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Citation: Dahle, H., C. L. Sarazin, L. A. Lopez, C. Kouveliotou, S. K. Patel, E. Rol, A. J. van der Horst, et al. "THE BURST CLUSTER: DARK MATTER IN A CLUSTER MERGER ASSOCIATED WITH THE SHORT GAMMA-RAY BURST, GRB 050509B." *The Astrophysical Journal* 772, no. 1 (July 2, 2013): 23. © 2013 American Astronomical Society.

As Published: <http://dx.doi.org/10.1088/0004-637X/772/1/23>

Publisher: Institute of Physics/American Astronomical Society

Persistent URL: <http://hdl.handle.net/1721.1/94494>

Version: Final published version: final published article, as it appeared in a journal, conference proceedings, or other formally published context

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THE BURST CLUSTER: DARK MATTER IN A CLUSTER MERGER ASSOCIATED WITH THE SHORT GAMMA-RAY BURST, GRB 050509B

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Received 2012 July 10; accepted 2013 May 19; published 2013 July 2

ABSTRACT

We have identified a merging galaxy cluster with evidence of two distinct subclusters. The X-ray and optical data suggest that the subclusters are presently moving away from each other after closest approach. This cluster merger was discovered from observations of the first well-localized short-duration gamma-ray burst (GRB), GRB 050509B. The *Swift*/Burst Alert Telescope error position of the source is coincident with a cluster of galaxies ZwCl 1234.0+02916, while the subsequent *Swift*/X-Ray Telescope localization of the X-ray afterglow found the GRB coincident with 2MASX J12361286+2858580, a giant red elliptical galaxy in the cluster. Deep multi-epoch optical images were obtained in this field to constrain the evolution of the GRB afterglow, including a total of 27,480 s exposure in the F814W band with *Hubble Space Telescope* Advanced Camera for Surveys, among the deepest imaging ever obtained toward a known galaxy cluster in a single passband. We perform a weak gravitational lensing analysis based on these data, including mapping of the total mass distribution of the merger system with high spatial resolution. When combined with *Chandra X-ray Observatory* Advanced CCD Imaging Spectrometer and *Swift*/XRT observations, we are able to investigate the dynamical state of the merger to better understand the nature of the dark matter component. Our weak gravitational lensing measurements reveal a separation of the X-ray centroid of the western subcluster from the center of the mass and galaxy light distributions, which is somewhat similar to that of the famous “Bullet cluster,” and we conclude that this “Burst cluster” adds another candidate to the previously known merger systems for determining the nature of dark matter, as well as for studying the environment of a short GRB. Finally, we discuss potential connections between the cluster dynamical state and/or matter composition, and compact object mergers, which is currently the leading model for the origin of short GRBs. We also present our results from a weak-lensing survey based on archival Very Large Telescope images in the areas of five other short GRBs, which do not provide any firm detections of mass concentrations representative of rich clusters.

Key words: galaxies: clusters: individual (ZwCl 1234.0+02916, NSC J123610+285901) – gamma-ray burst: individual (GRB 050509B)

Online-only material: color figures

1. INTRODUCTION

Dark matter comprises over 80% of the total mass in the universe, yet its nature is still unknown. Its presence is thus far inferred by its gravitational effects on galaxies and clusters of galaxies, and in particular cluster mergers where the galaxies and dark matter decouple from the hot gas component due to the effects of gas ram pressure as well as the weakness of the dark-matter–dark-matter and dark-matter–baryon interactions. This decoupling was first clearly seen in the Bullet cluster, 1E 0657–558, where the two cluster cores passed through each other ~ 100 Myr ago (Clowe et al. 2004, 2006). Weak-lensing mass reconstructions of the Bullet cluster have shown that the peaks of the mass distributions of the main cluster and the “bullet” subcluster coincide spatially with their corresponding galaxy concentrations and are offset from their X-ray halos at high significance. In 2007, another cluster merger,

MACS J0025.4+1222, was found with clear separation of its components, enhancing the case for further identification of the presence of dark matter in clusters (Bradač et al. 2008). Okabe & Umetsu (2008) analyzed seven nearby clusters at different merging stages using weak-lensing observations with Subaru together with archival *Chandra* and *XMM-Newton* data. Multi-wavelength studies of one of these clusters, A520, revealed a massive dark core coincident with the X-ray emission peak without a corresponding bright cluster galaxy, an observation not yet entirely understood (Mahdavi et al. 2007). In another cluster, A2163, a bimodal mass distribution is observed suggesting ram pressure stripping effects (Okabe et al. 2011; Jee et al. 2012) and contradictory results were presented by Clowe et al. (2012). One of the most actively merging clusters, A2744, exhibits a very complex and challenging phenomenology with a host of substructures (dark, ghost, bullet, and stripped; Merten et al. 2011). Recently, Ragozzine et al. (2012) reported the

detection of two separate mergers taking place at the same time in the cluster A1758, and Menanteau et al. (2012) reported multi-wavelength observations of ACT–CL J0102–4915, the most massive and most X-ray luminous cluster known at beyond redshift 0.6. Finally, analytic estimates of the time-since-collision for DLSCL J0916.2+2951 have revealed a merger at a highly evolved stage after merging (0.7 ± 0.2 Gyr; Dawson et al. 2012).

Given the lack of knowledge about the three-dimensional geometry of the mergers except for the Bullet cluster, a very simplified picture is that they contain two significant subclusters and a merger that occurs in the plane of sky. This geometry results in a maximum angular offset between the hot gas and dark matter, making the physical interpretation of the system relatively straightforward. In contrast, Jee et al. (2007) reported the discovery of a ring of dark matter around another cluster merger, CL 0024+17, where the dark matter distribution differs from the distribution of both the galaxies and the hot gas. In this case, the merger is believed to have occurred almost exactly along the line of sight toward the cluster. Cluster mergers also constitute powerful cosmic laboratories for the study of the particle physics nature of dark matter by providing constraints on the collisional cross-section for dark matter self-interaction (Markevitch et al. 2004; Randall et al. 2008) and on any high-energy emission from radiative decay of candidate dark matter particles (Boyanovsky et al. 2008; Riemer-Sørensen et al. 2006, 2007a, 2007b). Finally, when studying the mergers of massive clusters, we witness the formation of the largest, well-defined, and virialized structures that will ever form in our universe if the energy density is indeed dominated by dark energy.

In this paper, we report the discovery of a merging cluster of galaxies along the line of sight of the short gamma-ray burst, GRB 050509B, based on *Chandra* observations, coupled with a weak gravitational lensing analysis of *Hubble Space Telescope* (*HST*) and Very Large Telescope (VLT) data. GRB 050509B was discovered with the *Swift*/Burst Alert Telescope (Gehrels et al. 2004; Barthelmy et al. 2005). Subsequent X-ray observations with the *Swift*/X-ray Telescope (XRT; Burrows et al. 2005a) identified both the GRB afterglow and diffuse emission from a cluster of galaxies ZwCl 1234.0+02916, also known as NSC J123610+285901 (Gal et al. 2003), at redshift $z = 0.2214$ (Zwicky et al. 1963). An a posteriori probability of the chance coincidence of GRB 050509B with the host galaxy/cluster has been estimated as 0.1%–1% (Gehrels et al. 2005; Pedersen et al. 2005; Bloom et al. 2006; Berger et al. 2007). Our *Chandra* observations of the GRB, reported here, coupled with a weak gravitational lensing analyses of VLT and *HST* data, show that the cluster is undergoing a merger, with indications that two distinct subclusters are likely moving away from each other after their closest approach.

We report the X-ray, optical, and radio observations and data analysis of GRB 050509B and the cluster ZwCl 1234.0+02916 in Section 2. The results of our joint *Chandra* and *Swift* spectral and spatial analysis of the X-ray emitting gas from the merging cluster and a discussion on the cluster X-ray morphology are presented in Section 3. In Section 4, we derive the dimensionless surface mass density distribution and the total cluster mass using observations of the burst field with the VLT and the *Hubble Space Telescope*/Advanced Camera for Surveys (*HST*/ACS) to measure the weak shear due to gravitational lensing. In Section 5, we compute weak-lensing maps using public VLT data toward other well-localized short GRBs to determine the presence (if any) of cluster-size dark matter concentrations in

their fields. In Section 6, we discuss the possible scenarios that could account for the observed cluster structure and the putative association of GRB 050509B with this environment. Finally, in Section 7 we summarize our results. Throughout this work, we assume a cosmology with $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, and $H_0 = 70$ km s⁻¹.

2. OBSERVATIONS AND DATA REDUCTION

We present below the analyses of the multi-wavelength observations of the field of GRB 050509B, the first ever short GRB, which was accurately ($<10''$) and rapidly (minutes) located (Gehrels et al. 2005). As such, the event generated a large amount of interest resulting in multiple follow-up observations (in X-ray, optical, and radio wavelengths). The study of all these data sets is invaluable for understanding both the prompt emission mechanism of the burst, the properties of its afterglow emission, its bolometric energetics, as well as the properties of its putative host galaxy and its environment. In a serendipitous way, the combination of all these wavelengths has also allowed the discovery and study of the merger results presented here. We proceed in this section to discuss the various data sets and our analyses techniques as related to the understanding of the GRB phenomenon alone.

2.1. *Swift*/XRT X-Ray Observations

XRT observed the field of GRB 050509B intermittently starting 2005 May 9.17 for a total of 34.5 ks of useful X-ray imaging and spectral data in Photon Counting mode. The data were processed using standard *Swift* analysis tools in the HEASOFT (v6.0.4) package and the XRT CALDB (v20060104). The ZwCl 1234.0+02916 X-ray emission is clearly detected and imaged on the center of the detector. One X-ray point source was identified and removed from this region; this source was also detected in the *Chandra* observation and is discussed in Section 3.1. The total (cluster plus background) XRT count rate is 0.0158 count s⁻¹ (0.3–10.0 keV) and the average background, scaled to the same area and in the same energy band, is 0.0059 count s⁻¹. The XRT and *Chandra* cluster spectra are discussed in detail in Section 3.3.

