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2.626 Fundamentals of Photovoltaics
Fall 2008

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Quiz #1 Review: Fundamentals

Lecture 8 – 2.626

Tonio Buonassisi

General Matters

- Practice Exam:
 - Take 10 minutes to think about how you would solve the problems. Write notes.
 - *Then, we'll solve it together during today's review.*

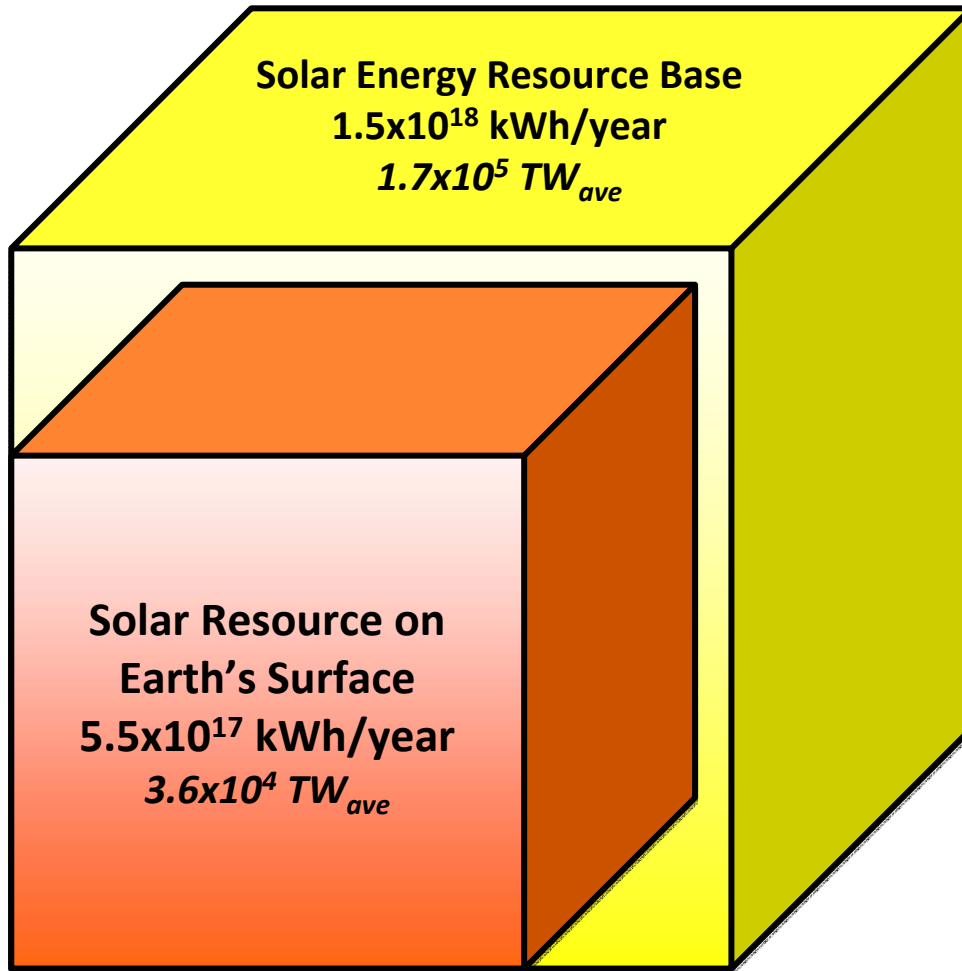
Fundamentals

1. Properties of Sunlight
2. Light Absorption (Charge Excitation and Conduction)
3. Charge Separation
4. Charge Collection

Fundamentals

1. Properties of Sunlight
 - a. Resource Base
 - b. Solar Spectrum
 - c. Atmospheric Absorption
2. Charge Excitation and Conduction
3. Charge Separation
4. Charge Collection

Solar Resource Base = Huge



**Wind Energy
Resource Base**
 6×10^{14} kWh/year
 72 TW_{ave}



**Human Energy Use
(mid- to late-century)**
 4×10^{14} kWh/year
 50 TW_{ave}

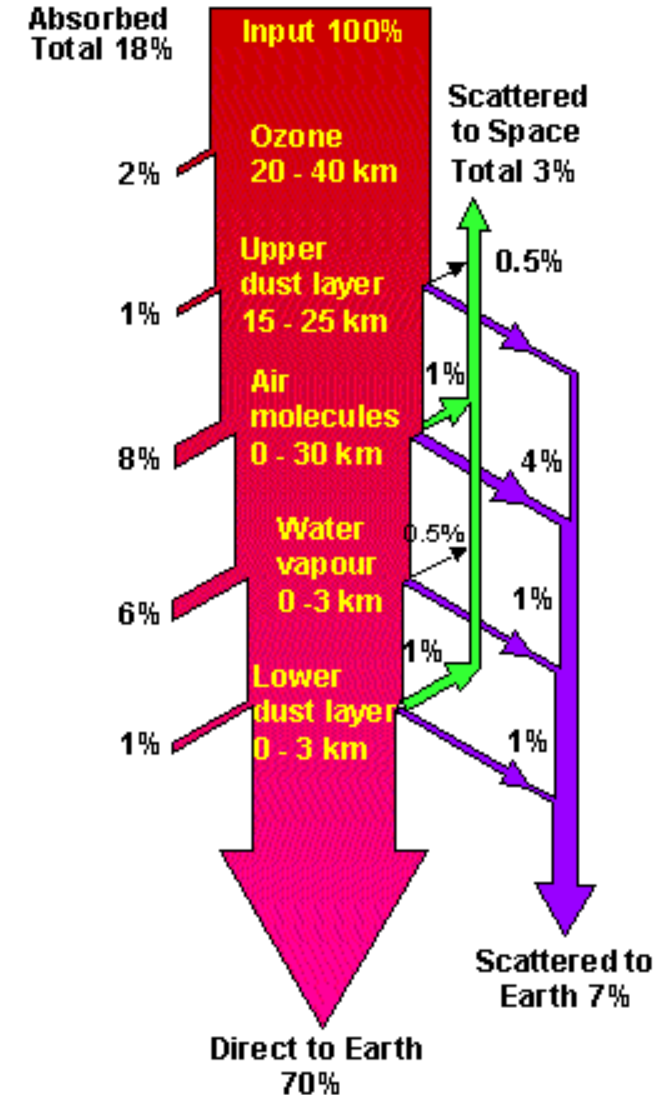
References:

Wind Energy: C.L. Archer and M.Z. Jacobson, *J. Geophys. Res.* **110**, D12110 (2005).

ATMOSPHERIC EFFECTS

Atmospheric effects have several impacts on the solar radiation at the Earth's surface. The major effects for photovoltaic applications are:

- A reduction in the power of the solar radiation due to absorption, scattering and reflection in the atmosphere;
- A change in the spectral content of the solar radiation due to greater absorption or scattering of some wavelengths;
- The introduction of a diffuse or indirect component into the solar radiation; and
- Local variations in the atmosphere (such as water vapor, clouds and pollution) which have additional effects on the incident power, spectrum and directionality.



Courtesy Christiana Honsberg and Stuart Bowden. Used with permission.

