owing to radioactivity in the AV and the H_2O shield; neutrons from atmospheric neutrino interactions and cosmogenic muons; and instrumental backgrounds with characteristics of those seen in strings J3 and N4.

The only observable used in signal extraction was the summed energy spectrum of the shaper-ADC in the NCD array readout (“shaper energy,” E_{NCD}), restricted to the range of $0.4 \text{ MeV} < E_{\text{NCD}} < 1.4 \text{ MeV}$. The shaper energy spectrum from a uniformly distributed ^{24}Na source was used as the NC PDF. The α background PDFs were derived from the simulation as discussed in Sec. VIII.

C. The extended log-likelihood function

In the extended log-likelihood function, two different approaches were used to handle systematic uncertainties. To include systematic uncertainties associated with the PMT array data and uncertainties associated with the neutron PDFs in the NCD array data, the respective PDFs were rebuilt from an array of event observables on each evaluation of the log-likelihood function. For the systematic uncertainties associated with the α background PDFs in the NCD array data, the shapes of those PDFs were modified by a multiplicative function.

Twenty-six systematic uncertainties were included in the fit. The systematic uncertainties owing to neutron detection efficiency were applied to the NC PDFs separately for the NCD and PMT data streams. The 12 systematic uncertainties associated with the PDF shapes for the PMT array were related to its energy scale and resolution, as well as the event vertex reconstruction algorithm's spatial accuracy, spatial resolution, and angular resolution. This last entity was applied to the ES signal only. The 12 systematic uncertainties associated with the PDF shapes for the NCD array were related to the energy scale ($a_1^{E_{\text{NCD}}}$) and resolution ($b_0^{E_{\text{NCD}}}$) of the counters, data reduction cut efficiency on alpha backgrounds, spatial distributions and intensity of different α background activities, energy spectrum of instrumental backgrounds, and parameters in the MC that could affect the observed signal, such as ion mobility. Details of the parametrization of these systematic parameters can be found in Appendix B.

A constraint term for each of the 26 systematic uncertainties was added to the log-likelihood function. The means and standard deviations of these constraint terms came from calibration measurements or MC studies and are listed in Table XIV in Sec. IX D.

The NLL function to be minimized was the sum of a NLL for the PMT array data $-\ln L_{\text{PMT}}$ and for the NCD array data ($-\ln L_{\text{NCD}}$):

$$-\ln L = -\ln L_{\text{PMT}} - \ln L_{\text{NCD}}. \quad (35)$$

The following subsections describe how these two NLLs were handled in this analysis.

1. PMT array: NLL with systematic uncertainties

For the PMT array data, the goal was to allow the solar neutrino data to tell us how their scales and resolutions differed from the simulations within the constraints from calibration data. We therefore modeled the differences as a possible remapping of the observables for the simulated events. The PDFs associated with the PMT array (P_{mc}) were rebuilt using the scaled and smeared value of T_{eff} , ρ , and $\cos\theta_{\odot}$. We then fitted for the remapped parameters to determine the allowable range of this remapping while still matching the data observables.

The negative log-likelihood for the PMT array data can therefore be written as

$$-\ln L_{\text{PMT}} = \sum_i^{M^{\text{PMT}}} f_i^{\text{PMT}} S_i^{\text{PMT}} \phi_i - \sum_j^{N^{\text{PMT}}} \ln \left(\sum_i^{M^{\text{PMT}}} f_i^{\text{PMT}} S_i^{\text{PMT}} \times \phi_i P_{\text{mc}}^i(\rho_j, \cos\theta_{\odot,j}, T_{\text{eff},j}) \right) - \ln L_{\text{constraints}}^{\text{PMT}}, \quad (36)$$

where there were M^{PMT} event classes (such as neutrino signals from different interactions and Cherenkov light backgrounds) and N^{PMT} PMT events. For the i th signal, the factor to convert from flux to the number of events is f_i , the fitted flux is ϕ_i , the fiducial volume correction factor is S_i , and the remapped

normalized PDF is P_{MC}^i . This expression is, in essence, a sum of the normalization (which is allowed to vary), the shape, and the constraint likelihoods. The constraint terms are

$$-\ln L_{\text{constraints}}^{\text{PMT}} = \frac{1}{2} \sum_i \left(\frac{p_i - \bar{p}_i}{\sigma_{p_i}} \right)^2 + \frac{1}{2} \sum_i \sum_j (b_i - \bar{b}_i)(b_j - \bar{b}_j)(V_b^{-1})_{ij}, \quad (37)$$

where p_i are all of the PMT systematic parameters other than two sets of correlated position resolution parameters denoted as b_i in the second term. V_b is the covariance matrix for these latter parameters. The parametrization of the nuisance parameters can be found in Appendix B.

2. NCD array: NLL with systematic uncertainties

For the NCD array data the probability distribution functions Q were one-dimensional functions of the shaper energy (E_{NCD}). Although the time separation of the direct and reflected pulse in the NCD strings can provide the location of the captured neutron, this position information was not used in the final solar neutrino analysis as the accurate calibration and determination of the z distributions of ionization owing to the neutron signal and the α backgrounds were not trivial, and would require resources, including a reduction to detector live time, for such calibrations. Also the manipulator system delivered calibration sources on two planes; thus, some of the NCD strings would require significant longer to calibrate with a point source.

For each evaluation of the likelihood, where the neutron PDF systematics were changed, the PDFs were rebuilt using E_{remap} :

$$E_{\text{remap}} = a_1^{E_{\text{NCD}}} E_{\text{NCD}} \times [1 + N(0, b_0^{E_{\text{NCD}}})], \quad (38)$$

where $N(\mu, \sigma)$ is a Gaussian distribution with mean μ and width σ . As mentioned previously, the simulated α background PDF was rebuilt by multiplying the unmodified PDF (Q^α) by a reweighting factor (α_i) and a multiplicative function in shaper energy $s_i(E_{\text{NCD}})$:

$$Q_{\text{mc}}^\alpha(E_{\text{NCD}}) = Q^\alpha(E_{\text{NCD}}) \left(1 + \sum_{i=1}^8 \alpha_i s_i(E_{\text{NCD}}) \right). \quad (39)$$

The eight reweighting functions s_i included systematic effects in the depth and intensity variations of ^{210}Po and other natural α emitters in the NCD string body, ion mobility, avalanche width and gradient offset, and drift time variation and data reduction cut efficiency. In addition, two instrumental background PDFs, based on instrumental background events observed in strings J3 and N4 (which were excluded from the neutrino candidate data set), were parameterized as skewed Gaussian distributions. A summary of the parametrizations of these systematic uncertainties can be found in Appendix B.

The negative log-likelihood for the NCD array data can therefore be written as

$$-\ln L_{\text{NCD}} = \sum_i^{M^{\text{NCD}}} f_i^{\text{NCD}} S_i^{\text{NCD}} \phi_i - \sum_j^{N^{\text{NCD}}} \ln \left(\sum_i^{M^{\text{NCD}}} f_i^{\text{NCD}} S_i^{\text{NCD}} \phi_i Q_{\text{MC}}^i(E_{\text{NCD},j}) \right) - \ln L_{\text{constraints}}^{\text{NCD}}, \quad (40)$$

where there are M^{NCD} event classes and N^{NCD} events in the NCD array data. For the i th event class, the flux-to-event conversion factor is f_i^{NCD} , the fitted flux is ϕ_i , the fiducial volume correction factor is S_i , and the remapped normalized PDF is Q_{MC}^i . The constraint terms are

$$-\ln L_{\text{constraints}}^{\text{NCD}} = \frac{1}{2} \sum_i \left(\frac{p_i - \bar{p}_i}{\sigma_{p_i}} \right)^2, \quad (41)$$

where p_i are all of the systematic parameters for the NCD array data.

In Eqs. (36) and (40) the factors S_i were used to propagate the fiducial volume uncertainty. They could change as the systematic parameters were changed and were calculated each time the PDFs were rebuilt. The other set of conversion factors f_i included live time and efficiencies that were used to convert the number of events in a fiducial volume into a flux above threshold. A full description of these conversion factors can be found in Ref. [5].

3. Neutron background and other input constraints

The neutron backgrounds in the PMT array data were determined from the neutron-background fits in the NCD array data. For N_i^{NCD} events fitted for a certain neutron background in the NCD array data, the number of background events from the same source is $f_i^{\text{PMT}} N_i^{\text{NCD}}$, where f_i^{PMT} is the ratio of the number expected in the NCD array data to that in the PMT array data. These conversion factors are summarized in Table X.

Several small Cherenkov light backgrounds required adjustments of the CC and ES fluxes. These backgrounds were Cherenkov events from β - γ decays in the D_2O and in

TABLE X. Multiplicative conversion factors for determining the number of neutron background events in the PMT array data from the number in the NCD array data. Photodisintegration is denoted as ‘‘PD’’ in this table.