2.2. *Chandra* X-Ray Observations

Chandra observed the XRT error box of GRB 050509B, including the cluster, starting 2005 May 11.167 for a total useful time of 49.3 ks with the ACIS-S3 CCD in Very Faint (VF) Timed Exposure mode (ObsID 5588). The X-ray data analysis software package CIAO¹² v4.1 and *Chandra* CALDB¹³ v4.2.0 were used to process the Advanced CCD Imaging Spectrometer (ACIS) data. We used the Repro III *Chandra* data release,¹⁴ which includes removal of background flares from the data. We then ran *acis_process_events* to ensure that the latest gains are applied and removed the pixel randomization. Since the data were collected in VF mode, which records a 5×5 pixel region per event, we were able to do additional cleaning and removal of particle background.¹⁵ We determined and removed the background and identified all point-like sources using *wavdetect* in the *Chandra* image (0.5–8.0 keV); a small region around each of the sources was excluded from the cluster region when

¹² <http://asc.harvard.edu/ciao/>

¹³ <http://asc.harvard.edu/caldb/>

¹⁴ http://cxc.harvard.edu/ciao4.4/repro_iii.html

¹⁵ http://cxc.harvard.edu/cal/Acis/Cal_prods/vfbkgrnd/

Table 1
Public VLT Data for Six Well-localized Short GRBs

| GRB Name | Galactic Extinction $E(B - V)$ | Date (UT) | Instrument Name | Filter | t_{exp} (s) | FWHM (arcsec) | N_{gal} |
|----------|-----------------------------------|----------------|--------------------|--------|-------------------------|------------------|------------------|
| 050509B | 0.019 | 2005 May 13.07 | FORS2 | V | 2700 | 0.63 | 2674 |
| | | 2005 Jun 01.02 | FORS2 | V | 1800 | 0.81 | 1883 |
| | | 2005 May 11.00 | FORS2 | R | 2160 | 0.85 | 1651 |
| | | 2005 May 31.96 | FORS2 | R | 3600 | 0.70 | 2257 |
| 050724 | 0.590 | 2005 Jul 24.99 | FORS1 | V | 480 | 0.82 | 543 |
| | | 2005 Jul 27.99 | FORS1 | V | 540 | 0.77 | 556 |
| | | 2005 Jul 25.00 | FORS1 | R | 540 | 1.08 | 277 |
| | | 2005 Jul 25.98 | FORS1 | R | 540 | 0.99 | 417 |
| | | 2005 Jul 27.97 | FORS1 | R | 540 | 0.77 | 537 |
| | | 2005 Jul 30.10 | FORS1 | R | 720 | 0.48 | 672 |
| | | 2005 Aug 22.99 | FORS1 | R | 780 | 1.45 | 359 |
| | | 2005 Aug 25.98 | FORS1 | R | 1140 | 0.68 | 815 |
| | | 2005 Jul 25.01 | FORS1 | I | 540 | 0.95 | 483 |
| | | 2005 Jul 25.97 | FORS1 | I | 360 | 0.77 | 483 |
| | | 2005 Jul 27.98 | FORS1 | I | 540 | 0.74 | 496 |
| | | 2005 Jul 30.11 | FORS1 | I | 720 | 0.50 | 733 |
| 050813 | 0.054 | 2005 Aug 19.06 | FORS2 | I | 1800 | 0.64 | 1048 |
| 050906 | 0.066 | 2005 Sep 25.28 | FORS1 | V | 1800 | 0.89 | 1541 |
| | | 2005 Sep 07.35 | FORS2 | V | 1800 | 0.65 | 2248 |
| | | 2005 Sep 12.31 | FORS2 | V | 1800 | 0.91 | 1611 |
| | | 2005 Sep 25.27 | FORS1 | R | 1800 | 0.88 | 2308 |
| | | 2005 Sep 07.33 | FORS2 | R | 1800 | 0.70 | 2088 |
| | | 2005 Sep 12.30 | FORS2 | R | 1800 | 0.79 | 2050 |
| | | 2005 Sep 25.29 | FORS1 | I | 1800 | 0.82 | 1759 |
| | | 2005 Sep 07.37 | FORS2 | I | 1920 | 0.73 | 1497 |
| 050911 | 0.010 | 2005 Sep 12.32 | FORS2 | I | 1440 | 0.86 | 1410 |
| | | 2005 Sep 12.19 | FORS1 | V | 540 | 0.57 | 1512 |
| | | 2005 Sep 12.41 | FORS1 | V | 540 | 0.77 | 1359 |
| | | 2005 Sep 12.18 | FORS1 | R | 540 | 0.56 | 1727 |
| 060313 | 0.064 | 2005 Sep 12.40 | FORS1 | R | 540 | 0.74 | 1061 |
| | | 2006 Mar 20–30 | FORS1 | R | 9190 | 1.21 | 2158 |

extracting spectra. To account for instrumental variations, we produced exposure-corrected images. To derive a clean image of the cluster’s diffuse emission (Section 3.2.2), we subsequently filled the point-source gaps with pixel count values selected from the Poisson distribution of the area surrounding the point sources using the CIAO tool *dmfilth*.

2.3. VLT Optical Data

We downloaded from the VLT archive all public data taken with the FOcal Reducer and low dispersion Spectrographs 1 and 2 (FORS1 and FORS2) for GRB 050509B (first published in Hjorth et al. 2005) and five other well-localized short GRBs. FORS1 has a pixel size of $0''.2$ and a field of view of $6''.8 \times 6''.8$, while the pixels of FORS2 were re-binned to a pixel scale $0''.25$, with a usable field of view of $6''.9 \times 6''.9$. The data were reduced in a standard fashion using IRAF¹⁶; overscan correction and bias subtraction were performed separately on the four read-out quadrants of the FORS1 chips. The reduced images were corrected for cosmic rays using the Laplacian cosmic-ray identification method of Van Dokkum (2001). Finally, individual images were interpolated to the same reference frame using a third-order polynomial and were added together. Images were not added across intervals spanning more than several hours, since in most cases, circumstances such as the seeing were significantly different between the various epochs.

¹⁶ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under a cooperative agreement with the National Science Foundation.

As shown in Table 1, the number of galaxies detected in the images of a given GRB field varied significantly between different epochs, due to variations in exposure time, sky brightness, and seeing. The Galactic extinction was also significantly different across the various GRB fields.

Finally, for GRB 050509B only, we aligned the VLT image to the Two Micron All Sky Survey (2MASS) coordinate system as the latter is used in Section 3.1 to register the X-ray images. To this end, we searched the central $6''.5 \times 3''.6$ region of the 3600 s R-band VLT/FORS2 image using LEXTRACT (with $\sigma_{\text{psf}} = 2.216$ pixels and signal-to-noise ratio (S/N) = 1.1) and detected 1515 sources. We then searched the 2MASS catalog and found 14 acceptable counterparts to our R-band sources, which we used to align the VLT image to the 2MASS coordinate system to an accuracy of $\sigma_{\text{rms}}(\text{R.A.}) = 0''.17$ and $\sigma_{\text{rms}}(\text{decl.}) = 0''.21$.

2.4. HST/ACS Optical Data

The field of GRB 050509B was observed for a total of 27,480 s, using the Wide Field Channel of the HST/ACS. The imaging data were obtained with the F814W filter on four different dates in 2005 May and June (Proposal ID 10119). Data were retrieved from the HST archive in the form of ACS association FITS images (i.e., each image was based on three individual dithered exposures), which had been calibrated using the *MultiDrizzle* pipeline software. This calibration pipeline produces cosmic-ray cleaned images of pixel scale $0''.05$, which have been re-sampled and combined using the drizzle software

Table 2
WSRT Measurements of the GRB 050509B Afterglow and Host Galaxy Flux

| Frequency (GHz) | Observation Date (UT) | Time Since GRB Trigger (days) | GRB Flux ^a (μ J) | Galaxy Flux (μ J) |
|-----------------|-------------------------|-------------------------------|----------------------------------|------------------------|
| 4.9 | May 09.63–10.12 | 00.46–00.96 | 31 ± 22 | 315 ± 61 |
| 4.9 | May 19.60–20.08 | 10.43–10.92 | 40 ± 33 | 325 ± 76 |
| 4.9 | May 9 and 19 (combined) | ... | 11 ± 20 | 321 ± 59 |
| 1.4 | May 19.63–20.10 | 10.46–10.93 | 1 ± 17 | 1183 ± 195 |

Note. ^a Formal flux measurements at the GRB position with 1σ rms noise around that position.

of Fruchter & Hook (2002), while correcting for the geometric distortion of ACS.

The resulting set of 12 images were offset by up to $2''$ in x and y (and also slightly rotated with respect to each other), and needed re-registration before combination into a single, deep image. This was done by running the IMCAT (Kaiser et al. 1995) peak finding algorithm (*hfindpeaks*) on the FITS image corresponding to each observation set, and performing aperture photometry (*apphot*) on the detected objects. Approximately 200 stars and bright, compact non-stellar objects were selected as reference objects for the re-registration of the images. A fiducial frame was chosen among the available images, and the transformation of each of the other images into this frame was calculated as a second-order polynomial. The solutions for these transformations were obtained iteratively by interactive rejection of outliers to the fit until a good solution was found with rms residuals less than 0.1 pixel in x and y . Based on the resulting transformation parameters, all images were then re-registered into the same coordinate system and combined into a single, deep frame. The image is comparable in depth to the combined F814W frames of the original Hubble Deep Fields (Williams et al. 1996), and is one of the deepest images ever obtained toward a galaxy cluster beyond Coma in a single passband.

2.5. Radio Observations

GRB 050509B was observed at centimeter wavelengths with the Westerbork Synthesis Radio Telescope (WSRT). We used the Multi Frequency Front Ends (Tan 1991) in combination with the IVC+DZB backend¹⁷ in continuum mode, with a bandwidth of 8×20 MHz. The observations were carried out on two epochs in 2005 (Table 2). On the first epoch, the source was observed for 12 hr at 6 cm (4.86 GHz); on the second epoch we switched between 6 cm and 21 cm (1.43 GHz), with 40 minutes dwell time per frequency. Gain and phase calibrations were performed with the calibrator 3C 286. Data reduction was performed using the MIRIAD software (Sault et al. 1995). In Table 2, we report the upper limits on the 4.9 and 1.4 GHz fluxes at the location of the *Swift*/XRT GRB afterglow. Preliminary results of the afterglow flux from the first epoch were published in Van der Horst et al. (2005).

We detected a radio source, G1, in the center of the GRB host galaxy. The source is unresolved at the WSRT resolution, with a synthesized beam size (FWHM) of $24''.7 \times 9''.5$ at 1.4 GHz and $17''.6 \times 8''.3$, in the combined image at 4.9 GHz. Using the flux of the combined data set at 4.9 GHz, we derived a spectral index of -1.07 ± 0.26 (1.4–4.9 GHz). Two possible origins for the radio source are either a mildly active galactic nucleus (AGN) or nuclear star formation. The latter is unlikely to be the correct explanation as our data would imply a star formation rate (SFR)

of $71 \pm 21 M_{\odot} \text{ yr}^{-1}$ (following Berger et al. 2003). This value is much higher than the 3σ upper limit to the current SFR of $0.1 M_{\odot} \text{ yr}^{-1}$ found for the host galaxy from the $H\alpha$ luminosity (Bloom et al. 2006). Note that if the radio emission is due to star formation, much of it is expected to come from supernova remnants (SNRs), and it would thus also depend on the SNR history. Finally, we do not expect to have a large fraction of obscured star formation by dust in this elliptical galaxy. On the other hand, the value of the radio spectral index is quite typical for optically thin AGN jets, although a bit at the high end of the distribution (Brown & Burns 1991). If the emission would be dominated by the optically thick core of the jet, the spectrum would be quite flat.

3. X-RAY DATA ANALYSIS AND RESULTS

3.1. Astrometry and X-Ray Source Catalog

We describe here the results of our X-ray source search in the *Chandra*/ACIS and *Swift*/XRT images. We searched the *Chandra* image for X-ray sources at the highest spatial resolution ($0''.492 \text{ pixel}^{-1}$) and within a radius of $\sim 8'$ of the nominal aimpoint. We used the point-source detection method (Gaussian point-spread function (PSF) fitting) described in Tennant (2006).¹⁸ Only events in the energy range 0.5–7.0 keV were used, to maximize the source to background count ratio. The search resulted in 43 X-ray sources with an $S/N > 2.8$. The sources are listed in Table 3, which gives their positions (after the *Chandra* image was aligned to the 2MASS catalog as described below), their net number of counts, the width of the *Chandra* PSF at each source location on the ACIS-S chip, and the S/N of each source.

We then searched for counterparts within $2''$ of each *Chandra* X-ray source in the 2MASS Point Source Catalog (Skrutskie et al. 2006). Only two counterparts were found (these are 6 and 12 in Table 3). We used these two sources and two different alignment methods (CIAO/wavdetect and LEXTRACT) to compute a mean shift between the *Chandra* and 2MASS images. The results from both methods were consistent and gave a shift of $< 0''.4$ in each coordinate. We also compared our X-ray source list to the USNO2B catalog and found nine potential matches. As the errors in the USNO source positions are known to be typically greater than the uncertainties in the 2MASS positions, we elected to tie all fields to the 2MASS coordinate system throughout this paper.