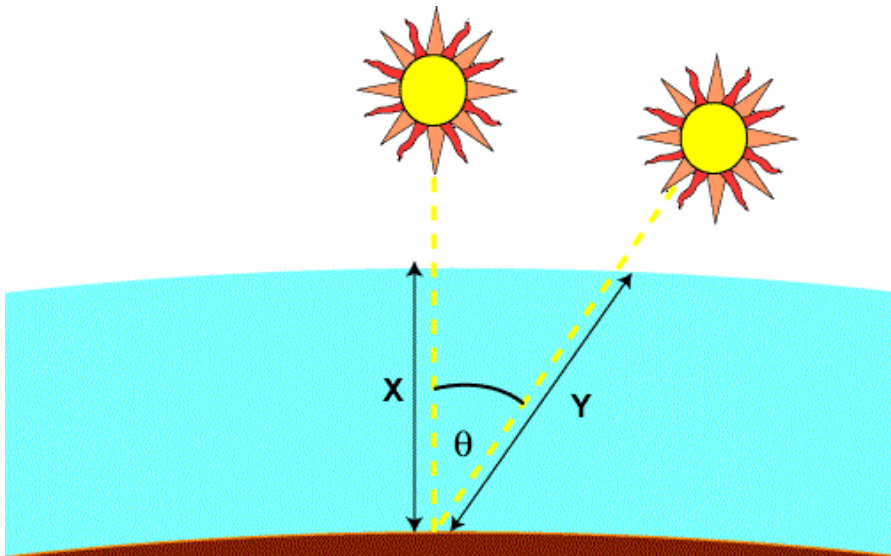
Typical clear sky absorption and scattering of incident sunlight (after Hu and White, 1983).

AIR MASS

The Air Mass is the path length which light takes through the atmosphere normalized to the shortest possible path length (that is, when the sun is directly overhead). The Air Mass quantifies the reduction in the power of light as it passes through the atmosphere and is absorbed by air and dust. The Air Mass is defined as:

$$AM = \frac{1}{\cos(\theta)}$$

Valid for small to medium θ



AM1: Sun directly overhead

AM1.5G: “Conventional”

G (Global): Scattered and direct sunlight

D (Direct): Direct sunlight only

AM0: Just above atmosphere (space applications)

Estimating System Output from Insolation Maps

Q: Let's say I have a 2.2 kW_p photovoltaic array. How much energy will it produce in a year?

A: Let's say our location receives, on average, 4 kWh/m²/day from the Sun. The calculation is then straightforward:



$$\text{Energy Output} = \frac{(2200 \text{ W}_p) \times (4.0 \text{ kWh/m}^2/\text{day})}{1000 \text{ W}_p/\text{m}^2} = 8.8 \text{ kWh/day} \approx 3200 \text{ kWh/year}$$

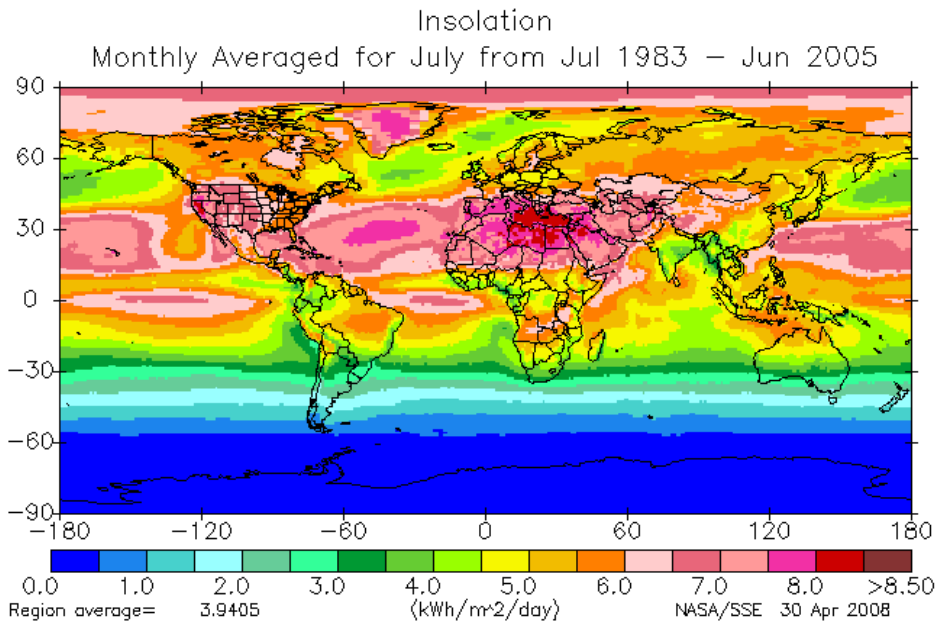


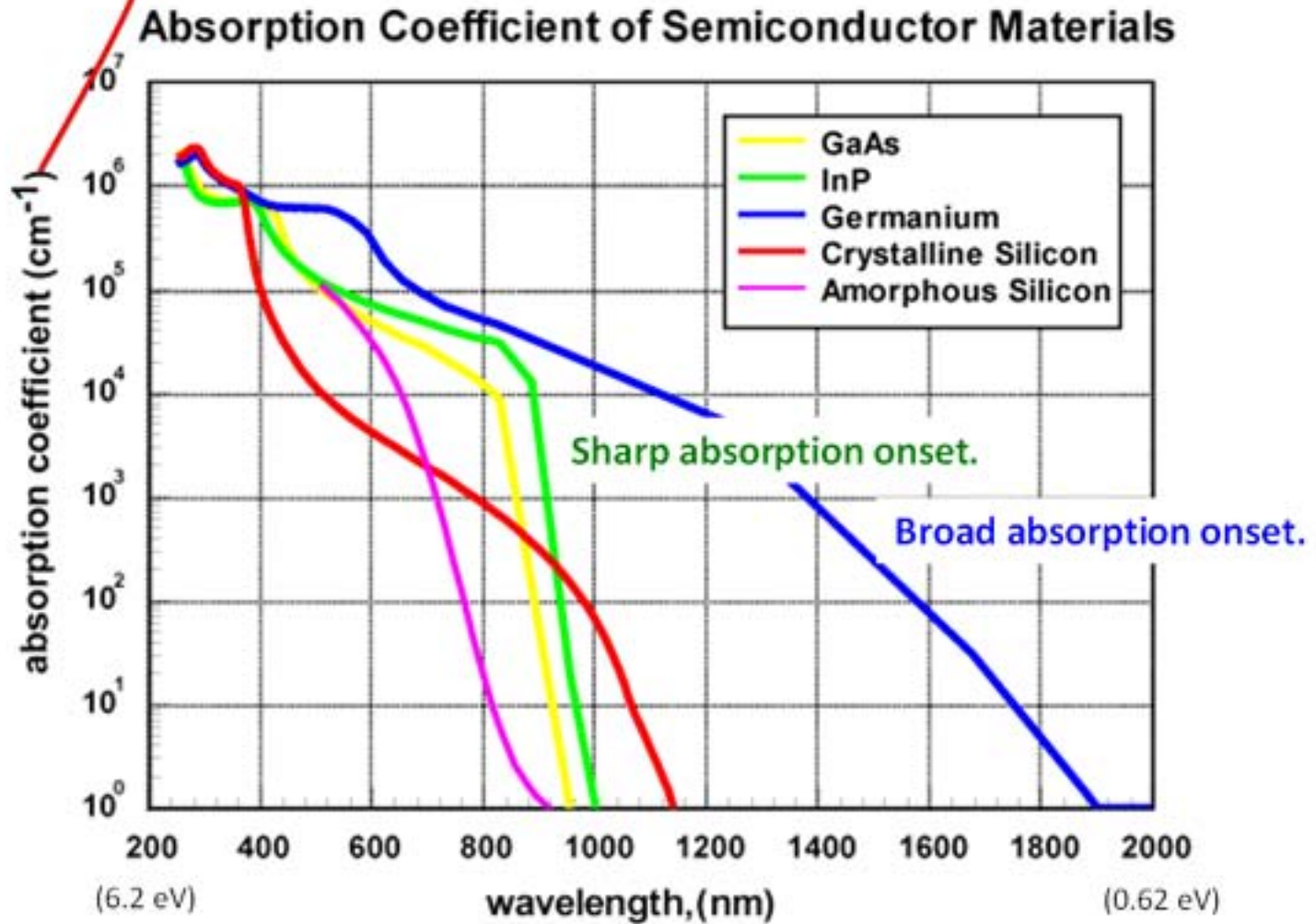
Image courtesy NASA Earth Observatory.

