## **PFC/JA-82-22**

# Fully Ionized and Total Abundances in the Alcator-C Tokamak

*R. Petrasso. F.H. Seguin, and N.G. Loter* American Science and Engineering, Inc. Cambridge, MA **02139**

*E. Marmar and J. Rice* Massachusetts Institute of Technology, Cambridge, MA **02139**

# **Fully Ionized and Total Silicon Abundances in the Alcator-C Tokamak**

**R. Petrasso, F. H.** Seguin, and **N. G.** Loter

*A merican Science and Engineering, Inc., Cambridge, Massachusetts 02139*

#### and

**E. Marmar and J. Rice**

*Massachusetts Institute of Technology, Cambridge, Massachusetts 02139* (Received **17** September **1982)**

With use of x-ray imaging arrays, spatially and temporally resolved measurements were made of fully ionized and total absolute Si  $(Z = 14)$  densities following injection of Si into Alcator **C.** The fully stripped ions were detected through x rays resulting from ground-state radiative recombination. Central ion densities were found to be close to coronal equilibrium values. In addition, radial profiles of **fully** Ionized and H-like **Si** showed large fluctuations during internal disruptions.

**PACS** numbers: 52.25.Fi, 32.30.Rj, **52.25.Gj,** 52.70.-m

Impurity transport in tokamak plasmas is an important process, since it greatly affects plasma properties such as radiative losses. $1^2$  To study this transport, workers have purposely injected nonperturbing, trace amounts of moderate- $Z$   $\left(\sim$  14) impurities into plasmas.<sup>3-5</sup> Until now such studies have been based on observed radiative emissions from partially stripped impurities, usually integrated along a single chord through the plasma cross section.<sup>3-5</sup> Here we report the first spatially and temporally resolved measurements of fully ionized and total Si  $(Z = 14)$  abundances, which we made following injection of trace Si into Alcator **C.** Measurements were also obtained of fluctuations of impurity ion densities due to internal disruptions, a topic addressed only recently. $6 - 8$  The method we have used is based in part on detection of continuum x rays, principally from ground-state recombination of the fully stripped ion. (In a different context, fully stripped low-Z impurities, abundant within and intrinsic to the plasma, have been previous**ly** measured through charge exchange between fully stripped species and neutral particles.<sup>3, 10</sup>

Very recently this method has been applied, with use of a diagnostic neutral beam, to obtain spatially resolved measurements of  $C^{+6}$  and  $O^{+6}$ , if one assumes plasma discharge repeatability.10)

Si was injected, via the laser-blowoff method,<sup>3-5</sup> into Alcator-C plasmas with central temperatures of 1.2 to **1.5** keV. Si transport was monitored **by** uv and x-ray spectrometers, and **by** two broadband, absolutely calibrated x-ray diode  $\ar{rays}$ <sup>11</sup> sensitive mainly to plasma radiation from inside a 10-cm radius (limiter radius is **16.5** cm). These arrays, the principal tools of this analysis, were usually filtered differently: The 17-detector "soft"-filtered array  $(E \ge 1)$ keV), filtered **by 10.0** mg/cm2 Be, responds mainly  $K\alpha$  lines from H- and He-like Si  $(2.00)$ and **1.86** keV, respectively); the 10-detector "hard"-filtered array  $(E \ge 3 \text{ keV})$ , filtered by **10.0** mg/cm2 Be and **32.7** mg/cm2 **C,** responds only to continuum x rays, primarily from groundstate radiative recombination of fully stripped and H-like Si. Even with allowance for the possibility of strong deviations from coronal equilibri**um12** (Fig. **1),** fully stripped, H- and He-like ions

**@ 1982 The** American Physical Society

### VOLUME 49, NUMBER 25 **PHYSICAL REVIEW LETTERS** 20 DECEMBER 1982





FIG. **1.** Coronal equilibrium calculations of Si ion abundances (Ref. 12).

comprise essentially all of the Si inside a radius **of 10 cm.**

Calculations of the response of the arrays to the different charge states of Si included contributions from allowed transitions to ground states plus the  $n = 2$  forbidden and intercombination lines of Si XIII (Refs. **13** and 14) and bremsstrahlung and radiative recombination, with free-free and free-bound Gaunt factors included.<sup>15, 16</sup> The hardfiltered response to the most significant sources **of** continuum radiation is shown in Fig. 2; for *T,*  $\geq 1.2$  keV (see also Fig. 1), the dominant term is ground-state radiative recombination of fully stripped ions, a term which has weak temperature dependence  $({\sim} T^{1/3})$  and which has been accurately calculated<sup>15</sup> ( $\leq 1\%$  error).

Each x-ray detector array provides (through Abel inversion) a time history of the radial profile of local plasma emissivity in a well-defined spectral passband. In each passband, the radiation of the injected Si may be isolated **by** subtracting the preinjection emissivity from the total emissivity at any post-injection time (as long as the injection occurs during the plateau stage), since it has been shown<sup>4,5</sup> that the injected impurities do not affect the macroscopic characteristics of the background plasma and its radiation; for example, the emissions of the most important intrinsic impurities-C, **N, 0,** and Mo (the limiter material)-as well as the electron cyclotron emissions were unchanged **by** the injections.