Factor	Description	Value
$f_{\text{ex}}^{\text{PMT}}$	AV, PD	0.5037
$f_{\text{d2opd}}^{\text{PMT}}$	D_2O , PD	0.2677
$f_{\text{ncdpd}}^{\text{PMT}}$	NCD, PD	0.1667
$f_{\text{k2pd}}^{\text{PMT}}$	K2, PD	0.2854
$f_{\text{k5pd}}^{\text{PMT}}$	K5, PD	0.2650
$f_{\text{cab}}^{\text{PMT}}$	NCD Cables, PD	0.1407
$f_{\text{atmos}}^{\text{PMT}}$	Atmospheric ν	1.8134

the external regions (AV, H_2O , and PMT support geodesic structure), and isotropic light background from the AV. To account for these small backgrounds at each step in the chain, their contributions were randomly drawn, assuming a Gaussian distribution with means and widths given in Table VIII. These contributions were then split between the CC and ES channels, with the latter assuming 10% of the total. This fraction was applied because the ES peak occupies a corresponding fraction in the $\cos \theta_{\odot}$ distribution for neutrino signal.

D. Results of fits to the full phase-III data set

Fits without systematic uncertainty evaluation were obtained by running a single-fit Markov chain with 320 000 steps where the first 20 000 steps were rejected to ensure convergence. The fits that allow systematic uncertainty evaluation were obtained by running 92 independent Markov chains with 6500 steps each. Each fit was started with different starting parameters near the best-fit point varied by a Gaussian distribution with width given by the estimated uncertainty on the parameter. The first 3500 steps in this fit were rejected to minimize the effect of initial values on the posterior inference, and the remaining 3000 steps of each fit were put into a histogram for each parameter. A total of 276 000 steps were used to estimate the parameter uncertainties.

To ensure the robustness of the MCMC signal extraction, an ensemble test with all systematic parameters floated was performed. Each of the mock data sets were assembled with PDFs that were regenerated with different randomly sampled systematic parameter values. To run this ensemble test of 100 runs with all systematic parameters included, each with a MCMC chain length of 22 000, a substantial amount of computational power was required. As a result, only the final fit configuration of the full phase-III data set was tested. The fits to the real data were performed with far more steps to ensure that the uncertainty estimates were more robust than those in the ensemble test. This ensemble test showed an acceptable level of bias and pull in the fit parameters given the available statistics and established the reliability of this signal extraction method.

In the signal extraction fit, the spectral distributions of the ES and CC events were not constrained to the ${}^8\text{B}$ shape, but were extracted from the data. In Table XI, the number of events in different signal and background classes determined from this ‘‘energy-unconstrained’’ fit are tabulated. The equivalent neutrino fluxes, derived from the fitted number of CC, ES, and NC events under the assumption of the ${}^8\text{B}$ neutrino spectrum in Ref. [58], were determined to be (in units of $10^6/\text{cm}^2/\text{s}$) [59,60]:

$$\begin{aligned} \phi_{\text{CC}} &= 1.67_{-0.09}^{+0.08}, \\ \phi_{\text{ES}} &= 1.77_{-0.23}^{+0.26}, \\ \phi_{\text{NC}} &= 5.54_{-0.46}^{+0.48}, \end{aligned} \quad (42)$$

where the uncertainties are the total uncertainties obtained from the posterior distributions. Their correlations are tabulated in Table XII. The ratio of the ${}^8\text{B}$ neutrino flux measured

TABLE XI. Results of fit to the full phase-III data set. The fitted number of signal and background counts, along with the integral neutrino flux values, are shown. In the data set, the total number of events in the PMT array and NCD array data sets were 2381 and 7302, respectively. The ^8B spectrum from Ref. [58] was used in deriving the equivalent neutrino fluxes from the fitted number of CC, ES, and NC events. All the fluxes are in units of $\times 10^6/\text{cm}^2/\text{s}$.

Fitted counts, PMT array	
CC	1867^{+91}_{-101}
ES	171^{+24}_{-22}
NC	267^{+24}_{-22}
Backgrounds	77^{+12}_{-10}
Total	2382^{+98}_{-107}
Fitted counts, NCD array	
NC	983^{+77}_{-76}
Neutron backgrounds	185^{+25}_{-22}
α backgrounds	5555^{+196}_{-167}
Instrumental backgrounds	571^{+162}_{-175}
Total	7295^{+82}_{-83}
Integral flux	
ϕ_{CC}	$1.67^{+0.08}_{-0.09}$
ϕ_{ES}	$1.77^{+0.26}_{-0.23}$
ϕ_{NC}	$5.54^{+0.48}_{-0.46}$

with the CC and NC reactions is

$$\frac{\phi_{\text{CC}}}{\phi_{\text{NC}}} = 0.301 \pm 0.033(\text{total}). \quad (43)$$

In Table XIII, the CC and ES electron differential energy spectra from the energy-unconstrained fit are tabulated.

The signal extraction fit was done with an extended maximum likelihood method, which does not yield an absolute goodness-of-fit measure readily. One means of investigating the goodness of fit is to calculate the χ^2 between the data and the best-fit sum of the weighted PDFs in each of the one-dimensional projections. This is shown in Fig. 22. The fitted systematic parameter values are provided in Table XIV, while the constraints and fit results for the amplitude of different neutron and instrumental backgrounds are given in Table XV. The uncertainties on the fitted values of the systematic parameters were, for the most part, the same as the width of the constraint that was used in the fit. There are three systematic parameters that have fit uncertainties that are considerably narrower than the constraint. The shaper energy scale is narrower, and this appears to be a real effect of the neutron energy peak setting the energy scale. The

TABLE XII. Statistical correlation coefficients for the CC, ES, and NC fluxes in the full SNO phase-III data set.

Signal	CC	ES	NC
CC	1.000	0.2376	-0.1923
ES	0.2376	1.000	0.0171
NC	-0.1923	0.0171	1.000

TABLE XIII. The CC and ES electron differential energy spectrum. The fluxes in each of the 13 T_{eff} bins are in units of $10^4/\text{cm}^2/\text{s}$. The uncertainties shown are total uncertainties with correlations between all systematic uncertainties described in the text included.

T_{eff} (MeV)	CC	ES
6.0–6.5	$19.0^{+2.3}_{-2.2}$	$33.4^{+10.6}_{-8.8}$
6.5–7.0	$23.7^{+2.0}_{-1.9}$	$10.5^{+9.2}_{-8.1}$
7.0–7.5	$21.2^{+2.0}_{-1.7}$	$33.2^{+9.3}_{-7.8}$
7.5–8.0	$18.9^{+1.7}_{-1.6}$	$28.1^{+8.4}_{-7.7}$
8.0–8.5	$17.2^{+1.6}_{-1.4}$	$12.4^{+6.7}_{-5.5}$
8.5–9.0	$14.2^{+1.4}_{-1.3}$	$16.3^{+7.3}_{-5.3}$
9.0–9.5	$13.3^{+1.4}_{-1.2}$	$17.8^{+6.9}_{-5.1}$
9.5–10.0	$10.0^{+1.3}_{-1.0}$	$9.1^{+5.8}_{-3.9}$
10.0–10.5	$10.3^{+1.2}_{-1.1}$	$0.2^{+4.8}_{-0.4}$
10.5–11.0	$6.6^{+1.0}_{-0.9}$	$1.2^{+3.3}_{-6.3}$
11.0–11.5	$4.1^{+0.7}_{-0.6}$	$3.16^{+3.2}_{-2.5}$
11.5–12.0	$3.3^{+0.6}_{-0.5}$	$2.32^{+3.1}_{-2.4}$
12.0–20.0	$5.3^{+1.1}_{-0.9}$	$9.00^{+4.8}_{-3.1}$

shaper energy resolution has fit out at +1.2% and is narrower only because of the combination of not being allowed to go negative and having a constraint of $^{+1.0}_{-0.0}\%$. The α Po depth have noticeably narrower fit uncertainties, most likely owing to constraints from the spectral information of α backgrounds above about 0.9 MeV in the data.

To better understand the contributions of the systematic uncertainties to the total uncertainties, we took the quadratic difference between the fit uncertainty including systematic parameter variations and the fit uncertainty without such variations. The equivalent neutrino fluxes with their statistical and systematic uncertainties are (in units of $10^6 \text{ cm}^{-2} \text{ s}^{-1}$)

$$\begin{aligned} \phi_{\text{CC}} &= 1.67^{+0.05}_{-0.04}(\text{stat.})^{+0.07}_{-0.08}(\text{syst.}), \\ \phi_{\text{ES}} &= 1.77^{+0.24}_{-0.21}(\text{stat.})^{+0.09}_{-0.10}(\text{syst.}), \\ \phi_{\text{NC}} &= 5.54^{+0.33}_{-0.31}(\text{stat.})^{+0.36}_{-0.34}(\text{syst.}). \end{aligned} \quad (44)$$

Table XVI is a summary of these systematic uncertainties, categorized by different sources.