Fundamentals

1. Properties of Sunlight
2. Charge Excitation and Conduction
 - a. Optical Absorption
 - b. Bandgap
3. Charge Separation
4. Charge Collection

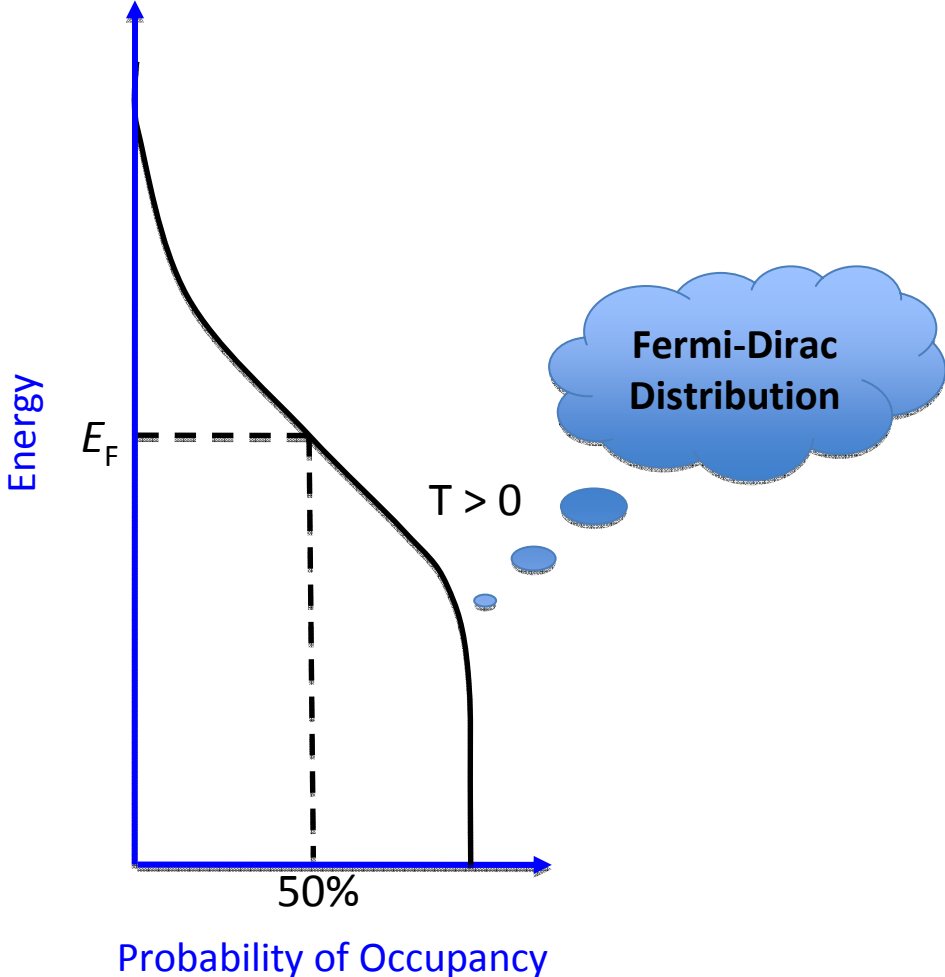
Absorption Coefficient (α) for different materials

$$I = I_0 \cdot e^{-\alpha \cdot l}$$

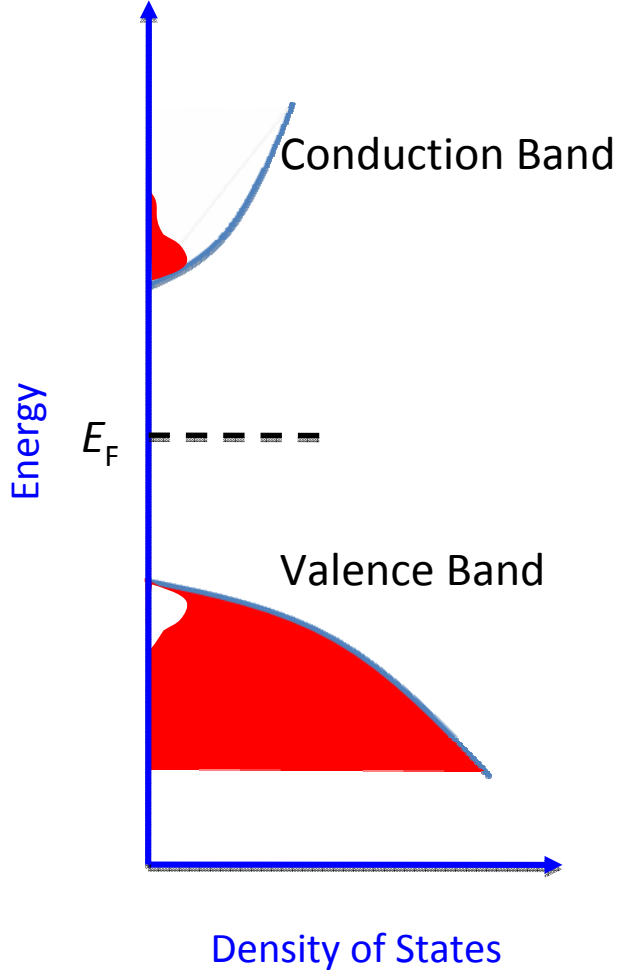


Defining the "Fermi Level"

Fermi-Dirac Probability Distribution Function at $T > 0$.