The observed local Si emissivities corresponding to soft and hard filtering  $(\epsilon_A \text{ and } \epsilon_C, \text{ respec-}$ tively) may be expressed as functions of the densities of the ionization states  $(N_{+i})$ :

 $\epsilon_A/n_e = A_{14} N_{+14} + A_{13} N_{+13} + A_{12} N_{+12}$  $\epsilon_c/n_e = C_{14}N_{+14} + C_{13}N_{+13} + C_{12}N_{+12}$ 

where  $n_e$  is the measured electron density, and the  $A_i$ 's and  $C_i$ 's depend only on temperature



FIG. 2. Calculated system response, in units of erg cm $\frac{3}{s}$ , to various components of Si ion continuum radiation (normalized to  $n_eN_{+i}$ ).

and filter response. (For example,  $C_{14}$  equals the sum of the **"+** 14" terms shown in Fig. 2, plus the tiny contribution-of order **1%--of** recombination terms with  $n \geq 3$ . At 1.5 keV,  $C_{14}$  is 2.3 and 16.0 times larger than  $C_{13}$  and  $C_{12}$ , respectively.) This gives us two equations with three unknowns, but by invoking one assumption-that coronal equilibrium exists between H- and He-like ions -we can solve for all three densities. This assumption can be justified *a posteriori* from our measurements, and it is also supported **by** our spectroscopic data. **Of** primary importance is the fact that the two quantities  $N_{+14}$  and  $N_{+14}$  $+N_{+13} + N_{+12} = N_{Si}$  (the total Si density) are extremely insensitive to this assumption.

Figure **3** shows a sample x-ray diode output for a hydrogen discharge into which Si was injected. The x-ray flux rises as the Si penetrates



**FIG. 3.** Soft-filtered x-ray flux, integrated along a central chord of the plasma, around the time of a **Si** injection.

the plasma (at about **232** ms), and then decreases as the Si leaks out. For comparison to coronal predictions, we choose a time shortly after the Si emission has peaked, and average the data over a sawtooth cycle.

Figure 4 shows the results of this analysis for a deuterium discharge, illustrating the deduced density profiles and the profiles predicted on the basis of coronal equilibrium.<sup>12</sup> The predictions are normalized to the measured total Si profile  $(N_{Si})$ , after accounting for the effects of the electron temperature on the ionization state ratios (see Fig. **1).** (The electron temperature profile was routinely determined from the electron cyclotron emissions.) At  $r = 0$  cm, the measured fully stripped abundance is **0.9 ±** 0.4 times the coronal prediction. The uncertainty stems from errors in the absolute temperature measurement **(± 10%),** and from the inverted x-ray data sets (the hardfiltered array had a partially obstructed field of view). We have not included the uncertainty associated with the coronal prediction, but the calculations unequivocally obtain their highest accuracy for the fully stripped state since, among other important reasons, the total recombination rate (for which no dielectronic component exists) has been calculated nearly exactly. **<sup>7</sup>**

From Fig. 4 we see that the coronal assumption for the H- to He-like ratio was thus justified; but even if this ratio were taken to be larger or smaller than coronal **by** a factor of 2, the derived values of  $N_{+14}$  and  $N_{\rm Si}$  in the plasma core would change **by** no more than **15%** and 2%, respectively. Thus, deviations of the fully ionized densities from coronal equilibrium are not large, at most a factor of **-2.** The proximity of our re-



Figure **5** shows the effects of an internal disruption on fully stripped ion densities, as derived from Si emissivity profiles just before and after an internal disruption, The post-disruption profile is seen to be strongly flattened in the plasma core, the central density decreasing by  $\sim$  50%. Regarding this result, several points should be noted. First, the changes in particle density are not a consequence of ionization or recombination, since the time scales **for** these processes **(-2** ms) are much longer than the time scale of the disruption **(-0.03** ms). Second, although disruption-induced changes in the electron temperature  $(-15\%)$  and density  $(-4\%)$  affect these calculations, their effects are small and the results are



FIG. 4. Solid curves: Observed fully stripped (+14), H-like **(+13),** and He-like (+12) ion density profiles, and their sum  $(N_{\text{Si}})$ , in units of  $10^{10}$  cm<sup>-3</sup>. Dashed curves: Coronal predictions normalized to *Nsi.* The central electron density was  $3.4 \times 10^{14}$  cm<sup>-3</sup>.



FIG. **5.** Fully stripped Si ion density profiles, before (solid line) and after an internal disruption, in units of  $10^{10}$  cm<sup>-3</sup>. (The central electron density was  $3.9 \times 10^{14}$ cm- **3.)** The error bar shows the uncertainty in the *difference* between these two profiles; given the uncertainty, the profile change is consistent with conservation of the fully stripped particles.