We used the PSF-fitting method described above to detect sources on the *Swift*/XRT image in the energy band 0.3–10 keV. Here, we adopted a fixed PSF FWHM of $3''.77$ (1.6 XRT pixels) based on the on-ground PSF calibration tests of the XRT using the Al K line at 1.49 keV (Moretti et al. 2004). We

¹⁷ Section 5.2 at <http://www.astron.nl/radio-observatory/astronomers/wsrt-guide-observations/wsrt-guide-observations>

¹⁸ Available as part of the image analysis package LEXTRACT (<http://www.wastro.msfc.nasa.gov/qdp>).

Table 3
Chandra X-Ray Source List

| Source ID | R.A. (h m s) | Decl. ($^{\circ}$ $'$ $''$) | Net Counts | σ_{psf} (pixel) | S/N |
|-----------|-----------------|----------------------------------|------------|----------------------------------|-------|
| 1 | 12 35 47.211 | 29 03 30.619 | 136.90 | 6.37 | 10.88 |
| 2 | 12 35 48.503 | 28 58 54.896 | 15.52 | 3.62 | 3.49 |
| 3 | 12 35 50.664 | 28 57 39.498 | 47.54 | 3.23 | 6.13 |
| 4 | 12 35 52.105 | 29 05 22.097 | 22.13 | 7.49 | 4.13 |
| 5 | 12 35 52.470 | 28 57 56.633 | 45.07 | 2.81 | 6.02 |
| 6* | 12 35 52.699 | 28 58 56.507 | 17.42 | 2.73 | 3.78 |
| 7 | 12 35 55.091 | 29 05 43.247 | 48.49 | 7.50 | 6.36 |
| 8 | 12 35 55.542 | 29 00 42.569 | 22.18 | 2.66 | 4.01 |
| 9 | 12 35 56.577 | 28 59 46.196 | 18.37 | 2.19 | 3.80 |
| 10 | 12 35 57.387 | 28 56 04.520 | 58.91 | 2.63 | 6.70 |
| 11 | 12 35 58.058 | 28 57 49.011 | 14.79 | 1.92 | 3.44 |
| 12* | 12 36 01.747 | 29 01 00.472 | 24.26 | 1.98 | 4.45 |
| 13 | 12 36 02.871 | 28 58 35.169 | 42.70 | 1.30 | 5.72 |
| 14 | 12 36 03.407 | 29 00 14.984 | 104.92 | 1.52 | 9.60 |
| 15 | 12 36 03.901 | 29 01 24.909 | 57.63 | 2.01 | 6.77 |
| 16 | 12 36 05.114 | 29 02 23.551 | 24.96 | 2.58 | 4.64 |
| 17 | 12 36 05.146 | 29 00 13.265 | 9.86 | 1.37 | 2.81 |
| 18 | 12 36 06.909 | 29 00 19.693 | 14.14 | 1.31 | 3.61 |
| 19 | 12 36 07.307 | 29 00 18.046 | 22.76 | 1.28 | 4.42 |
| 20 | 12 36 07.768 | 28 58 02.824 | 110.38 | 1.02 | 9.09 |
| 21 | 12 36 08.988 | 28 58 51.684 | 13.95 | 0.94 | 3.13 |
| 22 | 12 36 09.353 | 28 56 00.946 | 77.11 | 1.66 | 7.70 |
| 23 | 12 36 10.376 | 29 02 37.228 | 127.56 | 2.55 | 10.05 |
| 24 | 12 36 10.405 | 29 01 05.592 | 278.17 | 1.53 | 15.23 |
| 25 | 12 36 11.964 | 28 58 47.042 | 13.79 | 0.90 | 3.09 |
| 26 | 12 36 12.481 | 29 02 21.743 | 250.54 | 2.34 | 14.41 |
| 27 | 12 36 14.526 | 29 00 41.834 | 186.14 | 1.37 | 12.29 |
| 28 | 12 36 16.035 | 29 02 44.268 | 18.62 | 2.73 | 3.84 |
| 29 | 12 36 19.712 | 29 04 37.591 | 38.45 | 4.92 | 5.23 |
| 30 | 12 36 20.887 | 29 03 38.672 | 15.28 | 3.90 | 3.25 |
| 31 | 12 36 21.484 | 29 01 49.972 | 21.14 | 2.42 | 3.95 |
| 32 | 12 36 22.262 | 29 03 52.589 | 26.00 | 4.29 | 4.43 |
| 33 | 12 36 23.586 | 29 01 43.497 | 46.15 | 2.57 | 6.12 |
| 34 | 12 36 23.734 | 29 02 53.964 | 11.35 | 3.49 | 2.85 |
| 35 | 12 36 23.768 | 29 00 56.249 | 11.08 | 2.15 | 3.10 |
| 36 | 12 36 23.817 | 28 59 46.703 | 54.18 | 1.75 | 6.46 |
| 37 | 12 36 29.356 | 29 01 38.599 | 394.58 | 3.36 | 17.48 |
| 38 | 12 36 30.279 | 28 58 52.432 | 34.99 | 2.60 | 5.07 |
| 39 | 12 36 31.187 | 29 00 30.065 | 35.91 | 3.12 | 5.32 |
| 40 | 12 36 31.240 | 29 01 42.171 | 32.02 | 3.75 | 4.73 |
| 41 | 12 36 31.716 | 29 03 03.138 | 80.40 | 4.88 | 7.76 |
| 42 | 12 36 39.217 | 28 59 23.538 | 46.13 | 4.69 | 5.79 |
| 43 | 12 36 40.077 | 29 00 57.903 | 31.26 | 5.44 | 4.05 |

Notes. 1 pixel = $0''.492$. Coordinates are referenced to the 2MASS coordinate system, and R.A. and decl. are accurate to $0''.4$ (see Section 3.1). σ_{psf} is the width of the PSF at the pixel location of the point source on the ACIS-S3 chip, as given by LEXTRCT * 2MASS source.

detected 16 sources with $S/N > 2.8$ ($S/N_{\text{min}} = 2.80$ resulted in 9.4 net counts). Each source was individually inspected and all appeared to be real point-like X-ray sources and not merely background fluctuations. All the XRT-detected sources that fell on our *Chandra* field of view were within three XRT pixels of the *Chandra* position. We excluded one source, which was close to the detector edge and might have a less reliable position estimate, and used the remaining seven common sources between the *Chandra* and XRT images to derive a transformation from XRT to sky coordinates. The shifts in R.A. and decl. were $+0''.15$ and $-1''.20$, respectively. We attribute the difference between the offset reported here and the offsets reported in Burrows et al. (2005b) and Bloom et al. (2006) to

the refined boresight calibration (Moretti et al. 2006), which was not applied previously to the analyzed data.

We presented the earlier results of our re-analysis for the XRT position of the GRB 050509B afterglow in Burrows et al. (2005b). Our derived positions for the source centroid using PSF fitting with LEXTRCT and with CIAO/*wavdetect* are consistent with the location reported by Burrows et al. (2005b) at R.A. (J2000) = $12^{\text{h}}36^{\text{m}}13^{\text{s}}.58$, decl. (J2000) = $28^{\circ}59'01''.3$ (error radius of $9''.3$ at the 90% confidence level).

3.2. X-Ray Image Analyses

3.2.1. β Model Fits

To qualitatively characterize the distribution of the diffuse X-ray emission we utilized adaptive smoothing (using CIAO *csmooth*) to create the image shown in Figure 1. We note that the image shows that the gas distribution in ZwCl 1234.0+02916 is bimodal, indicative of a merger between the main cluster component to the east and a subcluster component to the west. The strongest emission is associated with the eastern subcluster, and roughly centered on the galaxies 2MASX J12362010+2859080 and 2MASX J12362094+2859290. There may be a region of concentrated X-ray emission centered on the brightest cluster galaxy 2MASX J12362094+2859290, and a more diffuse eastern component located to the southwest of this galaxy, and roughly centered on 2MASX J12362010+2859080. The western subclump in the X-ray image is elongated.

The X-ray surface brightness declines somewhat rapidly to the west, while there is a long tail of emission to the east. The brightest part of the western X-ray emission is displaced from the GRB host galaxy, 2MASX J12361286+2858580, even though this is the dominant galaxy in the western subcluster. The overall morphology of the western subclump resembles that seen in clusters with a merger ‘‘cold front’’ (Markevitch et al. 2000; Vikhlinin et al. 2001); unfortunately, the existing *Chandra* observation is not deep enough to allow either a sharp surface brightness edge or a temperature jump to be detected at a statistically significant level. The elongated tail on the east of the western subcluster, the drop in surface brightness to the west of the western peak, and the separation between the western X-ray peak and the dominant galaxy in this region, all suggest that the western subclump is moving to the west relative to the bulk of the gas in the eastern subcluster, and that the gas in the western subclump has been slowed and shaped by ram pressure. This morphology resembles that seen in numerical simulations of offset, slightly unequal mass mergers, just after first core passage (e.g., see Figures 4 and 18(b) in Poole et al. 2006).

We proceed below employing two methods to quantitatively describe the density distribution of the X-ray emitting gas in the cluster.

3.2.2. Wavelet-transform Analysis

The first method we used to quantitatively describe the morphology of the diffuse X-ray emission is the wavelet-transform analysis on the full-band (0.5–8.0 keV) ACIS image after point-source removal. Here we followed the method outlined in Lopez et al. (2009): we convolved every pixel of the *Chandra* image with Mexican-hat wavelet functions of different widths. A Mexican-hat function is the optimum wavelet to analyze X-ray images because it removes flat features, like noise, and has a shape similar to a Gaussian signal (Lopez et al. 2009). Wavelet-transformed images are produced by calculating

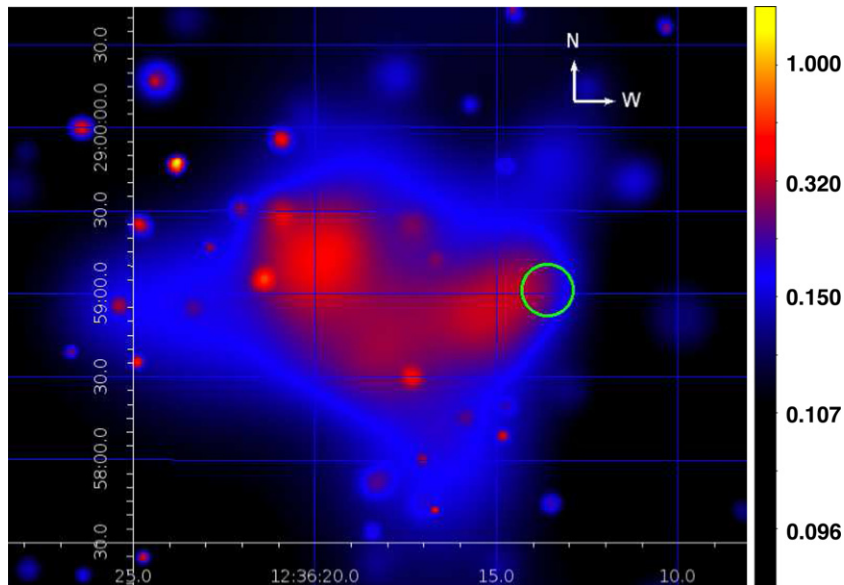


Figure 1. Adaptively smoothed *Chandra* X-ray image of ZwCl 1234.0+02916; image is $3'.4 \times 3'.0$ in size and the color bar is normalized to one. The detected point sources have been removed prior to smoothing. The point-like features seen in the image indicate the presence of slightly extended sources, or fluctuations in the background, which were exaggerated by the smoothing and exposure corrections. The green circle ($r = 9''.3$) is the 90% confidence error region of the XRT X-ray afterglow of GRB 050509B.

(A color version of this figure is available in the online journal.)