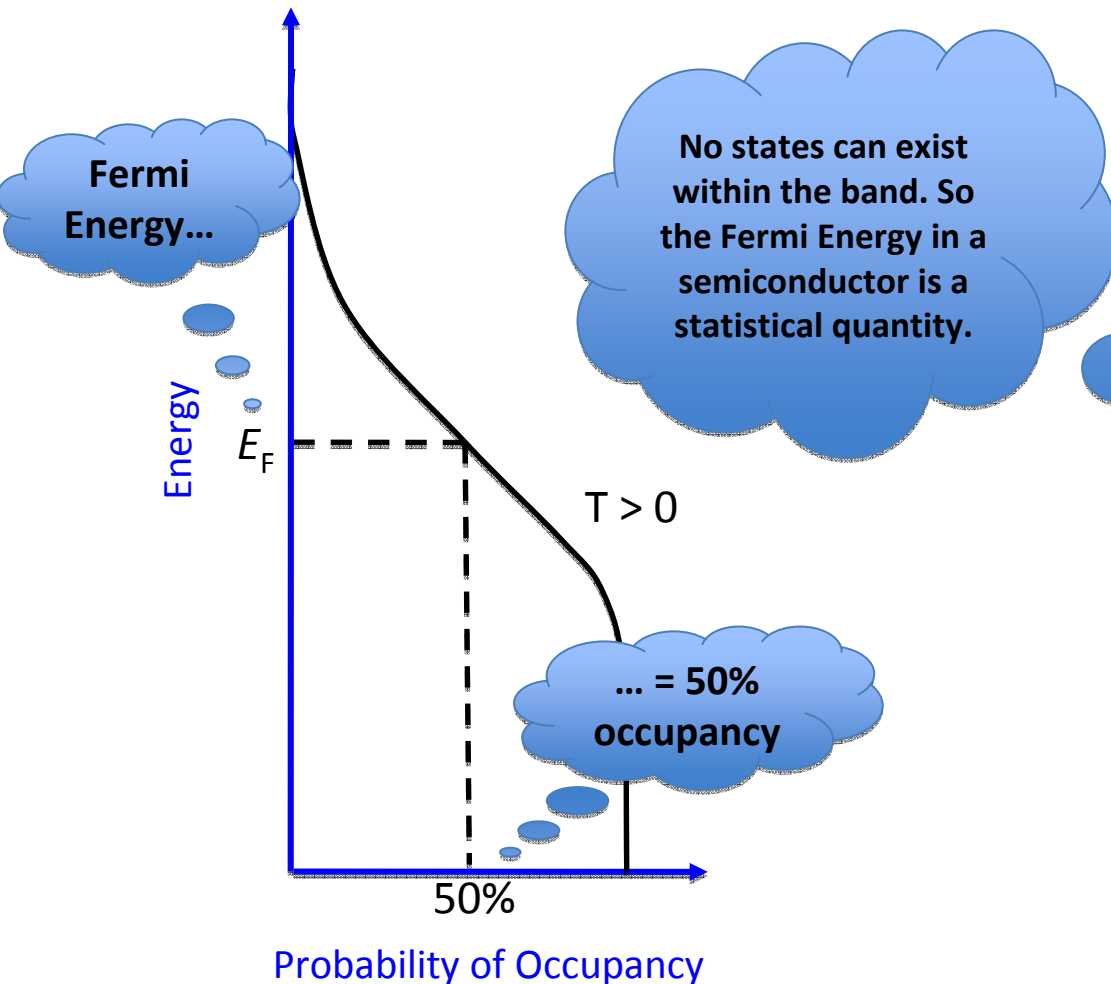


Density of States

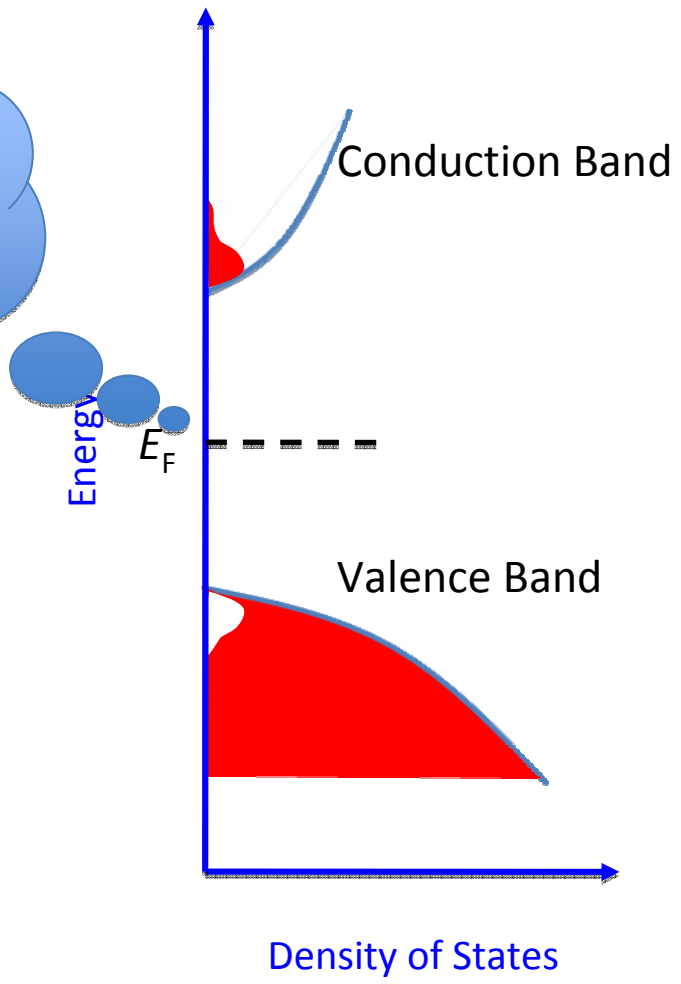


Defining the "Fermi Level"

Fermi-Dirac Probability Distribution Function at $T > 0$.