The MCMC fit results were checked against those from the other two independent methods described earlier in this section. A comparison of the results from these three analysis methods revealed two issues. A 10% difference between the NC flux uncertainties was found, and subsequent investigation revealed incorrect input parameters in two methods. After the inputs were corrected, the errors agreed, and there was no change in the fitted central values. However, the ES flux determined from the MCMC method was 0.5σ lower than those from the other two analyses. This difference was found to be from the use of an inappropriate algorithm to provide a point estimation of the ES posterior distributions. A better method of fitting the posterior distribution to a Gaussian with different widths on either side of the mode was implemented, and the ES flux results agreed with those from the other two analyses.

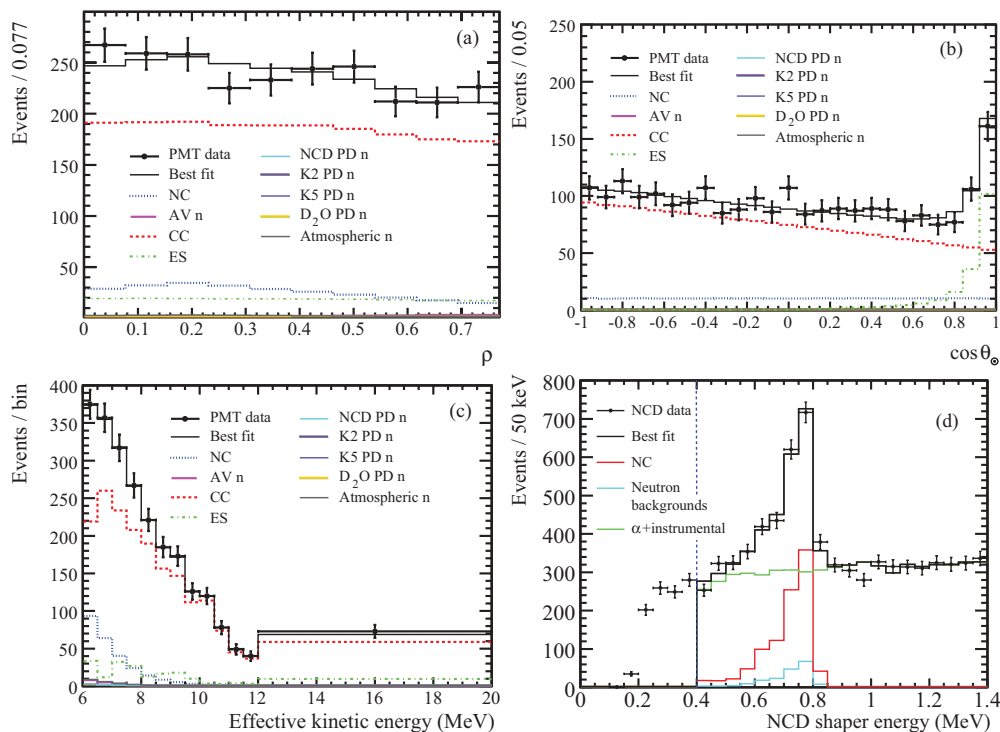


FIG. 22. (Color) One-dimensional projections of PMT array and NCD array data overlaid with best-fit results to signals. Instrumental background contributions from J3 and N4 are combined with the α background in (d). The χ^2 values for the (a) ρ , (b) $\cos \theta_{\odot}$, (c) T_{eff} , and (d) shaper energy distributions are 7.3, 11.9, 0.3, and 17.0, respectively. Because these are one-dimensional projections in a multidimensional space, these χ^2 values are quoted as a qualitative demonstration of goodness-of-fit and cannot be simply evaluated.

The ES flux presented here is 2.2σ lower than that found by Super-Kamiokande-I [62]. This is consistent with a downward statistical fluctuation in the ES signal, as evidenced in the shortfall of signals near $\cos \theta_{\odot} = 1$ in two isolated energy bins. The ${}^8\text{B}$ spectral shape [58] used here differs from that [63] used in previous SNO results. The CC, ES, and NC flux results are in agreement ($p = 32.8\%$ [64]) with the NC flux result of the first phase [2] and with the fluxes from the second phase [6]. Table XVII summarizes the CC, ES, and NC fluxes determined from energy-unconstrained fits in SNO's three phases, and Fig. 23 shows a comparison of these NC measurements.

X. NEUTRINO MIXING MODEL INTERPRETATION OF RESULTS

The SNO measurements of the NC and CC fluxes for neutrinos originated from ${}^8\text{B}$ decays inside the Sun unambiguously proved that neutrinos change their flavor while traveling to Earth. These results can be interpreted, as in previous SNO analyses [3,4,6], as neutrino flavor transitions owing to the mixing of the massive neutrino states via the MSW effect [17,18]. Neutrino mixing parameters can be extracted by comparing the experimental data from SNO and other experiments, with the model predictions from the neutrino mixing hypothesis. The neutrino-mixing analyses presented in the following do not include solar neutrino results from a recent low-threshold analysis that combined the phase-I and phase-II data sets [22].

A two-neutrino mixing analysis was performed because the mixing of the third mass eigenstate into ν_e is small [66] and $\Delta m_{\text{sol}}^2 \ll \Delta m_{\text{atm}}^2$ [67]. The two neutrino mixing parameters in this model are the squared-mass difference of the neutrino mass eigenstates, $\Delta m^2 \equiv \Delta m_{21}^2$, and the mixing angle between the appropriate mass eigenstates, $\theta \equiv \theta_{12}$. The mixing angle is also given as $\tan^2 \theta$ to compare its extracted value with our previous results, as well as with the results reported by others. The mixing model was used to propagate neutrinos inside the Sun, vacuum, and Earth, for each value of Δm^2 and $\tan^2 \theta$. The model predictions for each experiment used in the global solar analysis were computed with the neutrino fluxes from the BS05(OP) solar model [65], which is in good agreement with the helioseismological data, and the latest ${}^8\text{B}$ spectrum shape with its associated uncertainties from Winter *et al.* [58].

The default approach in our analyses was the covariance χ^2 method. The χ^2 function was minimized at each point in the $\tan^2 \theta - \Delta m^2$ plane with respect to the ${}^8\text{B}$ neutrino flux. The least-squares fit and the projection in the $\tan^2 \theta - \Delta m^2$ plane were then performed by allowing any values for the ${}^8\text{B}$ neutrino flux for a given value of $\tan^2 \theta$ and Δm^2 . At the minimum value, χ_{min}^2 , the best-fit values for the mixing parameters $\tan^2 \theta$ and Δm^2 were extracted, together with the corresponding value for the ${}^8\text{B}$ flux. Then the 68%, 95%, and 99.78% confidence level (CL) regions in the two-dimensional parameter space $\tan^2 \theta - \Delta m^2$ were drawn. The uncertainties on the mixing parameters were determined by projecting the χ^2 function passing through the best-fit point on the

TABLE XIV. Systematic parameters' constraints and fit results for the full phase-III data signal extraction. Details of the parametrization of these nuisance parameters in the fit are described in Appendix B. Those constraints marked with an asterisk were handled by the second term in Eq. (37), with the covariance matrices given in the same Appendix.

	Description	Gaussian constraint		Fit value
		Mean	σ	
Systematic Nuisance parameters, PMT array				
$f_{\text{NC}}^{\text{PMT}}$	NC flux to PMT NC events factor	0.467 25	0.006 03	$0.467\ 35 \pm 0.005\ 74$
a_0^x	x coordinate shift	0.0	4.0	1.0 ± 4.0
a_0^y	y coordinate shift	0.0	4.0	-1.0 ± 3.9
a_0^z	z coordinate shift	5.0	4.0	6.1 ± 3.6
a_1^x	Coordinate scale	0.000	0.006	-0.002 ± 0.008
b_0^{xy}	xy resolution constant term	0.065 46	0.028 60*	0.069 ± 0.029
b_1^{xy}	xy resolution linear term	$-0.000\ 055\ 01$	$0.000\ 060\ 51^*$	$-0.000\ 053 \pm 0.000\ 058$
b_2^{xy}	xy resolution quadratic term	3.9×10^{-7}	$0.2 \times 10^{-7}^*$	$0.000\ 000\ 38 \pm 0.000\ 000\ 20$
b_0^z	z resolution constant term	0.07096	0.028 05*	0.072 ± 0.027
b_1^z	z resolution linear term	0.000 115 5	0.000 082 51*	$0.000\ 12 \pm 0.000\ 082$
$b_0^{\theta a}$	PMT angular resolution	0.0	0.056	0.011 ± 0.059
a_1^E	PMT energy scale	1.000	0.0109	1.0047 ± 0.0087
b_0^E	PMT energy resolution (neutrons)	0.0119	0.0114	0.0121 ± 0.0104
Systematic Nuisance parameters, NCD array				
$f_{\text{NC}}^{\text{NCD}}$	NC flux to NCD NC events factor	1.7669	0.0590	1.7713 ± 0.0586
a_1^{NCDE}	NCD shaper energy scale	1.00	0.01	1.0047 ± 0.0035
b_0^{NCDE}	NCD shaper energy resolution	0.00	+ 0.01 -0.00	0.0124 ± 0.0065
α_0	α PDF, α Po depth	0	1	1.21 ± 0.62
α_1	α PDF, α bulk depth	0	1	0.25 ± 0.86
α_2	α PDF, drift time	0	1	-0.11 ± 0.97
α_3	α PDF, avalanche width	0	1	0.27 ± 0.94
α_4	α PDF, avalanche gradient	0	1	-0.06 ± 1.00
α_5	α PDF, Po/bulk fraction	0	1	0.13 ± 0.97
α_6	α PDF, ion mobility	0	1	-0.06 ± 0.96
α_7	α PDF, data reduction cuts	0	1	-0.49 ± 0.93
p_1^{J3}	J3-type background skew Gaussian mean	0.4584	0.0262	0.4724 ± 0.0231
p_1^{N4}	N4-type background skew Gaussian mean	0.0257	0.0138	0.0333 ± 0.0112