Table 4
Cluster X-Ray Substructure Properties

| Substructure | R.A. (h m s) | Decl. ($^{\circ}$ ' ") | Size (") |
|-----------------|-----------------|----------------------------|-------------|
| Eastern cluster | 12 36 20.29 | +28 59 14.74 | 17.2 |
| Western cluster | 12 36 14.75 | +28 58 55.84 | 10.8 |
| Bridge | 12 36 18.25 | +28 58 33.81 | 7.9 |

Notes. Derived using wavelet-transform analysis detailed in Section 3.2.2. The estimated uncertainties in the centroid values are 4 pixels $\approx 2''$.

the summed intensity enclosed by the area of the Mexican hat. The resulting transformed images thus filter the original signal intensity at the scale of each wavelet, producing a sequence of images over a range of physical scales. Consequently, the technique can be used to measure the size and distribution of X-ray emitting substructures.

Figure 2 shows the wavelet-transformed images of ZwCl 1234.0+02916 at five different scales. The $10''$ image shows numerous individual structures that have a similar distribution to the X-ray surface brightness contours shown in Figure 7. At larger scales ($15''$ and $20''$), the source has two identifiable emitting regions that are connected by a narrow, extended “bridge” of emission. Finally, at the largest scales, the source appears as a single emitting region. The bimodal structure evident in the transformed images supports the cluster merger interpretation of Section 3.2.1.

We calculated the centroids and sizes of the eastern and western subclusters and the bridge in the X-ray image using wavelet-transform analysis. Specifically, the sizes of the structures can be characterized by the scale at which the convolution of the wavelets and the structures are maximum; the centroids of the structures are given by the pixel locations where the transformed images have the largest values (Lopez et al. 2009). Table 4 lists their derived locations (R.A. and decl.) in the *Chandra* image and their sizes in arcseconds. We estimate the projected

separation between the eastern and western clusters (centroid to centroid) to be $73''$, corresponding to ~ 260 kpc at the cluster distance.

3.2.3. Power-ratio Method

Next, we used the power-ratio method (PRM) on the full-band (0.5–8.0 keV) ACIS image of the ZwCl 1234.0+02916 after point-source removal to quantify the X-ray morphology of ZwCl 1234.0+02916. This method was originally developed to probe the dynamical state of clusters observed with *ROSAT* (Buote & Tsai 1995, 1996) and was extended to *Chandra* by Jeltama et al. (2005). A detailed description of the method is given in Jeltama et al. (2005); below we present a qualitative description. The PRM measures asymmetries in an image using the multipole moments and the associated powers P_m of the surface brightness in a circular aperture. The higher-order terms measure asymmetries at successively smaller scales relative to the X-ray object size. We divided all powers by P_0 to form the power ratios, P_m/P_0 , to normalize with respect to flux. Morphological information is given primarily by the higher-order terms P_2/P_0 (quadrupole ratio) and P_3/P_0 (octupole ratio), which are sensitive to the source ellipticity and to deviations from mirror symmetry, respectively. Sources with high P_2/P_0 have elliptical or elongated morphologies; high P_3/P_0 ratios occur in sources with mirror asymmetry, such as bimodal structures of unequal sizes.

Figure 3 plots our derived P_3/P_0 versus P_2/P_0 using an aperture of 0.5 Mpc with the values for the clusters discussed in Jeltama et al. (2005); the ratio values are $P_3/P_0 = (19.9^{+28.8}_{-17.6}) \times 10^{-7}$ and $P_2/P_0 = (21.9^{+56.2}_{-21.0}) \times 10^{-7}$. The errors on the plot were estimated using the Jeltama et al. (2005) technique and are at the 90% confidence level. We note that the P_3/P_0 value of ZwCl 1234.0+02916 is the second largest in that group and a factor of ~ 30 higher than the average for $z < 0.5$, indicating a highly asymmetric (bimodal) structure. These results suggest that the merger in ZwCl 1234.0+02916 has occurred recently,

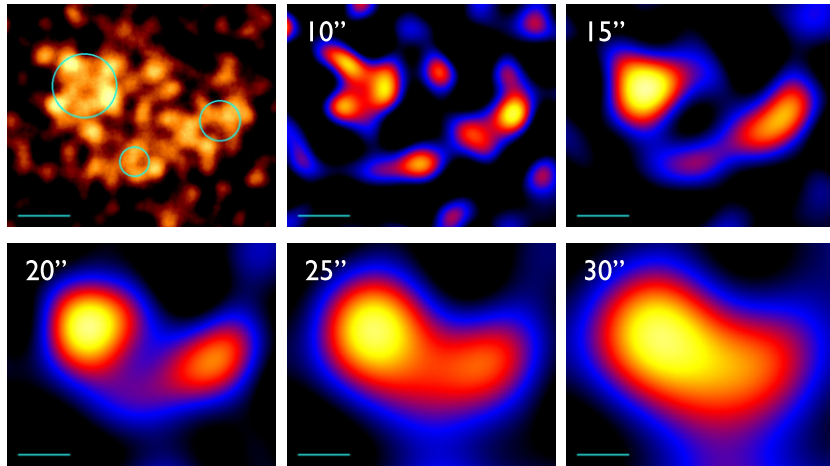


Figure 2. Wavelet-transformed images of ZwCl 1234.0+02916 at five different scales. Here east is to the left and north is up. Each image shows the emission at the scale of the wavelets. The scale bar is 100 kpc ($\approx 28''.5$) assuming an angular-size distance of 729.5 Mpc. The leftmost panel is the full X-ray band image smoothed with a $5''$ Gaussian. The three cyan circles indicate the centroids of the structures listed in Table 4. The following five panels are the wavelet-transformed images. At the smallest scale ($10''$) many structures are evident. At intermediate scales ($15''$ and $20''$), the source is characterized by two emitting regions bridged together. At the largest scales ($25''$ and $30''$), the source appears as a single emitting region.

(A color version of this figure is available in the online journal.)

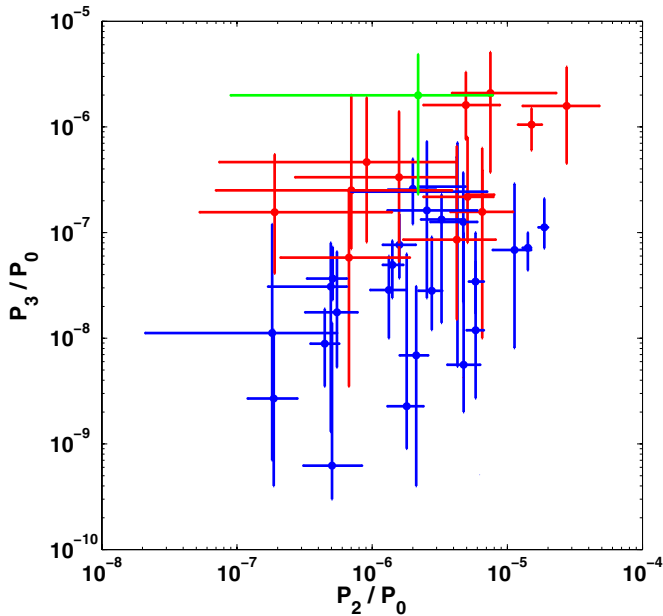


Figure 3. P_3/P_0 vs. P_2/P_0 of the clusters in the Jeltama et al. (2005) paper with the value of ZwCl 1234.0+02916 added in green. The red (blue) points are clusters with $z > 0.5$ (< 0.5), respectively.

(A color version of this figure is available in the online journal.)

as the cluster is quite asymmetric and is not yet relaxed (Jeltama et al. 2005). For example, the simulations shown in Poole et al. (2006, Figure 13) show that such large values primarily occur before second core passage.

3.3. Cluster X-Ray Spectral Analysis

We extracted a *Swift*/XRT source spectrum from a region $1'.7$ in radius centered on the cluster emission (at $12^{\text{h}}36^{\text{m}}18^{\text{s}}.848$, $+28^{\circ}59'12''.17$, J2000.0) and created the background spectrum using events from a $2'-7'$ annulus about the cluster center; all point sources were excluded from the source and the background regions. We restricted all XRT imaging and spectral analysis to events with energies between 0.3 and 10.0 keV and modeled

the cluster spectrum using a MEKAL model for X-ray emission from an optically thin plasma (Mewe et al. 1985). To justify use of the χ^2 statistic, we grouped spectral bins to obtain at least 10 counts per bin before background subtraction. We also accounted for absorption due to intervening gas and dust from the Galaxy using the Tuebingen–Boulder interstellar medium absorption model (XSPEC model *tbabs*) with the relative abundances and cross-sections from Wilms et al. (2000) and with the He cross-section from Yan et al. (1998); the inclusion of a redshifted absorption component (XSPEC model *ztbabs*) did not result in an improved fit and was not used further in the spectral analysis. We initially held the column density constant at the average Galactic value toward the cluster derived from radio measurements, $N_{\text{H}} = 1.52 \times 10^{20} \text{ cm}^{-2}$ (Dickey & Lockman 1990), and the metallicity at the mean value, $Z = 0.41$ solar, as derived from spectral fits to 38 clusters (Bonamente et al. 2006). The best-fit spectral model parameters are reported in Table 5. When we let the cluster metallicity vary, we found no statistically significant improvement in the fit and we were only able to place a 90% upper limit of $Z < 0.9$ solar. For the remainder of the XRT analysis, we therefore fixed this parameter to the fiducial value of 0.41 solar.

We collected a *Chandra*/ACIS source spectrum from a circular region $1'.7$ in radius centered on the cluster, and extracted a background spectrum from an annular region $1'.7-2'.8$ (1.5 times the source area) also centered on the cluster; all point sources were excluded from the source and the background regions. We generated a position-weighted ancillary response using the CIAO script *specextract* for each spectrum; while running this script, the energy redistribution matrix is generated using *mkacisrmf*. We analyzed the data by using the XSPEC (v12.3.1) spectral-fitting package (Arnaud 1996). To justify use of the χ^2 statistic, we grouped spectral bins to obtain at least 10 counts per bin before background subtraction. We fit the 0.5–8.0 keV energy band with the same absorbed MEKAL model that was applied to the XRT data. We first fit with N_{H} and metallicity fixed and obtained the fit described in Table 5. We then freed only the metallicity, which resulted in a best-fit abundance of $0.3^{+0.5}_{-0.3}$ solar. Next, we restored the metallicity to 0.41 solar and freed the column density, which resulted in

Table 5
Global X-Ray Cluster Spectral Properties

| | N_{H} (10^{20} cm^{-2}) | kT^{a} (keV) | Z (solar) | Unabs F_{x}^{b} ($10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2}$) | $L_{\text{x}}^{\text{b,c}}$ ($10^{43} \text{ erg s}^{-1}$) | χ^2/ν |
|----------------|---|--------------------------|----------------|--|---|--------------|
| Swift/XRT | 1.52 | $4.2_{-1.4}^{+1.7}$ | 0.41 | $1.9_{-0.4}^{+0.3}$ | $2.8_{-0.7}^{+0.4}$ | 54/41 |
| Chandra/ACIS-S | 1.52 | $2.9_{-0.7}^{+1.0}$ | 0.41 | $1.6_{-0.2}^{+0.3}$ | $2.3_{-0.3}^{+0.4}$ | 230/231 |

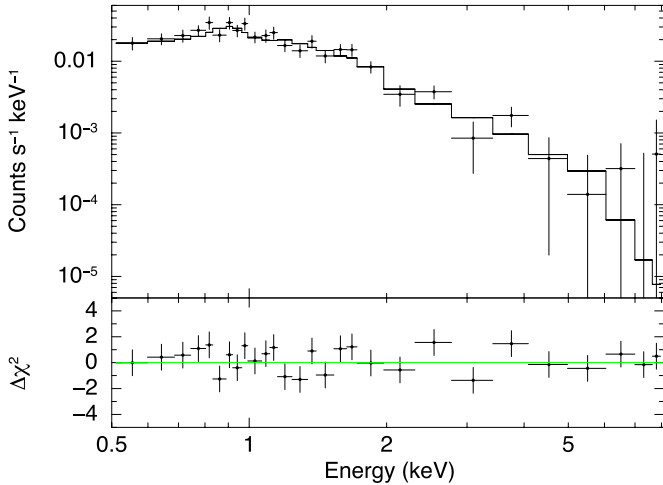
Notes.^a 90% confidence errors.^b Fluxes and luminosities are reported for the 0.5–10 keV band.^c Calculated assuming a luminosity distance of 1101.5 Mpc.