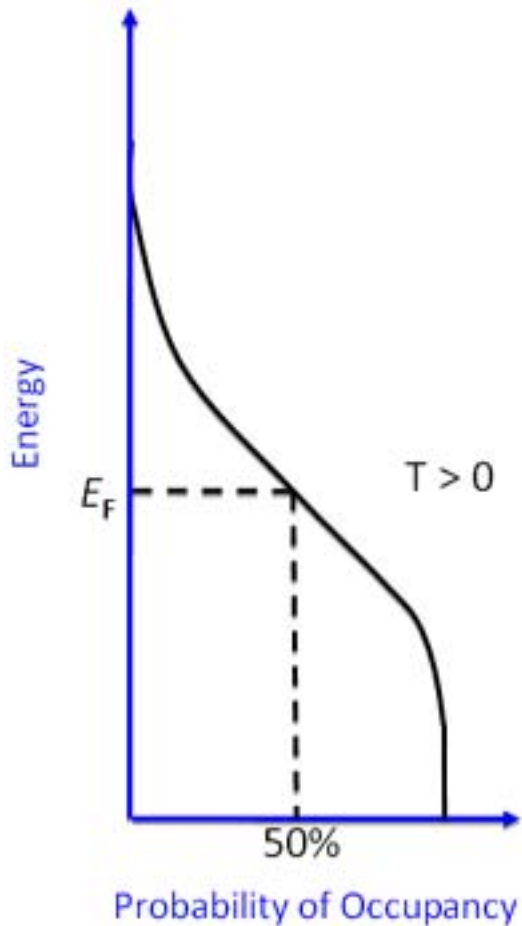


Density of States

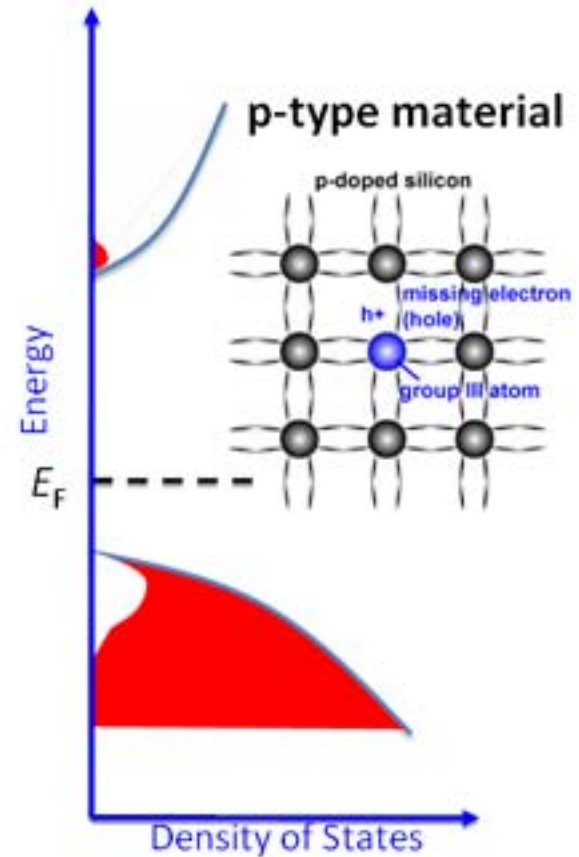


Fermi Level in p-type Material

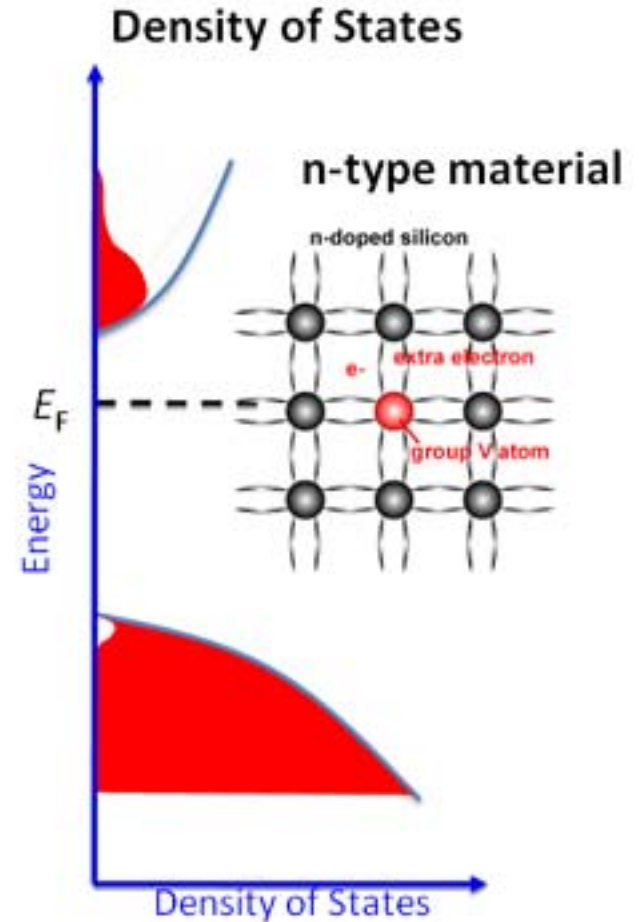
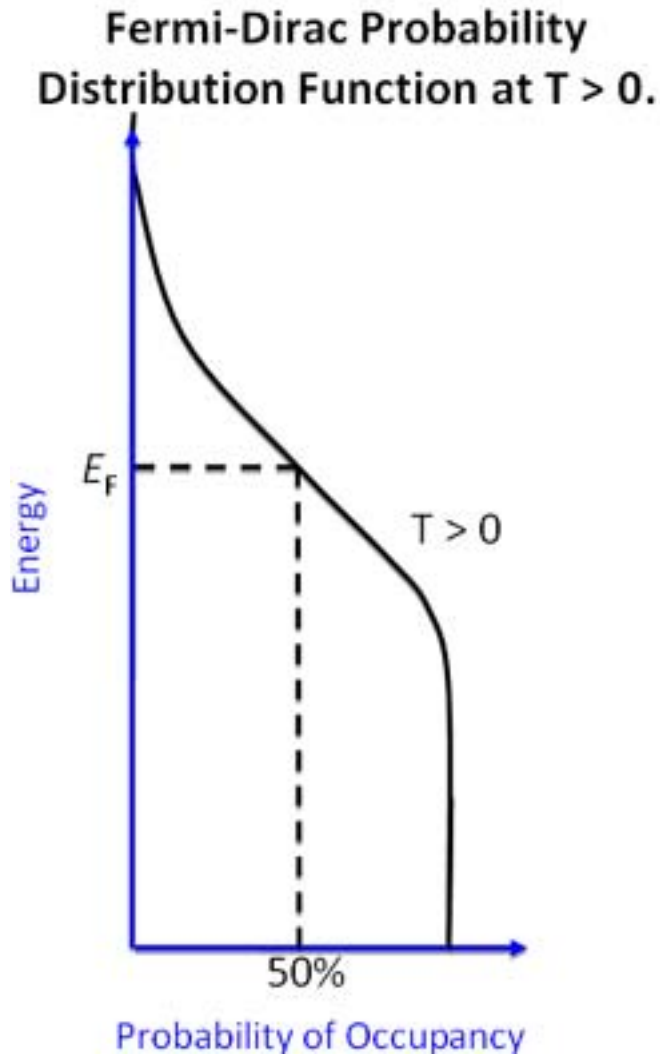
Fermi-Dirac Probability
Distribution Function at $T > 0$.