I

**INSTRUCTION** 

V

volume <del>1992, Number 25 october 25 pm/sical Review Letters and Columbation 2</del>0 pm/sical Review Letters and Columba<br>Prima 1982 pm/sical Review Letters and Columbation 2007 pm/sical Review Letters and Columbation 2007 pm/sic insensitive to uncertainties in these quantities. **(A** reduction **by** a factor of 2 in the assumed temperature fluctuation results in a **10%** change in the deduced  $N_{+14}$  fluctuation.) Finally, altering the H- to He-like ion ratio **by** a factor of 2 from coronal values changes the deduced  $N_{+14}$  fluctuation **by** only **6%.** H-like profiles were also observed to be flattened **by** disruptions, with central values decreasing **by ~25%** (again relatively insensitive to the parameter variations described above). The percentage change is smaller than for the fully stripped species because the profile is less peaked (see Fig. 4), and flattening has a smaller effect. In either case the flattening presumably occurs through the same momentary en-

hancement of central radial transport which flattens electron temperature and density profiles, **<sup>2</sup>** as recently discussed.<sup>7, 8</sup> In summary, a novel method has been used to measure-with spatial and temporal resolution -absolute abundances of fully stripped and total

Si following the injection of trace Si into Alcator plasmas; with these and related measurements, certain important aspects of impurity transport were addressed. Because this measurement technique has the potential of extremely high accuracy, it should ultimately provide a very sensitive means for studying the transport of injected impurities in magnetically confined plasmas.

The authors thank the Alcator staff for their encouragement and assistance, especially R. Gandy for discussions about the electron cyclotron data. We express our gratitude to **E.** Killne and **J.** Killne of Smithsonian Astrophysical Observatory, and **A.** Krieger of American Science and Engineering, Inc. for helpful discussions. This work was supported **by** the **U. S.** Department of Energy under Contracts No. **DE-AC02-77ET53068** and No. **DE-AC02-78ET51013,** and Massachusetts Institute of Technology under Contract No. MIT-**FC-A-246206.**

**. 'D. E.** Roberts, Nucl. Fusion 21, **215 (1981),** and references therein.

<sup>2</sup>R. Petrasso, M. Gerassimenko, F. H. Seguin R. Krogstad, and **E.** Marmar, Nucl. Fusion 21, **881 (1981),** and references therein.

**3S. A.** Cohen, **J.** L. Cecchi, and **E. S.** Marmar, Phys. Rev. Lett. **35, 1507 (1975). <sup>4</sup>**

**E. S.** Marmar, **J. E.** Rice, and **S.** L. Allen, Phys. Rev. Lett. 45, **2025 (1980).**

**1E. S.** Marmar, **J. E.** Rice, **J.** T. Terry, and F. H. Seguin, Massachusetts Institute of Technology Report No. **PFC/JA-82-12, 1982** (to be published).

**6E.** Hinnov *et al.,* Bull. Am. Phys. Soc. **25, 902 (1980).**

**7 F.** H. Seguin and R. Petrasso, in Proceedings of the Annual Sherwood Controlled Fusion Theory Conference, Santa Fe, New Mexico, 25-28 April 1982 (unpublished).

F. H. Seguin, R. Petrasso, and **E. S.** Marmar, Massachusetts Institute of Technology Report No. **PFC/JA-82-25, 1982** (to be published).

9R. **C.** Isler, *L.* **E.** Murray, **S.** Kasai, **J.** L. Dunlap, *S.* **C.** Bates, P. H: Edmonds, **E. A.** Lazarus, **C. H.** Ma, and M. Murakami, Phys. Rev. **A** 24, **2701 (1981).**

"R. **J.** Fonck, M. Finkenthal, R. **J.** Goldston, **D.** L. Hendron, R. **A.** Hulse, R. Kaita, and **D. D.** Meyerhofer, Phys. Rev. Lett. 49, **737 (1982).**

<sup>11</sup>R. Petrasso, M. Gerassimenko, F. H. Seguin,

**J.** Ting, R. Krogstad, P. Gauthier, W. Hamilton, **A.** T. Ramsey, P. Burstein, and R. Granetz, Rev. Sci. Instrum. **51, 585 (1980).**

**12J.** M. Shull and M. Van Steenberg, Astrophys. **J.,** Suppl. Ser. 48, **95 (1982),** and 49, **351(E) (1982),** and private communication.

**"A.** K. Pradhan, **D.** W. Norcross, and **D. G.** Hummer, Astrophys. **J.** 246, **1031 (1981).**

<sup>14</sup>R. Mewe, J. Schrijver, and J. Sylwester, Astron. Astrophys. **87, 55 (1980),** and references therein. **<sup>15</sup>**

**W. J.** Karzas and R. Latter, Astrophys. **J.,** Suppl. Ser. 6, 167 (1961).

<sup>16</sup>R. T. Brown, Astrophys. J. 170, 387 (1971). (Typographical errors need to be corrected before applying Brown's formula.)

**1 7 M. J..** Seaton, Mon. Not. Roy. Astron. Soc. **119** (2), **81 (1959).**

**18M.** Bitter, **S.** von Goeler, P. **S.** Efthimion, M. Goldman, K. W. Hill, and **N.** Sauthoff, Bull. Am. Phys. Soc. **26, 981 (1981),** and private communication.