Figure 4. Chandra/ACIS-S spectrum of the cluster ZwCl 1234.0+02916. The best-fit absorbed MEKAL model and residuals are also plotted.

(A color version of this figure is available in the online journal.)

$N_{\text{H}} = 1.0_{-1.0}^{+4.5} \times 10^{20} \text{ cm}^{-2}$; in both cases, however, the change in χ^2 was not significant, and the fit parameters were therefore not listed in Table 5. Figure 4 shows the ACIS spectrum.

Finally, we measured the temperatures/luminosities of the two X-ray subclusters over the 0.5–10 keV range and we find for the eastern subcluster: $kT = 3.2_{-0.9}^{+1.6}$, Unabs $F_{\text{x}} = 0.53_{-0.09}^{+0.12} \times 10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2}$, $L_{\text{x}} = 7.8_{-1.3}^{+1.7} \times 10^{42} \text{ erg s}^{-1}$, $\chi^2/\text{dof} = 67/72$, and for the western subcluster $kT = 3.3_{-1.2}^{+3.7}$, Unabs $F_{\text{x}} = 0.17 \pm 0.02 \times 10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2}$, $L_{\text{x}} = 2.5 \pm 0.3 \times 10^{42} \text{ erg s}^{-1}$, $\chi^2/\text{dof} = 23/25$. We note that their summed flux does not equal the total X-ray luminosity, indicating that there is non-negligible flux outside the two subclusters.

4. WEAK-LENSING RESULTS FOR ZwCl 1234.0+02916

To constrain the distribution and total amount of mass in ZwCl 1234.0+02916 and to search for massive concentrations of dark matter in the other short GRB fields discussed in Section 5, we measured the weak shear caused by gravitational lensing. Such measurements were made both from the ground-based VLT images and, in the case of ZwCl 1234.0+02916, from the deep *HST*/ACS image. Our methodology, described below, follows commonly used procedures for weak-lensing measurements. For ZwCl 1234.0+02916, the results are presented in maps of the reconstructed surface mass density distribution, based on the VLT and ACS data, and in the form of estimates of the integrated cluster mass. By combining these measurements of the (dark-matter-dominated) total cluster mass with the X-ray measurements of the hot intracluster gas component

discussed in Section 3, a clear physical interpretation of this system emerges, indicating that this is a cluster merger with obvious similarities to the Bullet cluster.

4.1. VLT Data

The weak gravitational lensing measurements were performed using standard IMCAT tools,¹⁹ based on the methodology introduced by Kaiser et al. (1995) and Luppino & Kaiser (1997), which has been extensively tested against simulated lensing data (Heymans et al. 2006). The weak-lensing measurements were made independently on each of the combined images listed in Table 1, and the results for different epochs were only combined in the end, as described below in Section 4.4. For all of these images, the sky background was first set to zero by subtracting a highly smoothed image of the background level as a function of position. Object detection was performed using the IMCAT peak finder (*hfindpeaks*), followed by estimation of the local sky background level and gradient around each object (*getsky*), aperture photometry (*apphot*), and object shape estimation (*getshapes*). The latter returned the ellipticity parameters e_{α} , calculated from the weighted quadrupole moments of each object (Kaiser et al. 1995).

Any anisotropies in the PSF will cause additive errors in the gravitational shear measurements, and such effects must be estimated and removed before the weak-lensing effect is measured. For seeing-limited images, this correction is normally well approximated by the expression

$$e_{\alpha}^{\text{cor}} = e_{\alpha}^{\text{obs}} - P_{\alpha\beta}^{\text{sm}} p_{\beta}, \quad (1)$$

where the smear polarizability P^{sm} given by Hoekstra et al. (1998) is measured by IMCAT (*getshapes*), and p is a measure of the PSF anisotropy. The latter quantity was estimated from 35 to 700 non-saturated stellar objects in each VLT field (the number of available stars depended mostly on Galactic latitude). A second-order polynomial was fit to the PSF anisotropy variation across the field, and the e_{α} values were corrected for PSF anisotropy according to Equation (1) (using the IMCAT routines *efit* and *ecorrect*, respectively).

The isotropic part of the PSF causes a multiplicative error in the shear measurements, which can be corrected using the pre-seeing polarizability tensor P^{γ} introduced by Luppino & Kaiser (1997):

$$e_{\alpha}^{\text{cor}} = e_{\alpha}^{\text{s}} + P_{\alpha\beta}^{\gamma} \gamma_{\beta}, \quad (2)$$

where γ is the weak gravitational shear and P^{γ} is given by

$$P_{\alpha\beta}^{\gamma} = P_{\alpha\beta}^{\text{sh}} - P_{\alpha\mu}^{\text{sm}} (P_{\mu\sigma}^{\text{sm}\star})^{-1} P_{\delta\beta}^{\text{sh}\star}. \quad (3)$$

¹⁹ <http://www.ifa.hawaii.edu/~kaiser/imcat/man/imcat.html>

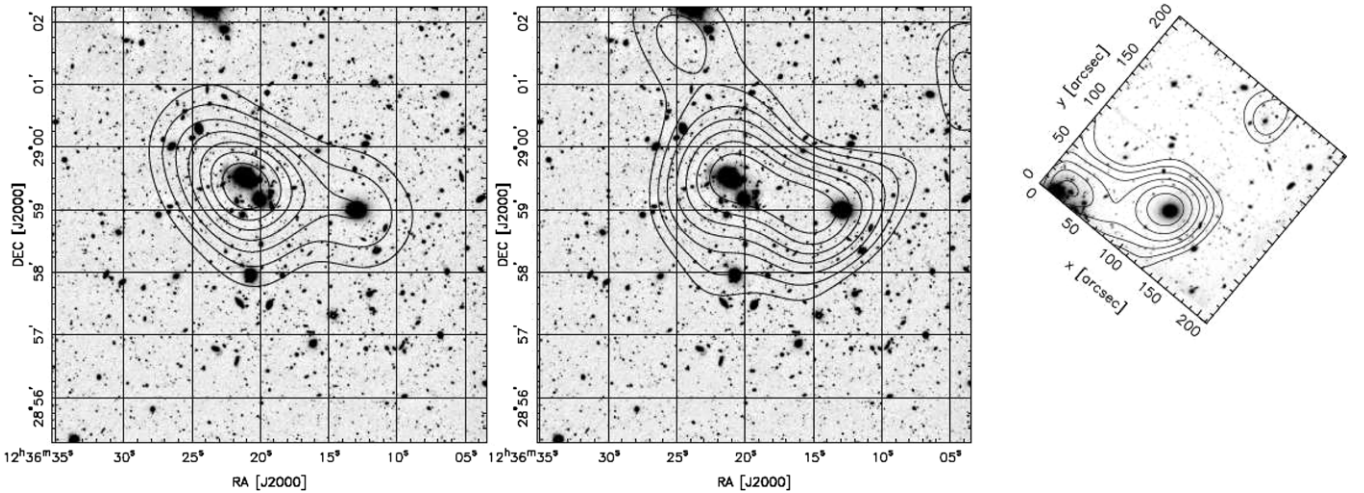


Figure 5. Distribution of light (left) and mass surface density κ derived from the VLT data (middle) and from the *HST*/*ACS* data (right) in the field of GRB 050509B; east is left and north is up. Both distributions from the VLT data have been smoothed with a Gaussian of scale $40''$. The light contours in the left panel are derived based on the light distribution of early-type galaxies that form the red sequence in a color–magnitude diagram, and the contour levels are scaled relative to the peak and average value in a similar manner as for the mass map. The contours are superposed on the 3600 s *R*-band FORS2 image (see Table 1). The *ACS* mass distribution map (right) has been smoothed with a Gaussian of scale $20''$ and rotated, scaled, and aligned with the VLT data. The density contours on the right panel are derived based on weak-lensing measurements from the *HST*/*ACS* data. The plotted contours start at $\kappa = 0.011$ (which corresponds to 1σ) and are spaced at an interval $\Delta\kappa = 0.011$. The contours are superposed on the 27,480 s combined *ACS* image.

Here, P^{sh} is the shear polarizability (from *getshapes*) given by Hoekstra et al. (1998), and asterisks denote the polarizabilities measured for stellar objects. The quantity P^γ will vary systematically as a function of the size (e.g., as measured by the radius r_g , determined by *hfindpeaks*) and magnitude of a galaxy, and the P^γ estimate will be very noisy for individual (small and faint) galaxies. Therefore, in practice these were calculated by binning galaxies in r_g -magnitude space (each bin containing typically 50 galaxies), and correcting the ellipticity of each galaxy according to the mean P^γ value of its bin. Since the P^γ tensor is close to diagonal, it is well approximated by a scalar value given by $1/2\text{Tr}(P^\gamma)$. By assuming that the net intrinsic ellipticities of galaxies average to zero, the resulting shear estimator is given by

$$\hat{\gamma}_\alpha = (P^\gamma)_{\alpha\beta}^{-1} [e_\beta^{\text{obs}} - P_{\beta\mu}^{\text{sm}} p_\mu]. \quad (4)$$

Finally, a weight was assigned to the $\hat{\gamma}_\alpha$ measurement from each galaxy, given by $\langle \gamma^2 \rangle^{-1}$ measured for galaxies in each bin in r_g -magnitude space (see above), and these weighted shear estimates were used in the further analysis. The sample of “background galaxies” used for the gravitational lensing measurements were selected on the S/N of their detection (the ν value returned by *hfindpeaks*), chosen to lie in the range $6 < \nu < 100$. The number of galaxies that passed this selection criterion is listed in the final column of Table 1, and the noise in each mass reconstruction is approximately inversely proportional to the square root of this number.

The projected matter distribution in each field was estimated from the measured shear values, using the maximum probability method introduced by Squires & Kaiser (1996). We used a regularization parameter of 0.05 and wave modes up to $k = 6$ (see Squires & Kaiser 1996 for details). This produced a 128^2 pixel image of the dimensionless surface mass density $\kappa = \Sigma/\Sigma_{\text{crit}}$, where Σ is the physical surface mass density and

$$\Sigma_{\text{crit}} = \frac{c^2}{4\pi G} \frac{D_s}{D_l D_{ls}}, \quad (5)$$

where D_l , D_s , and D_{ls} , are angular diameter distances to the lens, to the source, and between the lens and source, respectively.

Finally, each of these mass maps was smoothed with a Gaussian of scale $40''$ to dampen the small-scale noise in the mass reconstruction. The plotted κ contour levels in the right panel of Figure 5 start at $\kappa = 0.01$ and are spaced at intervals of $\Delta\kappa = 0.007$.

The significance of features in these mass maps was evaluated by generating randomized catalogues where the true object positions were kept, and an observed $\hat{\gamma}_\alpha$ value was randomly assigned to a given position. Mass maps were then calculated based on these randomized catalogues, and the resulting rms spread in dimensionless surface mass density, κ , values provided an estimate of the noise level at each position in the field, thus producing a noise map. An S/N map of each field was then generated by dividing the κ estimate by the noise estimate (Figure 5, middle panel).