^aThe 1σ width of the constraint for b_0^{θ} was input incorrectly in the analysis in Ref. [7]; the correct value should be 0.12. The fit value indicates that this error should not have any impact on the ES results.

TABLE XV. Constraints and fit results for the amplitude of different neutron and instrumental background classes in the full phase III data signal extraction.

Background	Description	Gaussian constraint		Fit value
		Mean	σ	
$N_{\text{ex}}^{\text{NCD}}$	External n (AV, H ₂ O backgrounds)	40.9	20.6	42.2 ± 19.3
$N_{\text{ncdpc}}^{\text{NCD}}$	NCD bulk, cable	35.6	12.2	35.2 ± 12.1
$N_{\text{k2pd}}^{\text{NCD}}$	K2	32.8	5.2	32.7 ± 5.1
$N_{\text{k5pd}}^{\text{NCD}}$	K5	31.6	3.7	31.7 ± 3.7
$N_{\text{d2opd}}^{\text{NCD}}$	D ₂ O photodisintegration	31.0	4.8	30.9 ± 4.8
$N_{\text{atmos}}^{\text{NCD}}$	Atmospheric ν and cosmogenic muons	13.6	2.7	13.6 ± 2.7
$N_{\text{J3}}^{\text{NCD}}$	J3-type instrumental background	Unconstrained	355.6 ± 192.3	
$N_{\text{N4}}^{\text{NCD}}$	N4-type instrumental background	Unconstrained	215.6 ± 170.5	

TABLE XVI. Sources of systematic uncertainties on CC, ES, and NC flux measurements. The total uncertainties differs from the individual uncertainties added in quadrature owing to correlations.

Source	CC uncertainty (%)	ES uncertainty (%)	NC uncertainty (%)
PMT energy scale	± 2.7	± 3.6	± 0.6
PMT energy resolution	± 0.1	± 0.3	± 0.1
PMT radial energy dependence	± 0.9	± 0.9	± 0.0
PMT radial scaling	± 2.7	± 2.7	± 0.1
PMT angular resolution	± 0.2	± 2.2	± 0.0
Background neutrons	± 0.6	± 0.7	± 2.3
Neutron capture	± 0.4	± 0.5	± 3.3
Cherenkov/AV backgrounds	± 0.3	± 0.3	± 0.0
NCD instrumentals	± 0.2	± 0.2	± 1.6
NCD energy scale	± 0.1	± 0.1	± 0.5
NCD energy resolution	± 0.3	± 0.3	± 2.7
NCD α systematics	± 0.3	± 0.4	± 2.7
PMT data reduction cuts	± 0.3	± 0.3	± 0.0
Total experimental uncertainty	± 4.0	± 4.9	± 6.5
Cross section [61]	± 1.2	± 0.5	± 1.1

$\tan^2 \theta$ and Δm^2 axes, separately. The one-dimensional (1D) projections were not a simple slice of the two-dimensional contour passing through χ^2_{\min} , but instead a projection in which $\Delta\chi^2 = \chi^2 - \chi^2_{\min}$ was computed for each 1D axis, allowing the other parameter to take any values. From these 1D projections the uncertainties on each parameter (1σ spreads) were determined from the values at $\chi^2_{\min} + 1$.

The solar neutrino data used in this analysis were SNO phase-I (SNO-I) summed kinetic energy spectra (CC + ES + NC + backgrounds) for day and night [4]; SNO phase-II (SNO-II) CC kinetic energy spectra, ES, and NC fluxes for day and night [6]; SNO phase III (SNO-III) CC, ES, and NC fluxes [7], Super-Kamiokande zenith binned energy spectra [62]; and the rate measurements from the Homestake [10], Gallex/GNO [68], SAGE [69] and Borexino experiments [70]. The global solar χ^2 function obtained by comparing these

data with the corresponding model predictions were then combined with the 2881-ton-year KamLAND reactor antineutrino results [71], assuming CPT invariance. While the results from SNO are highly sensitive to the mixing angle through the measured ratio $\phi_{\text{CC}}/\phi_{\text{NC}}$, the KamLAND measurement has a higher sensitivity to the allowed values of the Δm^2 parameter.

First, the constraint on the neutrino mixing parameters was placed by interpreting the measurements from the results of the three phases of SNO only. Detailed descriptions on the use of data sets from SNO-I and SNO-II to interpret neutrino mixing can be found in Refs. [6,72]. SNO-III data were obtained from signal extraction (Sec. IX) as integrated CC, ES, and NC fluxes (averaged over day and night), which are tabulated in Table XI. The statistical correlation coefficients between the three fluxes, which were needed in building the global χ^2 , are tabulated in Table XII.

TABLE XVII. Energy-unconstrained CC, ES and NC flux results (in units of $10^6 \text{ cm}^{-2} \text{ s}^{-1}$) from three phases of SNO. The T_{eff} thresholds for the PMT array data in phases I, II, and III were 5.0, 5.5, and 6.0 MeV, respectively. These thresholds were chosen to optimize the signal-to-background ratio near the threshold in each phase. For example, the U and Th backgrounds improved in phase III, but the loss of energy resolution in the PMT array (owing to shadowing of the NCD array) resulted in a broader spectrum from the β - γ background. By increasing the analysis threshold to 6 MeV, the effects of such backgrounds to the CC measurement was reduced in phase III. ‘‘Energy-constrained’’ flux results, in which ES and CC events were constrained to an undistorted ${}^8\text{B}$ spectrum in the fit, can be found in Refs. [5, 6] for phases I and II, respectively.

Data set	ϕ_{CC}	ϕ_{ES}	ϕ_{NC}
Phase I (306 live days)	–	–	$6.42^{+1.57+0.55}_{-1.57-0.58}$
Phase II (391 live days)	$1.68^{+0.06+0.08}_{-0.06-0.09}$	$2.35^{+0.22+0.15}_{-0.22-0.15}$	$4.94^{+0.21+0.38}_{-0.21-0.34}$
Phase III (385 live days)	$1.67^{+0.05+0.07}_{-0.04-0.08}$	$1.77^{+0.24+0.09}_{-0.21-0.10}$	$5.54^{+0.33+0.36}_{-0.31-0.34}$

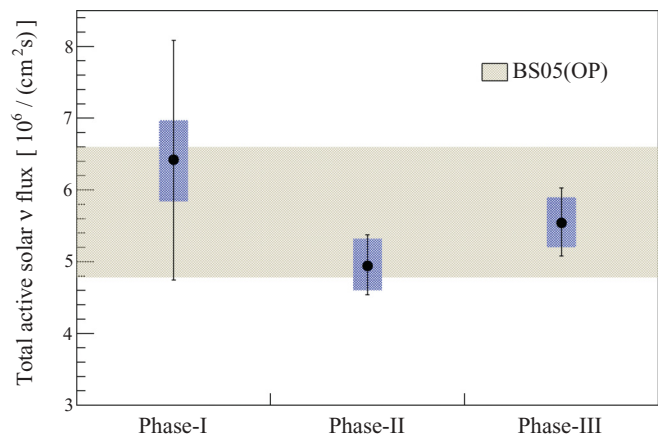


FIG. 23. (Color online) A comparison of the measured energy-unconstrained NC flux results in SNO’s three phases. The horizontal band is the 1σ region of the expected total ${}^8\text{B}$ solar neutrino flux in the BS05(OP) model [65].