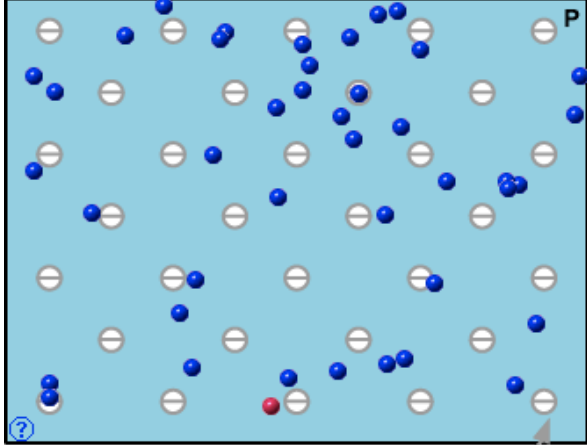


Density of States

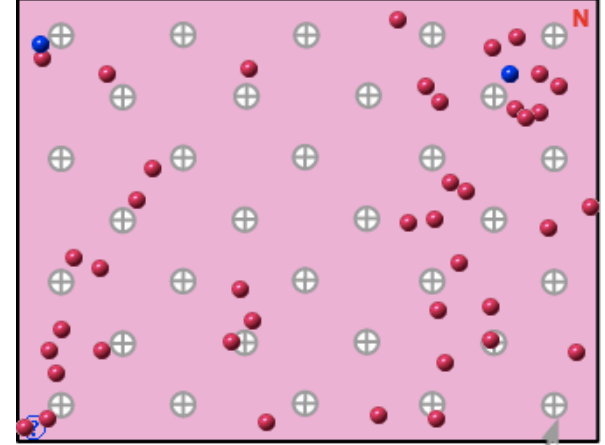


Fermi Level in n-type Material



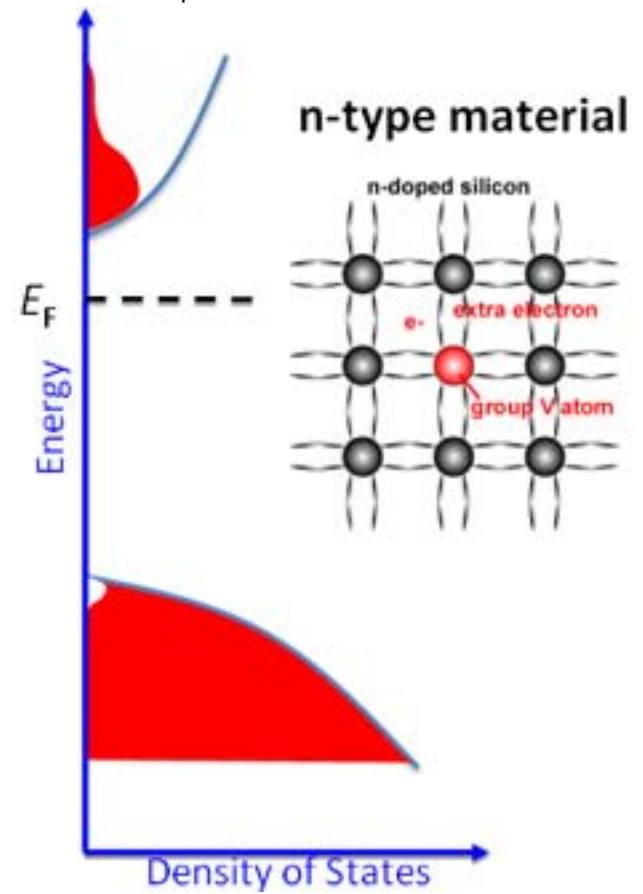
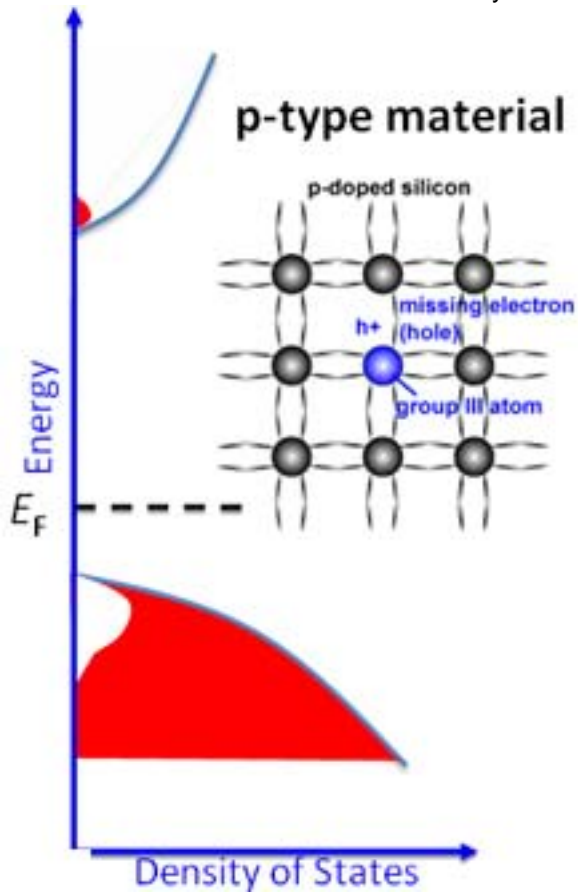


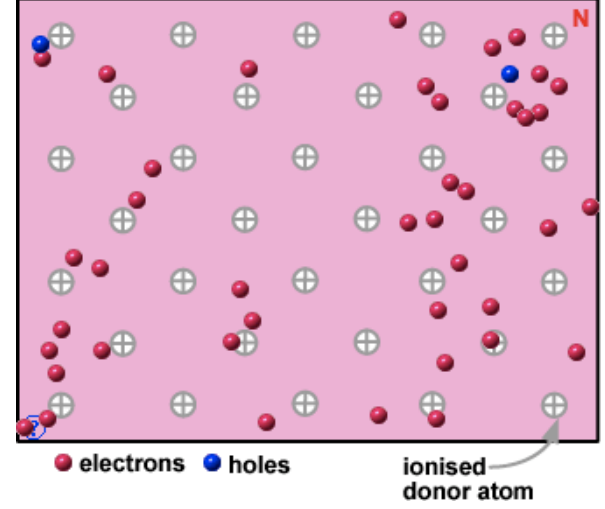
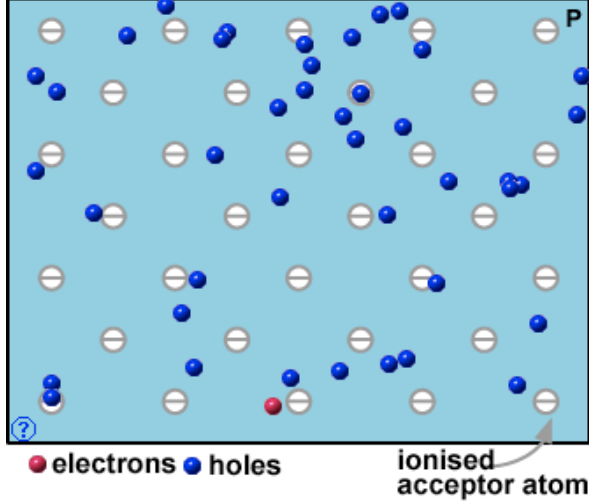
● electrons ● holes ionised acceptor atom



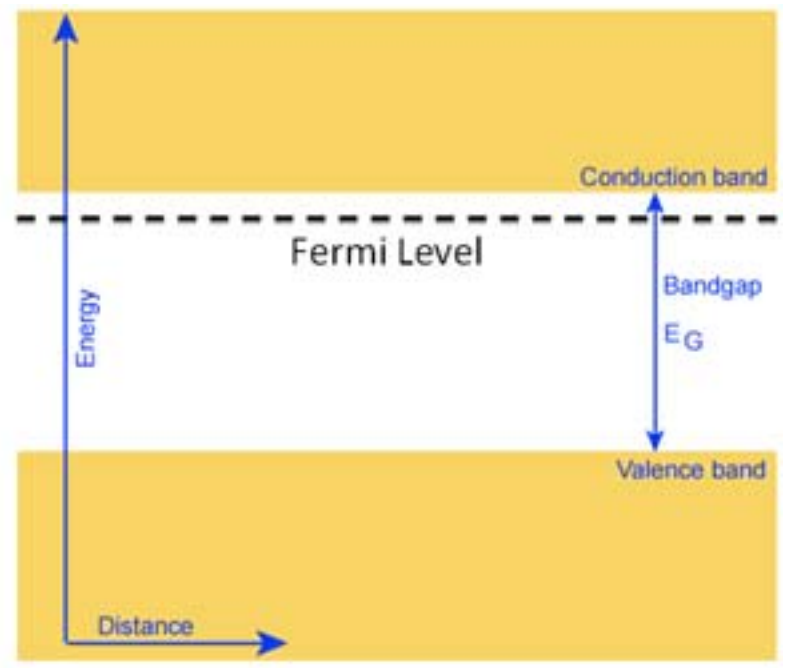
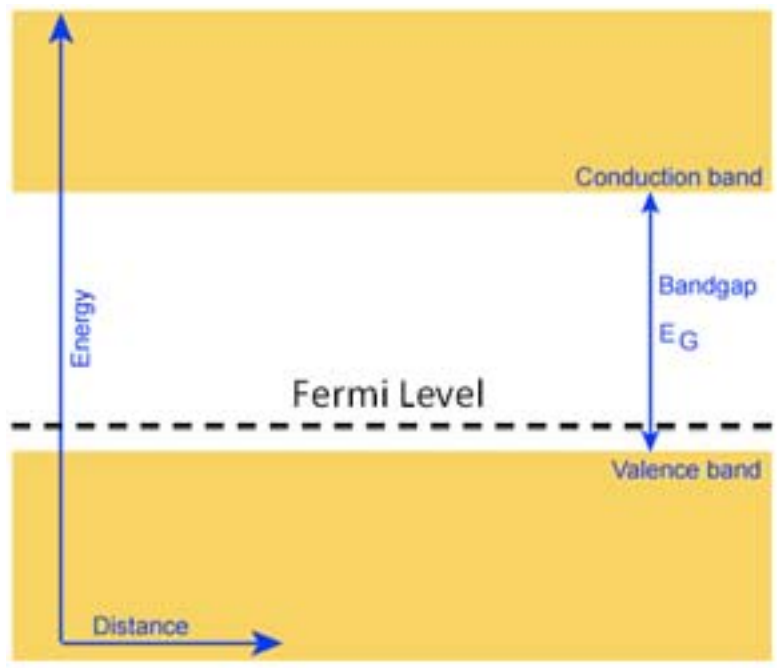
● electrons ● holes ionised donor atom

Courtesy Christiana Honsberg and Stuart Bowden. Used with permission.