4.2. *HST*/*ACS* Data

The weak-lensing measurements based on the *HST*/*ACS* data in the field of GRB 050509B generally followed the methodology described above for the VLT data, but with some minor deviations, detailed in this section.

Objects were detected in the combined F814W image using *hfindpeaks*, followed by sky estimation (*getsky*), aperture photometry (*apphot*), and shape estimation (*getshapes*). Their ellipticities were corrected for the effect of PSF anisotropy by fitting a second-order polynomial to the PSF variation across the field, based on 22 non-saturated stars (see Figure 6 for an illustration of the effect of this correction). Objects in a $\sim 2''$ wide region at the border of each *ACS* chip image (where cosmic-ray rejection was imperfect and the noise level was elevated because of dithering) were masked out. The IMCAT object finder tended to return multiple detections inside some of the brighter galaxies (e.g., from star-forming regions in spiral arms). Such detections were efficiently removed from the final background galaxy catalog, by excluding all sources which had a brighter neighbor within $1''$.

The area covered by the *ACS* image was 11 arcmin^2 ; ~ 4400 objects (detected at $6 < \nu < 250$) were used for the weak-lensing analysis, yielding a background galaxy density of

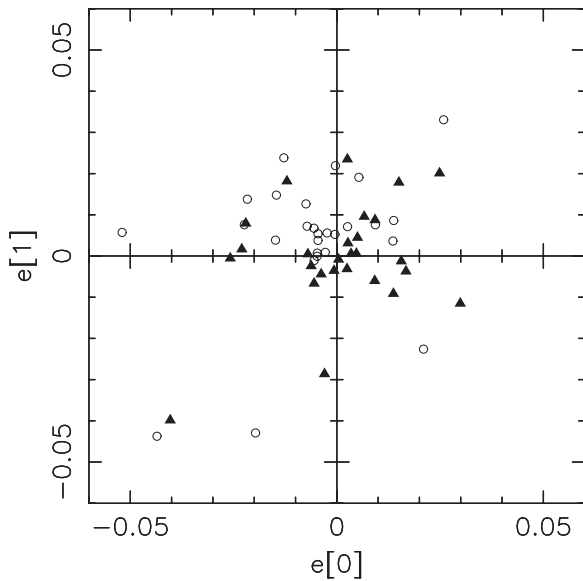


Figure 6. Stellar ellipticities in the ACS image of the field of GRB 050509B prior to (open circles) and after correction for PSF anisotropy (triangles).

400 objects arcmin^{-2} , compared to 40–50 objects arcmin^{-2} for the VLT data. Hence, weak-lensing mass measurements from the *HST*/ACS data are approximately $\sqrt{400/45} \simeq 3$ times more sensitive than the ones from VLT. The background galaxy samples generated from the VLT and *HST* data inevitably have some contamination from cluster galaxies; we describe below in Section 4.3 how this effect was corrected for when making quantitative cluster mass estimates.

We note that for a typical effective background galaxy redshift $z_{\text{bg}} \sim 1$, the critical surface mass density given in Equation (5) will be close to its minimum value at the redshift of ZwCl 1234.0+02916, making the gravitational lensing effect maximally sensitive to massive structures at this distance. In addition, a background object surface density of 400 objects arcmin^{-2} (of which ~ 270 objects arcmin^{-2} were detected at $\nu > 10$) is unprecedented for weak gravitational lensing studies of clusters. Within the area covered by the ACS field, our cluster mass reconstruction constitutes the most sensitive weak-lensing-based mass measurements thus far for any galaxy cluster.

The higher background galaxy density of the ACS data also enables a mass reconstruction at a finer resolution than from the VLT data for ZwCl 1234.0+02916. However, the mass distribution near the edges and corners of the ACS field is poorly constrained, given the lack of information about gravitational shear outside the ACS field of view. To overcome this problem, information on the shear external to the ACS field from the wider-field VLT data was incorporated into our analysis. The combination of *HST*/ACS-based shear measurements in the smaller area of the cluster core with ground-based data to assess the weak-lensing signal over a larger field is a technique used previously by Clowe et al. (2006) for the Bullet cluster. Even at the extreme corners of the ACS field, the *HST* data are the dominant contributor to the derived surface mass density estimate. The smoothing scale of $20''$ ensures that the contribution of the VLT data to the mass map becomes negligible within the ACS field.

The projected matter distribution was reconstructed from the measured shear values, using the maximum probability method (Squires & Kaiser 1996) with a regularization parameter of 0.05

and wave modes up to $k = 12$. This produced a 128^2 pixel image of the dimensionless surface mass density, κ , over the entire field covered by the VLT data, which was smoothed with a Gaussian of scale $20''$ (twice the resolution of the VLT-based mass maps derived in Section 4.1). The subarea of this mass map, which overlaps the area observed by ACS, is displayed as contours overlaid on the F814W-band ACS image in the right panel of Figure 5. The ACS image only covers about a quarter of the area of the FORS2 field, and the ACS field is rotated about 133° counterclockwise with respect to the principal north–south direction, such that the easternmost peak in the mass and light distribution is located at the corner of the field; the right panel of Figure 5 shows the ACS data aligned to the VLT image. The peak S/N value and the rms noise of this mass map is 6.0σ and $\kappa_{\text{rms}} = 0.011$, respectively, where the κ_{rms} value was calculated in a similar way as for the VLT data, as described in Section 4.1.

4.3. Cluster Mass Estimate

An estimate of the cluster mass can be derived from gravitational lensing measurements, e.g., by fitting the observed one-dimensional shear profile $g_T(r)$ to the predicted shear as a function of cluster-centric radius for a given spherically symmetric theoretical mass model, such as a Navarro–Frenk–White (NFW)-type mass density profile (Navarro et al. 1997). With two clearly separated components in its mass, light, and X-ray luminous gas distribution, ZwCl 1234.0+02916 is poorly represented by a single, monolithic structure, at least on scales smaller than or equal to the subcluster separation. However, by using only shear measurements at large radii that encompass both mass concentrations, a fair estimate of the total cluster mass can be made.

The galaxy shape distortions caused by gravitational lensing provide information on the reduced tangential shear, $g_T = \gamma_T / (1 - \kappa)$, where γ_T is the tangential component of the shear. The reduced g_T was measured in a set of non-overlapping annuli, each with a mean radius r centered on the optical cluster center (i.e., the highest peak in the contours displayed in the left panel of Figure 5). For the NFW model, the three-dimensional mass density has a radial dependence of the form, $\rho(r) \propto [(r/r_s)(1 + r/r_s)^2]^{-1}$, specified by a scale radius r_s and a concentration parameter $c_{\text{vir}} = r_{\text{vir}}/r_s$, where r_{vir} is the virial radius of the cluster. N -body simulations of the dark matter distribution in clusters in a Λ CDM universe give predictions for the dependency of c_{vir} on cluster mass and redshift (Bullock et al. 2001; Dolag et al. 2004). The mass of ZwCl 1234.0+02916 predicted from the X-ray-determined mass–temperature relation of Arnaud et al. (2005) is $M_{500c} \simeq 2 \times 10^{14} M_\odot$, where M_{500c} denotes the mass enclosed within a three-dimensional volume bounded by r_{500c} . Here, r_{500c} is the radius within which the mean cluster density is 500 times the critical density of the universe at the redshift of the cluster. Assuming the radial mass density distribution of the NFW mass model, this mass translates into $M_{180c} \simeq 3 \times 10^{14} M_\odot$, for which Bullock et al. (2001) predict a median halo concentration of $c_{\text{vir}} = r_{\text{vir}}/r_s = 1.14r_{180c}/r_s = 6.8/(1 + z)$ in a concordance model universe (here, M_{180c} and r_{180c} are defined in a similar manner as M_{500c} and r_{500c}). When making the fit, we fix the concentration parameter at this value. The lensing properties of the NFW model have been calculated by Bartelmann (1996) and Wright & Brainerd (2000).

As noted above, the shear measurements used for the NFW fit were only made at large cluster-centric radii ($133'' < r < 210''$), encompassing both subclusters of ZwCl 1234.0+02916. Only the VLT data were used for this fit, as they provided full

azimuthal coverage at all radii, unlike the smaller-field ACS data. Our photometric data in two pass-bands did not allow reliable discrimination between cluster galaxies and lensed background galaxies except for the small fraction of galaxies with observed $V-R$ colors redder than early-type cluster galaxies in ZwCl 1234.0+02916. The former must be at higher redshifts than the cluster galaxies to attain such red $V-R$ values for any normal galaxy spectral energy distribution (SED). The faint galaxy catalogs which were used to measure the gravitational shear are thus significantly contaminated by cluster galaxies that dilute the lensing signal by an amount given by their local sky density. To correct for this contamination, a radially dependent correction factor was applied to the shear, based on the estimation of the magnitude of this effect for a sample of clusters of similar richnesses and redshifts by Pedersen & Dahle (2007).

The distances to the background galaxies, entering into Equation (5), need to be known to convert the observed lensing signal into a cluster mass estimate in physical units. The background galaxy redshifts were estimated from spectroscopic and photometric redshifts in the Hubble Deep Field (for details, see Dahle et al. 2002). The average value of the ratio between the lens-source and observer-source angular diameter distances, $\beta \equiv D_{ls}/D_s$, was calculated given the weights assigned to shear measurement from each galaxy; we estimated a mean $\langle\beta\rangle = 0.694$ for the VLT data of ZwCl 1234.0+02916.

Our NFW fit yields a mass $M_{180c} = (4.08 \pm 1.56) \times 10^{14} M_{\odot}$ and $r_{180} = 1.44 \pm 0.20$ Mpc, or $M_{500c} = (2.81 \pm 1.07) \times 10^{14} M_{\odot}$ and $r_{500} = 0.91 \pm 0.13$ Mpc. The latter radius corresponds to 4.2 on the sky, i.e., the 6.9×6.9 FORS2 field virtually covers the entire cluster volume enclosed by r_{500} . Our cluster mass estimate is consistent with the value $M_{500c} = 2.2_{-0.6}^{+3.3} \times 10^{14} M_{\odot}$ derived by Pedersen et al. (2005), based on their analysis of the *Swift*/XRT data combined with the X-ray determined mass-temperature relation of Arnaud et al. (2005). Our mass is also consistent with the value $M_{500c} = (2.01 \pm 0.67) \times 10^{14} M_{\odot}$ derived from the *Chandra*/ACIS-S temperature measurement given in Table 5, using the X-ray-derived mass-temperature relationship of Vikhlinin et al. (2006; which is very similar to the Arnaud et al. 2005 relationship). A somewhat higher, but still consistent, value $M_{500c} = (3.14 \pm 1.42) \times 10^{14} M_{\odot}$ is estimated using the *Chandra* temperature and the weak-lensing-based mass-temperature relationship derived by Pedersen & Dahle (2007) (assuming a fixed slope $\alpha = 1.5$). Finally, our lensing-derived mass estimate implies that ZwCl 1234.0+02916 has about 25% the mass of the Bullet cluster (Clowe et al. 2006; Bradač et al. 2006).