For the SNO-III data sample, the χ^2 function is defined as [73]

$$\chi^2 = \sum_{i,j=1}^3 (Y_i^{\text{exp}} - Y_i^{\text{th}})^T [\sigma_{ij}^2(\text{tot.})]^{-1} (Y_j^{\text{exp}} - Y_j^{\text{th}}), \quad (45)$$

where Y_i^{exp} is the CC, ES, or NC averaged flux measurement and Y_i^{th} is the theoretical expectation obtained from the two-neutrino mixing model. The model prediction Y_i^{th} was calculated under the assumption of the two-neutrino oscillation hypothesis; thus, it depended on the number of free parameters n in the fit. In the physics interpretation presented here and in Ref. [7], the free parameters were the neutrino mixing parameters (Δm^2 and $\tan^2 \theta$) and the total flux of the ^8B neutrinos $\phi_{^8\text{B}}$. The shape of the ^8B spectrum was fully constrained by the mixing parameters.

The covariance error matrix $\sigma_{ij}^2(\text{tot.})$ was built as a sum of the squares of the statistical $\sigma_{ij}^2(\text{exp.})$ and systematic $\sigma_{ij}^2(\text{syst.})$ uncertainties:

$$\sigma_{ij}^2(\text{tot.}) = \sigma_{ij}^2(\text{stat.}) + \sigma_{ij}^2(\text{syst.}), \quad (46)$$

where the statistical covariance matrix is given by

$$\sigma_{ij}^2(\text{stat.}) = \rho_{ij} u_i u_j, \quad (47)$$

where u_i and ρ_{ij} are the statistical uncertainty for the measurement Y_i^{exp} and the statistical correlation coefficient between the observables Y_i^{exp} and Y_j^{exp} , respectively.

The effect of each systematic uncertainty S_k on the model expectation for the neutrino yield Y_i^{th} was estimated by computing the change in the expectation ΔY_{ik} with respect to the source of the uncertainty:

$$\gamma_{ik} = \frac{\Delta Y_{ik}}{Y_i}. \quad (48)$$

The relative uncertainties γ_{ik} were then used to construct the systematic error matrix $\sigma_{ij}^2(\text{syst.})$, which is defined as

$$\sigma_{ij}^2(\text{syst.}) = Y_i^{\text{th}} Y_j^{\text{th}} \sum_{k=1}^K r_{ij}^k \gamma_{ik} \gamma_{jk}, \quad (49)$$

where K is the number of systematic uncertainties affecting the observables i and j . A coefficient r_{ij}^k describes a correlation between the observables i and j induced by the systematic uncertainty k , within a single phase of a given experiment. Values of the correlation coefficients r_{ij}^k are summarized in Table XVIII for the SNO-III data sample. In our analyses, the correlations among the systematic uncertainties between the three phases of SNO was also accounted for. However, these correlations had little impact on the allowed regions in the $\tan^2 \theta - \Delta m^2$ plane. The relative errors for the most important energy related systematic uncertainties, such as PMT energy scale and resolution, and also for the ^8B spectrum shape uncertainty, were computed for each value of the mixing parameters as $\gamma_{ik} = \Delta Y_{ik}^{\text{th}} / Y_i^{\text{th}}$.

With the inclusion of SNO phase III results, the following best-fit neutrino mixing parameters were found for the SNO-only analyses: $\Delta m^2 = 4.57_{-1.22}^{+2.30} \times 10^{-5} \text{ eV}^2$ and $\tan^2 \theta = 0.447_{-0.048}^{+0.045}$. The flux of the ^8B neutrinos was floated with respect to the BS05(OP) model prediction of $5.69 \times$

TABLE XVIII. Correlation coefficients for the same sources of systematic uncertainties affecting different types of signals for the SNO-III data set.

Source of systematic	NC-CC	NC-ES	CC-ES
PMT energy scale	+1	+1	+1
PMT energy resolution	+1	+1	+1
PMT radial energy dependence	+1	+1	+1
PMT vertex resolution	+1	+1	+1
PMT vertex accuracy	+1	+1	+1
PMT angular resolution	+1	-1	-1
Background neutrons	+1	+1	+1
Neutron capture	+1	+1	+1
Cherenkov/AV backgrounds	+1	+1	+1
NCD instrumentals	+1	+1	+1
NCD energy scale	+1	+1	+1
NCD energy resolution	+1	+1	+1
NCD α systematics	+1	+1	+1
PMT data reduction cuts	0	0	+1

$10^6 \text{ cm}^{-2} \text{ s}^{-1}$, and the best-fit value was $\phi_{^8\text{B}} = 5.12 \times 10^6 \text{ cm}^{-2} \text{ s}^{-1}$. The flux of *hep* neutrinos was fixed at the BP05(OP) model value of $7.93 \times 10^3 \text{ cm}^{-2} \text{ s}^{-1}$. The minimum χ^2 at the best-fit point was 73.77 for 72 degrees of freedom. The allowed regions at 68%, 95%, and 99.73% confidence level (CL) in the $\Delta m^2 - \tan^2 \theta$ plane from this fit, shown in Fig. 24, were significant improvements compared to the phase II result [6]. The vacuum (“VAC”) oscillation region was ruled out at the 99.73% CL for the first time using SNO data only. The remaining regions in the oscillation plane are significantly smaller than those presented in Ref. [6], with reduced marginalized 1σ uncertainties. The two best-fit results are given in Table XIX, with a comparison of the effects of including the SNO-III data sample in the SNO-only oscillation analysis.

In our first report of phase-III results [7], the following changes in the global solar neutrino analysis were made with respect to our phase-II analysis in Ref. [6]: The model predictions for all solar neutrino experiments were computed using the BS05(OP) model and the ^8B neutrino spectrum shape

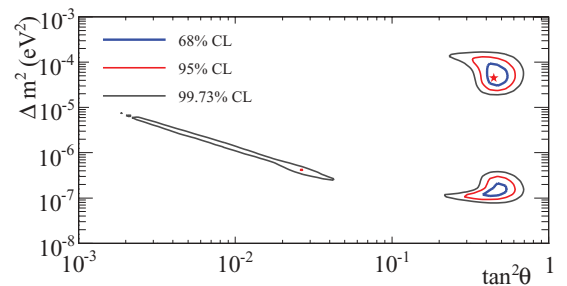


FIG. 24. (Color) After SNO-III. SNO-only neutrino oscillation confidence level contours published in Ref. [7]. This analysis includes the summed kinetic energy spectra from phase I (day and night); NC and ES fluxes, and CC kinetic energy spectra from phase II (day and night); and CC, NC, and ES fluxes from phase III. The best-fit point is at $\Delta m^2 = 4.57 \times 10^{-5} \text{ eV}^2$, $\tan^2 \theta = 0.447$, $\phi_{^8\text{B}} = 5.12 \times 10^6 \text{ cm}^{-2} \text{ s}^{-1}$. The *hep* neutrino flux was fixed at $7.93 \times 10^3 \text{ cm}^{-2} \text{ s}^{-1}$.

TABLE XIX. SNO-only neutrino oscillation best-fit parameters. This table shows the improvement on the neutrino mixing parameters after phase III.

Analysis	Δm^2 (10^{-5} eV 2)	$\tan^2 \theta$	${}^8\text{B}$ flux (10^6 cm $^{-2}$ s $^{-1}$)
Before SNO-III	$5.0^{+6.2}_{-1.8}$	$0.45^{+0.11}_{-0.10}$	5.11
After SNO-III	$4.57^{+2.30}_{-1.22}$	$0.45^{+0.05}_{-0.05}$	5.12

from Ref. [58], the inclusion of the 192 live-day results from Borexino [70], update of data from SK-1 using results from Ref. [62], and, most importantly, the new measurements from the third phase of SNO were incorporated. The global fit to these solar neutrino data led to the following neutrino mixing parameters: $\Delta m^2 = 4.90^{+1.64}_{-0.93} \times 10^{-5}$ eV 2 and $\theta = 33.5^{+1.3}_{-1.3}$, with ${}^8\text{B}$ flux of $\phi_{8\text{B}} = 5.21 \times 10^6$ cm $^{-2}$ s $^{-1}$. The minimum χ^2 at the best-fit point was 130.29 for 120 degrees of freedom. The allowed regions from this analysis are shown on the left panel in Fig. 25. The constraint on both neutrino mixing parameters was much better than our results in phase II [6]. When the 2881-ton-year KamLAND results were included in this analysis [71], the best-fit parameters became $\Delta m^2 = 7.59^{+0.21}_{-0.19} \times 10^{-5}$ eV 2 and $\theta = 34.4^{+1.3}_{-1.2}$ degrees, and the ${}^8\text{B}$ flux of $\phi_{8\text{B}} = 4.92 \times 10^6$ cm $^{-2}$ s $^{-1}$. The improvement in comparison with the former analysis was observed in the allowed regions from the combined fit shown in the right panel in Fig. 25. A summary of these results is given in Table XX.