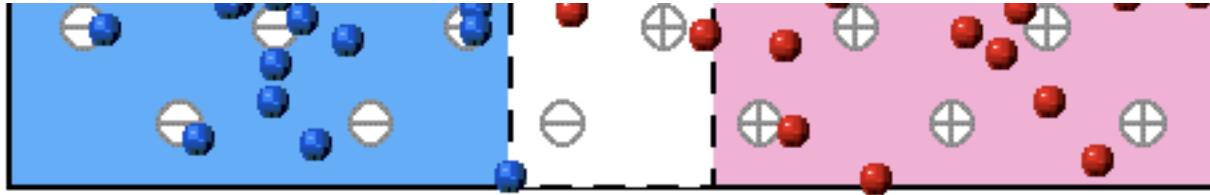
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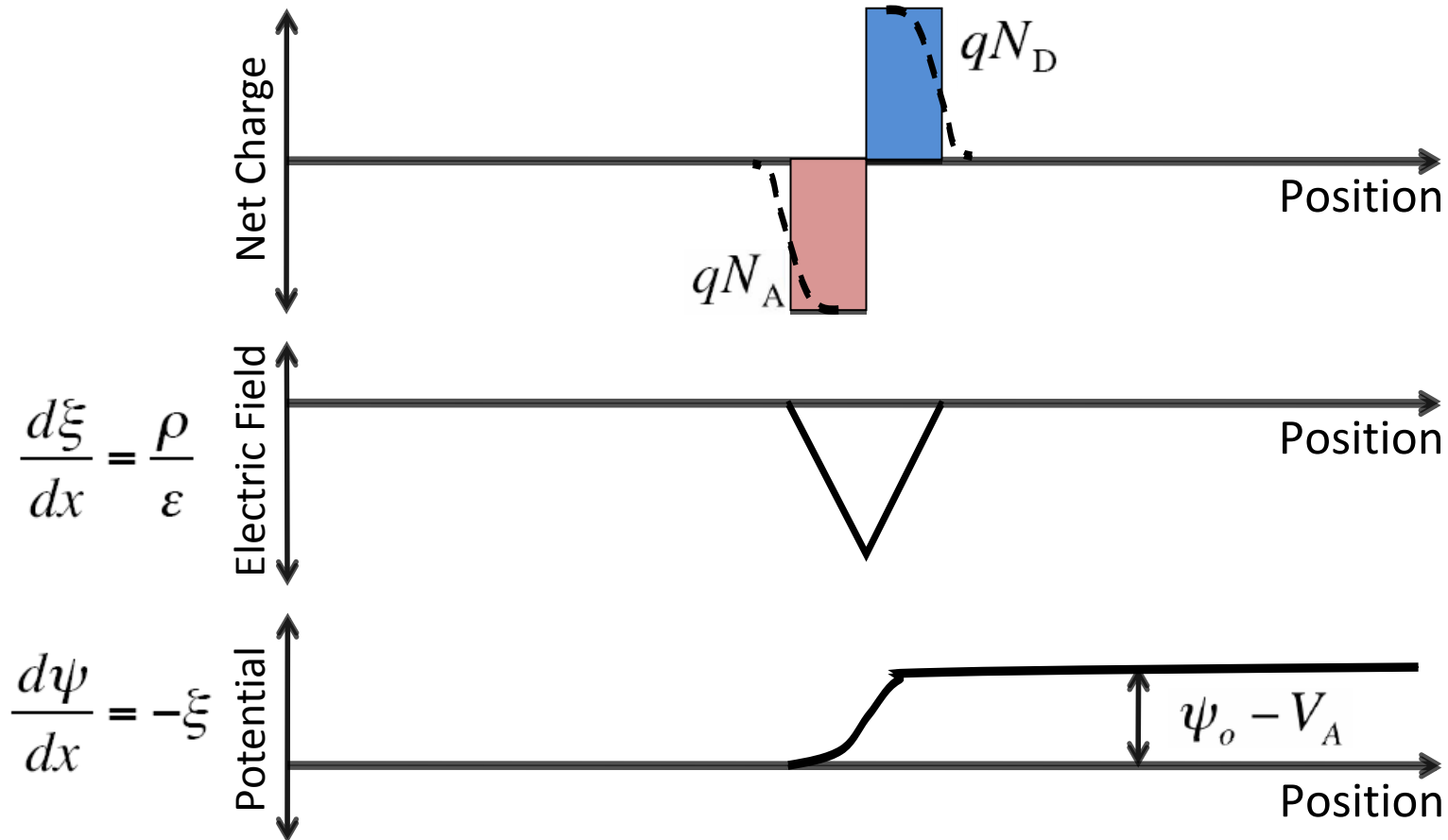
Fundamentals

1. Properties of Sunlight
2. Charge Excitation and Conduction
3. Charge Separation
 - a. Excitons
 - b. pn junction & IV curve
 - c. p-i-n junction
4. Charge Collection

How a pn-junction comes into being



Courtesy Christiana Honsberg and Stuart Bowden. Used with permission.



Dark Diodes: Forward and Reverse Bias

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http://people.seas.harvard.edu/~jones/es154/lectures/lecture_2/pn_junction/junc_dyn_1b.jpg.