We also made individual mass estimates of the two subclusters. This was done by measuring the projected two-dimensional masses using the ‘‘aperture densitometry’’ statistic of Fahlman et al. (1994). This statistic measures the mean projected density within an aperture, minus the mean density in a surrounding annulus, which provides a lower bound on the mass contained within the aperture. At the larger radii employed for our NFW model fit described above, it is impossible to unambiguously disentangle the relative mass contributions of the western and eastern subclusters. Hence, we choose an aperture of $28''$ ($=100$ kpc at the redshift of the cluster) and measured the mean reduced tangential shear within two annuli defined by ($28'' < r < 75''$); the first was centered on the GRB 050509B host galaxy and the second on the global maximum of the cluster light distribution, associated with the eastern subcluster (see Figure 5). The outer radius of $75''$ was chosen to be smaller

than the $\simeq 100''$ separation between the two clusters, in order to distinguish the mass of each subcluster. For the western subcluster containing GRB 050509B, we derive a mass estimate $M_{\text{ap}}(<100 \text{ kpc}) = (1.27 \pm 0.26) \times 10^{13} M_{\odot}$ from weak-lensing measurements based on the ACS data. The eastern subcluster is not adequately covered by the superior ACS data, and we therefore used the VLT data to similarly derive a mass $M_{\text{ap}}(<100 \text{ kpc}) = (2.06 \pm 1.27) \times 10^{13} M_{\odot}$ for this component. Although the uncertainty of the VLT-based mass estimate of the eastern subclump is fairly large, these values are consistent with the picture provided by the observed distributions of X-ray luminosity and stellar mass (see the following section), which both indicate that the eastern subcluster is the more massive of the two subclusters.

4.4. GRB 050509B and Host Cluster Results

We discuss below the possible association of GRB 050509B and ZwCl 1234.0+02916. The *Swift*/XRT error circle for GRB 050509B was centered in the outskirts of a giant elliptical galaxy (2MASX J12361286+2858580) within the cluster ZwCl 1234.0+02916 ($z = 0.2214$). As detailed above, we were able to recover a significant weak gravitational lensing signal from this cluster, using both ground- and space-based imaging data.

All the mass maps derived for this field, based on the four different combined VLT images (see Table 1), show evidence for a binary structure in the mass distribution, with one mass peak close to the optical light center of the cluster, and the other close to the giant elliptical galaxy which is the apparent host of GRB 050509B. The peak levels in these mass maps correspond to $\sim 4\sigma$. To further improve the accuracy of the mass reconstruction, the four different mass maps were averaged, and the resulting mass contours are plotted in the middle panel of Figure 5. The left panel in the figure shows the distribution of light from early-type galaxies inside the cluster, smoothed on the same scale as the mass distribution. This ‘‘light map’’ was derived by identifying the tight red sequence formed by early-type cluster galaxies in a $V-R$ versus R color-magnitude diagram, and selecting all galaxies within an interval of width $\Delta(V-R) = 0.25$, centered on the red sequence. The light map shows a binary structure similar to the mass map with the eastern peak more dominant, although the total masses of the two subclusters are similar as we discuss below. This indicates that the mass-to-light ratio is somewhat higher for the western peak, which coincides with the location of GRB 050509B. While the galaxy light from the eastern peak is dominated by two early-type galaxies (2MASX J12362010+2859080 and 2MASX J12362094+2859290), the light of the western peak is dominated by the apparent host of GRB 050509B.

Based on the 2MASS *JHK_s* photometry of these three galaxies, we estimate the stellar mass of the two subpeaks as $(2.5 \pm 1.2) \times 10^{11} M_{\odot}$ and $(5.2 \pm 1.9) \times 10^{11} M_{\odot}$ for the western and eastern peaks, respectively. These estimates were based on modeling of the SED, using the SED models of elliptical galaxies from Silva et al. (1998) for stellar populations ranging in age from 1.5 Gyr to 13 Gyr. The systematic and random errors of the stellar mass determinations from such SED modeling are estimated to be approximately a factor of two (Bell et al. 2007; Michałowski et al. 2008, 2010).

The mass map recovered from the ACS data (see Figure 5, right panel) shows a significant mass peak (6σ) centered on the GRB 050509B host galaxy, which coincides spatially

with the position of GRB 050509B (within the respective errors of the GRB 050509B position and the mass peak centroid). This mass peak has an elongation at a position angle similar to the GRB 050509B host galaxy, pointing toward the easternmost peak. Together, the weak-lensing mass reconstructions from the ACS data in Figure 5 (right panel) and from the VLT data in Figure 5 (middle panel) reveal a consistent picture of two subclumps of similar mass, separated by $\simeq 100''$. The morphology and position of the two mass peaks (particularly the western peak, which is located well within the ACS field) are most firmly constrained by the ACS data, given the significantly higher background galaxy density in the space-based images. We discuss the light, mass, and gas peak separations in detail in Section 6.

5. WEAK-LENSING SURVEY OF OTHER SHORT GRBs

The serendipitous discovery of a merger in the field of GRB 050509B motivated a search in the public data of other short GRBs for similar mergers. Establishing their presence in more than one GRB field would prove extremely significant for our understanding of the origin of the short GRB phenomenon. We therefore selected all public VLT data (up to the end of this study) taken during follow-up observations of well-localized short GRBs, resulting in five additional event fields. Using the same weak-lensing technique described above (Section 4.1), we estimated the mass concentration along their lines of sight in the archival VLT images. The results of our analysis are described below.

5.1. GRB 050724

The host galaxy of this burst has a measured redshift of $z = 0.258$ (Prochaska et al. 2006). Spectroscopy of other field galaxies reveal an apparent galaxy overdensity at $z = 0.30$, but no clustering of galaxies was seen in redshift space at the burst redshift (Berger et al. 2007). The mass map of this field does not show any significant peaks at $>2\sigma$. The field is at low Galactic latitude, with a Galactic extinction $E(B - V) = 0.59$, estimated from the dust maps of Schlegel et al. (1998). This produces ~ 2 mag of extinction in the V band and 1.1 mag of extinction in the I band. In addition, the field is heavily crowded by foreground stars, which further reduces the number of useful background galaxies. Hence, a very massive ($M \sim 10^{15} M_{\odot}$) GRB host cluster would be needed to produce a detectable lensing signal in this field. The absence of such a signal is consistent with the non-detection of a host cluster by Berger et al. (2007).

5.2. GRB 050813

The mass reconstruction of this field does not reveal any significant peaks at $>2\sigma$. The field was observed during poor conditions (clouds), and the high background sky level results in a somewhat reduced number of detected galaxies (see Table 1), and, consequently, a higher noise level in the mass reconstruction.

5.3. GRB 050906

Levan et al. (2008) reported a $z = 0.43$ galaxy cluster within the positional error circle of GRB 050906. Two mass peaks are visible in this field, a northern peak, visible at the $\sim 3\sigma$ level, and a southern peak at the 2σ – 3σ level. These peaks are at a level where they could easily be caused by random fluctuations in the orientations and ellipticities of the background galaxies, and hence, we are unable to report a firm detection of any mass concentration in this field.

5.4. GRB 050911

This GRB has a rather large error circle of radius $2'.8$ (Page et al. 2006) which coincides with the known cluster EDCC 493. Berger et al. (2007) measure a redshift $z \sim 0.1646$ and velocity dispersion $\sigma_v = 660_{-95}^{+135}$ kms $^{-1}$ for this cluster. They also estimate a mass of $M_{500c} \simeq 2.5 \times 10^{13} M_{\odot}$ for EDCC 493, based on an estimate of the X-ray temperature from diffuse emission seen in the *Swift*/XRT data, combined with the mass–temperature relation of Arnaud et al. (2005). This is an order of magnitude lower than the estimated mass of ZwCl 1234.0+02916, the host cluster of GRB 050509B, suggesting that the mass of EDCC 493 would be too small for a firm detection via weak lensing.

We did recover one peak in the mass map of this GRB field at the $\sim 3.5\sigma$ level; the rough structure of the mass map bears some similarity to the distribution of brighter galaxies in this field. However, the highest mass peak we recovered is offset from the peak of the X-ray emission (which is spatially coincident with the brightest cluster galaxy of EDCC 493) by $\sim 3'$. In addition, this mass peak is located in a region of fairly high projected galaxy density, and the redshifts of the brightest galaxies in this region show a wide scatter (Berger et al. 2007), suggesting that the mass peak, if real, may simply be a projection effect of unrelated structures over a wide range of redshifts. A second, lower peak (at $\sim 2.5\sigma$) is seen closer to the position of the X-ray and optical center of EDCC 493. The lensing detection of mass is thus at best at a marginal level in this field, and so we cannot firmly conclude that we have detected any mass overdensity associated with GRB 050911. It is, however, worth noting that this is the GRB field with the second most significant mass detection after GRB 050509B and also the only other field known to contain a galaxy cluster.

5.5. GRB 060313

The mass map of this field does not reveal any significant peaks at $>2\sigma$. The lensing analysis was based on a combined image from all available FORS1 and FORS 2 R -band data with a total exposure time of almost 5 hr.

6. DISCUSSION

Here, we discuss some implications from our X-ray and optical observations of the cluster ZwCl 1234.0+02916. In particular, we consider the implications of the association of this cluster and its galaxy environment for models of the origin of short GRBs. We also discuss the dynamical state of the cluster, and suggest that, with further observations, it might be a useful object to help constrain the properties of dark matter.

6.1. GRB 050509B–2MASX J12361286+2858580 Association

The association between GRB 050509B and 2MASX J12361286+2858580 is difficult to understand if the GRB resulted from any mechanism involving massive stars and recent star formation (Lee & Ramirez-Ruiz 2007), such as a collapsar (Woosley 1993). Such an association would instead suggest that there is at least a non-negligible fraction of short GRB progenitors with lifetimes greater than about 7 Gyr, which can outburst within a Hubble time. Among early-type galaxies approximately 50% of the stellar mass content is in galaxies with $M > 10^{11} M_{\odot}$ that typically reside in clusters (Zheng & Ramirez-Ruiz 2007). Since it is likely that galaxies in clusters shut off their star formation process early on, a long progenitor lifetime further supports

the tendency for short GRBs to happen in cluster galaxies. Indeed, 2MASX J12361286+2858580 shows no indications of UV emission or optical emission lines (Prochaska et al. 2005). The optical spectrum of the galaxy is just that of a typical giant elliptical, with no evidence for star formation within the last several Gyr.

On the other hand, 2MASX J12361286+2858580 is a very propitious site for the merger of neutron stars or any other mechanism involving an old, close binary containing at least one compact object. Although it is not possible to detect radio pulsars at the distances of even nearby giant elliptical galaxies, *Chandra* observations have shown that these galaxies have very large populations of low-mass X-ray binaries (LMXBs) containing accreting neutron stars and black holes (Sarazin et al. 2000; Angelini et al. 2001). This is particularly true of central dominant cluster galaxies. These systems are the likely progenitors to recycled millisecond pulsars. In giant ellipticals (particularly cDs), a high fraction ($\gtrsim 50\%$) of the LMXBs are located in globular clusters (GCs; e.g., Sarazin et al. 2003). This presumably occurs because close binary systems containing compact objects can be formed dynamically in GCs (Hut & Verbunt 1983; Grindlay et al. 2006; Lee et al. 2010). While there is less direct evidence that close double neutron star binaries can form easily in GCs, the double neutron star system PSR B2127+11C in the Galactic GC M15 (Anderson et al. 1990) is an example of such a system, and has a sufficiently short gravitational radiation lifetime to merge in $\sim 2 \times 10^8$ yr. The large bulge luminosities and very high specific frequencies of GCs in giant ellipticals and particularly cDs can explain their large populations of LMXBs. Thus, cluster-dominant giant elliptical galaxies like 2MASX J12361286+2858580 are very good locations for neutron star–neutron star or neutron star–black hole mergers (Zemp et al. 2009).

If GRB 050509B is due to a neutron star–neutron star merger (or other close binary mechanism) in 2MASX J12361286+2858580, then there is reasonable probability that it occurred in a GC in this galaxy (Lee et al. 2010). GCs have a broader radial distribution in elliptical galaxies than field stars (e.g., Harris 1991). Note that GRB 050509B is probably located in the outer regions of the galaxy, which might be more likely if it came from a binary merger progenitor. LMXBs tend to be associated preferentially with the more luminous GCs in a galaxy (Sarazin et al. 2003). If we assume an absolute magnitude $M_1 \gtrsim -11$ for the GC, its apparent magnitude at the distance of 2MASX J12361286+2858580 would be $m_1 \gtrsim 29$. Such a GC might be detectable as a nonvariable, red, point-like optical counterpart to GRB 050509B in our deep *HST* image. However, given the large positional uncertainty for the GRB, it would be very difficult to establish a reliable association at this magnitude level.