In comparison to phase-II results, this combined fit of the solar neutrino data and the 2881-ton-year results from KamLAND improved the constraints on the neutrino mixing parameters: mixing angle θ and Δm^2 by 45% and 60%, respectively, at the time of the publication of Ref. [7]. This improvement on the mixing angle was dominated by the SNO experiment and the phase-III results. The fitted values for the

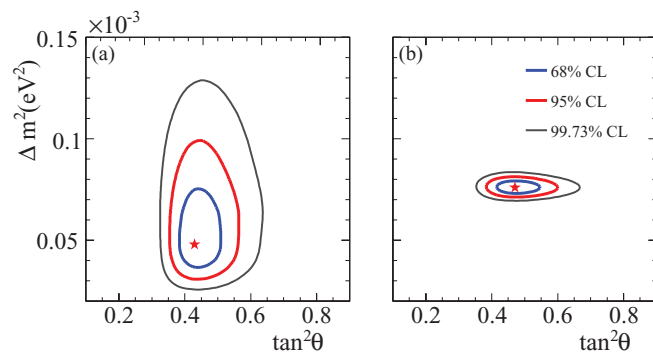


FIG. 25. (Color) After SNO-III. Neutrino oscillation parameters confidence level contours. (a) Global solar analysis including the rate measurements from Homestake, Gallex/GNO, SAGE, and Borexino; SK-I zenith-energy spectra from [62], summed kinetic energy spectra from SNO-I (day and night); NC and ES fluxes and CC kinetic energy spectra from SNO-II (day and night); and CC, ES, and NC fluxes from SNO-III. The best-fit point was at $\Delta m^2 = 4.90 \times 10^{-5}$ eV 2 , $\tan^2 \theta = 0.437$, $\phi_{8\text{B}} = 5.21 \times 10^6$ cm $^{-2}$ s $^{-1}$. The *hep* neutrino flux was fixed at 7.93×10^3 cm $^{-2}$ s $^{-1}$. (b) Including KamLAND data from [71], the best-fit point was at $\Delta m^2 = 7.59 \times 10^{-5}$ eV 2 , $\tan^2 \theta = 0.468$, $\phi_{8\text{B}} = 4.92 \times 10^6$ cm $^{-2}$ s $^{-1}$.

TABLE XX. Global solar only and global solar + KamLAND best-fit parameters. The global solar results before the SNO-III phase do not include data from Borexino. The global solar + KamLAND results after SNO-III include the latest data from Borexino [70] and KamLAND [71].

Analysis	Δm^2 (10^{-5} eV 2)	$\tan^2 \theta$	${}^8\text{B}$ flux (10^6 cm $^{-2}$ s $^{-1}$)
Before SNO-III phase			
Global solar	$6.5^{+4.4}_{-2.3}$	$0.45^{+0.09}_{-0.08}$	5.06
With KamLAND	$8.0^{+0.6}_{-0.4}$	$0.45^{+0.09}_{-0.07}$	4.93
After SNO-III phase			
Global solar	$4.90^{+1.64}_{-0.93}$	$0.44^{+0.05}_{-0.04}$	5.21
With KamLAND	$7.59^{+0.21}_{-0.19}$	$0.47^{+0.05}_{-0.04}$	4.92

${}^8\text{B}$ neutrino flux were in agreement with recent predictions from solar models.

XI. SUMMARY

We have presented a detailed description of the SNO phase III results that were published in Ref. [7]. Neutrons from the NC reaction were detected predominantly by the NCD array. The use of this technique, which was independent of the neutron detection methods in previous phases, resulted in reduced correlations between the CC, ES, and NC fluxes and an improvement in the mixing angle uncertainty.

Several techniques to reliably calibrate the PMT and NCD arrays were developed and are detailed in this paper. The presence of the NCD array changed the optical properties, and hence the energy response, of the PMT array. Extensive studies and evaluation of the techniques used in calibrating the PMT array in previous phases were performed and reported. Radioactive backgrounds associated with the NCD array, the D $_2$ O target and other detector components were precisely quantified by *in situ* and *ex situ* measurements. These measurements provided the constraints for the respective nuisance parameters in the determination of the ν_e and the total active neutrino fluxes.

The total flux of active neutrinos was measured to be $5.54^{+0.33}_{-0.31}$ (stat.) $^{+0.36}_{-0.34}$ (syst.) $\times 10^6$ cm $^{-2}$ s $^{-1}$ and was consistent with previous measurements and standard solar models. A global analysis of neutrino mixing parameters using solar and reactor neutrino results yielded the best-fit values of the neutrino mixing parameters of $\Delta m^2 = 7.59^{+0.19}_{-0.21} \times 10^{-5}$ eV 2 and $\theta = 34.4^{+1.3}_{-1.2}$ degrees.

A detailed paper that describes an analysis of data combined from all three phases of SNO is in preparation.

ACKNOWLEDGMENTS

This research was supported by the following: (Canada) Natural Sciences and Engineering Research Council, Industry Canada, National Research Council, Northern Ontario Heritage Fund, Atomic Energy of Canada, Ltd., Ontario Power Generation, High Performance Computing Virtual Laboratory, Canada Foundation for Innovation, Canada Research Chairs

program; (United States) Department of Energy, National Energy Research Scientific Computing Center, Alfred P. Sloan Foundation; (United Kingdom) Science and Technology Facilities Council (formerly Particle Physics and Astronomy Research Council); (Portugal) Fundação para a Ciência e a Tecnologia. We thank the SNO technical staff for their strong contributions and David Sinclair for careful review of the neutron efficiency analysis. We thank INCO (now Vale, Ltd.) for hosting this project in their Creighton mine.

APPENDIX A: INSTRUMENTAL BACKGROUND CUTS FOR NCD ARRAY DATA

Two independent sets of cuts were developed to remove instrumental backgrounds in the NCD array data. These cuts exploited the differences in the characteristics between ionization and nonionization events. One of these two sets inspected the characteristics of the digitized wave forms in the time domain, while the other did so in the frequency domain. Cuts in both sets were used in the selection of the candidate event data set in the solar neutrino analysis. A summary of these cuts is provided in the following.

1. Bursts and overflow cuts

Two burst cuts were developed to remove events that occurred within a very short time window. These events were believed to arise from electronic noise pickups. If there were four or more shaper events within 100 ms, all events

within this sequence were removed in the shaper burst cut. Similarly, if there were four or more MUX-scope events within 100 ms, such events were removed in the MUX burst cut. In the shaper-overflow cut, shaper events that arrived within a short time (15 μ s to 5 ms) after a previous event had saturated the shaper were removed.

2. Time correlation of shaper-ADC and MUX-scope events

For each real ionization in the proportional counter, there was a time-correlated pair of shaper and MUX-scope events. Instrumental background events exhibited a shorter time difference between the two events. A time-correlation cut was developed to remove shaper-MUX-scope event pairs from the same string that showed such an anomalous timing characteristic.

3. Fork event cuts

The end of each counter string was attached to an open delay line. Pulses were reflected at this open termination; thus, some physics pulses exhibited a double-peak structure. There were also instrumental background pulses that exhibited similar double-peak characteristics, but with much different pulse width, time separation, and amplitude ratio between the peaks. The time-domain fork cut removed these “fork” events by exploiting these differences. Some fork events also featured a third reflection at the tail of the wave form. These events were removed by a cut that was specifically designed to search for such reflection in the time domain. The frequency-domain

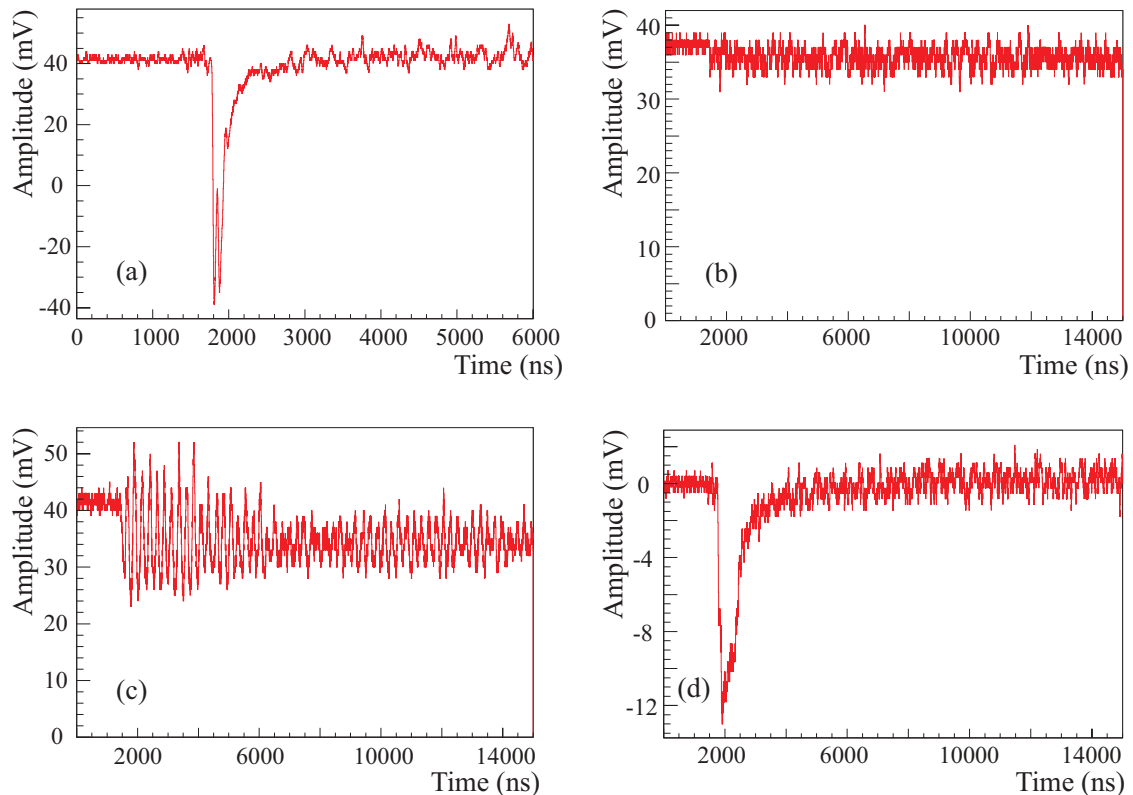


FIG. 26. (Color online) Examples of neutron signal and instrumental background pulses in the NCD array data. (a) A fork event, in which a small third reflection is seen in the tail of the pulse; (b) a flat trace. (c) an oscillatory pulse. (d) a neutron event.