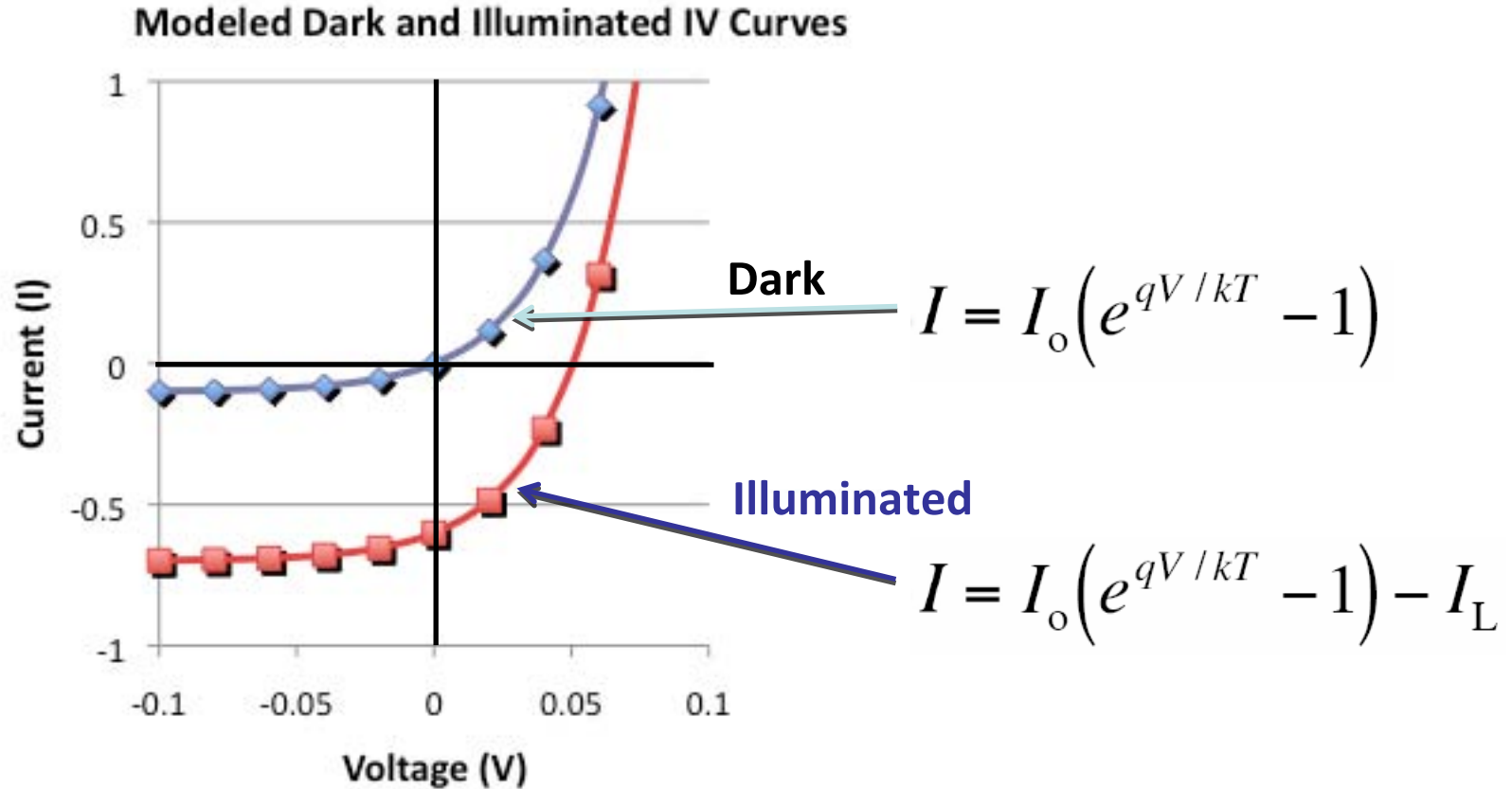
http://people.seas.harvard.edu/~jones/es154/lectures/lecture_2/pn_junction/junc_dyn_1a.jpg.

http://people.seas.harvard.edu/~jones/es154/lectures/lecture_2/pn_junction/junc_dyn_1c.jpg.

Study Tip: Read pp. 63, 65, 79 in Green,
and the PVCDROM.

http://people.seas.harvard.edu/~jones/es154/lectures/lecture_2/pn_junction/pn_junction.html

Ideal Diode Equation



Curves designed using ideal diode equation, with $I_o = 0.1$ (a.u.), and $I_L = 0.6$ (a.u.).

Fundamentals

1. Properties of Sunlight
2. Charge Excitation and Conduction
3. Charge Separation
4. Charge Collection
 - a. Types
 - b. Losses

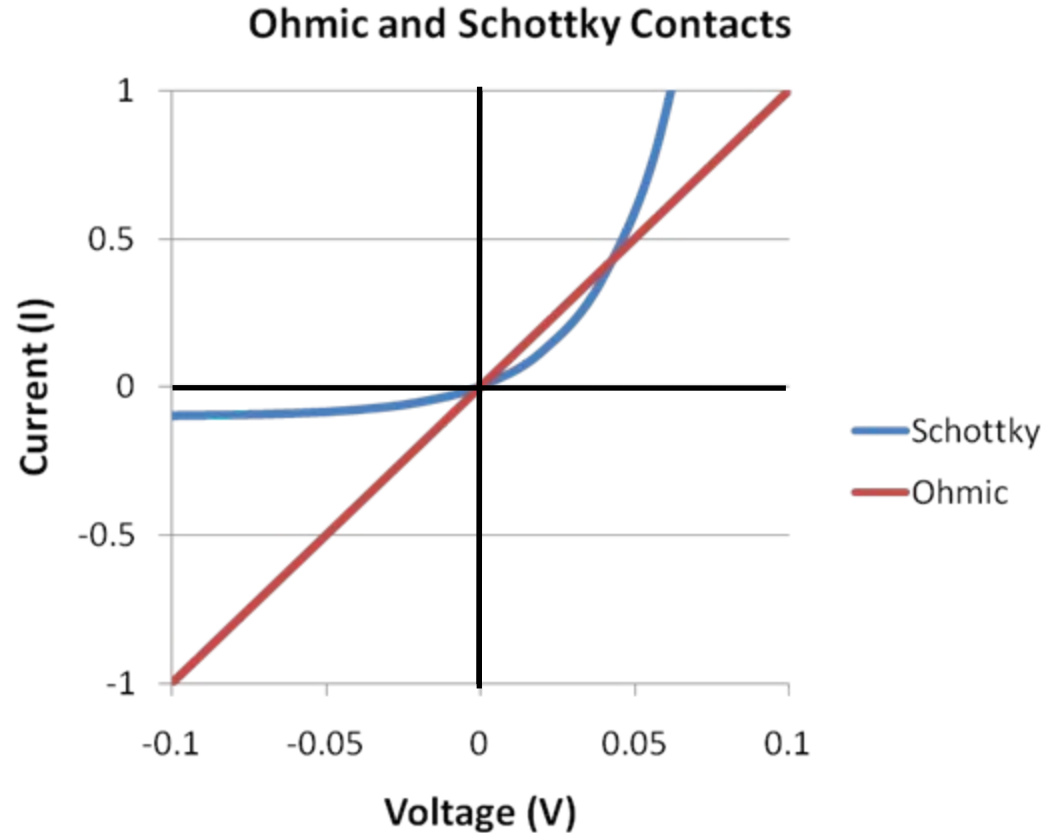
Classes of Contacts

- **Ohmic:**

- Electron barrier height ≤ 0 .
- Linear I-V curve.
- Typically used when charge separation is not a goal for metallization.

- **Schottky:**

- Electron barrier height > 0 .
- Exponential I-V curve.
- Used when charge separation is desired.



Materials Commonly Used for Contacts

- **Metals**
 - Optically opaque.
 - Electrically conductive.

- **Transparent Conducting Oxides (TCOs)**
 - Optically transparent.
 - Electrically conductive.

Evaluating Metals for Contacts - Theoretical

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<http://www.iue.tuwien.ac.at/phd/ayalew/node56.html>

Fundamentals

1. Properties of Sunlight
 - a. Solar Spectrum
 - b. Atmospheric Absorption
2. Charge Excitation and Conduction
 - a. Optical Absorption
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3. Charge Separation
 - a. Excitons
 - b. pn junction & IV curve
 - c. p-i-n junction
4. Charge Collection
 - a. Types
 - b. Losses

General Matters

- Practice Exam
- Concept Quiz Results