Assuming the association between GRB 050509B and 2MASX J12361286+2858580 is correct, then GRB 050509B occurred in a much more diffuse gaseous environment than GRBs in star forming regions of galaxies. In general, elliptical galaxies have very low density interstellar gas. Our *Chandra* image of 2MASX J12361286+2858580 and the cluster ZwCl 1234.0+02916 (Figure 1) shows that GRB 050509B probably occurred in a region dominated by intracluster gas, with a low ambient density likely effecting the strength and nature of the afterglow (Panaitescu et al. 2001). If the burst occurred instead in a GC, the resulting afterglow could then at least in part be due to the interaction of the relativistic ejecta with the stellar winds of the red giant cluster members. Due to the large

stellar density in the star cluster core, the external shock would then take place within a denser medium than the intergalactic medium (Lee et al. 2010).

6.2. Dark Matter in ZwCl 1234.0+02916

Our adaptively smoothed *Chandra* image (Figure 1), our wavelet transform images (Figure 2), and our power-ratio results (Figure 3) all show that the X-ray surface brightness distribution of the cluster ZwCl 1234.0+02916 is bimodal. The most luminous clump of X-ray emission is centered (Table 4) near the brightest cluster galaxy, 2MASX J12362094+2859290, which is about a half a magnitude brighter than the host of GRB 050509B, while the second clump of X-ray emission is centered just slightly to the east of the GRB host galaxy, 2MASX J12361286+2858580. The distribution of the galaxy light in the cluster shows a similar bimodal structure (left panel in Figure 5), again with the eastern subcluster containing more galaxies, galactic light, and stellar mass (see Section 4.4). Our VLT weak-lensing map of the mass surface density (middle panel in Figure 5) shows that the mass is also distributed with a bimodal distribution, although it appears to be somewhat more symmetric than the X-ray gas or galaxies. Moreover, our *HST/ACS* weak-lensing map (Figure 5, right panel) indicates that the bulk of the mass in the western subcluster is centered on the GRB host galaxy 2MASX J12361286+2858580, while the X-ray emission is displaced to the east.

Figure 7 shows the X-ray (gas), light, and mass contours overplotted on the *HST/ACS* field. The X-ray contours of the western subcluster associated with the GRB host galaxy 2MASX J12361286+2858580 shows an E–W elongated distribution of emission, with a drop in the X-ray surface brightness to the west of the X-ray peak and to the east of the galaxy. The galaxy itself is in a relatively faint region of cluster X-ray emission. This morphology suggests that the X-ray gas might be in a “cold front” (Markevitch et al. 2000; Vikhlinin et al. 2001). A possible interpretation is that the western X-ray clump was originally a cool core centered on the galaxy 2MASX J12361286+2858580, and that this galaxy was at the center of one of the two subclusters. In this model, the X-ray gas has been displaced by ram pressure due to the gas in the main subcluster.

This geometry may be similar to the merger in the Bullet cluster, 1E0657–558 (Markevitch et al. 2002), where the X-ray emission shows a cold front and a bow shock associated with the merging subcluster; it is somewhat different than the geometry in MACS J0025+4122 (Bradač et al. 2008) where there is only one gas (X-ray) peak between the light and mass peaks. Here and in the Bullet cluster, the X-ray gas is displaced to the east of the subcluster dominant galaxy and other galaxies associated with the subcluster. In all cases, lensing measurements indicate that the total mass and dark matter are centered on the galaxies in the subclusters, and do not follow the X-ray gas centroids.

We have used the wavelet transform results described in Section 3.2.2 to quantify the relative separation of the X-ray and mass centroids in both subclusters of ZwCl 1234.0+02916. We identified two X-ray and three mass centroids; Table 6 lists the coordinates of the mass centroids using the letter designation shown in Figure 7 and Table 4 gives the coordinates of the X-ray centroids. The separation between the two eastern mass centroids (*A*, *B*) and the eastern X-ray centroid (*C*) is $\delta S = 19''.0$ (68 kpc) and $50''.9$ (182 kpc), respectively, while the separation between the western mass centroid (*D*) and the western X-ray centroid (*E*) is $\delta S = 23''.4$ (84 kpc). We note that these

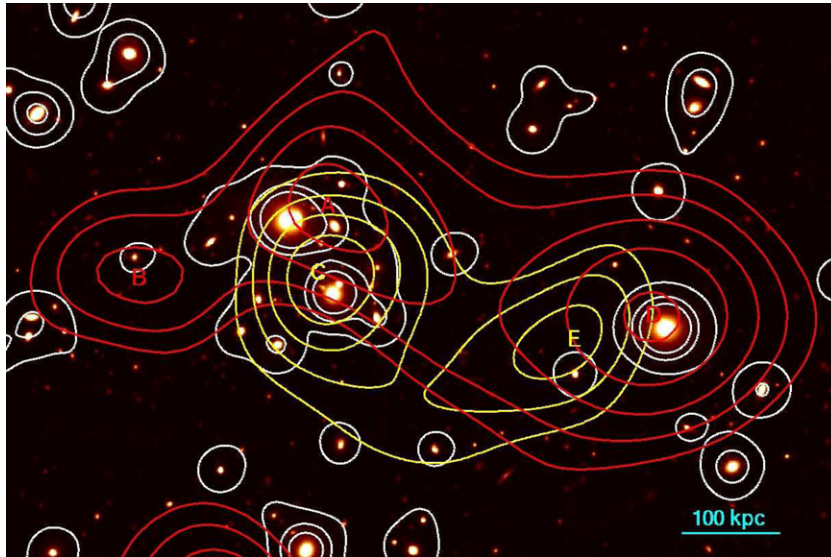


Figure 7. Distribution of light (white), mass surface density κ (red), and gas (yellow) in the field of GRB 050509B; east is left and north is up. The mass density contours are derived based on weak-lensing measurements from the VLT and *HST*/ACS data. The light contours are derived based on the light distribution of early-type galaxies that form the red sequence in a color–magnitude diagram, and the contour levels are scaled relative to the peak and average value in a similar manner as for the mass map. The gas (X-ray) contours are from the wavelet transform image filtered with a $20''$ Mexican hat. The contours are superposed on the 3600 s R-band FORS2 image (see Table 1). Letters A through E correspond to the peaks of these distributions as described in the text.

(A color version of this figure is available in the online journal.)

Table 6
Centroid Locations of the Cluster Main Mass Substructures

| Substructure | R.A. (h m s) | Decl. ($^{\circ}$ ' ") |
|--------------------|-----------------|----------------------------|
| Eastern cluster A* | 12 36 20.094 | +28 59 33.62 |
| Eastern cluster B | 12 36 24.162 | +28 59 13.24 |
| Western cluster D | 12 36 13.043 | +28 59 02.64 |

Notes. *We follow the same letter designation as shown in Figure 7. The estimated uncertainties in the centroid values are $\approx 3''$.

separations are much larger than the uncertainties in the X-ray and mass centroids ($2''$ and $3''$, respectively).

As noted in the case of the Bullet cluster (Clowe et al. 2004; Markevitch et al. 2004; Bradač et al. 2006; Clowe et al. 2006, 2007) and MACS J0025.4+1222 (Bradač et al. 2008) such systems provide a number of very strong tests of the existence and nature of non-baryonic dark matter. More recently, six additional merging clusters have been found with separated dark matter and baryonic components (e.g., Dawson et al. 2012 and references therein). The fact that the surface density of the total mass and of the dark matter aligns with the galaxy distribution is not in accord with the expectations of theories in which dark matter is actually the result of a different theory of gravity, but all of the mass is in baryonic material. In the cluster ZwCl 1234.0+02916, as in most rich clusters, the overwhelming majority of the baryonic matter is in the hot intracluster gas. Thus, in any theory in which baryons are the primary source of gravity, the gravitational potential and the apparent dark matter distribution should follow the gas in a cluster, and not the galaxies. The lensing observations of all cluster mergers thus far require the presence of non-baryonic dark matter.

The second test of the nature of dark matter comes from the agreement between the location of the centroid of the dark matter and that of the galaxies (Randall et al. 2008). The galaxies form a collisionless fluid, which is why they are not subject to the same ram pressure as the X-ray gas. The distribution of the dark

matter indicates that it also must be a nearly collisionless fluid. In the Bullet cluster, the implied upper limit for the dark matter self-interaction cross-section (Markevitch et al. 2004; Randall et al. 2008) is considerably smaller than the values required by models in which dwarf galaxy mass distributions are the result of cores in the dark matter due to self-interaction. Unfortunately, we cannot derive the same limits for ZwCl 1234.0+02916 since our X-ray observations are not deep enough to allow the detailed surface brightness profile and spectra in the possible cold front region of the western subcluster to be determined. Deeper imaging with *Chandra* and spectroscopy (with *Chandra* or *XMM-Newton*) are needed to determine the parameters of this merger, and allow ZwCl 1234.0+02916 to be used as a test of dark matter.

7. CONCLUSIONS

Since GRB 050509B was the first short GRB for which an X-ray afterglow was found, its putative association with ZwCl 1234.0+02916 motivated multiple follow-up observations in all wavelengths. Our *Chandra* observations revealed the merger nature of the cluster, which in turn provided the impulse for weak gravitational lensing measurements pointing to the effects of dark matter. The X-ray image analysis using the power-ratio technique placed ZwCl 1234.0+02916 among the most irregular of the clusters at the same redshift range and among clusters of $z > 0.5$, indicating that the merger is likely very recent (not relaxed) and highly disturbed. The centroids of light and mass in the cluster coincide and are offset from the gas centroids. These offsets range between 70 and 180 kpc strongly supporting the collisionless nature of dark matter. ZwCl 1234.0+02916 is now among the less than a dozen of dissociative systems providing a strong test of the nature of dark matter. However, ZwCl 1234.0+02916 is less massive and fainter in X-rays. A much longer dedicated *Chandra* observation is needed to determine the spectral properties of the gas accurately, and derive the merger velocity and kinematics for this

system. Since ZwCl 1234.0+02916 is at a relatively low redshift ($z = 0.22$), a deeper *Chandra* and/or *XMM-Newton* image might also detect a bow shock associated with the merger, as was seen in the Bullet cluster.

We thank A. von der Linden for her detailed and critical remarks on the paper. H.D. acknowledges support from the Research Council of Norway, including a postdoctoral research fellowship. C.L.S. was supported by the National Aeronautics and Space Administration through *Chandra* awards GO7–8129X, GO8–9085X, GO9–0135X, GO9–0148X, and GO1–12169X, and through *HST* awards HST–GO–10582.02–A, HST–GO–10835.01–A, HST–GO–11679.01, HST–GO–12012.02–A, and HST–GO–12202.0–A. Support for L.A.L. was provided by NASA through the Einstein Fellowship Program, grant PF1–120085; L.A.L. also acknowledges support by the Pappalardo Fellowship in Physics at MIT. C.K. thanks the Dark Cosmology Center, who provided the nourishing environment for the final stages of this paper. A.J.vdH. was supported by an appointment to the NASA Postdoctoral Program at the MSFC, administered by Oak Ridge Associated Universities through a contract with NASA. The Dark Cosmology Center is funded by the Danish National Research Foundation. The WSRT is operated by ASTRON (Netherlands Institute for Radio Astronomy) with support from the Netherlands foundation for Scientific Research. Some of the results were based on observations made with ESO Telescopes at the La Silla or Paranal Observatories under program IDs 075.D-0261, 075.D-0415, 075.D-0468, 075.D-0787, 076.D-0612.

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