fork cut removed events with a peak around 12 MHz, which was the characteristic frequency of the fork events, in the power spectrum. A fork event example is shown in the Fig. 26. A fork event does not exhibit the characteristic ion tail that is evidenced in the example neutron pulse that is also shown in the figure.

4. Flat, oscillatory, and narrow pulse event cuts

The flat-trace cut removed events that did not have a well-defined pulse profile. These pulses were mostly noise that crossed the trigger threshold. The dominant type of noise events during normal data taking was oscillatory pulses. These events were identified and removed by a cut on the number of times the wave form crossed the baseline. In the narrow pulse cut, pulses with widths that were too narrow to be ionization events were removed. These pulses were mostly discharges and some of them carried a large amount of charge. Spike events were also removed by identifying those with an abnormal ratio of their area and maximum amplitude. In the frequency domain, wave forms with unusual symmetry had a noncharacteristic zero-frequency intercept of the phase in their Fourier transform. Wave forms with little power at low frequency were mostly flat or oscillation events. Wave forms with a large peak in the power spectrum at a frequency above 8 MHz (3.7 MHz) were mostly fork (oscillation) events. Three cuts were implemented to reject events with such anomalous characteristics in the power spectrum. Examples of flat and oscillatory trace events are shown in Fig. 26.

APPENDIX B: PARAMETERIZATION OF SYSTEMATIC UNCERTAINTIES FOR THE PMT AND NCD ARRAY DATA

The handling of systematic uncertainties in the negative log-likelihood (NLL) for the PMT and NCD array data are

$$V_{b^{xy}} = \begin{pmatrix} 0.000\,818\,124 & -2.249\,84 \times 10^{-7} & -4.191\,31 \times 10^{-9} \\ -2.249\,84 \times 10^{-7} & 3.660\,98 \times 10^{-9} & 3.714\,23 \times 10^{-12} \\ -4.191\,31 \times 10^{-9} & 3.714\,23 \times 10^{-12} & 3.921\,18 \times 10^{-14} \end{pmatrix}, \quad (\text{B7})$$

and

$$V_{b^z} = \begin{pmatrix} 0.000\,786\,96 & 3.471\,88 \times 10^{-7} \\ 3.471\,88 \times 10^{-7} & 6.807\,61 \times 10^{-9} \end{pmatrix}. \quad (\text{B8})$$

2. Parameterization of systematic for the NCD array data

There were ten systematic uncertainties associated with the shaper energy PDFs. For the uncertainties associated with the shaper-ADC energy E_{NCD} and resolution, the energy PDFs were rebuilt using Eq. (38).

For the α -background-related PDFs, the simulated α background PDF was rebuilt by multiplying the unmodified PDF (Q^α) by a reweighting factor (α_i) and a multiplicative function in shaper energy $s_i(E_{\text{NCD}})$:

$$Q_{\text{MC}}^\alpha(E_{\text{NCD}}) = Q^\alpha(E_{\text{NCD}}) \left(1 + \sum_{i=0}^7 \alpha_i s_i(E_{\text{NCD}}) \right). \quad (\text{B9})$$

The reweighting functions, s_i , are

$$\text{Po } \alpha \text{ depth variation } s_0 = -2.06 + 6.58 E_{\text{NCD}} - 6.56 E_{\text{NCD}}^2 + 2.11 E_{\text{NCD}}^3 \quad (\text{B10})$$

described in Secs. IX C1 and IX C2, respectively. In this appendix, their parametrizations in the NLL are discussed.

1. Parameterization of systematic for the PMT array data

As discussed in Sec. IX C1, systematic effects were assessed by remapping observables in the simulated PMT array data. These remapped variables were the effective kinetic energy, reconstructed vertex position, and reconstructed direction of the event. Extensive calibrations were performed to constrain the parameters used in the parametrization of the remapping functions.

For the simulated values of the event kinetic energy (T_g), volume-weighted radius ρ_g , and angle relative to the vector from the Sun θ_g , the remapping functions are

$$T_{\text{remap}} = a_1^E T_0 + b_0^E (T_0 - T_g), \quad (\text{B1})$$

$$x_{\text{remap}} = a_0^x + (1 + a_1^x)x + (b_0^{xy} + b_1^{xy}z + b_2^{xy}z^2)(x - x_g), \quad (\text{B2})$$

$$y_{\text{remap}} = a_0^y + (1 + a_1^y)y + (b_0^{xy} + b_1^{xy}z + b_2^{xy}z^2)(y - y_g), \quad (\text{B3})$$

$$z_{\text{remap}} = a_0^z + (1 + a_1^z)z + (b_0^z + b_1^z z)(z - z_g), \quad (\text{B4})$$

$$\rho_{\text{remap}} = \left(\sqrt{x_{\text{remap}}^2 + y_{\text{remap}}^2 + z_{\text{remap}}^2} / 600 \text{ cm} \right)^3, \quad (\text{B5})$$

and

$$\cos \theta_{\text{remap}} = 1 + (1 + b_0^\theta)(\cos \theta - 1). \quad (\text{B6})$$

where T_0 , (x, y, z) , and θ are the nominal fitted effective kinetic energy, position, and angle of the simulated event. The remapping of $\cos \theta$ was applied to the ES channel only. In these expressions, the nuisance parameters were modeled as Gaussian distributions with means and widths, which are given in Table XIV. The nuisance parameters b_i^{xy} ($i = 0, 1, 2$) and b_i^z ($i = 0, 1$) are correlated, and their covariance matrices are

$$\text{bulk } \alpha \text{ depth variation } s_1 = -0.0684 + 0.0892 E_{\text{NCD}} \quad (\text{B11})$$

$$\text{drift time variation } s_2 = -0.131 + 0.252 E_{\text{NCD}} - 0.117 E_{\text{NCD}}^2 \quad (\text{B12})$$

$$\text{avalanche width offset variation } s_3 = -0.0541 + 0.0536 E_{\text{NCD}} \quad (\text{B13})$$

$$\text{avalanche gradient offset variation } s_4 = -0.0138 \quad (\text{B14})$$

$$\text{ion mobility variation } s_5 = -0.00930 \quad (\text{B15})$$

$$\text{Po/bulk fraction variation } s_6 = -0.00405 + 0.0386 E_{\text{NCD}} \quad (\text{B16})$$

$$\text{data reduction cut systematic } s_7 = 0.861 - 2.77 E_{\text{NCD}} + 2.72 E_{\text{NCD}}^2 - 0.870 E_{\text{NCD}}^3. \quad (\text{B17})$$

In addition, two instrumental background PDFs, based on events on strings J3 and N4, were parametrized as

$$Q_{\text{MC}}^{\text{J3}} = \exp \left[-\frac{1}{2} \left(\frac{E_{\text{NCD}} - p_1^{\text{J3}}}{0.34 p_1^{\text{J3}}} \right)^2 \right] [1 + \text{erf}(-2.0 (E_{\text{NCD}} - p_1^{\text{J3}}))] \quad (\text{B18})$$

and

$$Q_{\text{MC}}^{\text{N4}} = \exp \left[-\frac{1}{2} \left(\frac{E_{\text{NCD}} - p_1^{\text{N4}}}{19.6 p_1^{\text{N4}}} \right)^2 \right] [1 + \text{erf}(-1.59 (E_{\text{NCD}} - p_1^{\text{N4}}))]. \quad (\text{B19})$$

In both cases the number of instrumental background events was allowed to float freely in the fit.

The numerical values of the constraints α_i and the instrumental background parameters (p_1^{J3} and p_1^{N4}) are tabulated in Table XIV.

-
- [1] J. Boger *et al.* (SNO Collaboration), *Nucl. Instrum. Methods A* **449**, 172 (2000).
[2] Q. R. Ahmad *et al.* (SNO Collaboration), *Phys. Rev. Lett.* **87**, 071301 (2001).
[3] Q. R. Ahmad *et al.* (SNO Collaboration), *Phys. Rev. Lett.* **89**, 011301 (2002).
[4] Q. R. Ahmad *et al.* (SNO Collaboration), *Phys. Rev. Lett.* **89**, 011302 (2002).
[5] B. Aharmim *et al.* (SNO Collaboration), *Phys. Rev. C* **75**, 045502 (2007).
[6] B. Aharmim *et al.* (SNO Collaboration), *Phys. Rev. C* **72**, 055502 (2005).
[7] B. Aharmim *et al.* (SNO Collaboration), *Phys. Rev. Lett.* **101**, 111301 (2008).
[8] J. N. Bahcall, M. H. Pinsonneault, and S. Basu, *Astrophys. J.* **555**, 990 (2001).
[9] A. S. Brun, S. Turck-Chièze, and J. P. Zahn, *Astrophys. J.* **525**, 1032 (1999).
[10] B. T. Cleveland *et al.*, *Astrophys. J.* **496**, 505 (1998).
[11] J. N. Abdurashitov *et al.*, *Phys. Rev. C* **60**, 055801 (1999).
[12] W. Hampel *et al.*, *Phys. Lett. B* **447**, 127 (1999).
[13] S. Fukuda *et al.*, *Phys. Rev. Lett.* **86**, 5651 (2001).
[14] M. Altmann *et al.*, *Phys. Lett. B* **490**, 16 (2000).
[15] K. Eguchi *et al.*, *Phys. Rev. Lett.* **90**, 021802 (2003).
[16] B. Pontecorvo, *Zh. Eksp. Teor. Fiz.* **33**, 549 (1957) [*Sov. Phys. JETP* **6**, 429 (1957)].
[17] L. Wolfenstein, *Phys. Rev. D* **17**, 2369 (1978).
[18] S. P. Mikheyev and A. Smirnov, *Nuovo Cimento C* **9**, 17 (1986).
[19] J. Amsbaugh *et al.*, *Nucl. Inst. Meth A* **579**, 1054 (2007).
[20] G. Doucas *et al.*, *Nucl. Instrum. Methods A* **370**, 579 (1996).
[21] H. H. Chen, *Phys. Rev. Lett.* **55**, 1534 (1985).
[22] B. Aharmim *et al.* (SNO Collaboration), *Phys. Rev. C* **81**, 055504 (2010).
[23] B. A. Moffat *et al.*, *Nucl. Instrum. Methods A* **554**, 255 (2005).
[24] O. Simard, Ph.D. thesis, Carleton University, 2009.
[25] M. R. Dragowsky *et al.*, *Nucl. Instrum. Methods A* **481**, 284 (2002).
[26] N. Tagg *et al.*, *Nucl. Instrum. Methods A* **489**, 178 (2002).
[27] \mathcal{F}_T is derived for events in which only a single electron, of energy T_e , produces Cherenkov light. The energy of an event that follows from \mathcal{F}_T is therefore referred to as T_{eff} because it is the effective energy of the event assuming that it resulted from a single fast electron.
[28] R. J. Ford, Ph.D. thesis, Queen's University, 1998.
[29] W. R. Nelson, H. Hirayama, and D. W. O. Rogers, Stanford Linear Accelerator Center Report SLAC-265 (1985).
[30] G. A. Cox-Mobrand, Ph.D. thesis, University of Washington, 2008.
[31] K. Boudjemline *et al.*, *Nucl. Instrum. Methods A* **620**, 171 (2010).
[32] J. Loach, Ph.D. thesis, University of Oxford, 2008.
[33] T. C. Andersen *et al.*, *Nucl. Instrum. Methods* **501**, 399 (2003).
[34] T. C. Andersen *et al.*, *Nucl. Instrum. Methods A* **501**, 386 (2003).
[35] B. Aharmim *et al.*, *Nucl. Instrum. Methods A* **604**, 531 (2009).
[36] I. Blevins *et al.*, *Nucl. Instrum. Methods A* **517**, 139 (2004).
[37] H. M. O'Keefe, Ph.D. thesis, University of Oxford, 2008.
[38] L. C. Stonehill, Ph.D. thesis, University of Washington, 2005.
[39] H. M. O'Keefe *et al.*, *Nucl. Instr. Meth. A* **659**, 182 (2011).
[40] D. Casper, *Nucl. Phys. B, Proc. Suppl.* **112**, 161 (2002).
[41] B. Aharmim *et al.* (SNO Collaboration), *Phys. Rev. D* **80**, 012001 (2009).
[42] C. Kyba, Ph.D. thesis, University of Pennsylvania, 2007.
[43] B. Beltran *et al.*, *New J. Phys.* **13**, 073006 (2011).
[44] H. Wan Chan Tseung, Ph.D. thesis, University of Oxford, 2008.
[45] N. Oblath, Ph.D. thesis, University of Washington, 2009.

- [46] G. F. Wilkinson, *Ionization Chambers and Counters* (Cambridge University Press, Cambridge, 1950).
- [47] J. F. Ziegler, J. P. Biersack, and U. Littmark, *The Stopping and Range of Ions in Solids* (Pergamon Press, Oxford, 1985).
- [48] J. F. Ziegler, *TRansport of Ions in Matter* (TRIM), <http://www.srim.org>.
- [49] G. F. Knoll, *Radiation Detection and Measurement*, 3rd ed. (Wiley & Sons, New York, 2000).
- [50] International Commission on Radiation Units & Measurement Report ICRU-31, 1979.
- [51] R. Veenhof, *Garfield-Simulation of Gaseous Detectors*, <http://garfield.web.cern.ch/garfield>.
- [52] M. K. Kopp *et al.*, *Nucl. Instrum. Methods A* **201**, 395 (1982).
- [53] H. W. Ellis *et al.*, *At. Data Nucl. Data Tables* **17**, 177 (1976).
- [54] T. Quarles *et al.*, *SPICE*, <http://bwrc.eecs.berkeley.edu/Classes/icbook/SPICE/>.
- [55] F. James, CERN Program Library Long Writeup, 506, 1998.
- [56] N. Metropolis *et al.*, *J. Chem. Phys.* **21**, 1087 (1953).
- [57] W. K. Hastings, *Biometrika* **57**, 97 (1970).
- [58] W. T. Winter, S. J. Freedman, K. E. Rehm, and J. P. Schiffer, *Phys. Rev. C* **73**, 025503 (2006).
- [59] *hep* neutrinos could also be present in the measured fluxes; the SSM contribution would be 0.5%.
- [60] Electron-neutrino cross sections were used to calculate all fluxes. The ES cross section was from J. N. Bahcall, M. Kamionkowski, and A. Sirlin, *Phys. Rev. D* **51**, 6146 (1995).
- [61] Cross section uncertainty includes: g_A (0.5%), theoretical cross section [1% for CC and NC, 0.5% for CC/NC; S. Nakamura *et al.*, *Nucl. Phys. A* **707**, 561 (2002)], and radiative corrections (0.3% for CC, 0.1% for NC; A. Kurylov, M. J. Ramsey-Musolf, and P. Vogel, *Phys. Rev. C* **65**, 055501 (2002)).
- [62] J. Hosaka *et al.* (Super-Kamiokande Collaboration), *Phys. Rev. D* **73**, 112001 (2006).
- [63] C. E. Ortiz, A. Garcia, R. A. Waltz, M. Bhattacharya, and A. K. Komives, *Phys. Rev. Lett.* **85**, 2909 (2000).
- [64] A. Valassi *et al.*, *Nucl. Instrum. Methods A* **500**, 391 (2003).
- [65] J. N. Bahcall, A. M. Serenelli, and S. Basu, *Astrophys. J.* **621**, L85 (2005).
- [66] M. Apollonio *et al.* (CHOOZ Collaboration), *Eur. Phys. J. C* **27**, 331 (2003).
- [67] P. Adamson *et al.* (MINOS Collaboration), *Phys. Rev. Lett.* **101**, 131802 (2008).
- [68] M. Altmann *et al.* (GNO Collaboration), *Phys. Lett. B* **616**, 174 (2005).
- [69] J. N. Abdurashitov *et al.* (SAGE Collaboration), *Neutrino 2004 Conference* (2005).
- [70] C. Arpesella *et al.* (Borexino Collaboration), *Phys. Rev. Lett.* **101**, 091302 (2008).
- [71] S. Abe *et al.* (KamLAND Collaboration), *Phys. Rev. Lett.* **100**, 221803 (2008).
- [72] The SNO Collaboration, <http://www.sno.phy.queensu.ca/sno/prlwebpage/>.
- [73] G. Tešić, Ph.D. thesis, Carleton University, 2008.